

The bow shock, cold fronts and disintegrating cool core in the merging galaxy group RX J0751.3+5012

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ABSTRACT

We present a new *Chandra* X-ray observation of the off-axis galaxy group merger RX J0751.3+5012. The hot atmospheres of the two colliding groups appear highly distorted by the merger. The images reveal arc-like cold fronts around each group core, produced by the motion through the ambient medium, and the first detection of a group merger shock front. We detect a clear density and temperature jump associated with a bow shock of Mach number $M = 1.9 \pm 0.4$ ahead of the northern group. Using galaxy redshifts and the shock velocity of $1100 \pm 300 \text{ km s}^{-1}$, we estimate that the merger axis is only $\sim 10^\circ$ from the plane of the sky. From the projected group separation of $\sim 90 \text{ kpc}$, this corresponds to a time since closest approach of $\sim 0.1 \text{ Gyr}$. The northern group hosts a dense, cool core with a ram pressure stripped tail of gas extending $\sim 100 \text{ kpc}$. The sheared sides of this tail appear distorted and broadened by Kelvin–Helmholtz instabilities. We use the presence of this substructure to place an upper limit on the magnetic field strength and, for *Spitzer*-like viscosity, show that the development of these structures is consistent with the critical perturbation length above which instabilities can grow in the intragroup medium. The northern group core also hosts a galaxy pair, UGC 4052, with a surrounding IR and near-UV ring $\sim 40 \text{ kpc}$ in diameter. The ring may have been produced by tidal stripping of a smaller galaxy by UGC 4052 or it may be a collisional ring generated by a close encounter between the two large galaxies.

Key words: galaxies: groups: individual: RX J0751.3+5012 – intergalactic medium – X-rays: galaxies: clusters.

1 INTRODUCTION

A large fraction of the baryons in the nearby Universe resides in galaxy groups with X-ray luminosities $\sim 10^{41}–10^{43} \text{ erg s}^{-1}$ and temperatures $\sim 0.3–2 \text{ keV}$ (e.g. Mulchaey et al. 1993; Ponman & Bertram 1993; Mulchaey 2000). Galaxy groups are much more representative hosts than rarer rich cluster systems (e.g. Geller & Huchra 1983; Tully 1987) and important for understanding the grav-

itational and thermal evolution of most of the matter in the Universe. The effects of non-gravitational heating, such as supernova explosions, active galactic nuclei (AGN) feedback and merger shocks, are expected to be more pronounced in lower mass systems as the energy input from these sources is comparable to the binding energy of the group (e.g. Ponman et al. 1996; Ponman, Cannon & Navarro 1999; Helsdon & Ponman 2000). Major and minor group mergers, and their subsequent relaxation, govern the formation of the largest scale structures and can have a substantial impact on their constituent galaxies. However, studies of galaxy group mergers have been limited due to their faint X-ray emission and low

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galaxy densities and richer cluster mergers have instead been the focus of attention (but see e.g. Kraft et al. 2004, 2006; Machacek et al. 2010, 2011; Kraft et al. 2011).

Deep *Chandra* X-ray observations of the hot gas substructures generated by mergers have proved crucial to our understanding of the dynamics and thermal evolution. Narrow surface brightness edges are observed where the gas density drops across the discontinuity and the gas temperature increases so that the pressure is continuous (Markevitch et al. 2000; Vikhlinin, Markevitch & Murray 2001a; Markevitch & Vikhlinin 2007). These ‘cold fronts’ mark the boundary of a cool, dense core of gas that is moving through the warmer ambient medium. The ram pressure experienced by the central gas peak in a major merger produces a sharp cold front along the leading interface of the dense core (e.g. Vikhlinin et al. 2001a). Sloshing cold fronts can occur when a minor merger sets a dense core oscillating in the gravitational potential well (e.g. Markevitch, Vikhlinin & Mazzotta 2001; Ascasibar & Markevitch 2006). Shock fronts are also expected in mergers but there are only a handful of detections in clusters with a clear density and temperature increase (e.g. Markevitch et al. 2002, 2005; Russell et al. 2010; Macario et al. 2011). Shock front detections provide measurements of the gas bulk velocities in the plane of the sky and constrain the merger geometry. They are therefore key observational tools for studying the merger history (e.g. Markevitch & Vikhlinin 2007). The detailed structure of cold fronts has also been used to study the relatively unknown transport processes in the hot atmospheres of groups. Observations of Kelvin–Helmholtz instabilities along cold fronts, or their absence, can provide a measure of the gas viscosity and magnetic field strength in the intracluster medium (ICM; e.g. Vikhlinin, Markevitch & Murray 2001b; Roediger et al. 2011; ZuHone, Markevitch & Lee 2011). The sharp temperature and density changes across cold fronts also suggests that thermal conduction and diffusion are strongly suppressed (e.g. Ettori & Fabian 2000; Vikhlinin et al. 2001a).

X-ray observations with *ROSAT* and *XMM–Newton* of the nearby system RX J0751.3 + 5012 ($z = 0.022$; Ebeling et al. 1998; Watson et al. 2009) revealed two large galaxy groups with disturbed morphologies, which are in the process of merging. The archival *XMM–Newton* observation showed surface brightness edges around the two groups, each ~ 100 kpc across, and a long tail of stripped material extending behind the northern group core. The northern group core hosts luminous optical line emission, star formation and a striking IR bright ring of emission surrounding the two central group galaxies (O’Dea et al. 2008; Quillen et al. 2008). Here we present a new deep *Chandra* observation of the off-axis group merger RX J0751.3 + 5012 studying cold front and shock front structure, ram pressure stripping and turbulent mixing in detail in the intragroup medium. In Section 2, we present the imaging and spectroscopic analysis of the *Chandra* data and identify several surface brightness edges associated with cold fronts and a bow shock. The structure of the bow shock is analysed in Section 3 and the Mach number is calculated. In Sections 4 and 5, we discuss the formation of turbulent instabilities around the north-eastern (NE) group core and ram pressure stripped tail and determine an upper limit on the magnetic field strength. Finally, in Section 6, we show the detection of additional extended emission in the ACIS-I field of view likely corresponding to the galaxy cluster WHL J075052.4 + 500252 at $z = 0.417$.

We assume $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$ and $\Omega_\Lambda = 0.7$, translating to a scale of 0.45 kpc per arcsec at the redshift $z = 0.022$ of RX J0751.3+5012. All errors are 1σ unless otherwise noted.

2 CHANDRA DATA ANALYSIS

2.1 Data reduction

RX J0751.3 + 5012 was observed with the *Chandra* ACIS-I detector for a single 100 ks exposure on 2013 May 14 (ID 15170). There is also an archival ACIS-I observation of 5 ks taken on 2011 January 8 (ID 12811) but this contributed few extended source counts and was therefore not included in the analysis. The new observation was reprocessed with CIAO 4.5 and CALDB 4.5.9 provided by the *Chandra* X-ray Center. The level 1 event file was reprocessed to apply the latest gain and charge transfer inefficiency correction and then filtered to remove any photons detected with bad grades. The observation was telemetered in VFAINT mode which provided improved background screening. A background light curve was extracted from a source free region on chip 0 and filtered with the LC_CLEAN¹ script to identify periods in the observation affected by flares. There was a short flare towards the end of the observation reducing the final cleaned exposure to 79.8 ks.

Fig. 1 (left) shows an exposure-corrected image of RX J0751.3 + 5012 covering the 0.3–4 keV energy band. The raw image was corrected for exposure variation by dividing by a spectrum-weighted exposure map. Point sources were identified with WAVDETECT (Freeman et al. 2002), confirmed visually and excluded from the analysis. The background was subtracted from images and spectra using the appropriate blank-sky background observations. Standard blank-sky backgrounds were processed identically to the events file, reprojected to the corresponding sky position and normalized to match the count rate in the 9.5–12 keV energy range. The normalized blank-sky background spectrum for a source-free region of ACIS-S2 was compared to the observed background spectrum and found to be a close match over the whole energy range.

