

# Extreme AGN feedback in the MASSive Cluster Survey: a detailed study of X-ray cavities at $z > 0.3$

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## ABSTRACT

We present the first statistical study of X-ray cavities in distant clusters of galaxies ( $z > 0.3$ ). With the aim of providing further insight into how active galactic nucleus (AGN) feedback operates at higher redshift, we have analysed the *Chandra* X-ray observations of the MASSive Cluster Survey (MACS) and searched for surface brightness depressions associated with the brightest cluster galaxy (BCG). The MACS sample consists of the most X-ray luminous clusters within  $0.3 \leq z \leq 0.7$  (median  $L_{X,RASS} = 7 \times 10^{44} \text{ erg s}^{-1}$ ), and out of 76 clusters, we find 13 with ‘clear’ cavities and seven with ‘potential’ cavities (detection rate  $\sim 25$  per cent). Most of the clusters in which we find cavities have a short central cooling time below 3–5 Gyr, consistent with the idea that cavities sit predominantly in cool core clusters. We also find no evidence for evolution in any of the cavity properties with redshift, up to  $z \sim 0.6$ . The cavities of powerful outbursts are not larger (or smaller) at higher redshift, and are not able to rise to further (or lesser) distances from the nucleus. The energetics of these outbursts also remain the same. This suggests that extreme ‘radio mode’ feedback ( $L_{\text{mech}} > 10^{44} \text{ erg s}^{-1}$ ) starts to operate as early as 7–8 Gyr after the big bang and shows no sign of evolution since then. In other words, AGNs lying at the centre of clusters are able to operate at early times with extreme mechanical powers, and have been operating in such a way for at least the past 5 Gyr.

**Key words:** galaxies: clusters: intracluster medium – galaxies: evolution – galaxies: jets – X-rays: galaxies: clusters.

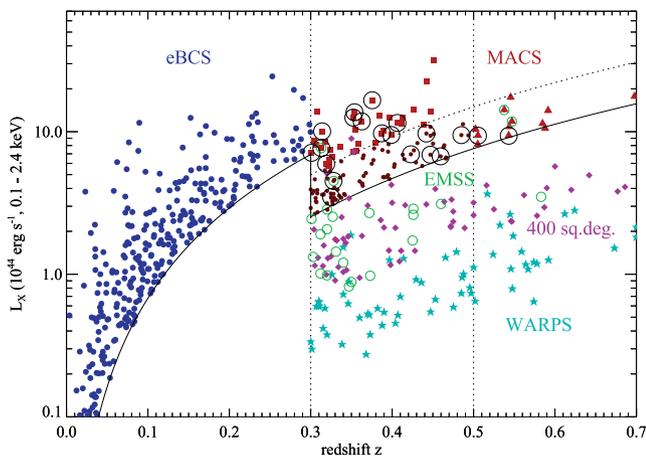
## 1 INTRODUCTION

Active galactic nucleus (AGN) feedback plays a major role in quenching star formation, enriching the surrounding medium with metals and fuelling the supermassive black hole (SMBH) of the host galaxy. Some of the most extreme examples of AGN feedback are seen in clusters of galaxies, where the central SMBH inflates large cavities filled with radio emitting particles through jets. These cavities appear as depressions in the X-ray image and provide a direct measurement of the energy being injected into the surrounding medium by the central SMBH (see reviews on the topic by Peterson & Fabian 2006; McNamara & Nulsen 2007).

Detailed studies of individual nearby clusters such as the Perseus Cluster (see Fabian et al. 2006, and references therein) have shown that the AGN lying at the centre can energetically offset cooling of the intracluster medium (ICM) not only by inflating these large

cavities, but also by inducing weak shocks and propagating energy through sound/pressure waves (Fabian et al. 2003, 2006; Bîrzan et al. 2004; Forman et al. 2005; Dunn & Fabian 2006, 2008; Rafferty et al. 2006; McNamara & Nulsen 2007; Sanders & Fabian 2007). Statistical studies of X-ray cavities have also provided a wealth of information on how AGN feedback operates (Bîrzan et al. 2004, 2008; Dunn, Fabian & Taylor 2005; Dunn & Fabian 2006, 2008; Nulsen et al. 2007; Cavagnolo et al. 2010; Dunn et al. 2010; Dong, Rasmussen & Mulchaey 2010; O’Sullivan et al. 2011). Essentially, these studies find that the power output of cavities is substantial, in most cases sufficient to prevent cooling of the ICM, and that cavities are common, especially in systems which require some form of heating to prevent the gas from cooling. However, all the statistical compilations of cavities have until now focused on the nearby Universe (only four objects are at  $z > 0.3$ ). The cavity properties at the higher redshift end therefore remain unexplored. It is also not clear if and how AGN heating evolves across time. According to cosmological simulations (Croton et al. 2006; Sijacki & Springel 2006), present-day black holes are in a ‘radio mode’

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**Figure 1.** The X-ray luminosity versus redshift distribution of different cluster samples. There are a total of 124 MACS clusters, all with  $z > 0.3$ . The MACS clusters outlined with black circles are those in which we identified cavities.

phase, which involves a black hole accreting at sub-Eddington rates and driving powerful outflows, whereas at earlier times ( $z \gtrsim 1$ ), ‘quasar mode’ feedback dominates and consists of a merger-driven phase where the black hole grows rapidly. When and how this transition occurs remains unclear.

We can start investigating how AGN feedback evolves with redshift by using targeted samples of  $z > 0.3$  clusters that have extensive follow-up X-ray observations, and in this case, the MASSive Cluster Survey (MACS) provides an ideal example. MACS was launched in 1999 and compiled the first large sample of very X-ray luminous, and therefore very massive, clusters in the distant Universe (Ebeling, Edge & Henry 2001; Ebeling et al. 2007, 2010). The MACS sample was based on the *ROSAT* All-Sky Survey (RASS; Trümper 1983, and references therein) Bright Source Catalogue (BSC; Voges et al. 1999). It consists of clusters above a redshift of  $z = 0.3$  and covers a solid angle over  $22\,000\text{ deg}^2$  in the extragalactic sky. In Fig. 1, we show the  $L_X$ - $z$  distribution for various samples, including the Brightest Cluster Survey (Ebeling et al. 1998, 2000), the *Einstein* Medium Sensitivity Survey (EMSS; Gioia & Luppino 1994), the Wide Angle *ROSAT* Pointed Survey (WARPS; Perlman et al. 2002), the  $400\text{ deg}^2$  cluster samples (Burenin et al. 2007) and the MACS sample. This figure shows that MACS selects on average clusters which are 10 to 20 times more luminous than the other samples in the same redshift range. The MACS sample now contains 124 spectroscopically confirmed clusters within  $0.3 \leq z \leq 0.7$ , for which more than two-thirds are new discoveries.

Since MACS contains only very X-ray luminous clusters (median  $L_{X,\text{RASS}} = 7 \times 10^{44}\text{ erg s}^{-1}$ ), those harbouring a cool core should require extreme feedback from their central AGN to prevent the ICM from cooling ( $L_{\text{mech}} \sim 10^{44}\text{--}10^{45}\text{ erg s}^{-1}$ ). We therefore expect to see some of the most extreme examples of AGN outbursts in this sample, which will allow us to extend our knowledge of X-ray cavities not only in the distant Universe ( $z > 0.3$ ), but also at the higher end of mechanical power arising from the AGN. To date, more than two-thirds of all MACS clusters have been targeted in *Chandra* follow-up observations, and the unique and exquisite spatial resolution of *Chandra* is essential to identify cavities at such redshifts. The MACS sample therefore provides a unique opportunity to study the cavity properties in clusters at  $z > 0.3$ , and determine if AGN feedback evolves with time.

We first present the data in Section 2, and then discuss how we proceeded in identifying systems with cavities in Section 3. Sections 4 and 5 outline the techniques used to calculate the energetics of the cavities, as well as their associated time-scales. In Section 6, we derive the properties of the cores and plot scaling relations in Section 7. Finally, we present the results in Section 8 and discuss them in Section 9. We adopt  $H_0 = 70\text{ km s}^{-1}\text{ Mpc}^{-1}$  with  $\Omega_m = 0.30$  and  $\Omega_\Lambda = 0.7$  throughout this paper. All errors are  $2\sigma$ , and the abundance ratios of Anders & Grevesse (1989) were used throughout this paper.

## 2 OBSERVATIONS AND DATA REDUCTION

### 2.1 X-ray observations

The X-ray data were obtained through the *Chandra* data archive. We found observations with the Advanced CCD Imaging Spectrometer (ACIS) for 83 of the MACS clusters, with exposure times varying between 10 and 100 ks. Some objects were also observed multiple times. For each cluster, we selected the deepest observational data set available. These data were then processed, cleaned and calibrated using the latest version of the CIAO software (CIAOV4.3, CALDB4.4.1), and starting from the level 1 event file. We applied both charge time interval (CTI) and time-dependent gain corrections, as well as removed flares using the LC\_CLEAN script, with a  $3\sigma$  threshold. When a cluster was observed multiple times with the target centred on the same detector and in the same observing mode (FAINT or VFaint), we combined the different observations only if this would improve the image quality significantly. We then exposure-corrected the images, using an exposure map generated with a monoenergetic distribution of source photons at 1.5 keV (which is almost the peak energy of the clusters). These images were used to identify systems with cavities (see Section 3). Out of the 83 clusters for which we have X-ray observations, two are clear mergers with multiple components merging together and five have a bright point source with no clear extended emission associated with it. We therefore only consider the remaining 76 in the analysis of X-ray cavities.

For the clusters in which we identified cavities (see Table 1), the spectra were analysed using XSPEC (v12.6.0 d.f; e.g. Arnaud 1996). For each object, the Galactic absorption was kept frozen at the Kalberla et al. (2005) value. Varying this parameter did not significantly improve the fit in any of our clusters. For targets observed multiple times, we fitted simultaneously the spectra. Finally, although some of the objects in our sample with cavities have a bright X-ray point source at their centres, we find no significant pile-up for any of these point sources. Pile-up occurs when two or more photons are detected as one event (see for more details Davis 2001; Russell et al. 2010), and the amount of pile-up can be estimated by comparing the fraction of good (grades 0,2,3,4,6) to bad grades (grades 1,5,7) for each point source. Typically, pile-up becomes problematic when the fraction of bad grades exceeds 10 per cent of the good grades, but we find a ratio less than 6 per cent for all of our central point sources.

### 2.2 Radio observations

For each cluster with cavities, we searched for radio emission associated with the central AGN by looking through different radio surveys available to the public. We first looked for radio emission associated with the AGN at 1.4 GHz by using the 1.4 GHz Very Large Array (VLA) Faint Images of the Radio Sky at Twenty-Centimetres survey (FIRST; average resolution of 5 arcsec; Becker,

**Table 1.** *Chandra* observations of the MACS clusters with cavities – (1) name; (2) alternate name; (3) redshift  $z$ ; (4) observation identification number; (5) exposure time; (6) chip name where the target was centred on; (7) references where the redshift was taken from: (i) Ebeling et al. (2007); (ii) Ebeling et al. (2010); (iii) Stocke et al. (1991); (iv) Wright, Ables & Allen (1983); (v) Mann & Ebeling (2011); (vi) Kleinmann et al. (1988); (vii) Dressler & Gunn (1992); (viii) Ebeling et al. (in preparation); (ix) Sloan Digital Sky Survey; (x) Smail et al. (2007).

(1) Cluster name	(2) Alternate name	(3) $z$	(4) Obs. ID	(5) Exposure (ks)	(6) Chip	(7) Reference
MACS J0159.8–0849	–	0.404	3265	14.5	ACIS-I3	(ii)
			6106	28.9	ACIS-I3	
			9376	18.0	ACIS-I3	
MACS J0242.5–2132	–	0.314	3266	8.6	ACIS-I3	(iv)
MACS J0429.6–0253	–	0.397	3271	21.4	ACIS-I3	(ii)
MACS J0547.0–3904	–	0.319	3273	19.2	ACIS-I3	(ii)
MACS J0913.7+4056	IRAS 09104+4109	0.442	10445	70.4	ACIS-I3	(vi)
MACS J0947.2+7623	RBS 0797	0.354	7902	38.8	ACIS-S3	(ii)
MACS J1411.3+5212	3C295	0.460	2254	76.4	ACIS-I3	(vii)
MACS J1423.8+2404	–	0.5449	4195	106.9	ACIS-S3	(i)
MACS J1532.8+3021	RX J1532.9+3021	0.3613	1649	9.5	ACIS-S3	(ii)
MACS J1720.2+3536	Z8201	0.3913	3280	17.6	ACIS-I3	(ii)
			6107	27.3	ACIS-I3	
			7718	6.8	ACIS-I3	
MACS J1931.8–2634	–	0.352	9382	95.0	ACIS-I3	(ii)
MACS J2046.0–3430	–	0.423	9377	35.9	ACIS-I3	(viii)
MACS J2140.2–2339	MS 2137.3–2353	0.313	4974	7.3	ACIS-S3	(iii)
			5250	34.7	ACIS-S3	
			928	34.8	ACIS-S3	
MACS J0111.5+0855	–	0.485	3256	15.2	ACIS-I3	(v)
MACS J0257.1–2325	–	0.5039	1654	18.0	ACIS-I3	(i)
			3581	16.1	ACIS-I3	
MACS J1359.1–1929	–	0.447	9378	46.1	ACIS-I3	(v)
MACS J1359.8+6231	MS 1358.4+6245	0.330	516	48.5	ACIS-S3	(iii)
MACS J1447.4+0827	RBS 1429	0.3755	10481	11.5	ACIS-S3	(ix)
MACS J2135.2–0102	1RXS J213515.7–010208	0.325	11710	24.7	ACIS-I3	(x)
MACS J2245.0+2637	–	0.301	3287	11.5	ACIS-I3	(ii)

White & Helfand 1994), or the 1.4 GHz NRAO (National Radio Astronomy Observatory) VLA Sky Survey catalogue (NVSS; average resolution of 45 arcsec; Condon et al. 1998). We then looked for radio emission associated with the AGN at lower frequency, either with the 326 MHz Westerbork Northern Sky Survey<sup>1</sup> (WENSS; resolution of  $\sim 50$  arcsec; Rengelink et al. 1997), the 74 MHz VLA Low-frequency Sky Survey (VLSS; resolution of  $\sim 80$  arcsec; Cohen et al. 2007), the 150 MHz Tata Institute of Fundamental Research (TIFR) Giant Metrewave Radio Telescope (GMRT) Sky Survey (TGSS,<sup>2</sup> resolution of  $\sim 20$  arcsec) or the 843 MHz Sydney University Molonglo Sky Survey (SUMSS; resolution of  $\sim 40$  arcsec; Bock, Large & Sadler 1999; Mauch et al. 2003). Finally, we searched for high-frequency radio emission associated with the central AGN at 5 GHz with the Parkes-MIT-NRAO (PMN; resolution of  $\sim 5$  arcmin) radio survey (Griffith & Wright 1993) or the EMSS (where the radio resolution of the survey lies between  $\sim 4$  and 14 arcsec; Gioia et al. 1990; Stocke et al. 1991), and then at 28.5 GHz with the Berkeley–Illinois–Maryland Association (BIMA; sensitive to angular scales around 1.5 arcmin; Coble et al. 2007). If no point source was seen within the central regions ( $r \lesssim 100$  kpc), we used the  $2\sigma_{\text{rms}}$  value within the beam area as an upper limit for the flux. The values we obtain are shown in Table 2. Some sources also have

flux densities available from the literature, and we include these in Table 2.

Errors are derived as the quadratic sum of the rms noise level in the map and the systematic uncertainty associated with the value. For some of the surveys used, the error quoted in the catalogue did not account for systematic uncertainties. This type of error varies with frequency, but is of the order of 5 per cent (see Carilli et al. 1991). For simplicity, we therefore choose to compute the total uncertainty assuming a 5 per cent systematic error and a  $2\sigma_{\text{rms}}$  noise level. MACS J1931.8–2634 has a complicated morphology at radio wavelengths, with a Narrow Angle Tail (NAT) source located 45 arcsec to the south that is contaminating the flux measurements of the central galaxy. The 1.4 GHz flux quoted in Table 2 is taken from Ehlert et al. (2011) and only includes the contribution of the central galaxy. However, for the 5 GHz and 150 MHz data, we had to remove the contribution of the NAT source. In this case, we used the peak flux value of the galaxy and subtracted the peak contribution of the NAT source.

Finally, by using the flux densities quoted in Table 2, we calculated a rough estimate of the total radio luminosity ( $L_{\text{radio}}$ ) with equation (1) and integrating between  $\nu_1 = 10$  MHz and  $\nu_2 = 10000$  MHz.

$$L_{\text{radio}} = 4\pi D_L^2 \int_{\nu_1}^{\nu_2} (S_\nu) d\nu. \quad (1)$$

Here,  $D_L$  is the luminosity distance to the source and  $S_\nu$  is the flux density ( $S_\nu \propto \nu^{-\alpha}$ , where  $\alpha$  is the spectral index). For MACS

<sup>1</sup> <http://www.astron.nl/wow/testcode.php?survey=1>

<sup>2</sup> <http://tgss.ncra.tifr.res.in>

**Table 2.** Radio properties of the MACS clusters with cavities – (1) name; (2) 74 MHz radio flux density from VLSS; (3) 326 MHz radio flux density from WENSS; (4) 1400 MHz radio flux density; (5) 5 GHz radio flux density; (6) 28.5 GHz radio flux density from BIMA (Coble et al. 2007); (7) other radio flux density; (8) integrated radio luminosity from 10 MHz to 10 GHz (see Section 2.2 for details); (9) References: (i) 1.4 GHz FIRST; (ii) 1.4 GHz NVSS; (iii) Parkes MIT-NRAO 4.85 GHz survey (Griffith & Wright 1993); (iv) 5 GHz (Hogan M. T., private communication); (v) 5 GHz from Hines & Wills (1993); (vi) 5 GHz from Cavagnolo et al. (2011); (vii) 5 GHz from Gregory & Condon (1991); (viii) EMSS 5 GHz (Gioia et al. 1990; Stocke et al. 1991); (ix) 843 MHz SUMSS (Bock et al. 1999; Mauch et al. 2003); (x) 15 GHz Arcminute Microkelvin Imager (AMI; Grainge K., private communication); (xi) 1.4 GHz from Ehlert et al. (2011); (xii) 150 MHz TIFR GMRT Sky Survey.

(1) Cluster name	(2) $S_{74\text{ MHz}}$ (mJy)	(3) $S_{326\text{ MHz}}$ (mJy)	(4) $S_{1.4\text{ GHz}}$ (mJy)	(5) $S_{5\text{ GHz}}$ (mJy)	(6) $S_{28.5\text{ GHz}}$ (mJy)	(7) $S_{\text{other}}$ (mJy)	(8) $L_{\text{radio}}$ ( $10^{42}\text{ erg s}^{-1}$ )	(9) Ref.
MACS J0159.8–0849	<112	–	$31.4 \pm 1.6$	$58 \pm 11$	–	–	3.56	(i,iii)
MACS J0242.5–2132 <sup>a</sup>	$890 \pm 145$	–	$1255 \pm 73$	$795 \pm 43$	–	–	27.1	(ii,iii)
MACS J0429.6–0253	<214	–	$138.8 \pm 8.1$	–	–	–	7.2	(ii)
MACS J0547.0–3904	–	–	$31.4 \pm 1.9$	$15.4 \pm 0.8$	–	$19.6 \pm 1.3$ [@843 MHz]	0.59	(ii,iv,ix)
MACS J0913.7+4056	<272	$54.0 \pm 7.1$	$8.3 \pm 0.5$	$1.6 \pm 0.1$	$0.69 \pm 0.12$	$0.80 \pm 0.04$ [@15 GHz]	1.33	(i,v,x)
MACS J0947.2+7623	<329	$91.0 \pm 6.1$	$21.7 \pm 1.3$	$4.0 \pm 0.2$	–	–	1.02	(ii,vi)
MACS J1411.3+5212 <sup>a</sup>	$120\,270 \pm 6022$	$61647 \pm 3082$	$22171 \pm 1109$	$7401 \pm 808$	–	–	1025.3	(i,vii)
MACS J1423.8+2404	<232	–	$5.2 \pm 0.4$	–	$1.49 \pm 0.12$	–	3.58	(i)
MACS J1532.8+3021	<222	$71.0 \pm 8.2$	$17.1 \pm 0.9$	$8.8 \pm 0.5$	$3.25 \pm 0.18$	–	0.94	(i,iv)
MACS J1720.2+3536	<266	$103.0 \pm 7.4$	$16.8 \pm 1.0$	–	–	–	1.23	(i)
MACS J1931.8–2634 <sup>b</sup>	–	–	$70 \pm 4$	$6.0 \pm 1.3$	–	$2799 \pm 161$ [@150 MHz]	80.4	(iv,xi,xii)
MACS J2046.0–3430	–	–	$8.1 \pm 0.6$	–	–	$13 \pm 1.3$ [@843 MHz]	2.49	(ii,ix)
MACS J2140.2–2339	<116	–	$3.8 \pm 0.5$	$1.0 \pm 0.1$	–	–	0.42	(ii,viii)
MACS J0111.5+0855	<124	–	<0.25	–	–	–	<3.20	(i)
MACS J0257.1–2325	<100	–	<0.89	–	–	–	<1.50	(ii)
MACS J1359.1–1929	<266	–	$9.5 \pm 0.8$	–	–	–	2.34	(ii)
MACS J1359.8+6231	–	<7	$2.8 \pm 0.6$	$3.8 \pm 0.2$	$1.67 \pm 0.08$	–	0.16	(i,viii)
MACS J1447.4+0827	<203	–	$39.0 \pm 2.0$	–	–	–	2.00	(i)
MACS J2135.2–0102	<192	–	<0.40	–	–	–	<1.90	(i)
MACS J2245.0+2637	<208	–	$5.8 \pm 0.6$	–	–	–	0.72	(ii)

<sup>a</sup>For the central galaxies of MACS J0242.5–2132 and MACS J1411.3+5212 (3C 295), we used the extensive follow-up radio observations available in the NED to determine  $L_{\text{radio}}$ .