2.2 Imaging analysis

The *Chandra* observation of RX J0751.3 + 5012 reveals an off-axis collision between two galaxy groups, where the NE group has recently passed to the north (N) of the primary, south-western (SW) group (Fig. 1 left). Ram pressure generated by the motion through the ambient medium is stripping material from the bright, dense core in the NE group. The tail of stripped material extends ~ 100 kpc (3.5 arcmin) and there is a trail of soft X-ray gas blobs ~ 30 kpc behind the core that appears to curve towards the SW group. There is a sharp drop in surface brightness around the N edge of the NE group core and this is likely to be a cold front marking the leading edge of the core. There is potentially a second, outer surface brightness edge ~ 15 kpc (30 arcsec) ahead of the NE group core (shown by arrows). The SW group appears more diffuse with a longer (~ 190 arcsec) surface brightness drop around its south-eastern (SE) and N edges. This edge has a ‘boxy’ morphology similar to that found in A496 (Dupke & White 2003; Dupke, White & Bregman 2007). This edge is likely disrupted by the passage of the NE group and interacting with the ram pressure stripped tail, which could explain the appearance of substructure around this interface.

Fig. 2 shows the merged MOS and PN image from the 12 ks archival *XMM–Newton* observation (obs. ID 0151270201). The broad structure is consistent with two clear surface brightness peaks associated with the two groups and surface brightness edges around

¹ See <http://cxc.harvard.edu/contrib/maxim/acisbg/>

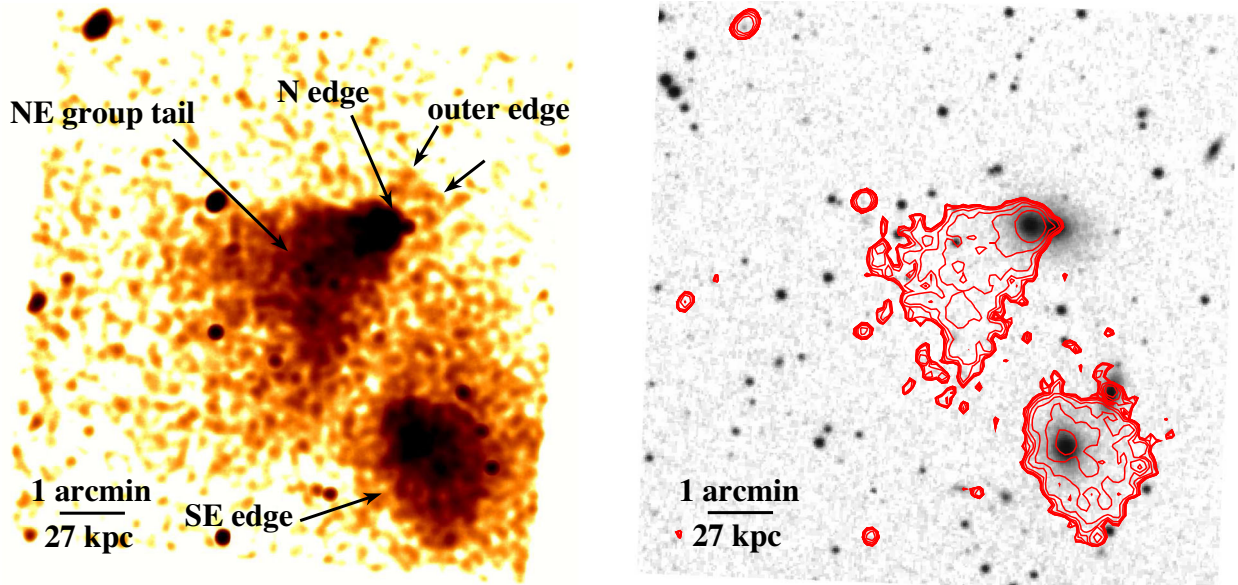


Figure 1. Left: exposure-corrected image in the 0.3–4.0 keV energy band smoothed with a 2D Gaussian $\sigma = 7$ arcsec (N is up and E is to the left). The arrows indicate a surface brightness edge ahead of the NE group. Right: DSS optical image covering the same field of view with X-ray contours.

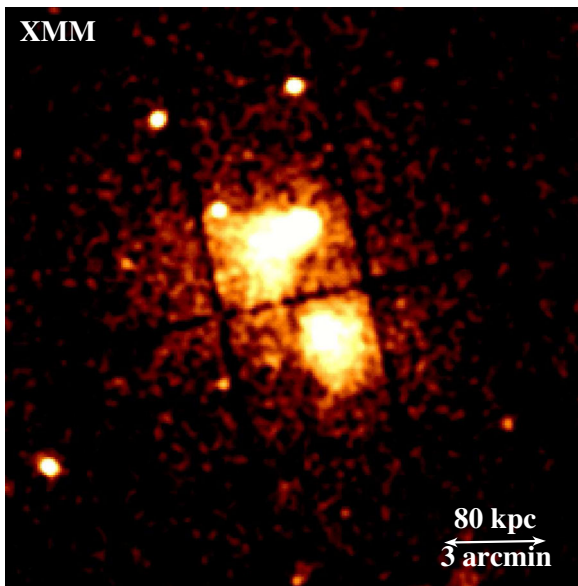


Figure 2. *XMM-Newton* MOS and PN merged image smoothed with a 2D Gaussian $\sigma = 12$ arcsec to show the large-scale environment of the galaxy groups.

each core. This image also reveals the larger scale hot atmosphere that the two group cores are moving through and clear diffuse emission ahead of the NE group. The ram pressure generated by the motion of the core through the hot ambient medium is sufficient to strip the outer gas layers from the NE group producing the observed long tail of debris. Similar structures have been observed behind dense, cool cores in major cluster mergers and in the Virgo cluster, where long plumes of gas extend behind infalling galaxies that are strongly interacting with the ICM (e.g. Forman et al. 1979; Nulsen 1982; Randall et al. 2008; Million et al. 2010; Sun et al. 2010; Kraft et al. 2011; Russell et al. 2012).

The DSS optical image (Fig. 1 right) shows the distribution of galaxies associated with the two groups. The NE group core hosts a pair of large elliptical galaxies (UGC 4052) separated by a projected distance of only 15 kpc (Crawford et al. 1999). The eastern galaxy of the pair, UGC 4052 east (E), is coincident with the X-ray surface brightness peak in the core and hosts luminous H α line emission ($L_{\text{H}\alpha} \sim 10^{40}$ erg s $^{-1}$; Crawford et al. 1999). UGC 4052 west (W) is located ahead of the NE group core and appears associated with a small spur of X-ray emission or ‘nose’ extending ~ 7 kpc (15 arcsec; Fig. 3 upper left). Another large elliptical galaxy (UGC 4051) is coincident with the X-ray surface brightness peak in the SW group. The redshifts for these three large ellipticals from Crawford et al. (1999) show that these galaxies are associated with the two merging groups and not a chance superposition along the line of sight. There is a hard X-ray point source spatially coincident with UGC 4051 in the SW group, which could indicate an AGN. The Very Large Array FIRST survey does not show a radio point source coincident with this source (Becker, White & Helfand 1995). There is a radio point source spatially coincident with UGC 4052 E in the NE group. Although there does not appear to be a hard X-ray point source associated with the radio point source, this could be simply too faint to detect in the existing data.

The ram pressure stripped tail behind the NE group core shows irregular structures or ‘wings’ (Fig. 4; Machacek et al. 2006; Kraft et al. 2011; Roediger et al. 2012a,b). The northern wing contains ~ 110 counts and the southern wing contains ~ 190 counts, which corresponds to $> 10\sigma$ above the surrounding ambient level for each structure. Distortions in cold fronts can be generated by the shear flow around them, which produces Kelvin–Helmholtz instabilities. The shear flow is expected to be maximum along the sides of the NE group tail, which explains the contrast between the narrow N edge of the cold front and the distorted structure at larger angles to the direction of motion. Non-viscous hydrodynamical simulations of cold fronts in the ICM commonly show the development of these instabilities (e.g. Murray et al. 1993; Heinz et al. 2003; ZuHone, Markevitch & Johnson 2010; Roediger et al. 2011) and

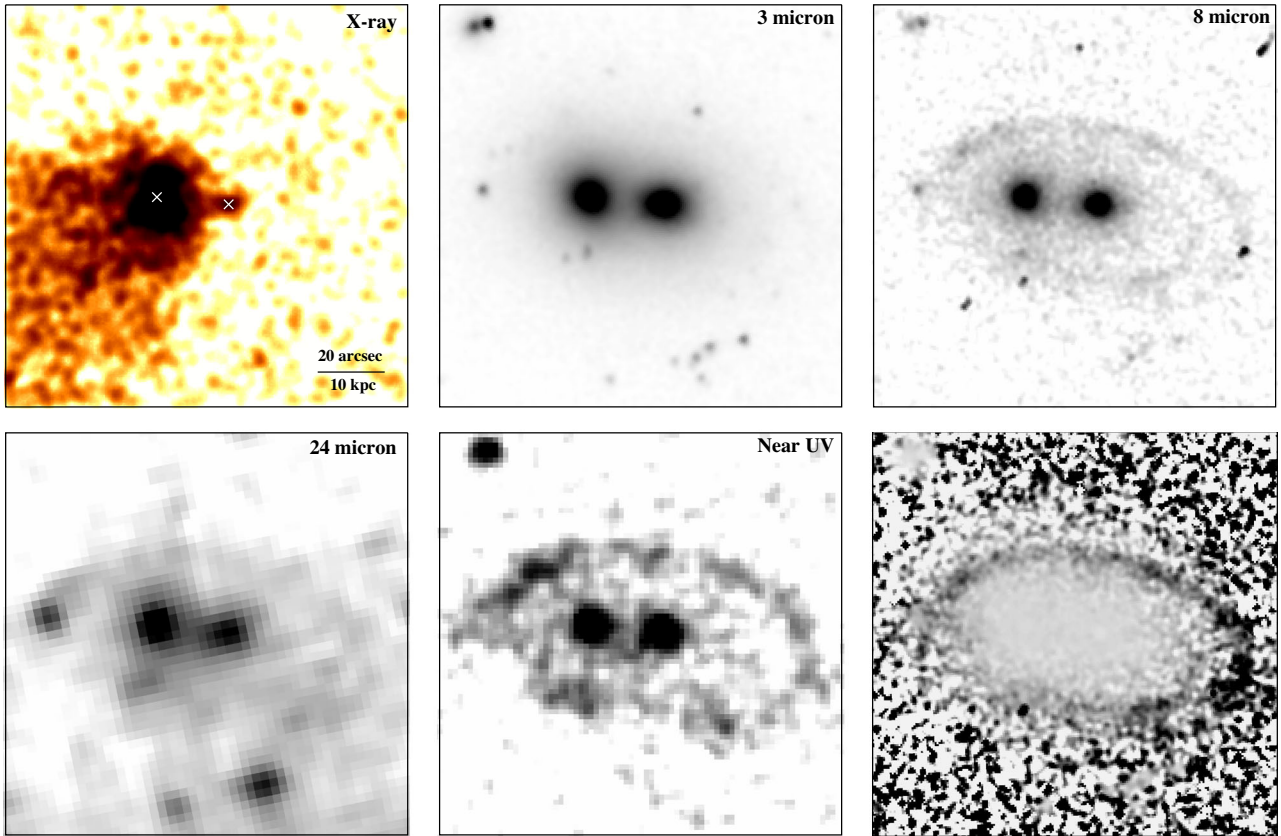


Figure 3. *Chandra* exposure-corrected image smoothed with a 2D Gaussian $\sigma = 6$ arcsec showing a $2 \text{ arcmin} \times 2 \text{ arcmin}$ region around the NE group core (upper left). The positions of the two elliptical galaxies are shown by the crosses. *Spitzer* IR images from IRAC at $3.6 \mu\text{m}$ (upper centre) and $8 \mu\text{m}$ (upper right) and MIPS at $24 \mu\text{m}$ (lower left) of the two central galaxies in the NE group (Quillen et al. 2008). The *GALEX* near-UV image is shown at lower centre. The sky-subtracted IRAC $3.6 \mu\text{m}$ image was also divided by the $8 \mu\text{m}$ image to show the ring emission more clearly (lower right). The *Spitzer* and *GALEX* images were smoothed with a 2D Gaussian $\sigma = 1$ arcsec. All images cover the same field of view.

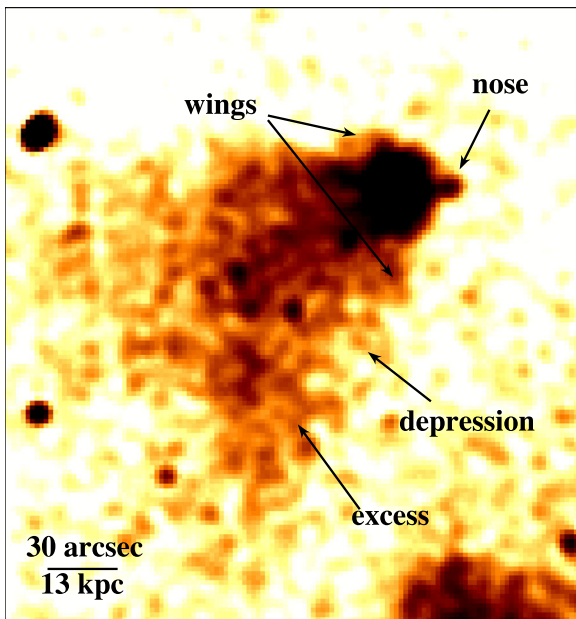


Figure 4. Exposure-corrected image showing structure in the ram pressure stripped tail of the NE group. The image covers the $0.3\text{--}4.0 \text{ keV}$ energy band and has been binned by a factor of 4 and smoothed with a 2D Gaussian $\sigma = 1.5$ arcsec.

can reproduce the structure of distortions in observed cold fronts (e.g. Roediger et al. 2012a,b). There is also an arc-like surface brightness excess and depression indicative of a hydrodynamical instability along the south (S) side of the tail (see Section 2.3; e.g. Mazzotta, Fusco-Femiano & Vikhlinin 2002; Roediger et al. 2013).

The ‘nose’ ahead of the NE group contains ~ 100 counts and is detected at more than 5σ above the ambient medium. This structure is more likely due to ram pressure stripping of an X-ray corona than a turbulent distortion of the cold front (Fig. 3 upper left; Sun et al. 2005a; Sun, Jerius & Jones 2005b). The N surface brightness edge appears very narrow and smooth on either side of this structure, rather than turbulent, and the ‘nose’ is spatially coincidence with a large elliptical galaxy, which is a likely host for an X-ray corona (Forman, Jones & Tucker 1985; Vikhlinin et al. 2001c; Sun et al. 2007). We discuss these structures in more detail in Sections 4 and 5.

UGC 4052 was included in the Quillen et al. (2008) *Spitzer* study of the mid-IR emission from the central galaxies in groups and clusters with optical emission lines. Quillen et al. (2008) found that UGC 4052 E has excess IR emission above that expected for a quiescent elliptical galaxy, which is likely due to hot dust powered by star formation. O’Dea et al. (2008) estimated a star formation rate for this galaxy from the IR luminosity of $0.7 - 1 M_{\odot}$. The *Spitzer* $8 \mu\text{m}$ image (Fig. 3 upper right) shows a striking bright oval ring of emission surrounding both galaxies with a diameter of ~ 40 kpc (90 arcsec). This feature is seen primarily at $8 \mu\text{m}$, which is likely