<sup>b</sup>MACS J1931.8–2634 has a complicated morphology at radio wavelengths (see Section 2.2 for details on how we obtained the flux densities).

J0242.5–2132 and MACS J1411.3+5212 (3C 295), we also use the extensive follow-up radio observations available in the NASA/IPAC Extragalactic Database (NED) to determine  $L_{\text{radio}}$ . Only MACS J1411.3+5212 has a complete coverage from  $\sim 10$  MHz to 10 GHz. For the remaining objects, we extrapolate the values at 10 MHz and 10 GHz based on the assumption that the flux density scales as  $S_\nu \propto \nu^{-\alpha}$ , and for each extremity (10 MHz and 10 GHz), we use the local spectral index as determined from the two nearest flux density data points. Using a simple trapezoid rule, we then integrate over the range 10 MHz to 10 GHz using the extrapolated values at 10 MHz and 10 GHz, as well as the other data points available for each source. Our results are shown in Table 2.

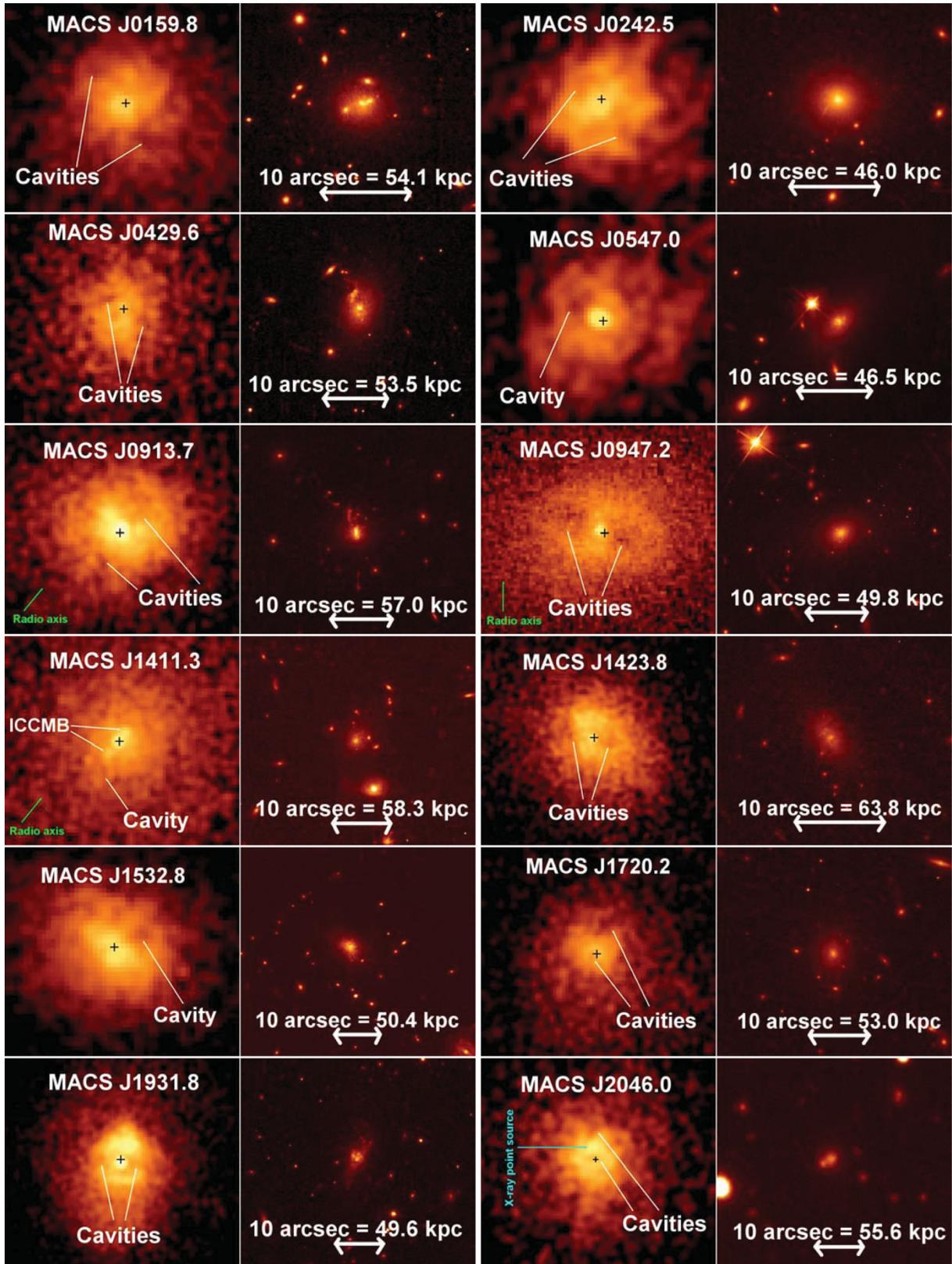
### 2.3 Optical observations

We also searched for optical observations of each cluster harbouring a cavity. These images were used to verify that the cavities identified were associated with the central dominant galaxy. For each cluster, we searched for *Hubble Space Telescope* (HST) archival images (many of which were acquired through snapshot campaigns; see Ebeling 2009), and if not available, we used the *R* passband images from Ebeling et al. (2007, 2010) taken with the University of Hawaii (UH) 2.2-m telescope (MACS J2046.0–3430 and MACS J0111.5+0855) or the red images from the Second Palomar Observatory Sky Survey (POSS II; Reid et al. 1991) for MACS J1447.4+0827. In Fig. 2, we show these optical images along with the 0.5–7.0 keV *Chandra* X-ray images.

### 3 IDENTIFYING SYSTEMS WITH CAVITIES

For each cluster, we visually searched for circular or ellipsoidal surface brightness depressions in the X-ray images associated with the central regions ( $r < 100$  kpc). Beyond 100 kpc, the count rate in most observations is not sufficient to identify cavities. As a first indicator, we computed unsharp-masked images for each source. This technique enhances deviations in the original image and consists of subtracting a strongly smoothed image from a lightly smoothed image, and is similar to subtracting an elliptical model of the cluster emission. Both the unsharp-masked and ellipse-subtracted images are shown in Fig. A1. However, although these figures can be very useful, they can also be misleading and cause false cavities to appear. Therefore, we only considered a cavity as being truly there if we could also see a hint of a depression in the original image. Once the systems with cavities were identified, we proceeded in classifying each cavity (see Column 2 of Table 3). First, if the cavity had a clear contrast in the raw image, as well as in the unsharp-masked and ellipse-subtracted image, we classified it as being ‘clear’. Out of these ‘clear’ cavities, some had surrounding bright rims in the raw image which we identified with the annotation 1 in Table 3. Those without the bright rims were given the annotation 2. If there was just a hint of a depression in the raw image, but the cavity could clearly be seen in the unsharp-masked or ellipse-subtracted image, then we classified it as being a ‘potential’ one.

In total, we find that 13 clusters have at least one ‘clear’ cavity and seven have ‘potential’ cavities. Therefore, out of the 76 MACS clusters with *Chandra* data, we find that 20 have X-ray cavities.



**Figure 2.** Images of the MACS clusters with cavities. There are two clusters in each horizontal line. For each cluster we show the 0.5–7 keV X-ray image along with an *HST* image or, if not available, the *R*-passband images from Ebeling et al. (2007, 2010) taken with the UH 2.2-m telescope (MACS J2046.0–3430 and MACS J0111.5+0855) or the Second Palomar Observatory Sky Survey image, POSS II, for MACS J1447.4+0827. The 10 arcsec scale for each object is shown with the vector in the optical images, and the black cross symbol shows the location of the BCG. The first 13 clusters have ‘clear’ cavities, while the remaining seven only have ‘potential’ cavities (see ‘potential cavities’ indicated in the figures). We emphasize that deeper observations are needed to confirm if the ‘potential’ cavities are real. For the clusters with published radio maps, we plot the radio axis with a green arrow.

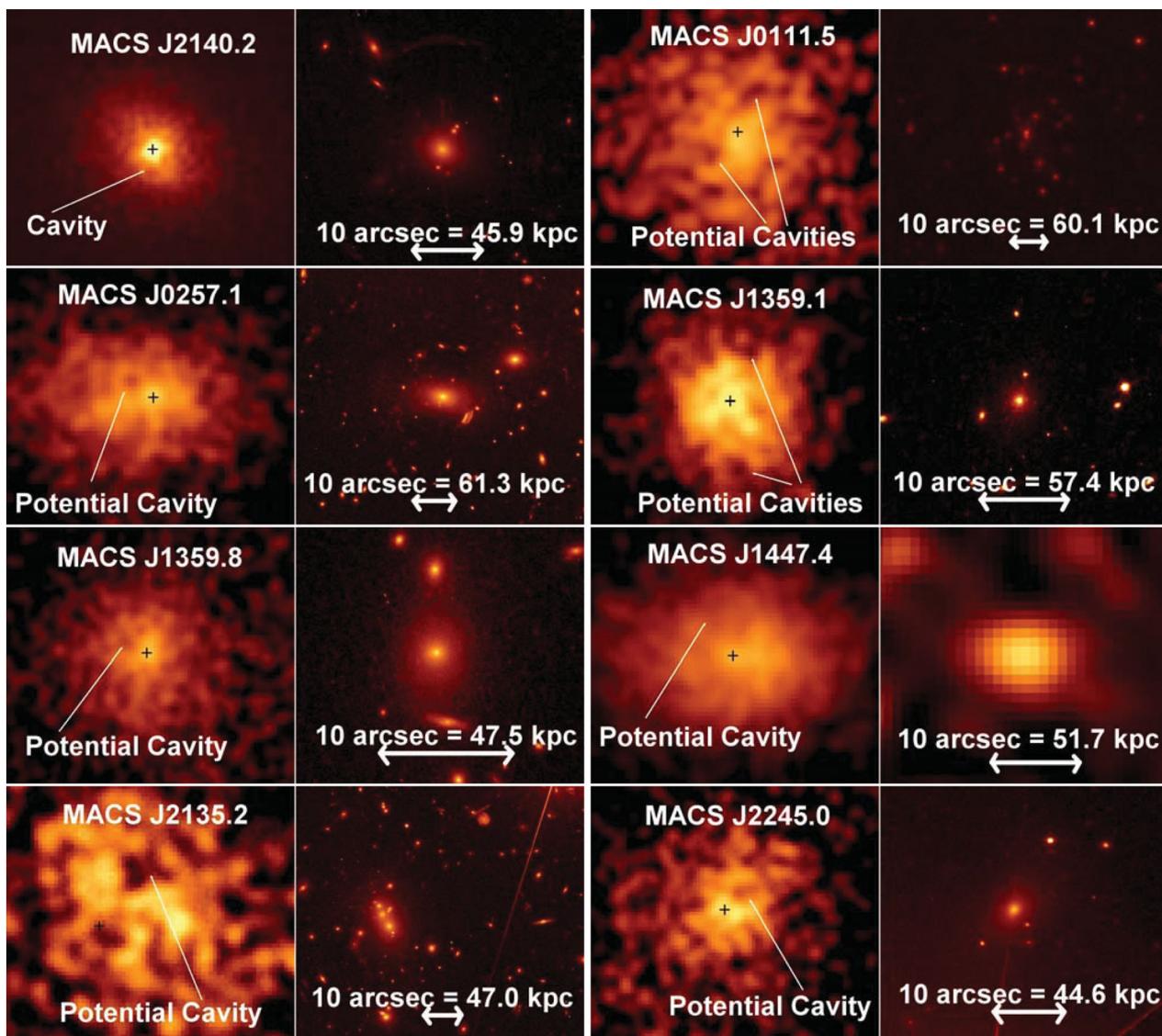


Figure 2 – continued

If we calculate the fractional difference between the counts within the cavity and the surrounding region, then we find that the decrements are of the order of 20–30 per cent for the ‘clear’ cavities and between 10 and 20 per cent for the ‘potential’ cavities. These fractions are roughly consistent with the decrements seen in nearby clusters of galaxies. In the calculation, we use the two immediate regions opposite the cavity, at the same radius and same size as the cavity, as the surrounding region. Note also that many of the MACS clusters have only been observed for 10–20 ks and some have as little as 500–1000 counts within 200 kpc. We could therefore be missing many cavities and our detection rate should only be considered as a lower limit to the number of MACS clusters with cavities.

Our sample of objects with cavities includes MACS 1411.3+5212 (3C 295). This is a particularly interesting cluster, which has a very powerful radio source at its centre and shows strong evidence for inverse-Compton scattered cosmic microwave background (ICM) photons along the central jet axis. Out of the

remaining clusters, MACS J2135.2–0102 has a central galaxy harbouring a double or triple nucleus (see the optical image in Fig. 2). Initially, we identified two cavities in this system, one to the north-west and another to the south-east. However, the brightest cluster galaxy (BCG) is located almost within the south-eastern cavity. Although this could be due to projection effects, we choose to discard the southern cavity and just consider the northern one as a ‘potential’ cavity. MACS J2135.2–0102 is more well known for its lensing arcs in the shape of a ‘Cosmic Eye’ and ‘Cosmic Eyelash’ (Smail et al. 2007; Siana et al. 2009; Swinbank et al. 2010).

#### 4 CAVITY ENERGETICS

We estimate the energy stored within each of the cavities using equation (2) (e.g. Birzan et al. 2004; Dunn et al. 2005; Dunn & Fabian 2006, 2008; Rafferty et al. 2006). Here,  $p$  is the thermal pressure of the ICM at the radius of the bubble and estimated from

**Table 3.** Cavity energetics. The annotations in Column 2 indicate the quality of the cavities detected. The superscript ‘c’ is for cavities with clear surface brightness depressions while the second annotation <sup>1,2</sup> indicates whether a cavity is surrounded by bright rims (<sup>1</sup>) or not (<sup>2</sup>). Column 3 shows the position angle of the major-axis along the cavity, and is measured counter-clockwise from the north.

Cluster	Lobe	PA (°)	$R_l$ (kpc)	$R_w$ (kpc)	$R$ (kpc)	$E_{\text{bubble}} = PV$ ( $10^{58}$ erg)	$t_{\text{cs}}$ ( $10^7$ yr)	$t_{\text{buoy}}$ ( $10^7$ yr)	$t_{\text{refill}}$ ( $10^7$ yr)	$P_{\text{cs}}$ ( $10^{44}$ erg s <sup>-1</sup> )	$P_{\text{buoy}}$ ( $10^{44}$ erg s <sup>-1</sup> )	$P_{\text{refill}}$ ( $10^{44}$ erg s <sup>-1</sup> )
MACS J0159.8–0849	NE <sup>c,2</sup>	29	16.4	8.8	25.6	7.81	2.1	3.3	6.8	4.73	2.99	1.45
	SW <sup>c,2</sup>	217	14.9	7.7	28.6	5.45	2.3	3.9	6.5	2.95	1.78	1.07
MACS J0242.5–2132	N <sup>c,2</sup>	21	6.1	5.5	12.9	1.88	1.3	3.0	5.2	1.85	0.81	0.46
	SW <sup>c,1</sup>	215	10.9	8.0	19.2	7.08	1.9	3.3	6.6	4.66	2.72	1.37
MACS J0429.6–0253	NE <sup>c,2</sup>	40	5.9	5.8	15.7	1.19	1.5	3.3	4.7	1.02	0.45	0.32
	W <sup>c,2</sup>	255	6.0	5.2	15.5	0.98	1.5	3.3	4.6	0.85	0.38	0.27
MACS J0547.0–3904	NE <sup>c,1</sup>	65	8.5	11.5	15.0	3.22	1.7	2.5	5.8	2.35	1.63	0.71
MACS J0913.7+4056	NW <sup>c,2</sup>	335	25.2	25.6	30.1	120.0	3.0	2.2	7.0	50.1	68.5	21.6
	SE <sup>c,2</sup>	155	23.7	17.2	46.1	30.0	4.2	5.4	9.6	9.01	7.06	3.94
MACS J0947.2+7623	NE <sup>c,1</sup>	60	20.9	16.3	24.0	89.1	2.1	1.7	5.1	53.0	68.0	22.0
	SW <sup>c,1</sup>	250	18.8	13.7	24.6	56.2	2.2	1.8	4.8	32.6	39.6	14.9
MACS J1411.3+5212	SE <sup>c,2</sup>	135	28.9	22.7	36.6	49.1	3.3	1.8	4.9	19.1	35.6	12.7
MACS J1423.8+2404	E <sup>c,2</sup>	90	12.6	16.5	23.7	38.7	2.3	4.2	9.0	21.3	11.7	5.49
	W <sup>c,1</sup>	280	14.0	11.9	19.3	22.5	1.9	3.2	8.5	15.2	8.83	3.36
MACS J1532.8+3021	W <sup>c,1</sup>	310	17.4	15.4	37.6	31.9	3.5	4.7	8.0	11.5	8.59	5.07
MACS J1720.2+3536	SE <sup>c,1</sup>	150	4.8	5.3	6.7	0.84	0.58	3.4	9.4	1.82	0.31	0.11
	N <sup>2</sup>	355	12.1	15.0	19.0	16.8	1.6	6.1	15.4	13.0	3.50	1.39
MACS J1931.8–2634	E <sup>c,2</sup>	90	15.6	12.8	29.5	41.7	2.6	1.5	2.8	20.7	36.0	19.0
	W <sup>c,2</sup>	270	15.6	12.8	20.6	41.7	1.8	1.0	2.8	29.6	51.6	19.0
MACS J2046.0–3430	S <sup>c,1</sup>	178	6.8	13.7	11.6	7.46	1.2	4.1	10.9	7.73	2.30	0.87
	N <sup>2</sup>	358	17.4	12.3	23.6	15.4	2.4	6.3	16.2	7.88	3.75	1.46
MACS J2140.2–2339	S <sup>c,2</sup>	180	7.1	6.8	14.4	3.43	1.4	2.4	4.5	3.06	1.79	0.98
MACS J0111.5+0855	N <sup>2</sup>	350	32.8	33.2	77.5	56.5	6.5	7.2	11.5	11.0	9.97	6.22
	SE <sup>2</sup>	90	28.0	23.7	60.7	24.5	5.1	6.1	10.2	6.12	5.10	3.05
MACS J0257.1–2325	NE <sup>2</sup>	58	17.5	15.3	39.5	16.2	2.5	3.6	5.9	8.08	5.68	3.51
MACS J1359.1–1929	N <sup>2</sup>	345	23.9	22.4	53.9	35.3	5.2	16.2	26.6	8.56	2.77	1.68
	S <sup>2</sup>	175	11.5	10.6	44.1	3.75	4.3	19.1	18.3	1.11	0.25	0.26
MACS J1359.8+6231	NE <sup>2</sup>	70	4.2	3.2	9.3	0.14	0.85	3.4	5.4	0.20	0.052	0.032
MACS J1447.4+0827	N <sup>2</sup>	5	21.0	20.4	26.9	123.4	2.4	3.5	10.3	64.0	44.2	15.1
MACS J2135.2–0102	NW <sup>2</sup>	335	27.2	28.2	75.3	42.5	5.1	12.9	17.7	10.6	4.18	3.04
MACS J2245.0+2637	NW <sup>2</sup>	325	8.5	7.8	15.5	1.8	1.3	2.9	5.9	1.76	0.79	0.39

X-ray data,  $V$  is the volume of the cavity and for a relativistic fluid,  $\gamma_1 = 4/3$ , therefore  $E_{\text{bubble}} = 4pV$  (this is also supported observationally; see Graham, Fabian & Sanders 2008):

$$E_{\text{bubble}} = \frac{\gamma_1}{\gamma_1 - 1} pV. \quad (2)$$

We assume that the cavities have a prolate shape. The volume is then given by  $V = 4\pi R_w^2 R_l/3$ , where  $R_l$  is the semi-major axis along the direction of the jet, and  $R_w$  is the semi-major axis perpendicular to the direction of the jet. We define ‘jet’ as the line that connects the nuclear X-ray source to the centre of the cavity. The dimensions of the cavities we derive are given in Table 3.