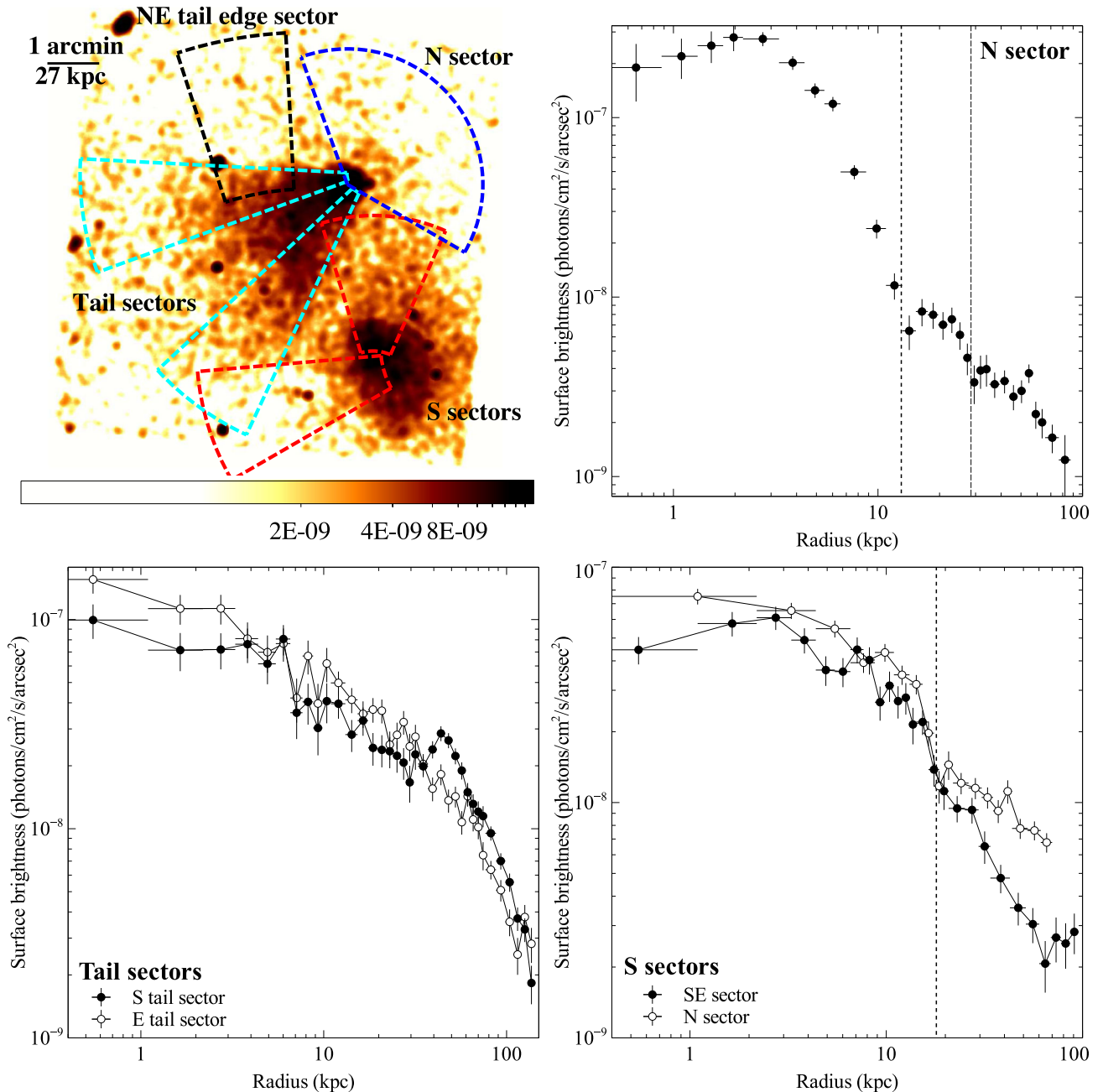


Figure 5. Background-subtracted surface brightness profiles in the energy band 0.3–4.0 keV for the N sector, Tail sectors and S sectors shown in the exposure-corrected image (upper left). Upper right: N sector from the centre of the NE subcluster across the leading edge. Lower left: profiles for sectors through the NE group subcluster tail. Lower right: profiles for the S sectors in the SW group. The surface brightness edges are marked by dotted and dashed lines. The surface brightness profile for the NE tail edge sector is shown in Fig. 10.

due to a polycyclic aromatic hydrocarbon (PAH) emission feature at $7.7 \mu\text{m}$, but is also seen faintly at $24 \mu\text{m}$ and in *GALEX* near-UV observations (Fig. 3 lower left and centre). The ring appears broken in places to the E and W and contains several knots of enhanced emission.

2.3 Radial profiles

We extracted projected surface brightness profiles across the edges identified in Section 2.2 and through the ram pressure stripped NE group tail. Fig. 5 (upper left) shows the regions selected, where the

centre of each sector was selected to coincide with the centre of curvature of the surface brightness edges. Radial bins were 1 arcsec in width for the bright, core regions and increased in width at larger radii to ensure a minimum of 30 source counts per spatial bin. The surface brightness profile for the N sector (Fig. 5 upper right) shows the clear N edge at 14 kpc where the surface brightness drops by a factor of ~ 10 over a radial distance of ~ 7 kpc across the leading edge of the NE group core. The N sector includes the ‘nose’ structure, which broadens the edge. The outer surface brightness edge tentatively identified in Fig. 1 is clearly detected as a drop by a factor of ~ 2 at a radius of 30 kpc. The location of this edge ahead of

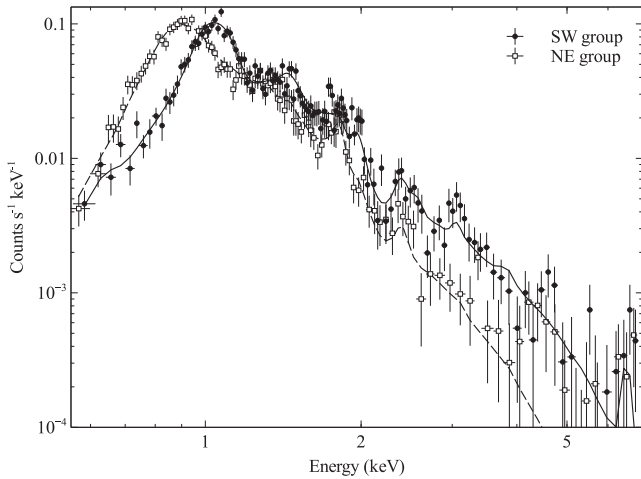


Figure 6. The total background-subtracted spectrum for the NE group (squares) and the SW group (circles) and the corresponding best-fitting absorbed *APEC* models (solid lines).

the NE group core suggests it could be a bow shock. In Section 2.4, we extract the temperature profiles across these edges to determine if there is a corresponding temperature jump.

Fig. 5 (lower left) shows surface brightness profiles through the ram pressure stripped tail of the NE group. The S and E tail sectors show a similar decline in surface brightness with radius. However, there is a clear excess in the S tail sector from 30 to 60 kpc radius, which corresponds to the excess surface brightness identified in Fig. 4. This surface brightness excess is preceded by a region of reduced surface brightness, by comparison with the E sector, from 10 to 30 kpc radius. The profiles confirm that these are significant structures and could correspond to a dynamical instability resulting from the merger. The location of this feature appears similar to that of the large instability identified in the merging cluster A3667 (Mazzotta et al. 2002; see also Roediger et al. 2013) and we discuss this further in Section 5.

The surface brightness profiles across the edge around the SW group core are shown in Fig. 5 (lower right). The surface brightness drops by a factor of ~ 2 at a radius of 20 kpc in both of the sectors analysed and the edge appears consistently narrow for ~ 90 kpc around the N and SE sides of the core. However, this analysis was limited due to the overlap with the ram pressure stripped material from the NE group core and the low number of counts did not allow us to place a strong upper limit on the width.

2.4 Spatially resolved spectroscopy

The total spectrum for each group was extracted from a circular region centred on each X-ray peak with a radius of 1 arcmin, which contained the bulk of the emission and excluded point sources. A corresponding background spectrum was extracted from the normalized blank-sky data set and appropriate responses and ancillary responses were generated. Each spectrum was restricted to the energy range 0.5–7 keV with a minimum of 20 counts per spectral channel. The spectra were fitted in *XSPEC* 12.8 (Arnaud 1996) with an absorbed single temperature *APEC* model version 2.0.2 (Smith et al. 2001). The cluster redshift was fixed to $z = 0.022$ (Ebeling et al. 1998; Crawford et al. 1999) and abundances were measured assuming the abundance ratios of Anders & Grevesse (1989). The temperature, metallicity and normalization parameters were left free.

The absorbing column density was determined using the spectrum extracted from the region covering the SW group, which excluded the cooler gas components in the NE group core and tail. The absorption parameter was left free and the best-fitting value was $N_{\text{H}} = 0.08 \pm 0.02 \times 10^{22} \text{ cm}^{-2}$. This is a little higher than the Galactic absorption found in H I surveys (Kalberla et al. 2005) but a similar analysis for the earlier *Chandra* observation also indicates absorption above the Galactic value. We therefore fixed the absorption parameter to this best-fitting value for all other spectral fits.

The total spectrum for each group is quite different (Fig. 6). The SW group has a global temperature of $1.60^{+0.02}_{-0.03}$ keV and a metallicity of $0.71^{+0.10}_{-0.09} Z_{\odot}$. The NE group is much cooler and metal poor with a temperature of 0.84 ± 0.01 keV and a metallicity of $0.16 \pm 0.01 Z_{\odot}$. We also used a two temperature spectral model to determine if the lower metallicity in the NE group was a result of the Fe bias at low temperatures (Buote 2000). However, the single temperature model was found to be a good fit to both total spectra and we did not find any significant additional temperature component in either galaxy group. This observed difference in metallicity follows the trend of higher metal abundances in the central regions of higher temperature groups (e.g. Rasmussen & Ponman 2009). We repeated the single temperature spectral fits leaving the redshift parameter free and obtained values of $z = 0.022^{+0.001}_{-0.002}$ and $0.026^{+0.001}_{-0.003}$ for the SW and the NE groups, respectively. Although this is not a significant difference, these results are consistent with the galaxy redshifts $z = 0.0218$ for UGC 4051 in the SW group and $z = 0.0236$ and 0.0228 for UGC 4052 E and W, respectively (Crawford et al. 1999). From the global temperature of each group and the mass–temperature relation (e.g. Finoguenov, Reiprich & Böhringer 2001; Sanderson et al. 2003; Vikhlinin et al. 2006), we estimate that the SW group is roughly three times more massive than the NE group. The global temperature of the NE group is likely to have been underestimated as ram pressure stripping has removed most of the warmer, lower density gas. However, adiabatic compression due to the ram pressure and static pressure of the larger group will counteract this to some extent. The mass ratio is therefore likely to be an upper limit.

Maps of the projected gas properties were produced using the Contour Binning algorithm (Sanders 2006), which generates spatial bins that closely follow the X-ray surface brightness. For the temperature, normalization and pressure maps, regions were selected to provide a signal-to-noise ratio of 32 (~ 480 counts). Larger regions with a signal-to-noise ratio of 22 were used for the metallicity map. The maps are 8 arcmin \times 8 arcmin covering the two groups. The length of the spatial regions was constrained to be at most twice their width. For each region, a spectrum was extracted from the *Chandra* data set and restricted to the energy range 0.5–7 keV. A corresponding background spectrum was extracted from the normalized blank-sky data set and appropriate responses and ancillary responses were generated. An absorbed single temperature *APEC* model was fitted to each spectrum with the temperature and normalization parameters left free and minimizing the *C*-statistic (Cash 1979). The metallicity was fixed to $0.3 Z_{\odot}$ for the temperature, normalization and pressure maps that were produced with fewer counts per spatial region.

Fig. 7 shows detailed maps of the best-fitting values for the projected emission measure per unit area, temperature and metallicity. The projected emission measure traces the square of the gas density in the group. The projected ‘pressure’ map was produced by multiplying the temperature map and the square root of the normalization per unit area map. The uncertainty in the normalization and

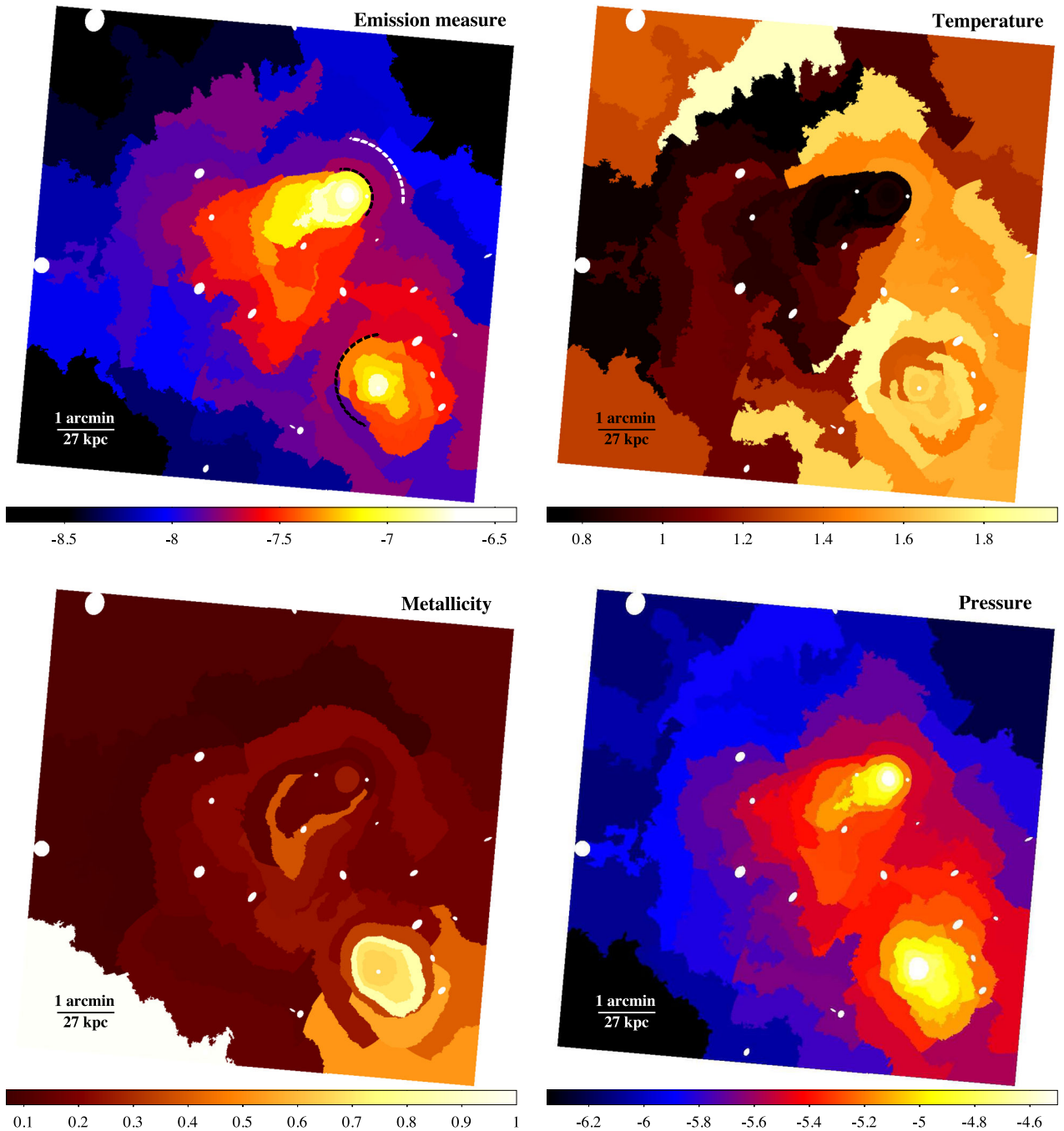


Figure 7. Upper left: projected emission measure per unit area map (units are $\log_{10} \text{cm}^{-5} \text{arcsec}^{-2}$). The emission measure is the *xspec* normalization of the *APEC* spectrum $K = \text{EI}/(4 \times 10^{14} \pi D_A^2 (1+z)^2)$, where EI is the emission integral $\text{EI} = \int n_e n_H dV$. Upper right: projected temperature map (units keV). Lower left: projected metallicity map (Z_{\odot}) generated using larger spatial bins. Lower right: projected pseudo-pressure map per unit area ($\log_{10} \text{keV cm}^{-5/2} \text{arcsec}^{-2}$). The excluded point sources are visible as small white circles. The positions of the cold fronts and shock front are shown on the emission measure map as black and white dashed lines, respectively.

temperature values at the centre of each group are ~ 4 and ~ 6 per cent, respectively. The temperature is below 1.5 keV here and Fe L line emission provides a good temperature diagnostic. At large radii, the uncertainty increases to ~ 10 per cent for the temperature and ~ 7 per cent for the normalization. The uncertainty in the metallicity values is ~ 20 per cent.

There is a clear peak in the emission measure at the centre of each of the groups. The surface brightness drops to the NW of the

NE group and along the SE edge of the SW group can be seen in the emission measure map. The emission measure declines smoothly through the tail of ram pressure stripped material behind the NE group core. The trail of densest gas in the tail appears to bend towards the S approximately 25 kpc (55 arcsec) behind the core. At this point, the tail also appears to flare out rapidly, increasing in width from 20 kpc (40 arcsec) to at least 70 kpc (160 arcsec). There is a region of increased emission measure to the N of the SW

group core suggesting there could be some ram pressure stripping. However, the SW group core does not appear strongly disrupted by the collision, which is on the borderline between a major and minor merger.

The temperature map shows that the temperature in the NE group drops below 0.8 keV in the dense core. The S group is much hotter, peaking at $1.71^{+0.09}_{-0.07}$ keV and the temperature increases across the surface brightness edge from $1.34^{+0.07}_{-0.06}$ keV to 1.7 ± 0.1 keV, indicative of a cold front. The ‘boxy’ morphology of the cold front, similar to A496 (Dupke & White 2003; Dupke et al. 2007), and the lack of strong disruption to the SW group suggests that this is a young sloshing cold front. The merger has displaced the hot atmosphere in the gravitational potential minimum and it is now starting to oscillate, producing the observed cold fronts (e.g. Markevitch & Vikhlinin 2007). Roediger et al. (2012a) found that the boxy structure in A496 arises due to the development of Kelvin–Helmholtz instabilities along cold fronts. The cold front appears strongest around the N edge of the SW group core where the outer hot atmospheres of each group are colliding. The entropy of the gas in this region appears comparable to that located to the SE of the SW group core suggesting that the gas has been compressed here by the collision between the two groups.