We determine the local thermal pressure by selecting a set of annuli containing roughly the same signal-to-noise ratio and centred on the X-ray peak. For some clusters, the data quality was sufficient to allow  $\sim 4000$  counts per annuli, but for others we could only consider annuli with  $\sim 900$  counts each. Once the regions were selected, we chose a background region for each cluster within the same chip but far from any cluster emission. Since all of our clusters are above a redshift of  $z = 0.3$ , they only occupy a small portion of the chip, allowing ample area for selecting a background within the same chip. We then proceeded in deprojecting the data using the Direct Spectral Deprojection method of Sanders & Fabian (2007). For each spectrum, we fitted an absorbed MEKAL model using  $\chi^2$  statistics and left the temperature, abun-

dance and normalization parameter (related to the electron density) free to vary. Here, the absorption accounts for Galactic absorption and we keep it frozen at the Kalberla et al. (2005) value. In some annuli, it was difficult to constrain the abundance. In this case, we kept the abundance value frozen at the cluster average which we determined by selecting a region within the entire cluster ( $r < 200$  kpc).

For MACS J0111.5+0855 and MACS J2135.2–0102, the count rate was so low that we could only extract regions containing  $\sim 400$  counts each, and the data quality was too poor to correct for deprojection effects. Here, we proceeded in the same way, but did not deproject the spectra and used C-statistics instead. For MACS J1411.3+5212 (3C 295), the central jet axis ( $r < 25$  kpc) and nucleus are dominated by non-thermal emission. During the spectral analysis, we therefore remove the regions along the jet axis, so that we are only left with the thermal emitting gas. In all systems, we start the inner-most annulus at  $r = 1$  arcsec (which corresponds to the *Chandra* point spread function, PSF), so that the central point source is not included.

Following this technique, we were able to derive the temperature ( $kT$ ) and electron density ( $n_e$ ) as a function of radius for each cluster. We then calculate the deprojected electron pressure ( $p_e = n_e kT$ ) and entropy ( $S = kT n_e^{-2/3}$ ) profiles, but only projected quantities for the two clusters mentioned in the previous paragraph. Our results are shown in Fig. B1 of Appendix B. When deriving the cavity

energetics, we use the local thermal pressure of the gas as an estimate of  $p$  in equation (1) ( $p = nkT \sim 1.92 \times n_e kT$ ).

Note that deprojection relies on the cluster having a ‘relaxed’ morphology, i.e. that it can be approximated as being spherical. 10 of the systems in which we find surface brightness depressions are part of the X-ray brightest MACS clusters and have been regarded as having a ‘relaxed’ morphology since they either have a pronounced cool core and very good alignment with a single central dominant cD galaxy or a good optical/X-ray alignment and concentric contours (Ebeling et al. 2010, MACS J0159.8–0849, MACS J0242.5–2132, MACS J0429.6–0253, MACS J0547.0–3904, MACS J0947.2+7623, MACS J1532.8+3021, MACS J1720.2+3536, MACS J1931.8–2634, MACS J2140.2–2339 and MACS J2245.0+2637). Two others are part of the 12 most distant MACS clusters (Ebeling et al. 2007, MACS J0257.1–2325 and MACS J1423.8+2404), and have also been considered as having a ‘relaxed’ morphology. We can therefore approximate these 12 systems as being spherical and deproject the spectra. The remaining eight systems are not part of the Ebeling et al. (2007, 2010) studies and have not yet been classified. MACS J2046.0–3430 shows a slight offset between the central X-ray peak and galaxy, but the cluster still has a pronounced cool core and concentric X-ray contours. We therefore choose to pursue with spectral deprojection for this cluster. All of the remaining clusters have concentric X-ray contours and good optical/X-ray alignment, apart for MACS J2135.2–0102. However, as mentioned in the previous paragraphs, the count rate was so low in this cluster that we could not deproject the spectra and used the projected quantities to calculate the cavity energetics.

## 5 CAVITY AGES

The power capabilities of the cavities are determined by dividing the energy in the bubble with its age. The latter is given by the buoyancy rise time, the refill time or the sound crossing time. See, respectively, equation (3) (Churazov et al. 2001), equation (4) (McNamara et al. 2000) and equation (5) (McNamara & Nulsen 2007). Here,  $R$  is the distance from the central BCG to the middle of the cavity (projected),  $S$  is the cross-sectional area of the bubble ( $S = \pi R_w^2$ ),  $C_D=0.75$  is the drag coefficient (Churazov et al. 2001),  $g$  is the local gravitational acceleration [ $g = GM(<R)/R^2$ ],  $r$  is the bubble radius [ $r = (R_1 R_w)^{1/2}$  for an ellipsoidal bubble], and  $c_s$  is the sound crossing time [ $c_s = \sqrt{\gamma_2 kT / (\mu m_H)}$ , where  $kT$  is the plasma temperature at the radius of the bubble,  $\gamma_2 = 5/3$  and  $\mu = 0.62$ ].

$$t_{\text{buoyancy}} = R \sqrt{\frac{SC_D}{2gV}}, \quad (3)$$

$$t_{\text{refill}} = 2\sqrt{\frac{r}{g}}, \quad (4)$$

$$t_{\text{cs}} = \frac{R}{c_s}. \quad (5)$$

The buoyancy rise time is the time it takes a bubble to reach its terminal buoyancy velocity and depends on drag forces. This is a good estimate of a bubble’s age if it has clearly detached from its AGN and has risen, which is the case for most of our cavities. The refill time is the time it takes a bubble to rise buoyantly through its own diameter starting from rest. The sound crossing time is the time it takes a bubble to travel the projected distance from the central AGN to its current location at the sound speed. The latter is used

under the assumption that the bubbles travel at subsonic speeds. Between the three time estimates, it is still not clear which is the best to use, although they should not vary significantly.

The values we derive in the spectral deprojecting technique are used to estimate the thermal properties of the gas (see Fig. B1 in Appendix B). To determine  $M(<R)$ , we assume hydrostatic equilibrium (equation 6), where  $G$  is the gravitational constant,  $\rho_{\text{gas}} \sim 1.92\mu m_H n_e$  and  $dp/dr$  is the pressure gradient.

$$M(<r) = -\frac{1}{G} \frac{r^2}{\rho_{\text{gas}}} \frac{dp}{dr}. \quad (6)$$

Table 3 shows the values we obtain for the power stored in each cavity using  $t_{\text{cs}}$ ,  $t_{\text{buoy}}$  and  $t_{\text{refill}}$ .

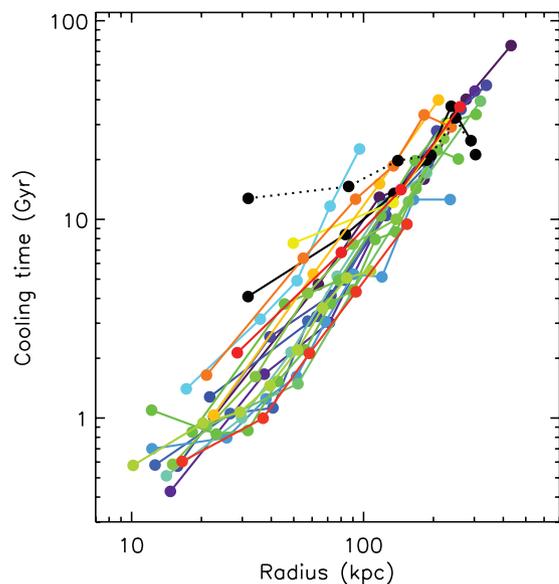
## 6 COOL CORES

### 6.1 Clusters in the MACS sample with cavities

Statistical studies of cavities in clusters of galaxies have shown that cavities sit predominately in cool core clusters (Dunn et al. 2005; Dunn & Fabian 2006). In order to compare our results with theirs, we have computed detailed cooling time profiles for all of our clusters with cavities. More precisely, we compute the cooling time ( $t_{\text{cool}}$ ) by using equation (7) and the thermal gas properties derived in Sections 4 (shown in Fig. B1).

$$t_{\text{cool}} = \frac{5}{2} \frac{1.92 n_e kTV}{L_X}. \quad (7)$$

Here,  $n_e$  is the electron density,  $kT$  is the gas temperature,  $L_X$  is the gas X-ray luminosity and  $V$  is the gas volume contained within each annulus. We use the deprojected spectral quantities as an estimate of the plasma parameters and show our results in Fig. 3. However, we



**Figure 3.** Cooling time profiles for all clusters in our sample for which we were able to identify cavities. Each colour illustrates a different cluster. The clusters shown with the black lines are those in which the *Chandra* data were too poor to deproject the spectra. In this case, we deprojected the surface brightness profiles to obtain an estimate of the deprojected  $n_e$  and  $L_X$ . Using these values, combined with the average cluster temperature within  $r < 200$ , we determined  $t_{\text{cool}}$ . MACS J2135.2–0102 is the only object in our sample with cavities that has a central cooling time larger than 10 Gyr (shown with the black dotted line).

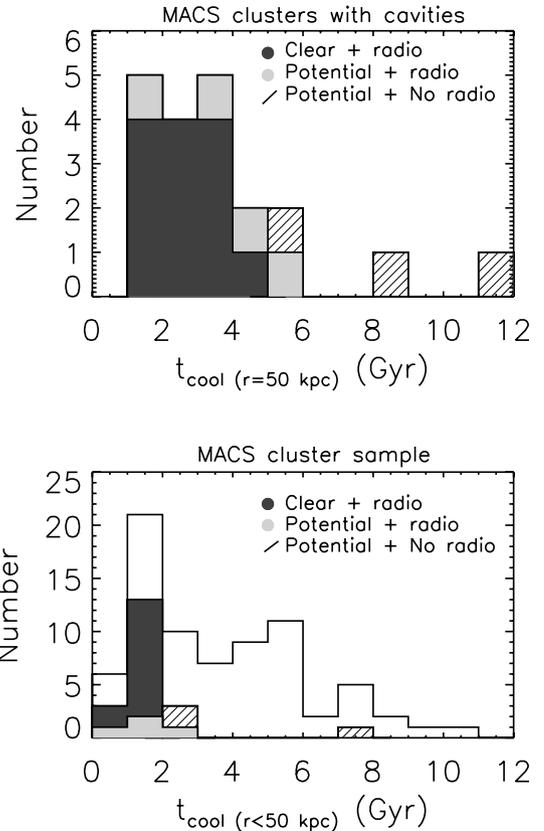
**Table 4.** Cooling flow properties of the MACS clusters with cavities – (1) name; (2) bolometric X-ray luminosity (0.01–50 keV) estimated from the *Chandra* data; (3) cooling time at  $r = 50$  kpc; (4) cooling radius estimated as the radius at which  $t_{\text{cool}} = 7.7$  Gyr; (5) bolometric X-ray luminosity (0.01–100 keV) of the clusters within the cooling radius.

(1) Cluster name	(2) $L_X$ (0.01–100 keV) ( $10^{44}$ erg s $^{-1}$ )	(3) $t_{\text{cool}}$ (50 kpc) (Gyr)	(4) $r_{\text{cool}}$ (kpc)	(5) $L_X (< r_{\text{cool}})$ (0.01–100 keV) ( $10^{44}$ erg s $^{-1}$ )
MACS J0159.8–0849	$23.1^{+0.7}_{-0.6}$	3.0	86	$11.8^{+0.4}_{-0.3}$
MACS J0242.5–2132	$21.1^{+0.8}_{-0.8}$	2.3	110	$16.5^{+0.6}_{-0.6}$
MACS J0429.6–0253	$16.8^{+0.7}_{-0.6}$	3.1	105	$11.6^{+0.5}_{-0.5}$
MACS J0547.0–3904	$6.9^{+0.3}_{-0.3}$	2.7	100	$4.8^{+0.2}_{-0.2}$
MACS J0913.7+4056	$19.3^{+0.4}_{-0.4}$	2.1	107	$15.2^{+0.3}_{-0.3}$
MACS J0947.2+7623	$37.6^{+0.7}_{-0.7}$	1.6	120	$29.9^{+0.6}_{-0.6}$
MACS J1411.3+5212	$8.9^{+0.3}_{-0.3}$	4.8	60	$4.0^{+0.2}_{-0.2}$
MACS J1423.8+2404	$23.3^{+0.5}_{-0.4}$	2.1	95	$16.3^{+0.3}_{-0.3}$
MACS J1532.8+3021	$33.2^{+1.1}_{-1.1}$	1.5	115	$24.9^{+0.9}_{-0.8}$
MACS J1720.2+3536	$15.1^{+0.4}_{-0.4}$	3.8	100	$8.9^{+0.3}_{-0.3}$
MACS J1931.8–2634	$9.0^{+0.1}_{-0.1}$	1.8	112	$6.50^{+0.09}_{-0.09}$
MACS J2046.0–3430	$13.4^{+0.4}_{-0.4}$	3.1	81	$8.3^{+0.3}_{-0.3}$
MACS J2140.2–2339	$18.6^{+0.2}_{-0.2}$	1.9	107	$14.2^{+0.2}_{-0.2}$
<sup>a</sup> MACS J0111.5+0855	$5.3^{+0.6}_{-0.5}$	5.7	78	$2.3^{+0.5}_{-0.3}$
MACS J0257.1–2325	$17.4^{+1.6}_{-1.2}$	8.0	54	$3.4^{+0.6}_{-0.4}$
MACS J1359.1–1929	$7.5^{+0.3}_{-0.3}$	3.9	76	$4.2^{+0.2}_{-0.2}$
MACS J1359.8+6231	$8.8^{+0.4}_{-0.3}$	5.9	67	$2.9^{+0.1}_{-0.1}$
MACS J1447.4+0827	$44.0^{+1.4}_{-1.3}$	1.3	125	$36.2^{+1.2}_{-1.1}$
<sup>a</sup> MACS J2135.2–0102	$4.1^{+0.4}_{-0.3}$	11.9	–	–
MACS J2245.0+2637	$11.9^{+0.6}_{-0.6}$	4.0	89	$6.3^{+0.4}_{-0.4}$

<sup>a</sup>The *Chandra* data of these clusters were too poor to deproject the spectra. In this case, we deprojected the surface brightness profiles to obtain an estimate of the deprojected  $n_e$  and  $L_X$ . Using these values, combined with the average cluster temperature within  $r < 200$ , we determined  $t_{\text{cool}}$ .

show in black the two clusters for which we were not able to deproject the spectra (MACS J0111.5+0855 and MACS J2135.2–0102). In this case, we deprojected the surface-brightness distribution, and then converted the counts per second of each annulus into the predicted electron density and flux of a *MEKAL* model for which the temperature and abundance were frozen at the average value within  $r < 200$  kpc. We then used these values of  $n_e$  and  $L_X$  combined with the average cluster temperature within  $r < 200$  kpc to determine  $t_{\text{cool}}$  using equation (7). This is possible since X-ray observations in clusters give very precise estimates of the surface brightness profiles. It is therefore much easier to deproject the surface brightness profiles than the spectra, allowing us to obtain an estimate of the ‘deprojected’ cooling time profiles for these two objects.

Fig. 3 shows that many of the clusters in which we detected cavities are strong cool cores, with cooling times reaching down to  $\sim 0.5$  Gyr in the central regions. The estimates we obtain for the central cooling times are shown in Table 4. We define the central cooling time as the cooling time we derive at a radius of 50 kpc and not the cooling time of the inner most bin. This is because the radius of the inner most bin depends not only on the cluster distance, but more importantly on the data quality, and is therefore not necessarily the best estimate of the central cooling time (see also Bauer et al. 2005). For almost all of the clusters in our sample, we have cooling time estimates below 50 kpc, allowing us to obtain an



**Figure 4.** Top panel: distribution of  $t_{\text{cool}}$  at 50 kpc for the MACS clusters in which we identified cavities. Here, we derived the central cooling time by extracting a series of annuli and determining the cooling time at  $r = 50$  kpc. Bottom panel: distribution of  $t_{\text{cool}}$  within 50 kpc ( $r < 50$  kpc) for all 76 MACS clusters with *Chandra* observations. Among these, we outline those with clear bubbles in black (all of these have a radio source associated with the central galaxy) and those with potential cavities and a radio source in grey, as well as those with potential cavities but no radio source in diagonal lines.

accurate central cooling time at 50 kpc. For the remaining clusters, we determine a rough estimate by extrapolating the values down to 50 kpc.

Adopting this definition of the central cooling time, we find that 19 of the 20 clusters in which we identified cavities are mild cool cores with  $t_{\text{cool}} < 8$  Gyr and that all of the systems with ‘clear’ cavities are strong cool cores with  $t_{\text{cool}} < 3$ –5 Gyr. We also illustrate this in the top panel of Fig. 4, where we plot the distribution of cooling times in our sample of clusters with cavities.

Since many clusters at  $z = 1$  appear to have similar properties to present-day ones, they can be regarded as relaxed systems, and more importantly, we can assume that a cooling flow should have had the time to establish itself since then. We therefore choose to define the cooling radius of each cluster as the radius at which the cooling time is equal to the  $z = 1$  look-back time. For our cosmology, this corresponds to 7.7 Gyr. This definition of the cooling radius was also adopted by Rafferty et al. (2006), allowing us to directly compare our results with theirs. Although there are other statistical studies of cavities in the literature (e.g. Bîrzan et al. 2004, 2008; Dunn et al. 2005; Dunn & Fabian 2006, 2008), many of the objects have already been included in Rafferty et al. (2006).

Finally, we derive the bolometric X-ray luminosity of each cluster [ $L_X(0.01$ –100 keV)], and the cooling luminosity which

we define as the bolometric X-ray luminosity (0.01–100 keV) within the cooling radius,  $L_X(< r_{\text{cool}})$ . When deriving these quantities, we exclude the central point source ( $r < 1$  arcsec) for clusters with point-like nuclei at their centres, since some have important non-thermal emission associated with them. These include MACS J0547.0–3904, MACS J0913.7+4056, MACS J0947.2+7623, MACS J1411.3+5212, MACS J1423.8+2404, MACS J1931.8–2634 and MACS J2046.0–3430.