There is also a cold front around the leading edge of the NE group where the temperature increases much more steeply from 0.79 ± 0.04 to $1.5^{+0.2}_{-0.1}$ keV. The temperature increases through the ram pressure stripped tail with the gas along the S edge warmer than that along the NE edge. The NE edge of the tail appears much narrower, suggesting that there is much less mixing with the hotter, ambient medium, consistent with the more gradual increase in temperature. The outer surface brightness edge ahead of the NE group core appears coincident with a temperature drop from $1.7^{+0.4}_{-0.3}$ to 1.3 ± 0.1 keV. This could therefore be a bow shock and we analyse this structure in detail in Section 3.

The SW group has a strong, central metallicity peak of $0.8^{+0.3}_{-0.2} Z_{\odot}$. The NE group appears metal poor by comparison with a peak of only $0.35^{+0.14}_{-0.09} Z_{\odot}$ in the dense core. The metallicity is lowest at $0.10^{+0.03}_{-0.02} Z_{\odot}$ for 27 kpc (60 arcsec) behind the NE group core and then shows a modest increase, particularly towards the S tail edge where there is likely to be mixing with metal-rich material stripped from the SW group. The projected ‘pressure’ map shows the extent of the shock heating between the trajectory of the two groups. There is also a decrease in pressure ahead of the NE group core consistent with a bow shock.

2.5 Hardness ratio

Although there are too few counts to produce a spectrum for the NE group ‘nose’, we have produced a hardness ratio profile across this region to look for evidence of temperature variations. We have used the Bayesian method of calculating hardness ratios detailed by Park et al. (2006), which is particularly useful for the low counts regime. The hardness ratio was calculated as the fractional difference between the number of counts in the soft (0.5–1.5 keV) and hard (1.5–5.0 keV) energy bands,

$$HR \equiv \frac{H - S}{H + S}. \quad (1)$$

The source counts in each energy band were extracted from a series of regions covering a narrow sector across the two galaxies. The blank-sky background counts were also modelled as independent Poisson random variables.

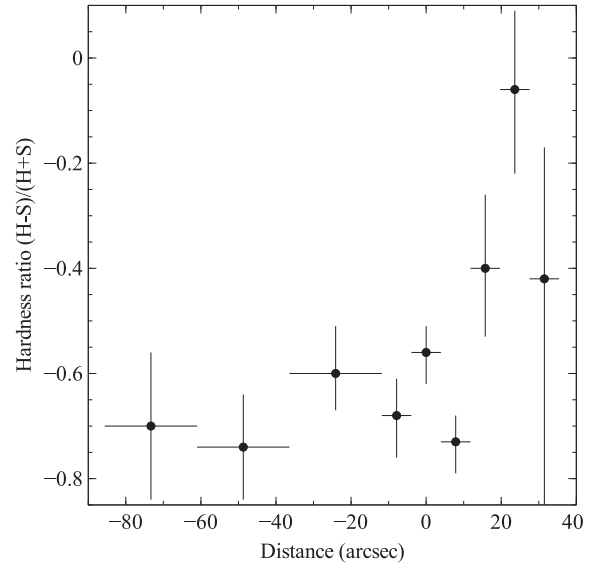


Figure 8. Hardness ratio profile using a soft energy band 0.5–1.5 keV (S) and a hard energy band 1.5–5.0 keV (H). The profile is centred on the E elliptical galaxy in the N group and the W galaxy is located +22 arcsec.

Fig. 8 shows that the hardness ratio is roughly constant at -0.65 through the NE group cool core then appears to increase at the location of the ‘nose’ where the W galaxy is located. However, there are large uncertainties on the individual values. This could indicate a few keV temperature increase in the ‘nose’ caused by heating from the surrounding post-shock gas. Alternatively, there could be a hard X-ray point source at the W galaxy centre consistent with an AGN. Although the FIRST survey does not show a radio point source coincident with the W galaxy (Becker et al. 1995), there also does not appear to be radio emission associated with the hard X-ray point source at the centre of the SW group. The increase in hardness ratio appears to extend over several neighbouring regions, which suggests a temperature increase rather than a compact point source.

3 BOW SHOCK DETECTION

Fig. 5 (upper left) shows the N sector selected for the analysis of the possible bow shock ahead of the NE group. The sector was positioned to align with the curvature of this outer surface brightness edge and covers the angular range over which the edge is clearly defined. This sector was divided into five radial bins that were positioned to determine the gas properties on either side of the surface brightness edges with a minimum of 300 source counts per region. Spectra extracted from these regions were each fitted with an absorbed single temperature APEC model, as detailed in Section 2.4. The metallicity was fixed to the average of $0.2 Z_{\odot}$ found for the N sector (Fig. 7). The projected temperature profile for the N sector is shown in Fig. 9. The two central temperature bins at 0.8–0.9 keV cover the NE group cool core. The temperature increases ahead of the core to $2.0^{+0.5}_{-0.2}$ keV and then drops to 1.5 ± 0.1 keV beyond the outer surface brightness edge at a radius of 27 kpc. The temperature and surface brightness, and therefore the gas density, decrease significantly with radius across this outer edge showing that it is a bow shock front ahead of the NE group core.

The XSPEC deprojection model PROJECT was used to subtract the projected contributions from the outer cluster layers off the inner spectra and produce deprojected temperature and density profiles.

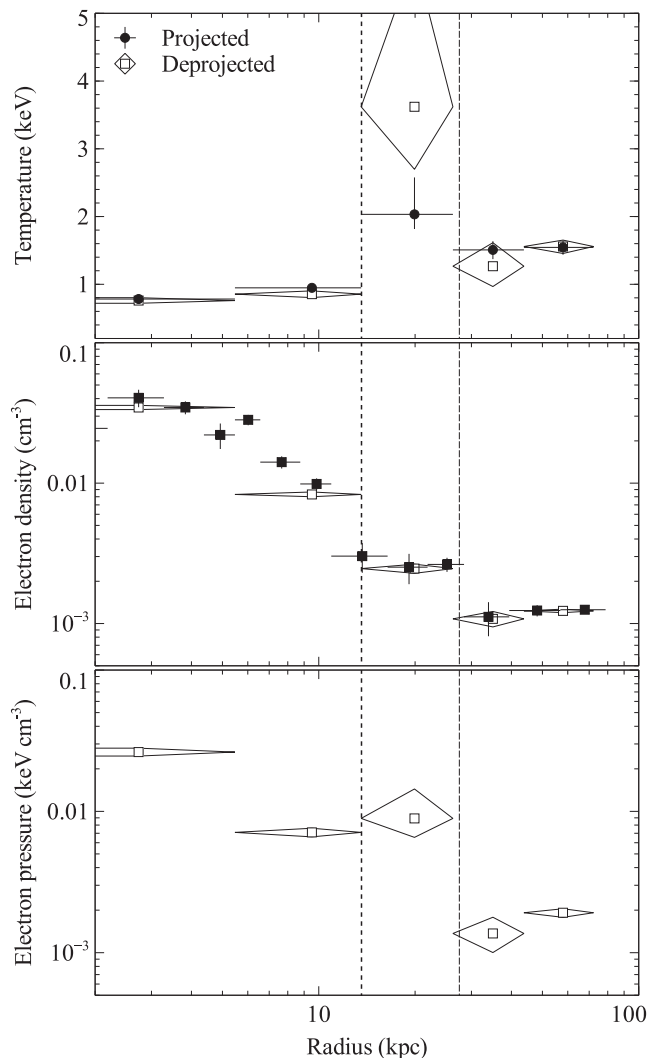


Figure 9. Temperature (upper), electron density (middle) and electron pressure (lower) profiles for the N sector (see Fig. 5). The position of the surface brightness edge ahead of the N group core is shown by the dashed line. The square symbols denote deprojected points. The deprojected electron density profile was produced with `PROJECT` (open squares) and a surface brightness deprojection (filled squares). The dotted line at 14 kpc radius marks the merger cold front across the leading edge and the dashed line at 28 kpc radius marks the shock front.