## 6.2 MACS cluster sample

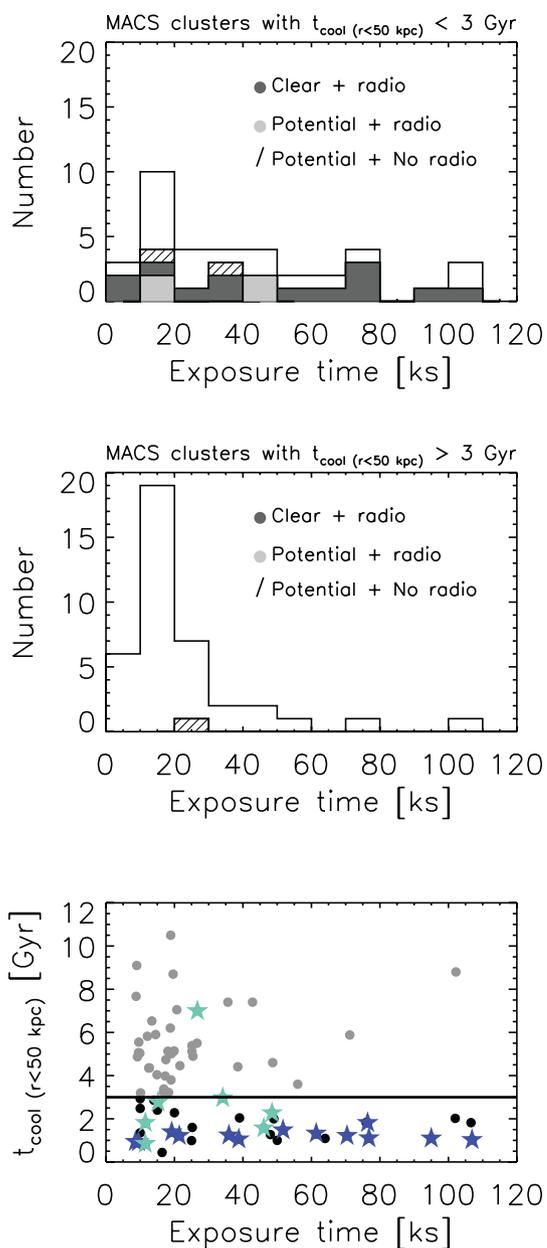
In order to derive the fraction of cool core clusters with cavities in the MACS sample, we computed the central cooling time within 50 kpc for all 76 MACS clusters with *Chandra* archival data. Here, we extracted a spectrum for each cluster within  $r < 50$  kpc and fitted a MEKAL model to the spectrum while letting the temperature, abundance and normalization parameter free to vary. For some clusters, the counts were not sufficient to constrain the fit. In this case, we kept the abundance and temperature frozen at the value derived for the entire cluster ( $r < 200$  kpc).

We could not compute detailed cooling time profiles for all the MACS clusters since the data quality in many is not sufficient to derive deprojected or even projected profiles. We therefore choose to use the cooling time within the central 50 kpc [ $t_{\text{cool}}(r < 50 \text{ kpc})$ ] as an estimate of the central cooling time. This definition of the central cooling time is not the same as the one used previously and defined as the cooling time derived at a radius of 50 kpc [ $t_{\text{cool}}(r = 50 \text{ kpc})$ ]. To determine the fraction of cool core clusters with cavities in the MACS sample, we therefore recalculated the central cooling time of clusters with cavities adopting the new definition [ $t_{\text{cool}}(r < 50 \text{ kpc})$ ]. Our results are shown in the bottom panel of Fig. 4.

We find that 37 of the 76 MACS clusters have a strong central cool core [ $t_{\text{cool}}(r < 50 \text{ kpc}) < 3 \text{ Gyr}$ ], corresponding to a cool core fraction of 50 per cent. Bauer et al. (2005) derive detailed cooling time profiles for 38 X-ray luminous clusters from the *ROSAT* BCS survey ( $0.15 < z < 0.4$ ) and find that 11 have a central cooling time less than 3 Gyr (and 55 per cent less than 10 Gyr). They use the cooling time derived at a radius of 50 kpc as their definition of the central cooling time, whereas we use the cooling time derived *within* 50 kpc for the central cooling time which gives smaller cooling times. Nonetheless, we find that a significant fraction in our sample are cool core clusters and, out of these 37 clusters, we find that 19 have cavities, corresponding to a cavity detection rate in cooling systems of 50 per cent. Note, however, that more than half of the remaining cool cores [ $t_{\text{cool}}(r < 50 \text{ kpc}) < 3 \text{ Gyr}$ ] with no cavities have an exposure time below 30 ks (see top panel of Fig. 5), which could explain why we do not detect cavities in these systems. We also plot the central cooling time as a function of *Chandra* exposure time for all MACS clusters in the bottom panel of Fig. 5, and discuss these results in Section 9.1.

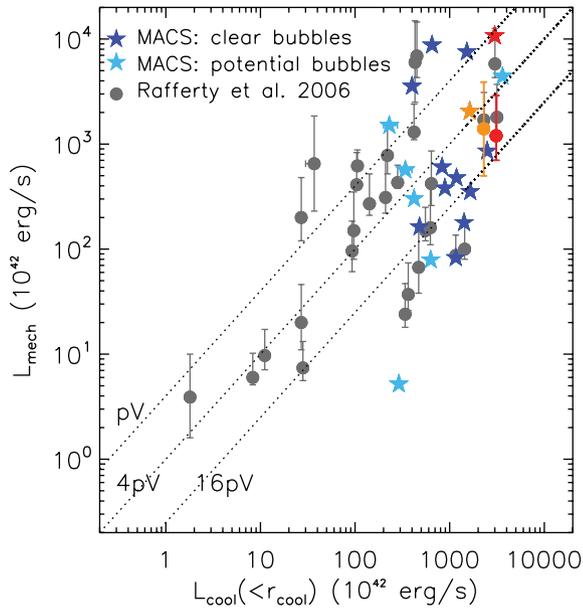
## 7 CAVITY SCALING RELATIONS

We now address the various scaling relations known for the cavity population in groups and clusters. It has been known for quite some time that AGN feedback is finely tuned to the energy needed to offset cooling of the ICM. If the input of energy were too low, then the core would cool rapidly and cause massive quantities of gas to condense and form stars at rates much higher than what we observe. However, if the input of energy were too high, then this would cause the core to overheat and we would not observe the low central cooling times in the cores of some clusters.



**Figure 5.** *Chandra* exposure time distribution for the 37 MACS clusters with  $t_{\text{cool}}(r < 50 \text{ kpc}) < 3 \text{ Gyr}$  (top panel) and for the 39 MACS clusters with  $t_{\text{cool}}(r < 50 \text{ kpc}) > 3 \text{ Gyr}$  (middle panel). Among these, we outline the 13 with clear bubbles in black (all of these have a radio source associated with the central galaxy) and those with potential cavities and a radio source in grey, as well as those with potential cavities but no radio source in diagonal lines. Note that more than half of the cool cores [ $t_{\text{cool}}(r < 50 \text{ kpc}) < 3 \text{ Gyr}$ ] with no cavities have an exposure time below 30 ks, which could explain why we do not detect cavities in these systems. Also shown in the bottom panel is the central cooling time as a function of *Chandra* exposure time. The MACS clusters with clear bubbles are shown in dark blue, and those with potential bubbles are shown in light blue. The straight line shows the cut-off between those with  $t_{\text{cool}}(r < 50 \text{ kpc}) > 3 \text{ Gyr}$  (light grey circles) and those with  $t_{\text{cool}}(r < 50 \text{ kpc}) < 3 \text{ Gyr}$  (black circles).

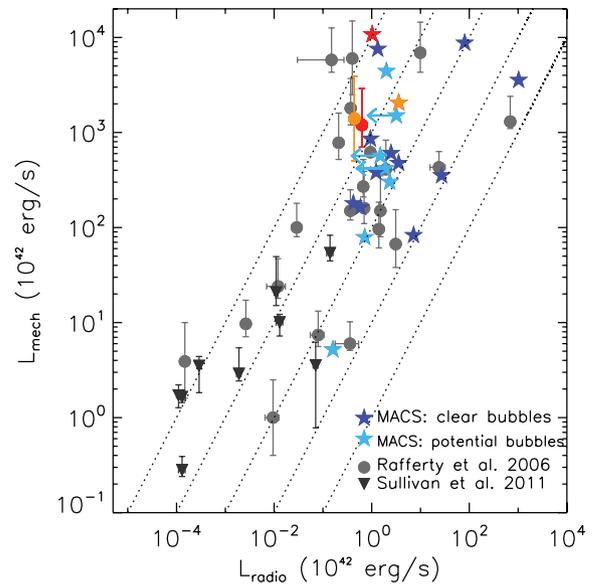
To illustrate the fine tuning of AGN feedback, we plot the mechanical power arising from the central AGN as a function of the bolometric X-ray luminosity within the cooling radius [ $L_{\text{cool}}(0.01\text{--}100 \text{ keV})$ ] in Fig. 6. Within this radius, the gas should have had



**Figure 6.** Mechanical energy being injected into the ICM by the central AGN (estimated from the energetics of the cavities) as a function of energy needed to offset cooling of the ICM by the central galaxy. Those common to both samples are shown in red and yellow.

the time to cool if no additional heating sources were present. Although there is some cooling happening in the cores of clusters, studies suggest it is only of the order of 10 per cent (Bîrzan et al. 2004; Rafferty et al. 2006) or even less than 10 per cent according to the more recent studies with high-resolution spectroscopy (Sanders et al. 2008, 2010), and therefore does not significantly affect the relation plotted in Fig. 6. The mechanical power is estimated using the energy stored within the cavities of each system ( $4pV$ ) and the buoyancy rise time ( $t_{\text{buoy}}$ ). We also include the cavities of Rafferty et al. (2006) and illustrate in red the two systems common to both samples (MACS J1423.8+2404 and MACS J0947.2+7623). For MACS J1423.8+2404, our results agree well with those of Rafferty et al. (2006), but there is a large discrepancy between our estimate of the mechanical power and that of Rafferty et al. (2006) in MACS J0947.2+7623. However, the more recent study on MACS J0947.2+7623 by Cavagnolo et al. (2011) finds an outburst power consistent with our estimate of  $L_{\text{mech}}$  ( $\sim 6 \times 10^{45} \text{ erg s}^{-1}$ ). We do not include the data points of Dunn et al. (2005) or Dunn & Fabian (2006, 2008) as they define the cooling radius as the radius at which  $t_{\text{cool}} = 3 \text{ Gyr}$ , nor the data points of Cavagnolo et al. (2010) as they only calculate the cooling luminosity within 0.3–2.0 keV. Fig. 6 shows that the MACS sample contains some of the most powerful outbursts known, and that even at these extreme energies, the central AGN is able to offset cooling of the ICM at a fine level.

We also look at the properties of the mechanical power arising from the central AGN as a function of radio luminosity. As these cavities are being created by the central AGN and being filled with relativistic electrons, one would expect a correlation between the mechanical power and radio luminosity, which is a measure of the radio power arising from the AGN. This has been looked at in detail for clusters, ellipticals and more recently groups of galaxies (Bîrzan et al. 2004, 2008; Cavagnolo et al. 2010; O’Sullivan et al. 2011), although the scatter remains large and extends over several orders of magnitude ( $>3$ ). We have computed the radio luminosities for our sample (see Table 2) and show in Fig. 7 the mechanical power



**Figure 7.** Mechanical energy being injected into the ICM by the central AGN (estimated from the energetics of the cavities) as a function of total integrated radio luminosity (10 MHz to 10 GHz) of the central source. Those common to both samples are shown in red and yellow.

arising from the central AGN ( $4pV/t_{\text{buoy}}$ ) as a function of radio luminosity ( $L_{\text{radio}}$ ). We include the data points of O’Sullivan et al. (2011) for galaxy groups, as well as those in Rafferty et al. (2006), but use the estimate of  $L_{\text{radio}}$  given in Bîrzan et al. (2008). We do not include the data points of Cavagnolo et al. (2010) as they do not include an estimate of  $L_{\text{radio}}$  in their paper. We discuss this plot in Section 9.2.

## 8 RESULTS

In this section we review the results of each system. We begin by addressing the systems with ‘clear’ cavities, and then address those with ‘potential’ cavities. Many of the MACS clusters have been studied in terms of their mass distribution and cosmological implications (e.g. Puetzfeld, Pohl & Zhu 2005; Allen et al. 2008; Ettori et al. 2009), as well as in terms of how the cluster metallicity content evolves with redshift (Balestra et al. 2007; Maughan et al. 2008). These include MACS J0159.8–0849, MACS J0242.5–2132, MACS J0429.6–0253, MACS J0947.2+7623, MACS J1423.8+2404, MACS J1532.8+3021, MACS J1720.2+3536, MACS J1931.8–2634, MACS J2046.0–3430, MACS J2140.2–2339 and MACS J1359.8+ 6231.

### 8.1 Clusters with ‘clear’ cavities

#### 8.1.1 MACS J0159.8–0849

We find two clear X-ray cavities in this system. They are located symmetrically opposite with respect to the central AGN, and at a distance of  $\sim 25 \text{ kpc}$  from the core. The power stored within them ( $P_{\text{buoy}} \sim 5 \times 10^{44} \text{ erg s}^{-1}$ ) falls short by a factor of 2 from the cooling luminosity. The *HST* optical image shows interesting features, including a potential double or triple nucleus with streams of material leading back to the central nucleus.

### 8.1.2 MACS J0242.5–2132

We find two cavities in this cluster, with the south-west one being clearer and somewhat larger than its northern counterpart. The optical image shows that there is a bright central isolated galaxy at the centre, which coincides with a very strong radio source of flat spectrum ( $\sim 1$  Jy bright, the second strongest in our sample). The power stored within the cavities ( $P_{\text{buoy}} \sim 4 \times 10^{44} \text{ erg s}^{-1}$ ) falls short by a factor of 4 from the cooling luminosity. This cluster is part of the 34 most luminous MACS clusters (Ebeling et al. 2010).

### 8.1.3 MACS J0429.6–0253

This cluster has two clear cavities being inflated by the central AGN. The cavities are quite small ( $r \sim 5$  kpc), and the power stored within them falls short by almost an order of magnitude when compared to the cooling luminosity ( $P_{\text{buoy}} \sim 10^{44} \text{ erg s}^{-1}$ ). There is a hint of another potential cavity to the north-east, but deeper X-ray observations are needed to confirm this. The optical image shows a dominant central galaxy harbouring a bright nucleus. This cluster is also part of the 34 most luminous MACS clusters (Ebeling et al. 2010).

### 8.1.4 MACS J0547.0–3904

We find one cavity in this cluster of radial dimensions  $8.5 \times 11.5 \text{ kpc}^2$ , surrounded by bright rims. The power stored within it alone ( $P_{\text{buoy}} \sim 2 \times 10^{44} \text{ erg s}^{-1}$ ) can prevent the gas from cooling to almost half of the cooling luminosity. The dominant central galaxy harbours a bright point-like nucleus which shows hints of a jet in the optical, aligned with X-ray cavity jet axis. The X-ray image also shows a bright nucleus dominating in emission. This cluster is part of the 34 most luminous MACS clusters (Ebeling et al. 2010).

### 8.1.5 MACS J0913.7+4056

The central galaxy in this cluster, more commonly known in the literature as IRAS 09104+4109, is a hyperluminous infrared cD galaxy ( $L_{\text{IR}} \sim 12 \times 10^{12} L_{\odot}$ ; Kleinmann et al. 1988) hosting a heavily obscured Type 2 AGN ( $\sim 1\text{--}5 \times 10^{23} \text{ cm}^{-2}$ ; Vignali et al. 2011). The *Chandra* X-ray spectrum of the nucleus shows that there is a prominent Fe  $K\alpha$  line at the redshift of the source (see also Iwasawa, Fabian & Ettori 2001).

We find two cavities in this system. The cavity to the north-west was first noticed by Iwasawa et al. (2001), but the 9.1 ks *Chandra* observations were too faint to confirm whether this cavity was real. By using a much deeper observation of 70.4 ks, we confirm the existence of the first cavity and find evidence of a second located at a similar radius to the south-east. Although larger, the cavities carved out in the X-ray image coincide with the location of the radio hotspots seen at 20 cm (Hines & Wills 1993, also see Fig. 2 where we show the radio jet axis), and the optical image shows a spectacular set of filaments spreading towards the north-east and south. These filaments run along both cavities that we have identified. However, one of these filaments extends further out [ $r_{\text{max}(\text{filament})} = 51$  kpc and  $r_{\text{max}(\text{cavity})} = 55$  kpc] and is directed more eastwards than the northern cavity. As seen in many nearby clusters, including M87 in the Virgo cluster (Bohringer et al. 1995; Churazov et al. 2001) and NGC 1275 in the Perseus Cluster (Fabian et al. 2003), filaments usually trail behind cavities, and could indicate that there is another, older cavity to the north-east, although it is difficult to tell with the current X-ray observations. The power stored within the cavities is

substantial ( $P_{\text{buoy}} \sim 8 \times 10^{45} \text{ erg s}^{-1}$ ) and is more than sufficient to prevent the ICM from cooling. This cluster harbours the third most powerful outburst in our sample (in terms of total power), and has the cavity with the largest enthalpy in those classified as ‘clear’ cavities ( $pV \sim 10^{60} \text{ erg s}^{-1}$ ). MACS J0913.7+4056 has one of the most powerful AGN outbursts known in clusters.

### 8.1.6 MACS J0947.2+7623

MACS J0947.2+7623 (RBS 0797) is another cluster known to have one of the most powerful AGN outbursts (see Cavagnolo et al. 2011, for a detailed analysis on the AGN outburst energetics). It has a very strong cool core, requiring extreme mechanical feedback from its central AGN ( $L_{\text{cool}} \sim 3 \times 10^{45} \text{ erg s}^{-1}$ ), and the X-ray image shows two large X-ray cavities surrounded by bright rims, as well as a point-like nucleus. This cluster is also part of the 34 most luminous MACS clusters (Ebeling et al. 2010). In the optical, we see a large central galaxy hosting a complex and disturbed core. The power stored within each cavity [ $P_{\text{buoy}} \sim (4\text{--}7) \times 10^{45} \text{ erg s}^{-1}$ ] is more than sufficient to prevent the surrounding gas from cooling. This cluster has the most powerful outburst in our sample and the energetics we derive are consistent with Cavagnolo et al. (2011), except that we estimate the size of the north-east cavity to be slightly larger.

Interestingly, Gitti, Feretti & Schindler (2006) found that the 4.8 GHz radio jet axis is perpendicular to the jet axis running through the cavities (see Fig. 2 for the direction of the 4.8-GHz radio jet axis). This was later confirmed by Bîrzan et al. (2008). Cavagnolo et al. (2011) also suggested that there is a hint of a small, newly formed cavity at the tip of the southern 4.8 GHz jet. These authors furthermore find that the bright rims surrounding the cavities could be driving weak shocks into the ICM and suggest that the unusually deep surface brightness depressions of the cavities, combined with the ambiguous correlation between the X-ray and radio morphologies, are more consistent with these cavities being elongated along the line of sight. These cavities could therefore be larger than measured in the plane of the sky.

### 8.1.7 MACS J1411.3+5212

MACS J1411.3+5212 (3C295) hosts a very powerful central AGN embedded in a giant cD galaxy. This galaxy has a bright radio core and is one of the most powerful Fanaroff–Riley type II (FR II) sources known. The cluster shows evidence of non-thermal emission filling the inner radio lobes, consistent with ICCMB photons (e.g. Harris et al. 2000; Allen et al. 2001; Brunetti et al. 2001). The radio maps show GHz radio emission filling the inner  $\sim 25$  kpc of the core, with a jet axis aligned in the north-west to south-east direction (Perley & Taylor 1991; Taylor & Perley 1992; Harris et al. 2000; Allen et al. 2001; Brunetti et al. 2001, see also Fig. 2 for the direction of the GHz radio jet axis). Brunetti et al. (2001) also find a prominent Fe K line at the redshift of the source associated with the nucleus of the central galaxy, consistent with this object being a powerful absorbed AGN.

During the spectral analysis of this source (Sections 4–6), we omitted the central regions containing the ICCMB X-ray emission since our aim was to extract the properties of the thermal gas. The cavity that we find is located at the edge of south-east ICCMB X-ray lobe, and could be an older cavity carved in the past by the radio source but which has now risen buoyantly. The radio maps at GHz frequencies show that the emission is limited to the inner  $r < 25$  kpc (Perley & Taylor 1991; Taylor & Perley 1992; Harris et al. 2000;

Allen et al. 2001; Brunetti et al. 2001). The cavity we find is beyond this radius ( $r_{\text{cavity}} \sim 35$  kpc), and is consistent with being older since the GHz radio observations do not seem to show emission extending out to this radius. This cavity is large and the power stored within it ( $P_{\text{buoy}} \sim 4 \times 10^{45}$  erg s $^{-1}$ ) is more than sufficient to prevent the gas from cooling.

#### 8.1.8 MACS J1423.8+2404

This cluster hosts a very large central galaxy harbouring a bright nucleus seen both at optical and X-ray wavelengths. In the optical, the *HST* image shows that the central galaxy is heavily disturbed: the bright point-like nucleus is surrounded by a thick envelope in the inner regions ( $r < 2.6$  kpc) and a more diffuse and disrupted envelope out to  $r \sim 20$  kpc. The optical image also shows two depressions, most likely due to dust: one to the west which coincides with the location of the cavity seen at X-ray wavelengths and a hint of a second to the north-east which coincides with a very small depression seen in the X-rays (although we have not included this inner smaller cavity in our sample). The power stored in the cavities ( $P_{\text{buoy}} \sim 2 \times 10^{45}$  erg s $^{-1}$ ) is sufficient to prevent the gas from cooling within the cooling radius. This cluster is also the most distant object in the MACS sample in which we have identified cavities ( $z = 0.5449$ ), and is part of the 12 most distant clusters detected by MACS (Ebeling et al. 2007). The bright nucleus does not show any evidence of a hidden quasar in the form of Fe K lines at the redshift of the source, as opposed to other MACS clusters mentioned previously with a bright X-ray nucleus. Kartaltepe et al. (2008) looked at the large-scale structure of the 12 most distant MACS clusters by tracing the surface brightness density of galaxies near the red sequence, and singled out that MACS J1423.8+2404 was the most relaxed cluster in the sample. Limousin et al. (2010) also confirm that the cluster seems to be nearly fully virialized as suggested by gravitational lensing analysis. Having such a virialized cluster only  $\sim 7$  Gyr after the big bang puts constraints on structure formation and evolution in a cosmological context.

#### 8.1.9 MACS J1532.8+3021

MACS J1532.8+3021 (RX J1532.9+3021) has the fourth most infrared luminous BCG in the Quillen et al. (2008) sample of 62 X-ray luminous clusters, enough that it can be classified as a luminous infrared galaxy (LIRG) with  $L_{\text{IR}} = 22.6 \times 10^{44}$  erg s $^{-1}$ . The central galaxy also has one of the most massive molecular gas detections known ( $2.5 \times 10^{11} M_{\odot}$ ; the most massive detection in Edge 2001) and is the second most optically line-luminous central cluster galaxy in the sample of Crawford et al. (1999) with an optical line emission of  $4.2 \times 10^{42}$  erg s $^{-1}$ . The galaxy is also particularly blue, suggesting star formation, and has the largest star formation region in the O’Dea et al. (2008) cluster sample, with a scale length of  $\sim 70$  kpc. The cluster is also particularly massive  $M > 10^{15} M_{\odot}$ , as found by the weak lensing study of Dahle et al. (2002).

We find one clear cavity associated with the central galaxy, located to the west. It is quite large ( $r \sim 20$  kpc) and situated almost 40 kpc from the nucleus. This cluster has the third strongest cooling luminosity in our MACS sample with cavities, requiring around  $3 \times 10^{45}$  erg s $^{-1}$  to prevent the gas from cooling. We find that the power stored within the cavity falls short by a factor of  $\sim 2$ – $3$  of the cooling luminosity. The *HST* image suggests that there is some filamentary structure surrounding the central galaxy (although it

is not as clear as in MACS 0913.7+4056), and one of the northern filaments seems to sweep around the bottom portion of the cavity.

#### 8.1.10 MACS J1720.2+3536

We find two cavities associated with the central regions of MACS J1720.2+3536 (Z8201), one to the north and another, though very small and not as clear, to the south-east. If we consider the power stored within the cavities using the buoyancy time-scale, then the power falls short by a factor of 2 of the cooling luminosity. If we consider the sound-crossing time, then the power is more than sufficient to offset cooling of the ICM. The *HST* image shows that the central galaxy has a bright core surrounded by a large diffuse envelope. Rafferty, McNamara & Nulsen (2008) also find that the central galaxy has blue emission extending out to 2.5 arcsec ( $\sim 14$  kpc). This cluster is part of the 34 most luminous MACS clusters (Ebeling et al. 2010).

#### 8.1.11 MACS J1931.8–2634

MACS J1931.8–2634 is a strong cool core cluster harbouring one of the most powerful AGN outbursts known. The central nucleus has carved two large X-ray cavities to the west and east, greatly disrupting the core (Ehlert et al. 2011). Diffuse 1.4 GHz radio emission fills the central regions and extends in the west–east direction, but no clear radio lobes or jets are seen. The X-ray image shows two bright ridges formed to the north and south, some 25 kpc from the nucleus. The north ridge is also made of cool, metal-rich gas rich in H $\alpha$  filaments and young stars. Ehlert et al. (2011) propose that these ridges were formed by the AGN outburst, as well as sloshing of the core along the north to south direction induced by a past merger. Ehlert et al. (2011) found that the temperature profile decreased with radius, but that in the innermost regions ( $r < 30$  kpc), the temperature increased once more. This could indicate shock heating. In Fig. B1, we show the spectral profiles of MACS J1931.8 we obtained during our analysis. Our temperature profile also increases in the innermost regions ( $r < 30$  kpc).

The cavity edges are not well defined, and the radio emission shows no clear radio lobes. Determining a precise measurement of the cavity sizes is therefore difficult. Taking this into account, we estimate that the  $4pV$  energy stored within the cavities is roughly  $\sim 3 \times 10^{60}$  erg and the power input is  $P_{\text{buoy}} \sim 8$ – $9 \times 10^{45}$  erg s $^{-1}$  or using the time-crossing time,  $P_{\text{cs}} \sim (4$ – $5) \times 10^{45}$  erg s $^{-1}$ . Ehlert et al. (2011) find a total energy of  $\sim (1$ – $8) \times 10^{60}$  erg ( $4pV$ ) and  $P_{\text{cs}} \sim (4$ – $14) \times 10^{45}$  erg s $^{-1}$ , both consistent with our results. MACS J1931.8–2634 therefore has one of the most powerful AGN outbursts known. This cluster is also part of the 34 most luminous MACS clusters (Ebeling et al. 2010).

#### 8.1.12 MACS J2046.0–3430

We find two cavities associated with this cluster. The first is located to the south, and has a flat shape with bright rims surrounding it. The second cavity is located to the north but is less well defined. We were not able to find any *HST* archival data and show in Fig. 2 the  $R$  passband UH 2.2-m image. The optical image shows that the central dominant galaxy is offset from the X-ray point source by  $\sim 2.7$  arcsec (15 kpc). The power stored in the cavities ( $P_{\text{buoy}} \sim 6 \times 10^{44}$  erg s $^{-1}$ ) is comparable to the amount of energy needed to prevent the gas from cooling.

### 8.1.13 MACS J2140.2–2339

MACS J2140.2–2339 (MS 2137.3–2353) is a luminous cluster, dominated by a large central galaxy. The most interesting feature seen in the optical image is a large radial lensing arc some 70 kpc from the nucleus, which was the first of its kind to be discovered (Fort et al. 1992) and has been looked at in detail by many authors (e.g. Gavazzi et al. 2003; Gavazzi 2005; Comerford et al. 2006; Sand et al. 2008; Merten et al. 2009). This cluster is part of the 34 most luminous MACS clusters (Ebeling et al. 2010). We find a small cavity to the south of the nucleus, only capable of preventing the gas from cooling to 10 per cent of the cooling luminosity. Although we classify this cavity as being ‘clear’, it is very small and does not have any bright rims surrounding it.

## 8.2 Clusters with ‘potential’ cavities

We now address the systems with ‘potential’ cavities. Although we found evidence for cavities in these clusters, many of the X-ray observations have a low number of counts, and deeper observations are needed to confirm if these cavities are real.

### 8.2.1 MACS J0111.5+0855

We find two cavities associated with the dominant central galaxy in this cluster. They are both quite large (30 kpc in length) and located some 60–70 kpc from the nucleus. However, the data quality for this system is quite poor, and we were not able to deproject the spectra. Furthermore, we were only able to extract two annuli to derive the variations of the thermal properties with radius (see Fig. B1). Although the central cooling time we derive at  $r = 50$  kpc is below 10 Gyr, it is not clear whether this cluster has a cool core because the data quality does not allow us to obtain a detailed temperature profile in the inner regions. This cluster also has no central radio source associated with it. If we assume that this cluster has a small cool core, then the cooling luminosity within the cooling radius ( $r_{\text{cool}} = 78$  kpc) is  $\sim 2 \times 10^{44}$  erg s $^{-1}$ . The cavity power is of the order of  $P_{\text{buoy}} \sim 15 \times 10^{44}$  erg s $^{-1}$ , but since these cavities are not very well defined (and only classified as potential cavities), this value could be quite uncertain.

### 8.2.2 MACS J0257.1–2325

This cluster is located at a redshift of  $z = 0.5039$  and is part of the 12 most distant clusters detected by the MACS survey at  $z > 0.5$  (Ebeling et al. 2007). Kartaltepe et al. (2008) looked at the large-scale structure in the 12 most distant MACS clusters (including MACS J1423.8+2404), and found that MACS J0257.1–2325 has the most complex system in terms of large-scale structure (on Mpc scales), with infalling systems along cosmic filaments. The *HST* image in Fig. 2 shows that the cluster is dominated by a large central galaxy, hosting a bright core and diffuse envelope. We also find many lensed arcs surrounding the central regions ( $r < 70$  kpc), including a very bright one to the south-west. We find no radio source associated with the central galaxy, and only a mild cooling flow with a central cooling time of  $\sim 8$  Gyr. Although, we were able to deproject the spectra, it is not clear if the temperature profile is decreasing in the inner regions because the error bars are still quite large. The cavity we find is located to the north-east some 40 kpc from the nucleus and the power stored within ( $P_{\text{buoy}} \sim 6 \times 10^{44}$  erg s $^{-1}$ ) is sufficient to prevent the gas from cooling.

### 8.2.3 MACS J1359.1–1929

This cluster hosts a large central galaxy, surrounded by a diffuse envelope as seen in the *HST* image, and has a radio source associated with the central regions. We also find a central cooling time of 4 Gyr and a temperature profile steadily declining in the inner regions, with a cooling luminosity of the order of  $4 \times 10^{44}$  erg s $^{-1}$ . We were able to identify two cavities in this system, although not well defined and deeper data are needed to confirm their existence. If these cavities are real, then the power stored within them is  $P_{\text{buoy}} \sim 3 \times 10^{44}$  erg s $^{-1}$ . The X-ray image also shows three small depressions ( $\sim 6$  kpc in size) immediately surrounding the central X-ray peak, which could also contribute to the heating.

### 8.2.4 MACS J1359.8+6231

MACS J1359.8+6231 (MS 1358.4+6245 or EMSS 1358+6245) is a relaxed cool core cluster known for its lensing properties (Franx et al. 1997; Allen 1998; Hoekstra et al. 1998; Zitrin et al. 2011) and detailed mass estimates (Arabadjis, Bautz & Garmire 2002).

We have identified one small cavity to the north-east located near the central galaxy. The optical image shows that the central regions are dominated by a large galaxy with a bright core and diffuse envelope. The power capabilities of the cavity are quite small with  $P_{\text{buoy}} \sim 5 \times 10^{42}$  erg s $^{-1}$  and are not capable of preventing the cool core from cooling by more than an order of magnitude. This cluster stands out in the  $L_{\text{mech}}-L_{\text{cool}}$  plot shown in Fig. 6 since the mechanical energy within the cavity fails to meet the power needed to offset cooling.

### 8.2.5 MACS J1447.4+0827

MACS J1447.4+0827 (RBS 1429) has been recently added to the MACS cluster sample. It is an ultraluminous cluster (the most X-ray luminous in our sample with cavities;  $L_{\text{X}} \sim 5 \times 10^{45}$  erg s $^{-1}$ ) and has a very strong cool core, therefore requiring extreme mechanical feedback from its central AGN [ $L_{\text{cool}} \sim (3-4) \times 10^{45}$  erg s $^{-1}$ ]. With just  $\sim 10$  ks on the source, we find evidence of a large cavity, although not well defined. The energetics of the cavity in itself are substantial, and the power stored within this outburst is enough to prevent the gas from cooling. This cluster has the most energetic outburst in our sample in terms of enthalpy ( $\sim 10^{60}$  erg), making it one of the most powerful outbursts known.

### 8.2.6 MACS J2135.2–0102

MACS J2135.2–0102 (1RXS J213515.7–010208) is a luminous cluster known for its lensing properties. An earlier study found a series of arcs in the shape of a ‘Cosmic Eye’ and ‘Cosmic Eyelash’ associated with this cluster (see Smal et al. 2007; Siana et al. 2009; Swinbank et al. 2010). However, the data quality of the *Chandra* X-ray image is very poor, and we were not able to deproject the spectra. We find no radio source associated with the central regions and that the central cooling time is more than 10 Gyr. There is, however, at least one depression in the X-ray image within the inner regions (see Fig. 2). Initially, we identified two cavities, one to the north-west and another to the south-east. However, as mentioned in Section 3, the BCG is located almost within the south-eastern cavity. This could be due to projection effects, but we choose to discard this cavity and only consider the north-western one. If the cavity to the north-west is real, then the power capability is quite large

( $P_{\text{buoy}} \sim 4 \times 10^{44} \text{ erg s}^{-1}$ ). The optical image also shows an interesting structure with a central galaxy hosting two or three potential nuclei.

### 8.2.7 MACS J2245.0+2637

Finally, we present the results for MACS J2245.0+2637. The *HST* image shows that the cluster hosts a central dominant galaxy dominated by a bright core. This cluster has a radio source associated with the central regions, as well as a short central cooling time of  $\sim 4$  Gyr ( $L_{\text{cool}} \sim 6 \times 10^{44} \text{ erg s}^{-1}$ ). We find one cavity in this system, but the data quality remains poor. The cavity is also quite small ( $\sim 8$  kpc in size) and the power stored within it is only  $P_{\text{buoy}} \sim 1 \times 10^{44} \text{ erg s}^{-1}$  or  $P_{\text{cs}} \sim 2 \times 10^{44} \text{ erg s}^{-1}$ . This cluster is also part of the 34 most luminous MACS clusters (Ebeling et al. 2010).

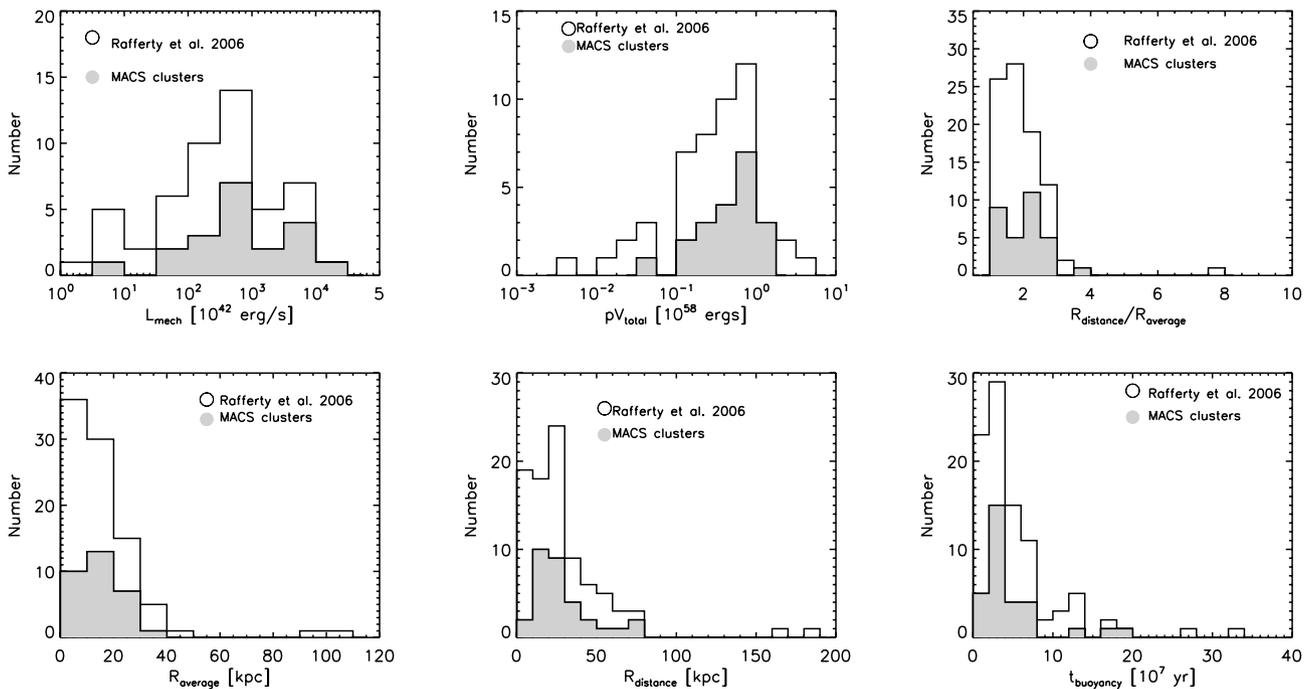
## 9 DISCUSSION

### 9.1 Selection effects

We have analysed the *Chandra* X-ray observations of the MACS cluster sample and searched for surface brightness depressions in these systems. The MACS cluster sample consists of luminous, massive clusters all above a redshift of  $z = 0.3$ . Out of the 76 clusters, we find that 13 have ‘clear’ surface brightness depressions associated with their central AGN. We interpret these depressions as cavities being inflated by the central SMBH, which then rise buoyantly into the outskirts. If we include the seven systems in which we find ‘potential’ surface brightness depressions, then the detection rate goes up to 26 per cent. We stress that many of the clusters in which we identified cavities were those that had the ‘deepest’ *Chandra* observations (more than 30 ks), and that many of the remaining MACS clusters only have 10–20 ks observations (or as little as 500–1000 counts within a radius of 200 kpc). Our detection rates should therefore be regarded as lower limits.

The 76 clusters that have been observed with *Chandra* were mostly selected on a flux-limited basis (Ebeling et al. 2007, 2010), and should therefore not be more biased towards cool cores than non-cool cores. As shown in Fig. 4, we find the same percentage of cool core clusters in the MACS sample (with *Chandra* data) than non-cool core clusters. However, the bottom panel of Fig. 5 shows that most of the clusters with deep observations are cool core clusters. This is because many of those that were initially observed for short exposure times and that showed interesting features, such as a peaked X-ray distribution, were followed-up with longer observations. We could therefore be missing many cavities in the non-cool core systems which do not have long observations. However, as shown by earlier studies (Bîrzan et al. 2004; Dunn et al. 2005; Dunn & Fabian 2006, 2008; Rafferty et al. 2006), X-ray cavities seem to lie predominantly in systems that require some form of heating, i.e. cool core clusters. Note also the *Chandra* observations targeted the brightest clusters in the sample. We might therefore be missing many of the less energetic outflows found in the less luminous clusters. This does not affect our results, since we draw our conclusions only on the powerful AGN outbursts and compare these to the powerful outbursts at lower redshift, both of which have  $L_{\text{mech}} \gtrsim 10^{44} - 10^{45} \text{ erg s}^{-1}$ .

Finally, we mention that although the MACS sample lies within  $0.3 \leq z \leq 0.7$ , the number of clusters with *Chandra* data decreases steadily with redshift and there are only four MACS clusters with *Chandra* data at  $z > 0.55$ . The probability of detecting cavities in systems beyond  $z = 0.55$  is therefore low considering the limited sample. According to our cosmology, 1 arcsec corresponds to 4.45 kpc at  $z = 0.3$  and 7.15 kpc at  $z = 0.7$ . Cavities have typical sizes of 20 kpc in radius (see Fig. 8), whereas *Chandra* has a pixel size of  $\sim 0.49$  arcsec and a PSF of  $\sim 1$  arcsec. A cavity of 20 kpc in radius located at  $z = 0.3$  would therefore spread across 4.5 arcsec (or 9 pixels), and if the same cavity were located at  $z = 0.7$ , it would spread across 3 arcsec (or 6 pixels). Although deeper observations would be needed, and we would be pushing the unique resolution



**Figure 8.** From left to right and top to bottom: distribution of the mechanical power being injected by the central AGN, the total energy output stored in the cavities, the ratio of the distance from the core and the average radius of the bubble, the average radius, the distance from the core and the buoyancy rise time.

of *Chandra* to its limit, we could in principle identify cavities in  $z = 0.6$ – $0.7$  clusters, and might therefore be missing cavities in these further systems.