Assuming spherical symmetry, `PROJECT` simultaneously fits to all the projected spectra accounting for projected emission by adding suitably scaled model components from the outer annuli to the inner spectral fits. The X-ray emissivity depends strongly on the gas density ($\propto n^2$) and only weakly on the temperature. Therefore, it is possible to produce density profiles with higher spatial resolution by converting the surface brightness to electron density and using modest corrections for the temperature and exposure variation. Fig. 9 shows a finer resolution deprojected electron density profile overplotted on the `PROJECT` result. The electron density drops by a factor of 2.2 ± 0.4 across the bow shock at a radius of 27 kpc. This coincides with a decrease in the temperature by a factor of 3_{-1}^{+2} . Applying the Rankine–Hugoniot shock jump conditions (e.g. Landau & Lifshitz 1959), we can calculate the Mach number of the shock, $M = v/c_s$, where v is the pre-shock gas velocity with respect to the shock surface and c_s is the speed of sound in the pre-shock gas. The

Mach number can be calculated from the density jump,

$$M = \left(\frac{2 \left(\frac{n_{e,2}}{n_{e,1}} \right)}{\gamma + 1 - \left(\frac{n_{e,2}}{n_{e,1}} \right) (\gamma - 1)} \right)^{1/2}, \quad (2)$$

where $n_{e,1}$ and $n_{e,2}$ are the electron densities upstream and downstream of the shock, respectively. An adiabatic index of $\gamma = 5/3$, appropriate for a monatomic gas, is assumed. The Mach number for the bow shock is therefore $M = 1.9 \pm 0.4$. Although the uncertainty on the post-shock temperature is large, the temperature jump appears broadly consistent with the density jump. The sound speed in the pre-shock gas is $580 \pm 70 \text{ km s}^{-1}$ and therefore the shock velocity is $1100 \pm 300 \text{ km s}^{-1}$.

We used the shock velocity to estimate the time since the closest passage between the two groups. The galaxy redshifts suggest a line of sight velocity difference of $\sim 400 \text{ km s}^{-1}$ between the two groups (section 2.4). Therefore, we estimate that the merger axis is $\sim 10^\circ$ from the plane of the sky. The detection of a shock front is also consistent with a merger axis close to the plane of the sky because projection effects would otherwise obscure the surface brightness edge. From the projected separation between the two groups of $\sim 90 \text{ kpc}$, we estimate an age of $\sim 0.1 \text{ Gyr}$.

The deprojected temperature and electron density were also multiplied together to produce the electron pressure profile (Fig. 9). This shows the expected sharp drop in pressure across the bow shock front. The pressure across the leading edge of the NE group core is continuous within the error showing that this is a merger cold front. The pressure then peaks at the centre of the NE group core.

4 NE GROUP CORE STRUCTURE

The N edge of the NE group core (Fig. 1 left) represents a sharp transition between the 1.5 keV ambient gas and the cool core gas at only 0.7 keV. Fig. 7 (upper right) shows that this steep temperature gradient is preserved over $\sim 35 \text{ kpc}$ distance behind the core along the NE edge of the ram pressure stripped tail. However, the tail appears to broaden rapidly at a distance of $\sim 45 \text{ kpc}$ behind the dense core. We compared a surface brightness profile across this edge with the surface brightness profile across the N edge (Fig. 5 upper right) to demonstrate this broadening (Fig. 10). The surface brightness profile for the leading edge excluded the ‘nose’. The profiles were divided into radial bins of 5 arcsec width, increasing to 25 arcsec at larger radii to ensure a minimum of 30 source counts per spatial bin. The energy range was restricted to 0.3–4.0 keV to maximize the signal-to-noise ratio. The background was subtracted using the normalized blank-sky background data set (see Section 2.1). Fig. 10 shows the final background-subtracted surface brightness profiles with the characteristic shape of a projected density discontinuity. The leading edge of the dense core is particularly narrow and appears broader in the sector along the NE edge of the subcluster’s tail.

The surface brightness edges were fitted with a model for a projected spherical density discontinuity (e.g. Markevitch et al. 2000, 2002; Owers et al. 2009). The radial density model consists of two power laws either side of a density jump, which is then projected under the assumption of spherical symmetry. The model is then convolved with a Gaussian function, to account for any intrinsic width, and fitted to the surface brightness profiles. The free parameters are the gradients and normalizations of the power-law models, the radius of the density jump and the width of the density jump. The best-fitting model for the leading edge and the NE tail edge are

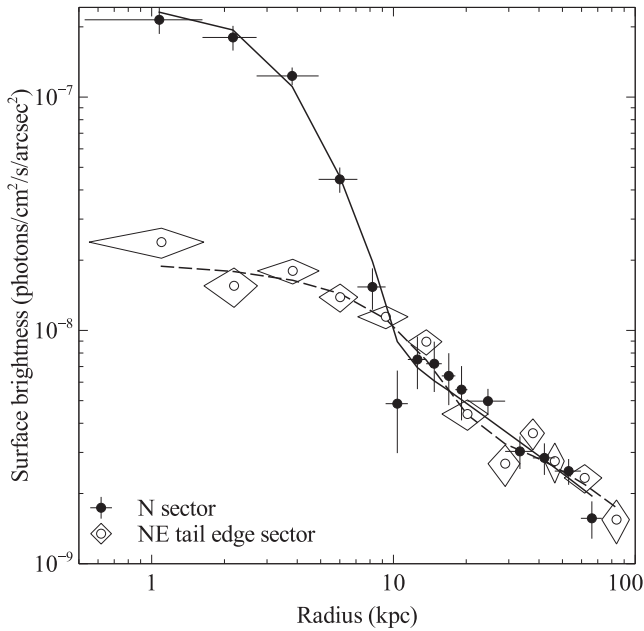


Figure 10. Surface brightness profiles across the N sector (excluding the ‘nose’) and the NE tail edge. These regions are shown in Fig. 5 (upper left). The best-fitting projected density discontinuity models are shown overlaid (solid and dashed lines).

shown overlaid on the surface brightness profiles in Fig. 10. The leading edge of the NE group core has a best-fitting width of 2.4 ± 0.7 kpc. We note that this is an upper limit as any variations in the shape of the edge over the sector used will broaden the profile. The edge widens to 6 ± 1 kpc in the sector across the NE edge of the subcluster’s tail, which coincides with the observed broadening of the temperature jump. This is likely caused by the development of turbulent instabilities along this sheared interface and will promote mixing of the cold ram pressure stripped gas into the ambient medium.

The ‘nose’ of the NE group core is likely to be a ram pressure stripped X-ray corona hosted by the large elliptical galaxy UGC 4052 W. Following Nulsen (1982), the mass-loss rate due to ram pressure is $\dot{M}_{\text{loss}} \sim \pi r^2 \rho_h v \sim 0.3 - 1 M_{\odot}$, where ρ_h is the gas density in the ambient medium, r is the cross-sectional radius of the corona and v is the gas velocity. Assuming a cylindrical volume and a density comparable to the cool core, we estimate the total gas mass in the corona $M \sim 3 \times 10^7 M_{\odot}$ and therefore the survival time of the corona is ~ 0.1 Gyr (see also Sun et al. 2005a,b). The corona may survive longer if there is significant mass injection from the stellar population or ICM cooling (e.g. Acreman et al. 2003) or if the transport processes are suppressed by a magnetic field (Vikhlinin et al. 2001c; Sun et al. 2005a).

The striking IR and near-UV ring surrounding the galaxy pair UGC 4052 is ~ 40 kpc across and the major axis appears aligned with the central galaxies (Fig. 3). This alignment suggests that the formation of the ring could be linked to these galaxies. The group merger may have caused collisions between the constituent galaxies and the ring could correspond to the orbit of a smaller galaxy being stripped of its gas in the interaction. However, there is not a clear progenitor galaxy and the apparent spiralling of the orbit to the W complicates this interpretation (Fig. 3). The ring may have been produced by the stripping of multiple galaxies around UGC 4052. An explosion origin appears less likely as the centre of the ring falls between the pair of galaxies but cannot be ruled

out. Another possibility is that the ring was produced by a close encounter between the two central galaxies (e.g. Lynds & Toomre 1976; Theys & Spiegel 1976; Appleton & Struck-Marcell 1987, 1996). The passage of the companion galaxy close to the centre of the parent galaxy produces an outwardly travelling density wave in the gas and stellar discs and triggers star formation in the expanding ring. However, the two galaxies are aligned along the major rather than the minor ring axis in contrast to other collisional ring galaxies, such as the Cartwheel ring galaxy (e.g. Higdon 1995), Auriga’s Wheel (Conn et al. 2011) and AM 1724–622 (Wallin & Struck-Marcell 1994). There is a velocity difference of ~ 250 km s $^{-1}$ along the line of sight between the two galaxies (Crawford et al. 1999) and, if this interaction was triggered by the group merger, there could be a greater velocity difference in the plane of the sky. UGC 4052 E was targeted in the SDSS (Abazajian et al. 2009) and the spectrum shows a possible velocity offset between the gas emission lines and absorption lines in the stellar continuum. There could therefore also be a difference in the gas and stellar dynamics in this complex system.