## 9.2 Detection rates

In more recent years, there has been considerable interest in studying the correlation between jet power, as determined from cavity energetics, and radio luminosities (e.g. Bîrzan et al. 2004, 2008; Cavagnolo et al. 2010; O’Sullivan et al. 2011, and see Fig. 7 here). This relation extends from individual elliptical galaxies to groups and even clusters of galaxies, and is interesting for two reasons. First, it provides insight into the nature of jets (e.g. Willott et al. 1999), and secondly, it provides a means to determine jet powers from the radio properties alone (e.g. Best et al. 2007; Magliocchetti & Brüggen 2007). Although the scatter remains large and more effort is needed to fully understand its nature, this method could in principle be used to determine jet powers in more distant clusters where the X-ray image quality remains too poor to identify surface brightness depressions (see Ma et al. 2011, who looked at AGN heating in the 400SD survey using this relation). However, deriving jet powers directly from X-ray cavities remains the most direct and reliable method.

Bîrzan et al. (2004) initially compiled a study of 80 systems with *Chandra* X-ray archival data, in which the authors visually searched for surface brightness depressions associated with the central regions. They found 18 systems with well-defined cavities (16 clusters of galaxies; one galaxy group, HCG 62; and one giant elliptical, M84), corresponding to a detection rate of the order of 20 per cent. Rafferty et al. (2006) then followed up this study, by adding 13 more objects where cavities had been recently discovered. Although many targets overlap between the various samples, Dunn & Fabian (2006) also provided an in depth study of X-ray cavities in the Brightest 55 sample of galaxy clusters (Edge et al. 1990), which was then followed by a more extensive study including the Brightest Cluster Sample (Ebeling et al. 1998) in Dunn & Fabian (2008). The total sample of Dunn & Fabian (2006, 2008) consists of 71 clusters in the redshift range  $0 \leq z \leq 0.4$ . The authors find 22 systems with clear cavities, similar to our detection rate if we include the ‘potential’ cavities ( $\sim 25$  per cent). However, Dunn & Fabian (2006) concluded that cavities sit mostly in clusters which require some form of heating, i.e. cool core clusters. They find that at least 14 of 20 (70 per cent) cool core clusters have cavities, making the detection rate in these systems very high. Since the majority of their clusters have an X-ray luminosity above  $10^{44}$  erg s $^{-1}$ , similar to the MACS clusters, we can directly compare our rates with theirs. We find that 20 (26 per cent), or at least 13 (17 per cent), of the 76 MACS clusters have cavities. We also find that the majority of our systems with cavities have a short central cooling time (see Figs 3 and 4). If we consider only the 37 MACS clusters that require some form of heating with  $t_{\text{cool}}(r < 50 \text{ kpc}) < 3 \text{ Gyr}$ , then we find that 19 have cavities. The detection rate is therefore 50 per cent in cooling systems, but could be higher considering that more than half of the remaining 18 MACS clusters with  $t_{\text{cool}}(r < 50 \text{ kpc}) < 3 \text{ Gyr}$  have less than 30 ks observations (Fig. 5). Note also cool core clusters are more likely to be the ones with bright X-ray cores, since the central densities are high. When a cavity is present, the contrast it creates in these peaked cores is higher, which makes the identification of cavities easier. We might therefore be missing cavities in non-cool core systems.

More recently, studies have focused on extending our knowledge of AGN feedback to the general population of giant ellipticals (e.g.

Finoguenov & Jones 2001; Jones et al. 2002; Machacek et al. 2006; Nulsen et al. 2007; Dunn et al. 2010). Nulsen et al. (2007) searched through an extensive sample of 160 nearby giant ellipticals, and found 109 with significant emission from surrounding hot gas. From these, they find that 27 have X-ray cavities. This corresponds to a detection rate of only 25 per cent in these cooling systems (compared to 70 per cent in cool core clusters), suggesting that duty cycles are greater in larger systems (see Jones et al., in preparation).

## 9.3 Cavity properties

In the top panel of Fig. 4 we show that most of the systems where we find cavities have a strong cool core with  $t_{\text{cool}}(r=50 \text{ kpc}) < 3$ – $5 \text{ Gyr}$ , therefore requiring extreme feedback from their central AGN to prevent the gas from cooling ( $L_{\text{mech}} \sim 10^{44-45}$  erg s $^{-1}$ ). Although we expect to see some of the most extreme AGN outbursts in this sample, the cavities we find are not larger or more powerful than the counterparts found in nearby ( $z < 0.3$ ) clusters with similar  $L_{\text{cool}}$ .

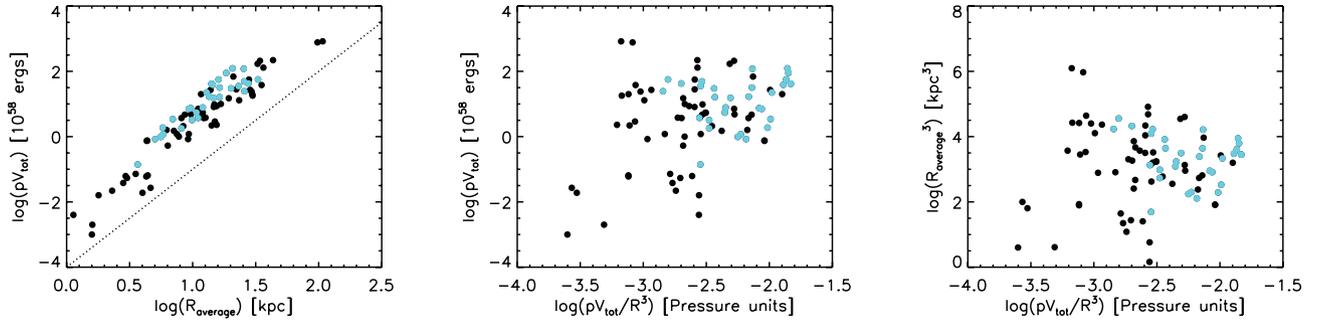
Fig. 6 shows that all but one<sup>3</sup> of the MACS clusters in which we find cavities are consistent with the  $L_{\text{mech}}$ – $L_{\text{cool}}$  correlation, implying that highly luminous clusters are able to prevent the gas from cooling at the same level than the less luminous clusters and groups. By contrast, Nulsen et al. (2007) provided evidence of a turnover at the lower luminosity end, with outbursts being more powerful for a given cooling luminosity. This is consistent with the idea that in smaller systems, the duty cycles are smaller (only  $\sim 25$  per cent of the time, compared to  $\sim 70$  per cent in cool core clusters), requiring that the outbursts be more energetic for a given cycle. Dunn & Fabian (2008) looked at the cavity properties of six distant clusters ( $0.1 \leq z \leq 0.4$ ) and found that they all sit above the line of exact balance in the  $L_{\text{mech}}$ – $L_{\text{cool}}$  correlation. Our more complete study of distant clusters finds differently. Essentially, our clusters agree well with the  $L_{\text{mech}}$ – $L_{\text{cool}}$  correlation with some above and some below the line of exact balance. There is therefore no turnover at the higher luminosity end.

In Fig. 8, we show the distribution of mechanical power ( $L_{\text{mech}}$ ), cavity energetics ( $pV_{\text{tot}}$ ) and buoyancy rise time ( $t_{\text{buoy}}$ ), as well as the cavity radius ( $R_{\text{average}}$ ) and distance from the nucleus ( $R_{\text{distance}}$ ). According to these figures, we find that the MACS cluster outbursts follow the same trends as the clusters in Rafferty et al. (2006). The Rafferty et al. (2006) sample consists of clusters that are  $z \lesssim 0.3$ , some of which require the same heating powers as the MACS clusters to prevent the gas from cooling ( $L_{\text{mech}} \sim 10^{44}$ – $10^{45}$  erg s $^{-1}$ ). In other words, the MACS outbursts are powerful, but not more powerful than the lower redshift counterparts.

It is also important to emphasize that we can only interpret the results concerning the largest and most powerful outbursts. We are probably missing a number of smaller cavities which could not be detected since our clusters are at a relatively high redshift and the resolution of *Chandra* can only go so far. We will have to wait for future missions with better spatial resolution than *Chandra* to study the small cavities at high redshift.

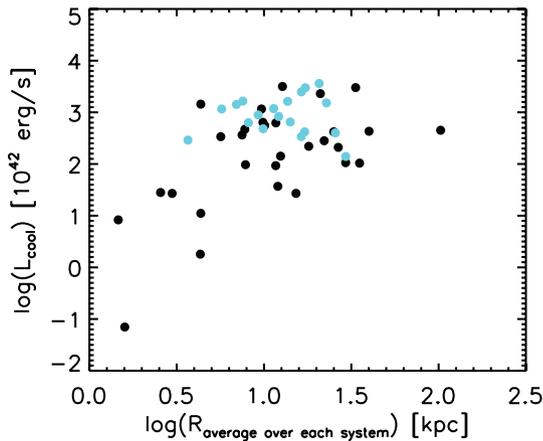
Fig. 8 also shows that the  $pV_{\text{tot}}$  distribution rises smoothly, peaks at  $10^{59}$ – $10^{60}$  erg and then decreases steeply. By contrast,  $L_{\text{mech}}$  seems to rise and decrease more smoothly, although this distribution would be altered if we included the giant ellipticals for which  $L_{\text{mech}}$  peaks at  $10^{41}$ – $10^{42}$  erg s $^{-1}$ . Since  $pV_{\text{tot}}$  depends only on the thermal pressure

<sup>3</sup> The outlier refers to MACS J1359.8+6231, where we found a cavity to the north-east, although very small (some 3–4 kpc in size) and therefore not energetically capable of preventing the gas from cooling.



**Figure 9.** Left-hand panel: energy of *each cavity* as a function its average radius ( $R_{\text{average}}$ ). The dotted line shows the slope of three one would expect between  $pV_{\text{tot}}$  and  $R_{\text{average}}$  since  $pV_{\text{tot}} \propto R_{\text{average}}^3$ . Middle panel: energy of *each cavity* as a function of the pseudo-pressure, the latter determined by dividing  $pV_{\text{tot}}$  by the average radius cubed ( $R_{\text{average}}^3$ ). Right-hand panel: average radius cubed ( $R_{\text{average}}^3$ ) of *each cavity* as a function of the pseudo-pressure. The light blue points are for the MACS clusters and the black points are for the Rafferty et al. (2006) objects.

and cavity size, the sudden drop at  $10^{59}$ – $10^{60}$  erg might be telling us something about how cavities form. As they are being inflated by the central AGN, it could be that they reach a maximum size in which they no longer can expand no matter the physical conditions. There could therefore be a maximum  $pV_{\text{tot}}$  value that can be reached in a single outburst. To provide further insight into this idea, we plot in the left-hand panel of Fig. 9 the energy in each cavity ( $pV_{\text{tot}}$ ) as a function of the average radius [ $R_{\text{average}} = (R_w \times R_l)^{0.5}$ ]. Since  $pV_{\text{tot}} \propto R_{\text{average}}^3$ , we expect a correlation between these two quantities, and in a log–log plot, the slope should be 3 (this is shown with the dotted line). However, this plot also shows that the most powerful outflows can only be created by the largest cavities. If a small cavity was being created in a high-pressure environment, the resulting  $pV_{\text{tot}}$  could be large. However, Fig. 9 shows that we do not observe such cavities. Furthermore, the middle and right-hand panels of Fig. 9 show that powerful outbursts are created in a variety of pressure environments, and that the size of each cavity has no dependence on the surrounding pressure. In other words, large cavities are created both in high- and low-pressure environments and the maximum allowed cavity radius in clusters does not depend on the surrounding pressure. By contrast, one could argue that large cavities can only be created if the surrounding thermal pressure is not large. Fig. 9 suggests otherwise. Finally, in Fig. 10, we plot



**Figure 10.** Shown is the cooling luminosity of the system as a function of the average cavity radius in each system (if there are two cavities in one system, then we calculate and plot the average radius between the two cavities). The light blue points are for the MACS clusters and the black points are for the Rafferty et al. (2006) objects.

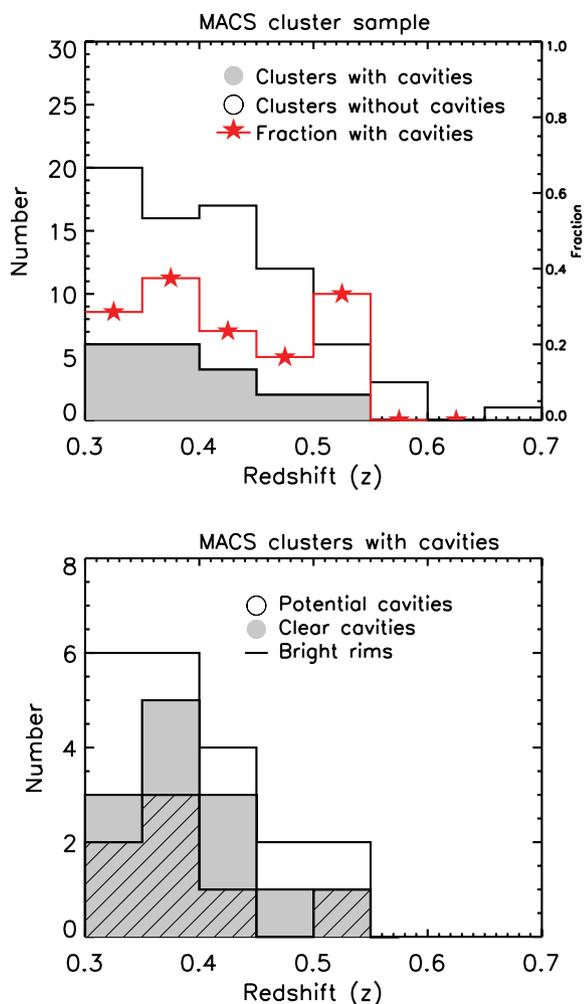
the cooling luminosity as a function of the average cavity radii in each system. Dunn & Fabian (2008) provided some evidence for a trend between these two quantities, and although the scatter remains large, our data support this idea (the larger the X-ray cooling, the larger the cavities). Again, this emphasized that somehow the central AGN knows about the X-ray cooling and provides the cavities needed to offset cooling of the ICM down to a fine level.

As shown by Fig. 8, the distribution of cavity radii and distances from the nucleus also suggest that extremely large and distant cavities, such as those found in MS 0735.6+7421 (McNamara et al. 2005; Gitti et al. 2007), are rare (these cavities are the outliers seen in the  $R_{\text{average}}$  and  $R_{\text{distance}}$  distributions of Fig. 8). Furthermore,  $R_{\text{distance}}/R_{\text{average}}$  peaks strongly at a value of 2 and then falls off quickly. In other words, cavities travel up to a distance of  $2 \times R_{\text{average}}$  (i.e. their diameter) and then either disintegrate or, more likely, become too difficult to detect (McNamara & Nulsen 2007). More distant cavities have been detected in the Perseus cluster, but only through very deep X-ray observations ( $\sim 500$  ks; Sanders & Fabian 2007; Fabian et al. 2011).

Finally, we mention that the shapes of the cavities we find vary from almost perfectly circular to elongated and even flattened (e.g. MACS J2046.0–3430). Two clusters also seem to have a misalignment between the X-ray peak and central galaxy (MACS J2046.0–3430, MACS J2135.2–0102), which may simply be due to projection effects. If a cavity is viewed along a different line of sight, projection effects can alter its shape and make it difficult to detect. The cavities we find may therefore be much larger and even more numerous than what we find (see also Sanders, Fabian & Taylor 2009). There are also two MACS clusters with long observations ( $> 100$  ks) and short central cooling times [ $t_{\text{cool}}(r < 50 \text{ kpc}) < 3 \text{ Gyr}$ ], but no cavities (MACS J1311.0–0311 at  $z = 0.49$  and MACS J1621.3+3810 at  $z = 0.465$ ; see the top panel in Fig. 5). Statistically, we expect that some systems will have a jet axis aligned with the line of sight, rendering the cavities almost impossible to detect. This could explain why we do not see cavities in these two systems.

#### 9.4 Evolution of cavity properties

The MACS sample consists of clusters within  $0.3 \leq z \leq 0.7$ , for which we were able to find surface brightness depressions in objects up to  $z \sim 0.6$ . In the top panel of Fig. 11 we show the fraction of MACS clusters with cavities as a function of redshift, and in the bottom panel we show the redshift distribution of those with cavities. This fraction remains constant around 0.25–0.30 up to  $z \sim 0.6$ ,



**Figure 11.** In the top panel, we show the distribution of clusters in the MACS sample as a function of redshift. We outline those in which we identified cavities (also shown in more detail in the bottom panel). The fraction of clusters with cavities for each redshift bin is shown in red. The fraction remains around 25–30 per cent until a redshift of 0.55.

and is consistent with the fraction found in previous cavity studies by Birzan et al. (2004,  $\sim 30$  per cent) and Rafferty et al. (2006,  $\sim 25$  per cent). Therefore, we find no significant evolution in the fraction of clusters with cavities as a function of redshift.

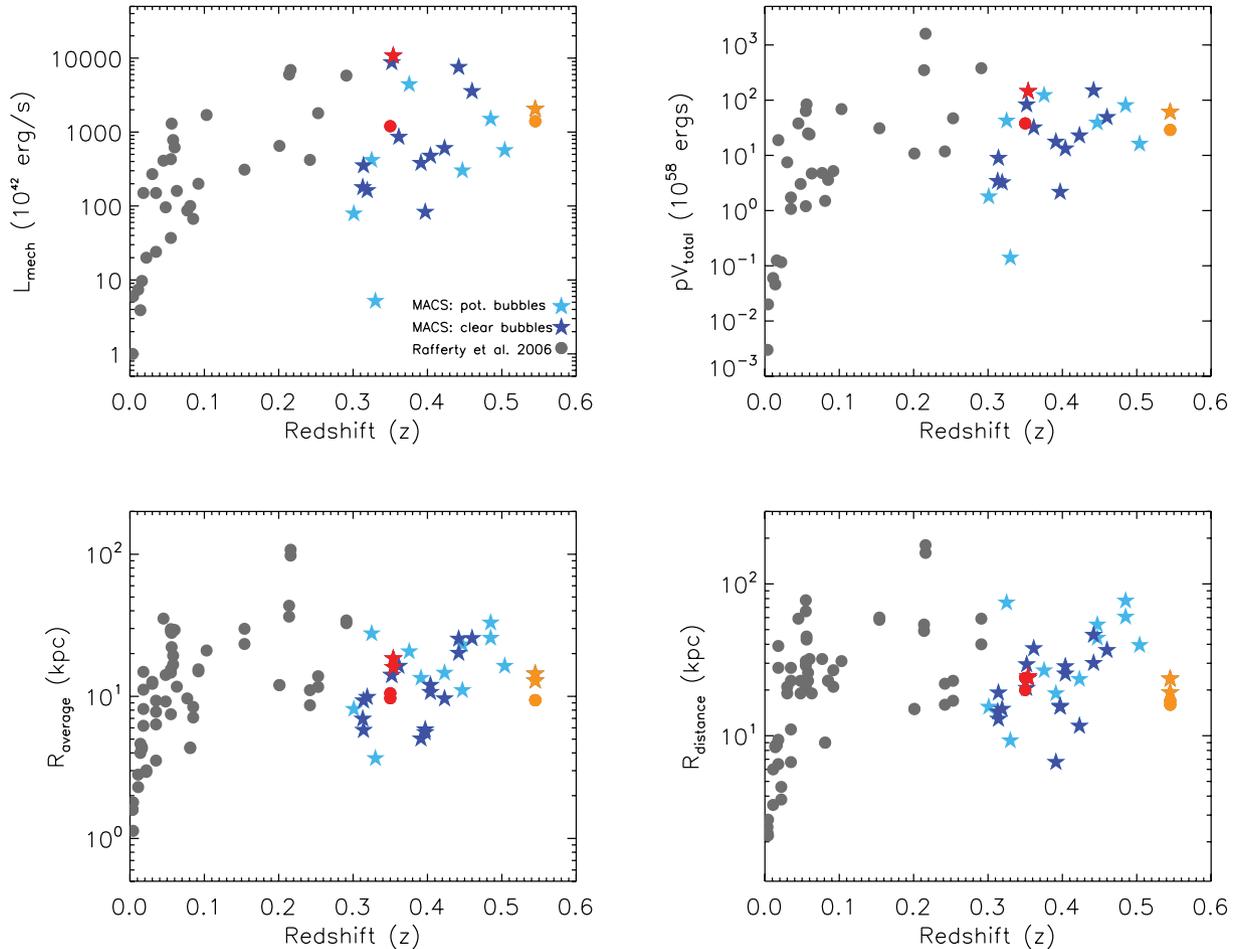
An interesting finding in this study has been the lack of any significant evolution in the cavity properties. This is shown in Fig. 12, where we plot the mechanical and energy outputs ( $L_{\text{mech}}$  and  $pV_{\text{tot}}$ ), as well as average cavity radius ( $R_{\text{average}}$ ) and distance from the nucleus ( $R_{\text{distance}}$ ) for each cavity, as a function of redshift. This figure shows that at least for the largest and most powerful outbursts, there is no evidence for evolution in the cavity properties. Cavities are not larger (or smaller) at higher redshift, and the distance to the nucleus remains on average the same, i.e. cavities at higher redshifts are not able to rise more (or less) than their lower redshift counterparts. The cavity energetics of the most powerful outbursts also remain the same.

Therefore, AGNs lying at the centres of clusters have not only been able to power extreme mechanical outbursts as early as  $\sim 7$ – $8$  Gyr after the big bang (corresponding to a redshift of  $z \sim 0.6$ ), but they have also been able to maintain these extreme outbursts up to

present day [ $L_{\text{mech}} \sim 10^{45} \text{ erg s}^{-1}$ ]. In a cosmological context, this means that ‘radio mode’ feedback has been in place for at least the past 5 Gyr and shows no sign of evolution since then. Based on the cold dark matter cosmological simulations of Croton et al. (2006), ‘radio mode’ feedback occurs at late times and involves a phase where the AGN is accreting at sub-Eddington rates and driving powerful outflows that suppress cooling flows from  $z \sim 1$  to present. In this phase, it is thought that black holes do not grow significantly in mass. However, in order to maintain the powerful MACS outbursts for the past 5 Gyr, the black hole mass must have been initially very high. If we equate the energy required to power the outflows over the last 5 Gyr (assuming an average output of  $\sim 10^{45} \text{ erg s}^{-1}$ ), to the energy released by a black hole ( $\eta M_{\text{BH}} c^2$ , assuming an efficiency of  $\eta = 0.1$ ) then it takes at least a  $10^9 M_{\odot}$  black hole to power such outflows. Here we assume that the duty cycles of dominant cluster galaxies are high ( $> 70$ – $90$  per cent), as supported by the observation that most or even all of the clusters that require some form of heating to counterbalance cooling have jetted outflows (Dunn & Fabian 2006). Note also, at least 50 per cent of our clusters requiring feedback from their central AGN seem to have jetted outflows, therefore supporting the idea that duty cycles are high (see Section 9.2). Even with a duty cycle of 50 per cent, the black hole masses required to power the outflows are of the order of  $10^9 M_{\odot}$  for an efficiency of  $\eta \sim 0.1$ . This rivals the largest black hole masses known, and suggests that some, especially those lying at the centres of the most extreme cooling flows, may be very massive and even ultramassive with  $> 10^{10} M_{\odot}$ .

Likewise, we could argue that to power outflows of  $10^{45} \text{ erg s}^{-1}$  for the past 5 Gyr, this requires the accreted mass ( $M_{\text{acc}} = E_{\text{cav}}/\eta c^2$ ) to be at least of the order of  $10^9 M_{\odot}$ . If the initial mass of the black hole was  $\sim 10^8 M_{\odot}$ , then by accreting some  $10^9 M_{\odot}$  in mass, we should have seen some variation in the cavity properties with redshift. Indeed, the mass of the black hole is one of the fundamental properties that regulates the power capabilities (another would be the spin). Since we do not see any evidence for evolution in the outflow properties, this strongly suggests that the  $10^9 M_{\odot}$  increase in mass does not significantly affect the AGN, and therefore that the black holes are initially very massive with  $M_{\text{BH}} > 10^9$ – $10^{10} M_{\odot}$ .

Samuele et al. (2011) looked at nebular emission in 77 BCGs belonging to the 160SD X-ray cluster survey (Vikhlinin et al. 1998; Mullis et al. 2003) and found no  $[\text{O III}]\lambda 3727$  or  $\text{H}\alpha$  emission stronger than 15 and 5  $\text{\AA}$ , respectively, in any of the BCGs. The authors suggest that this may be due to a significant decrease in the number of strong cooling flow clusters from  $z \sim 0.5$  to today (which does not mean that there are no strong cool cores, just less). They also suggest that this decline could be due to overheating of the core by AGN (as supported by Ma et al. 2011) or cluster mergers, which happen more frequently at higher redshift. By using the relation between jet power and radio power (Birzan et al. 2004, 2008; Cavagnolo et al. 2010; O’Sullivan et al. 2011), Ma et al. (2011) find that the power injected by radio sources within  $r < 250$  kpc is significant compared to the power radiated by the cluster atmosphere. However, as shown in Fig. 6, at least for the central AGN in the MACS clusters, the outbursts we find are able to balance cooling down to a fine level up to  $z \sim 0.6$ . The central AGN in clusters therefore does not seem to overheat the core at higher redshift and cannot account for the decrease in strong cool cores with redshift. Other studies in the literature also find a decrease in the number of strong cooling flow clusters with redshift (see Vikhlinin et al. 2007; Santos et al. 2008), but suggest that it is most likely due to the higher major merger rate of the past.



**Figure 12.** Profiles of the mechanical power being injected by the central AGN (top left-hand panel), the total energy output stored in the cavities (top right-hand panel), the average radius of a bubble (bottom left-hand panel) and distance from the core to the cavity centre (bottom right-hand panel) as a function of redshift. The MACS clusters are shown with the star symbols (dark blue for ‘clear’ cavities and light blue for ‘potential’ cavities), and those of Rafferty et al. (2006) are shown with the grey circles. Objects common to both samples are shown in red and yellow.

As a last note, we mention that many of the clusters in which we find cavities have a bright X-ray nucleus associated with the central AGN. These include MACS J0547.0–3904, MACS J0913.7+4056, MACS J0947.2+7623, MACS J1411.3+5212, MACS J1423.8+2404, MACS J1931.8–2634 and MACS J2046.0–3430. Most of these either have an Fe K line in the spectrum of the nucleus or a spectrum that can best be fitted by a power law, indicating the presence of an active AGN. BCGs with bright X-ray cores are rare at low redshift, and many that have similar  $L_{\text{mech}}$  as those in the MACS clusters do not have any evidence of an X-ray nucleus associated with their central AGN (Hlavacek-Larrondo & Fabian 2011). This may be indicating that we are seeing some evolution, not in terms of the mechanical properties of AGN feedback, but in terms of the radiative properties of the central AGN, thus suggesting that we are starting to see the transition between ‘quasar mode’ and ‘radio mode’ feedback. This will be analysed in detail in a forthcoming paper (Hlavacek-Larrondo et al., in preparation).

## 10 CONCLUDING REMARKS

We have analysed the *Chandra* X-ray observations of the MACS cluster sample and searched for surface-brightness depressions associated with the central AGN in each cluster. The MACS sample

consists of very X-ray luminous clusters within  $0.3 \leq z \leq 0.7$  (median  $L_{\text{X,RASS}} = 7 \times 10^{44} \text{ erg s}^{-1}$ ). Out of the 76 clusters, we find 13 with ‘clear’ cavities and seven with ‘potential’ cavities, bringing the detection rate to  $\sim 25$  per cent. Most of the clusters in which we find cavities have a short central cooling time below 3–5 Gyr, consistent with the idea that cavities sit predominantly in cool core clusters. By combining our results with those of previous cavity studies, the latter focusing on systems at  $z < 0.3$ , we find no evidence for evolution in any of the cavity properties. Although we expect to see some of the most extreme outbursts ( $L_{\text{mech}} \sim 10^{44}–10^{45} \text{ erg s}^{-1}$ ), the cavities are not larger (or smaller) at higher redshift, and they are not able to rise to further (or lesser) radii. The energy and power capabilities of the most powerful outbursts also remain the same. Extreme ‘radio mode’ AGN feedback therefore starts to operate as early as 7–8 Gyr after the big bang and shows no sign of evolution since then. In other words, AGNs lying at the centre of clusters are able to operate at early times and with extreme powers.

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## REFERENCES

- Allen S. W., 1998, *MNRAS*, 296, 392  
 Allen S. W. et al., 2001, *MNRAS*, 324, 842  
 Allen S. W., Rapetti D. A., Schmidt R. W., Ebeling H., Morris R. G., Fabian A. C., 2008, *MNRAS*, 383, 797  
 Anders E., Grevesse N., 1989, *Geochim. Cosmochim. Acta*, 53, 197  
 Arabadjis J. S., Bautz M. W., Garmire G. P., 2002, *ApJ*, 572, 66  
 Arnaud K. A., 1996, in Jacoby G. H., Barnes J., eds, *ASP Conf. Ser. Vol. 101, Astronomical Data Analysis Software and Systems V. XSPEC: The First Ten Years*. Astron. Soc. Pac., San Francisco, p. 17  
 Balestra I., Tozzi P., Ettori S., Rosati P., Borgani S., Mainieri V., Norman C., Viola M., 2007, *A&A*, 462, 429  
 Bauer F. E., Fabian A. C., Sanders J. S., Allen S. W., Johnstone R. M., 2005, *MNRAS*, 359, 1481  
 Becker R. H., White R. L., Helfand D. J., 1994, in Crabtree D. R., Hanisch R. J., Barnes J., eds, *ASP Conf. Ser. Vol. 61, Astronomical Data Analysis Software and Systems III. The VLA's FIRST Survey*. Astron. Soc. Pac., San Francisco, p. 165  
 Best P. N., von der Linden A., Kauffmann G., Heckman T. M., Kaiser C. R., 2007, *MNRAS*, 379, 894  
 Birzan L., Rafferty D. A., McNamara B. R., Wise M. W., Nulsen P. E. J., 2004, *ApJ*, 607, 800  
 Birzan L., McNamara B. R., Nulsen P. E. J., Carilli C. L., Wise M. W., 2008, *ApJ*, 686, 859  
 Bock D. C.-J., Large M. I., Sadler E. M., 1999, *AJ*, 117, 1578  
 Bohringer H., Nulsen P. E. J., Braun R., Fabian A. C., 1995, *MNRAS*, 274, L67  
 Brunetti G., Cappi M., Setti G., Feretti L., Harris D. E., 2001, *A&A*, 372, 755  
 Burenin R. A., Vikhlinin A., Hornstrup A., Ebeling H., Quintana H., Mescheryakov A., 2007, *ApJS*, 172, 561  
 Carilli C. L., Perley R. A., Dreher J. W., Leahy J. P., 1991, *ApJ*, 383, 554  
 Cavagnolo K. W., McNamara B. R., Nulsen P. E. J., Carilli C. L., Jones C., Birzan L., 2010, *ApJ*, 720, 1066  
 Cavagnolo K. W., McNamara B. R., Wise M. W., Nulsen P. E. J., Brügggen M., Gitti M., Rafferty D. A., 2011, *ApJ*, 732, 71  
 Churazov E., Brügggen M., Kaiser C. R., Böhringer H., Forman W., 2001, *ApJ*, 554, 261  
 Coble K. et al., 2007, *AJ*, 134, 897  
 Cohen A. S., Lane W. M., Cotton W. D., Kassim N. E., Lazio T. J. W., Perley R. A., Condon J. J., Erickson W. C., 2007, *AJ*, 134, 1245  
 Comerford J. M., Meneghetti M., Bartelmann M., Schirmer M., 2006, *ApJ*, 642, 39  
 Condon J. J., Cotton W. D., Greisen E. W., Yin Q. F., Perley R. A., Taylor G. B., Broderick J. J., 1998, *AJ*, 115, 1693  
 Crawford C. S., Allen S. W., Ebeling H., Edge A. C., Fabian A. C., 1999, *MNRAS*, 306, 857  
 Croton D. J. et al., 2006, *MNRAS*, 365, 11  
 Dahle H., Kaiser N., Irgens R. J., Lilje P. B., Maddox S. J., 2002, *ApJS*, 139, 313  
 Davis J. E., 2001, *ApJ*, 562, 575  
 Dong R., Rasmussen J., Mulchaey J. S., 2010, *ApJ*, 712, 883  
 Dressler A., Gunn J. E., 1992, *ApJS*, 78, 1  
 Dunn R. J. H., Fabian A. C., 2006, *MNRAS*, 373, 959  
 Dunn R. J. H., Fabian A. C., 2008, *MNRAS*, 385, 757  
 Dunn R. J. H., Fabian A. C., Taylor G. B., 2005, *MNRAS*, 364, 1343  
 Dunn R. J. H., Allen S. W., Taylor G. B., Shurkin K. F., Gentile G., Fabian A. C., Reynolds C. S., 2010, *MNRAS*, 404, 180  
 Ebeling H., 2009, HST Proposal A Snapshot Survey of The Most Massive Clusters of Galaxies. Proposal ID. 12166, NASA/STScI  
 Ebeling H., Edge A. C., Bohringer H., Allen S. W., Crawford C. S., Fabian A. C., Voges W., Huchra J. P., 1998, *MNRAS*, 301, 881  
 Ebeling H., Edge A. C., Allen S. W., Crawford C. S., Fabian A. C., Huchra J. P., 2000, *MNRAS*, 318, 333  
 Ebeling H., Edge A. C., Henry J. P., 2001, *ApJ*, 553, 668  
 Ebeling H., Barrett E., Donovan D., Ma C.-J., Edge A. C., van Speybroeck L., 2007, *ApJ*, 661, L33  
 Ebeling H., Edge A. C., Mantz A., Barrett E., Henry J. P., Ma C. J., van Speybroeck L., 2010, *MNRAS*, 407, 83  
 Edge A. C., 2001, *MNRAS*, 328, 762  
 Edge A. C., Stewart G. C., Fabian A. C., Arnaud K. A., 1990, *MNRAS*, 245, 559  
 Ehlert S. et al., 2011, *MNRAS*, 411, 1641  
 Ettori S., Morandi A., Tozzi P., Balestra I., Borgani S., Rosati P., Lovisari L., Terenziani F., 2009, *A&A*, 501, 61  
 Fabian A. C., Sanders J. S., Allen S. W., Crawford C. S., Iwasawa K., Johnstone R. M., Schmidt R. W., Taylor G. B., 2003, *MNRAS*, 344, L43  
 Fabian A. C., Sanders J. S., Taylor G. B., Allen S. W., Crawford C. S., Johnstone R. M., Iwasawa K., 2006, *MNRAS*, 366, 417  
 Fabian A. C. et al., 2011, *MNRAS*, 418, 2154  
 Finoguenov A., Jones C., 2001, *ApJ*, 547, L107  
 Forman W. et al., 2005, *ApJ*, 635, 894  
 Fort B., Le Fevre O., Hammer F., Cailloux M., 1992, *ApJ*, 399, L125  
 Franx M., Illingworth G. D., Kelson D. D., van Dokkum P. G., Tran K.-V., 1997, *ApJ*, 486, L75  
 Gavazzi R., 2005, *A&A*, 443, 793  
 Gavazzi R., Fort B., Mellier Y., Pelló R., Dantel-Fort M., 2003, *A&A*, 403, 11  
 Gioia I. M., Luppino G. A., 1994, *ApJS*, 94, 583  
 Gioia I. M., Maccacaro T., Schild R. E., Wolter A., Stocke J. T., Morris S. L., Henry J. P., 1990, *ApJS*, 72, 567  
 Gitti M., Feretti L., Schindler S., 2006, *A&A*, 448, 853  
 Gitti M., McNamara B. R., Nulsen P. E. J., Wise M. W., 2007, *ApJ*, 660, 1118  
 Graham J., Fabian A. C., Sanders J. S., 2008, *MNRAS*, 386, 278  
 Gregory P. C., Condon J. J., 1991, *ApJS*, 75, 1011  
 Griffith M. R., Wright A. E., 1993, *AJ*, 105, 1666  
 Harris D. E. et al., 2000, *ApJ*, 530, L81  
 Hines D. C., Wills B. J., 1993, *ApJ*, 415, 82  
 Hlavacek-Larrondo J., Fabian A. C., 2011, *MNRAS*, 413, 313  
 Hoekstra H., Franx M., Kuijken K., Squires G., 1998, *ApJ*, 504, 636  
 Iwasawa K., Fabian A. C., Ettori S., 2001, *MNRAS*, 321, L15  
 Jones C., Forman W., Vikhlinin A., Markevitch M., David L., Warmflash A., Murray S., Nulsen P. E. J., 2002, *ApJ*, 567, L115  
 Kalberla P. M. W., Burton W. B., Hartmann D., Arnal E. M., Bajaja E., Morras R., Pöppel W. G. L., 2005, *A&A*, 440, 775  
 Kartaltepe J. S., Ebeling H., Ma C. J., Donovan D., 2008, *MNRAS*, 389, 1240  
 Kleinmann S. G., Hamilton D., Keel W. C., Wynn-Williams C. G., Eales S. A., Becklin E. E., Kuntz K. D., 1988, *ApJ*, 328, 161  
 Limousin M. et al., 2010, *MNRAS*, 405, 777  
 Ma C., McNamara B. R., Nulsen P. E. J., Schaffer R., Vikhlinin A., 2011, *ApJ*, 740, 51  
 Machacek M., Nulsen P. E. J., Jones C., Forman W. R., 2006, *ApJ*, 648, 947  
 McNamara B. R., Nulsen P. E. J., 2007, *ARA&A*, 45, 117  
 McNamara B. R. et al., 2000, *ApJ*, 534, L135  
 McNamara B. R., Nulsen P. E. J., Wise M. W., Rafferty D. A., Carilli C., Sarazin C. L., Blanton E. L., 2005, *Nat*, 433, 45

Magliocchetti M., Brüggén M., 2007, MNRAS, 379, 260  
Mann A. W., Ebeling E., 2011, arXiv:1111.2396  
Mauch T., Murphy T., Buttery H. J., Curran J., Hunstead R. W., Piestrzynski B., Robertson J. G., Sadler E. M., 2003, MNRAS, 342, 1117  
Maughan B. J., Jones C., Forman W., Van Speybroeck L., 2008, ApJS, 174, 117  
Merten J., Cacciato M., Meneghetti M., Mignone C., Bartelmann M., 2009, A&A, 500, 681  
Mullis C. R. et al., 2003, ApJ, 594, 154  
Nulsen P. E. J., Jones C., Forman W. R., David L. P., McNamara B. R., Rafferty D. A., Bîrzan L., Wise M. W., 2007, in Böhringer H., Pratt G. W., Finoguenov A., Schuecker P., eds, Heating versus Cooling in Galaxies and Clusters of Galaxies. ESO Astrophysics Symposia. Springer-Verlag, Berlin, p. 210  
O’Dea C. P. et al., 2008, ApJ, 681, 1035  
O’Sullivan E., Giacintucci S., David L. P., Gitti M., Vrtilik J. M., Raychaudhury S., Ponman T. J., 2011, ApJ, 735, 11  
Perley R. A., Taylor G. B., 1991, AJ, 101, 1623  
Perlman E. S., Horner D. J., Jones L. R., Scharf C. A., Ebeling H., Wegner G., Malkan M., 2002, ApJS, 140, 265  
Peterson J. R., Fabian A. C., 2006, Phys. Rep., 427, 1  
Puetzfeld D., Pohl M., Zhu Z.-H., 2005, ApJ, 619, 657  
Quillen A. C. et al., 2008, ApJS, 176, 39  
Rafferty D. A., McNamara B. R., Nulsen P. E. J., Wise M. W., 2006, ApJ, 652, 216  
Rafferty D. A., McNamara B. R., Nulsen P. E. J., 2008, ApJ, 687, 899  
Reid I. N. et al., 1991, PASP, 103, 661  
Rengelink R. B., Tang Y., de Bruyn A. G., Miley G. K., Bremer M. N., Roettgering H. J. A., Bremer M. A. R., 1997, A&AS, 124, 259  
Russell H. R., Fabian A. C., Sanders J. S., Johnstone R. M., Blundell K. M., Brandt W. N., Crawford C. S., 2010, MNRAS, 402, 1561  
Samuele R., McNamara B. R., Vikhlinin A., Mullis C. R., 2011, ApJ, 731, 31  
Sand D. J., Treu T., Ellis R. S., Smith G. P., Kneib J.-P., 2008, ApJ, 674, 711  
Sanders J. S., Fabian A. C., 2007, MNRAS, 381, 1381  
Sanders J. S., Fabian A. C., Allen S. W., Morris R. G., Graham J., Johnstone R. M., 2008, MNRAS, 385, 1186  
Sanders J. S., Fabian A. C., Taylor G. B., 2009, MNRAS, 393, 71  
Sanders J. S., Fabian A. C., Frank K. A., Peterson J. R., Russell H. R., 2010, MNRAS, 402, 127  
Santos J. S., Rosati P., Böhringer H., Ettori S., Bignamini A., 2008, A&A, 483, 35  
Siana B. et al., 2009, ApJ, 698, 1273  
Sijacki D., Springel V., 2006, MNRAS, 366, 397  
Smail I. et al., 2007, ApJ, 654, L33  
Stoche J. T., Morris S. L., Gioia I. M., Maccacaro T., Schild R., Wolter A., Fleming T. A., Henry J. P., 1991, ApJS, 76, 813  
Swinbank A. M. et al., 2010, Nat, 464, 733  
Taylor G. B., Perley R. A., 1992, A&A, 262, 417  
Trümper J., 1983, Adv. Space Res., 27, 1404  
Vignali C. et al., 2011, MNRAS, 416, 2068  
Vikhlinin A., McNamara B. R., Forman W., Jones C., Quintana H., Hornstrup A., 1998, ApJ, 502, 558  
Vikhlinin A., Burenin R., Forman W. R., Jones C., Hornstrup A., Murray S. S., Quintana H., 2007, in Böhringer H., Pratt G. W., Finoguenov A., Schuecker P., eds, Heating versus Cooling in Galaxies and Clusters of Galaxies. ESO Astrophysics Symposia. Springer-Verlag, Berlin, p. 48

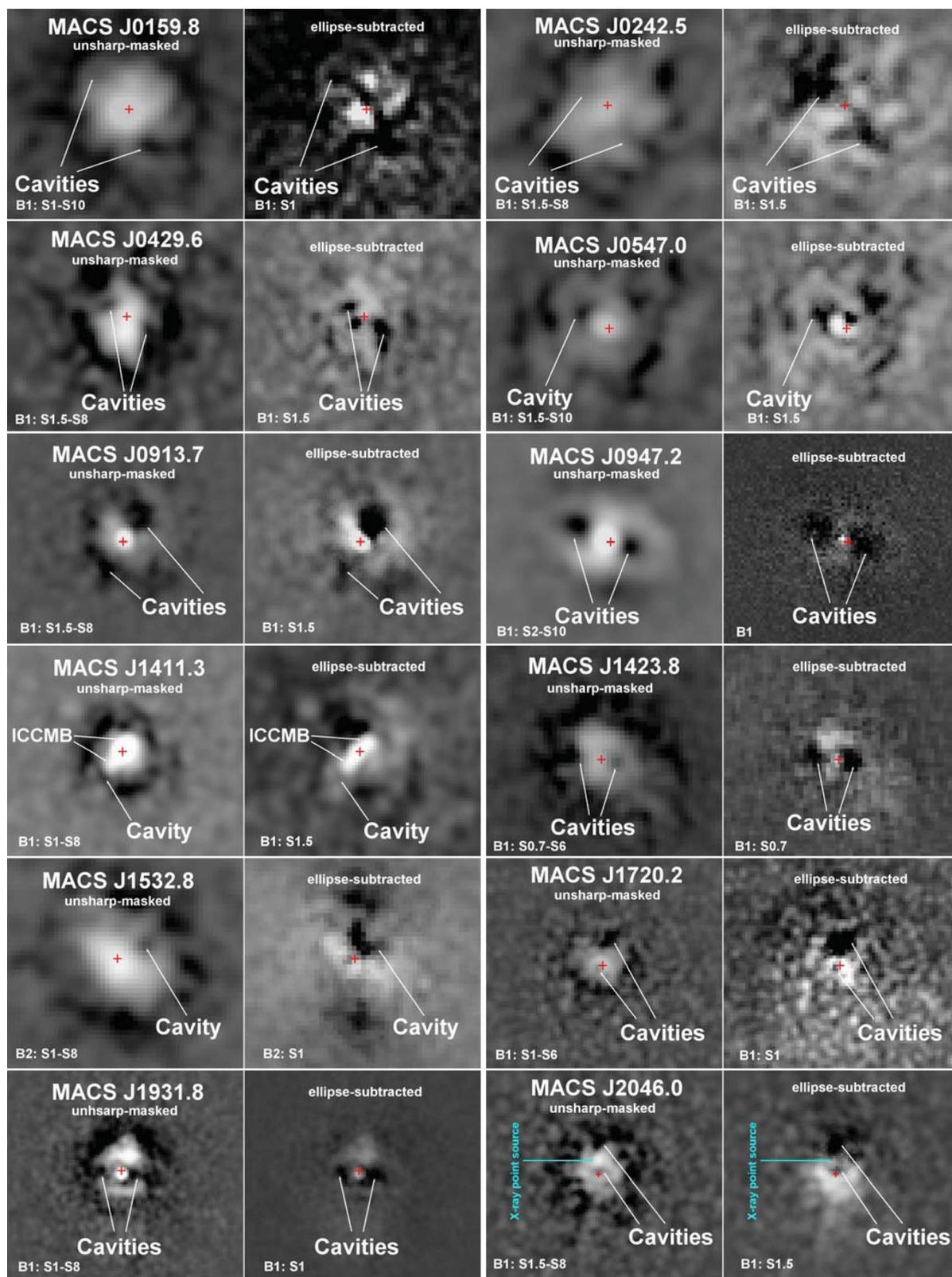
Voges W. et al., 1999, A&A, 349, 389  
Willott C. J., Rawlings S., Blundell K. M., Lacy M., 1999, MNRAS, 309, 1017  
Wright A. E., Ables J. G., Allen D. A., 1983, MNRAS, 205, 793  
Zitrin A., Broadhurst T., Coe D., Liesenborgs J., Benítez N., Rephaeli Y., Ford H., Umetsu K., 2011, MNRAS, 413, 1753

## APPENDIX A: UNSHARP-MASKED AND ELLIPSE-SUBTRACTED X-RAY IMAGES

We show the unsharp-masked and ellipse-subtracted images for all clusters in which we identified cavities (Fig. A1). The first consists of subtracting a strongly smoothed image from a lightly smoothed image, and the second consists of subtracting an elliptical model of the cluster emission from the original image. These techniques enhance deviations in the original image, and the images were used as a first step to identify cavities. However, we only considered that a cluster had a cavity if we could also see a faint depression in the original 0.5–7 keV *Chandra* image. There are two clusters for each horizontal line in Fig. A1, with the unsharp-masked 0.5–7 keV X-ray image shown in the left-hand panel and the ellipse-subtracted 0.5–7 keV X-ray image shown in the right-hand panel. The red cross shows the location of the BCG nucleus. We also indicate in the lower left-hand corner of each image the binning ( $B_X$ ) and smoothing ( $S_X$ ) scales adopted for creating the images. For  $B_X$ ,  $X$  indicates the binning factor where  $X = 1$  corresponds to no binning and  $X = 2$  gives an image where each of pixel corresponds to four pixels in the original image. For  $S_X$ ,  $X$  indicates the smoothing factor and corresponds to the sigma of a Gaussian in units of pixels. For each unsharp-masked image, we show the two smoothing scales used to create the image ( $S_{X_1} - S_{X_2}$ ), and for each ellipse-subtracted image, we show the smoothing scale adopted ( $S_{X_3}$ ) for the original image before creating and subtracting an elliptical model. The first 13 clusters have ‘clear’ cavities, while the remaining seven only have ‘potential’ cavities. Deeper observations are needed to confirm if the ‘potential’ cavities are real.

## APPENDIX B: SPECTRAL PROFILES

Fig. B1 shows the projected (black symbols) and deprojected (red symbols) temperature, electron density, electron pressure and entropy profiles for each cluster in which a cavity was identified. We used  $\chi^2$  statistics and each annulus had a minimum of 900 counts, but for clusters in which the data were of good quality we allowed the annuli to have up to 4000 counts. For MACS J0111.5+0855 and MACS J2135.2–0102, we could not deproject the spectra because of the limiting data quality. In this case, we used C statistics and each annulus only had  $\sim 400$  counts.



**Figure A1.** Unsharp-masked and ellipse-subtracted images for all clusters in which we identified cavities. The same layout and scaling as Fig. 2: two clusters for each horizontal line with the unsharp-masked 0.5–7 keV X-ray image shown in the left-hand panel and the ellipse-subtracted 0.5–7 keV X-ray image shown in the right-hand panel. We identify the cavities in each system, and show with the red cross symbol in the X-ray image the location of the BCG. We also show in the lower-left corner of each image the binning ( $B_X$ ) and smoothing ( $S_X$ ) scales adopted for creating the images. For each unsharp-masked image, we show the two smoothing scales used to create the image ( $S_{X_1} - S_{X_2}$ ), and for each ellipse-subtracted image, we show the smoothing scale adopted ( $S_{X_3}$ ) for the original image before creating and subtracting an elliptical model. If no  $S_X$  is shown, then no smoothing was applied to the image. The first 13 clusters have ‘clear’ cavities, while the remaining seven only have ‘potential’ cavities (see ‘potential cavities’ indicated in the figure). We emphasize that deeper observations are needed to confirm if the ‘potential’ cavities are real.

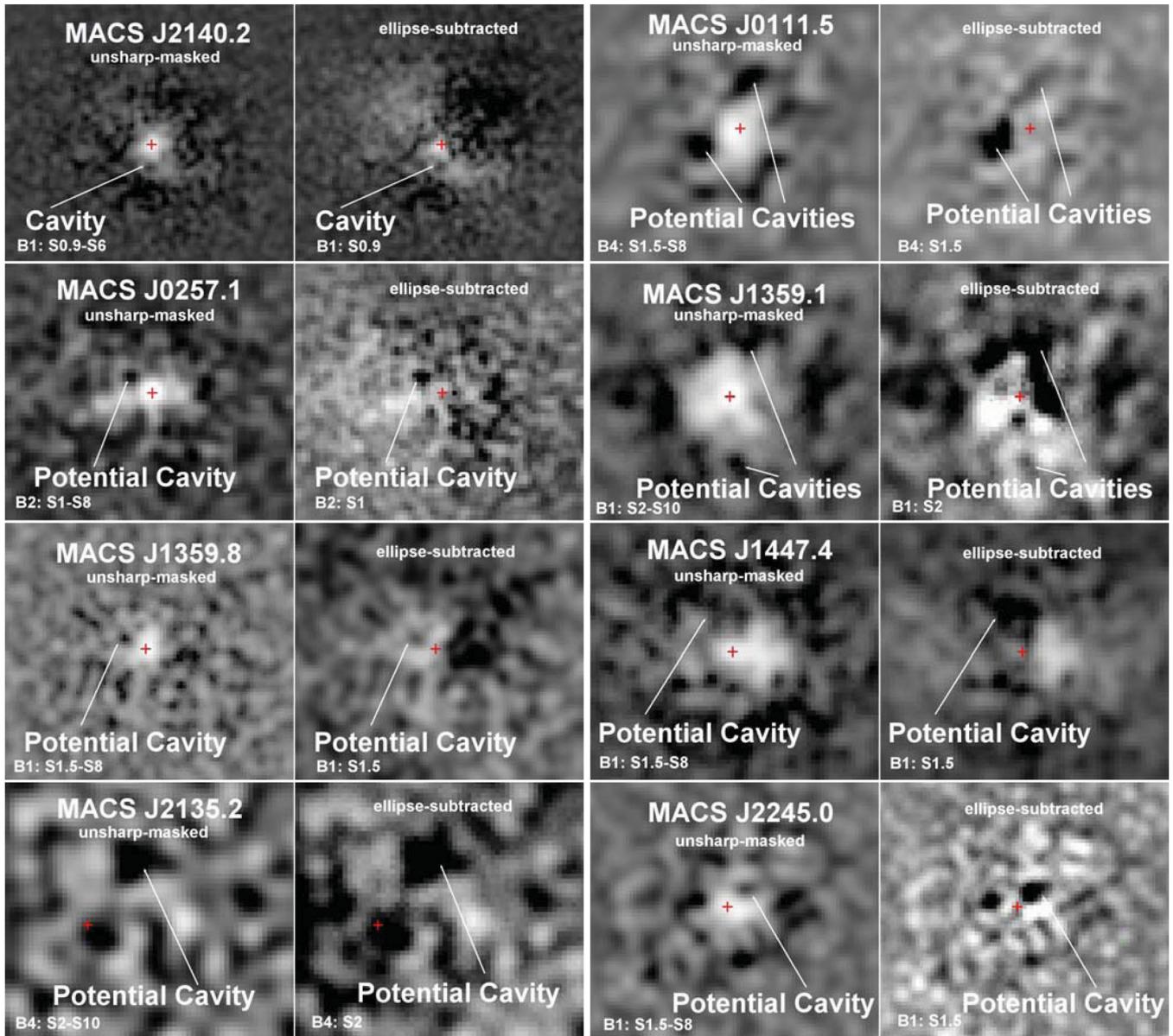
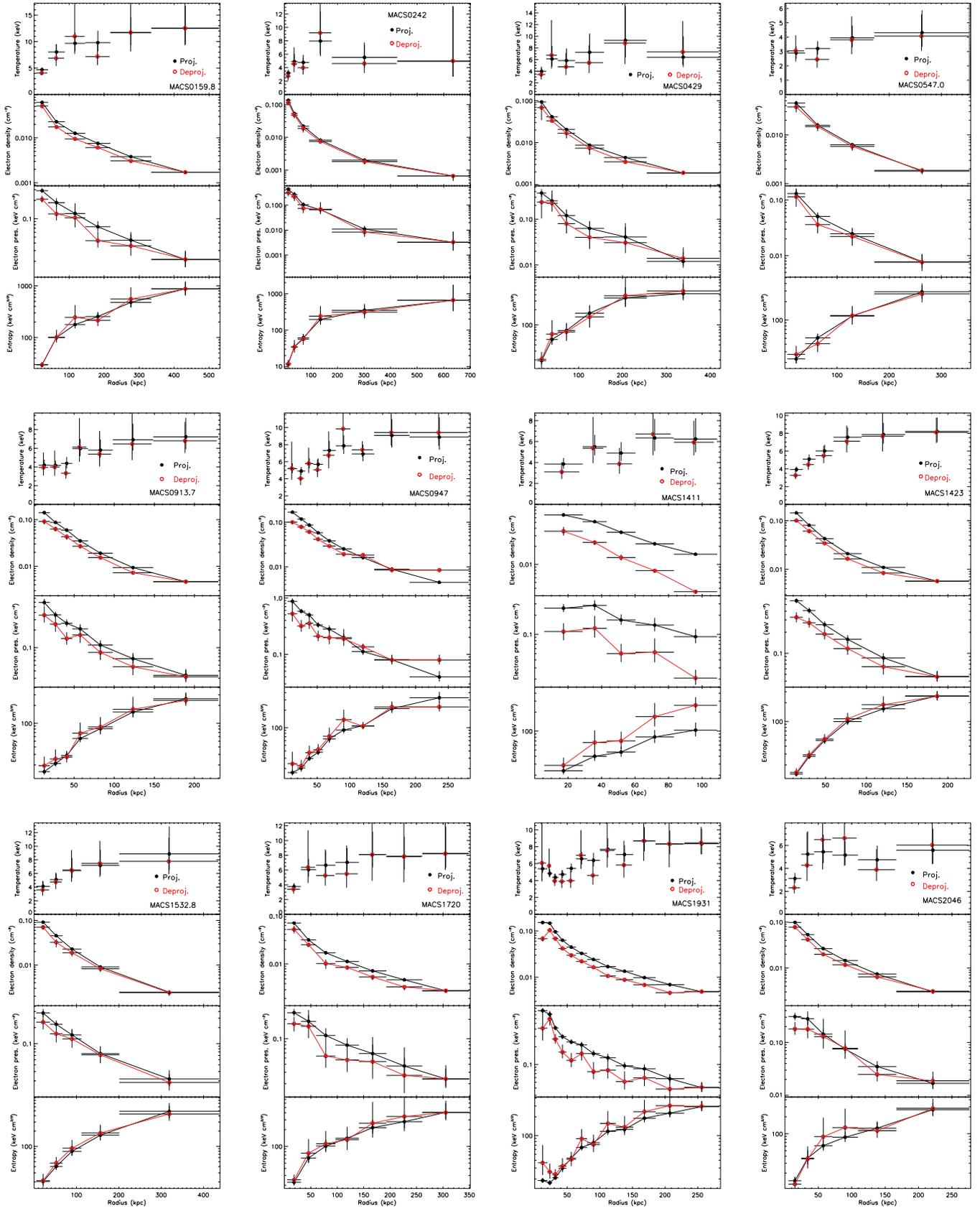


Figure A1 – *continued*



**Figure B1.** Projected (black symbols) and deprojected (red symbols) temperature, electron density, electron pressure and entropy profiles for each cluster in which a cavity was identified.

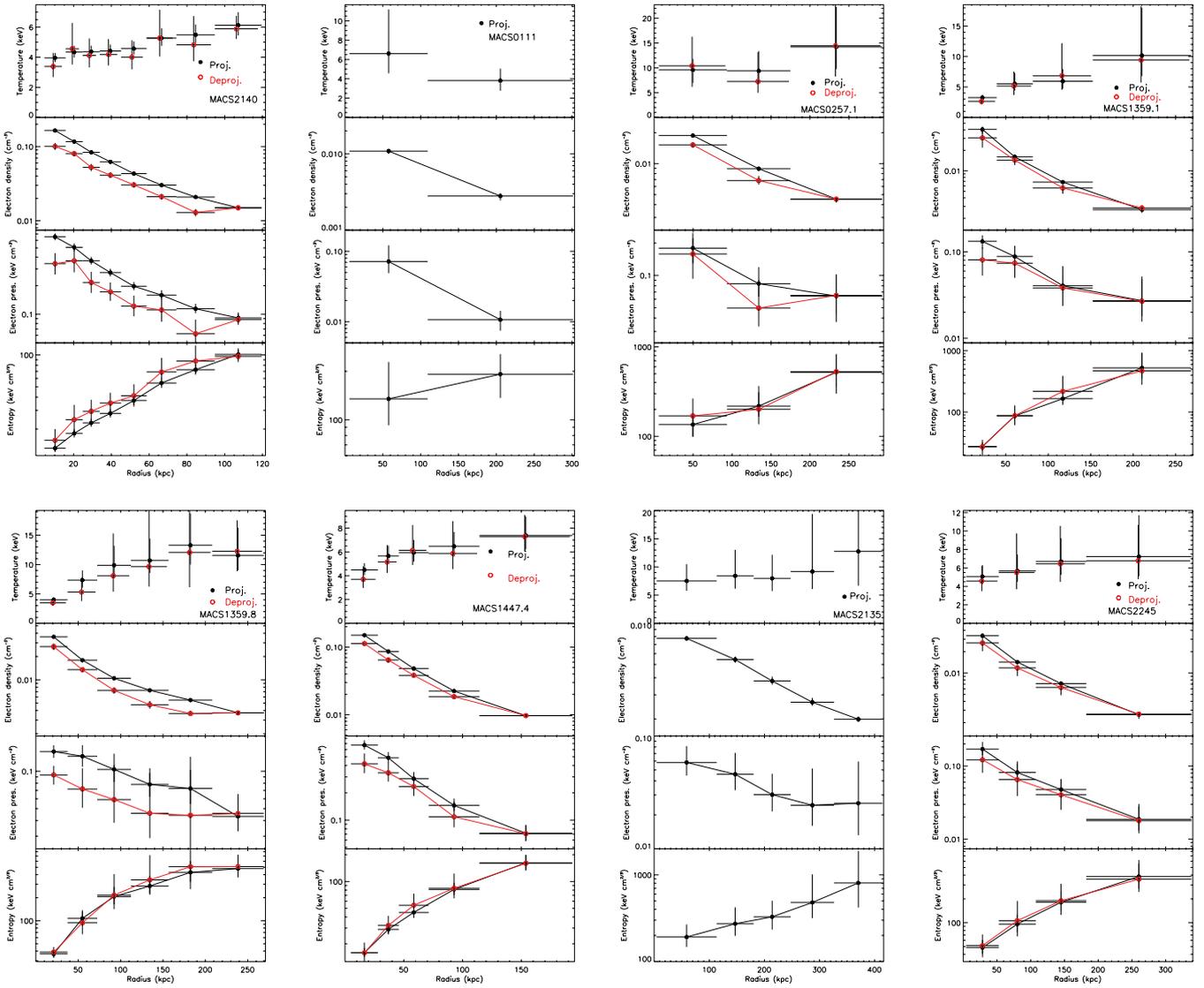


Figure B1 – *continued*

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