Fig. 3 shows that at least part of the ring, depending on the projection, is located in the shock-heated ambient medium. PAHs, which produce the strong emission feature at $7.7 \mu\text{m}$, are easily destroyed by ionizing UV and X-rays and by shocks and therefore must be shielded somehow in this environment (e.g. Voit 1992).

5 RAM PRESSURE STRIPPING OF THE NE GROUP TAIL

Fig. 4 shows distortions around the sheared sides of the NE group’s ram pressure stripped tail and a potential large-scale hydrodynamic instability (e.g. Heinz et al. 2003; Iapichino et al. 2008). The shear flow around the NE group core will be greatest along the S edge where the two groups are moving past each other and the speed of the ambient medium is maximum (e.g. Mazzotta et al. 2002). Fig. 7 shows that there is increased temperature and pressure in this region between the two groups. We compare the E and S sectors of the ram pressure stripped tail to determine if there is potentially mixing between the cool, dense core gas and the hot, ambient medium (Fig. 11). The temperature profiles increase through both sectors from ~ 0.75 keV to ~ 1.0 keV, with a higher temperature in the S sector out to 65 kpc. The temperature in the S sector increases significantly at a radius of 25 kpc to $1.01^{+0.03}_{-0.04}$ keV, which is comparable to the ambient gas. This corresponds to the surface brightness depression identified in Fig. 4 (see Section 2.3). The surface brightness excess in the neighbouring region at a radius of 37 kpc has a temperature consistent with the NE group tail. The excess is therefore a colder, denser structure extending into the ambient medium while the depression is a hotter, lower density filament penetrating from the ambient medium into the cool core. The temperature and surface brightness structure are therefore consistent with the shearing along this edge producing turbulent mixing of the gas between the two merging groups.

Turbulent mixing is the most natural explanation for this structure observed in a group merger and similar features have been found in other merging systems. The dense, cool cores in the cluster mergers in A3667 and A2146 show comparable ram pressure stripped structures and turbulent instabilities generated by the shear flow (Mazzotta et al. 2002; Russell et al. 2012). Observations of galaxies moving through the ICM also show long tails of stripped gas, which can be distorted by Kelvin–Helmholtz instabilities (Machacek et al. 2005, 2006; Randall et al. 2008; Sun et al. 2010; Kraft et al. 2011). The growth of these instabilities can be suppressed by the surface

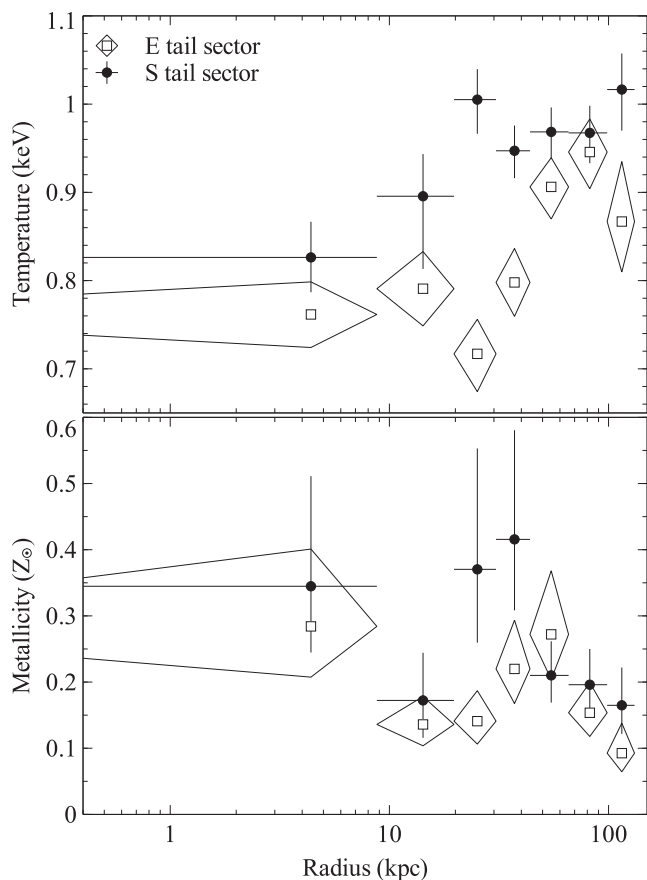


Figure 11. Temperature (upper) and metallicity (lower) profiles for the NE group tail. Note that these sectors are similar to the tail sectors shown in Fig. 5 but also extend over the centre of the tail.

tension of a magnetic field and the viscosity of the gas, which may explain why some cold fronts appear smooth and undisturbed (e.g. Machacek et al. 2005). Gravity is not expected to be important for suppressing Kelvin–Helmholtz instabilities in these systems (e.g. Chandrasekhar 1961; Roediger et al. 2012b). The presence or absence of distortions to cold fronts can therefore be used to constrain the local ICM viscosity and magnetic field strength in the ICM.

Tangled magnetic fields may be amplified and stretched by the gas flow around cool cores in a merger (e.g. Vikhlinin et al. 2001b; Dursi & Pfrommer 2008). Kelvin–Helmholtz instabilities can be suppressed by this magnetic layer that is aligned with the interface and the flow direction. Therefore, observations of instabilities put an upper limit on the magnetic field strength at this location in RX J0751.3+5012. Following Vikhlinin & Markevitch (2002), the criterion for a cold front *stabilized* by the magnetic pressure is

$$\frac{B_c^2 + B_h^2}{8\pi} > \frac{1}{2} \frac{\gamma M^2}{1 + T_c/T_h} P_{\text{ICM}}, \quad (3)$$

where B_c and B_h (μG) are the magnetic field strengths in the cold and hot gas either side of the discontinuity, T_c and T_h are the corresponding gas temperatures, M is the Mach number for the gas flow and P_{ICM} is the ambient pressure. We assume the adiabatic index $\gamma = 5/3$ for a monatomic gas. Therefore, for a maximum post-shock gas velocity with $M \sim 0.6$, gas temperatures either side of the discontinuity of $T_h = 2.0 \text{ keV}$ and $T_c = 0.8 \text{ keV}$, and $P_{\text{ICM}} \sim 0.01 \text{ keV cm}^{-3}$, the total magnetic field at this discontinuity is less than $9 \mu\text{G}$. This is consistent with observa-

tional constraints from Faraday rotation measurements for other galaxy groups (e.g. Laing et al. 2008; Guidetti et al. 2010).

Roediger et al. (2012a) show that viscosity is not expected to significantly suppress the development of Kelvin–Helmholtz instabilities in galaxy groups. Following Roediger et al. (2013) and assuming *Spitzer*-like viscosity, we calculate the critical perturbation length-scale below which viscosity will increase the growth time of perturbations

$$\lambda_{\text{crit}} = 5 \text{ kpc} \left(\frac{U}{500 \text{ km s}^{-1}} \right)^{-1} \left(\frac{n_e}{3 \times 10^{-3} \text{ cm}^{-3}} \right)^{-1} \left(\frac{k_B T_h}{2 \text{ keV}} \right)^{5/2} \quad (4)$$

where U is the shear velocity and n_e is the ambient gas density. Therefore, for RX J0751.3 + 5012, we estimate that full *Spitzer* viscosity will suppress the growth of Kelvin–Helmholtz instabilities on length-scales below 5 kpc. This appears consistent with the observed distortions on scales of $\sim 10 \text{ kpc}$ around the NE group core but there may be smaller scale instabilities that cannot be detected significantly in the existing observation.

6 WHL J075052.4+500252

Fig. 12 (left) shows the full ACIS-I field of view. The two merging subgroups of RX J0751.3+5012 are visible to the NE on ACIS-I3 but there is also extended emission clearly detected above the background to the SW on ACIS-I0. The X-ray emission peak for this additional extended source is located at RA $07^{\text{h}}50^{\text{m}}52.4 \text{ Dec.} +50^{\circ}02'55''$ (J2000). This corresponds to the galaxy cluster WHL J075052.4+500252 at $z = 0.417$ (Hao et al. 2010; Wen, Han & Liu 2012). Fig. 12 (right) shows the background-subtracted X-ray spectrum for this source and the background spectrum. Using an absorbed single temperature *APEC* model, we found a best-fitting temperature of $5.2^{+1.3}_{-0.6} \text{ keV}$ and total X-ray flux of $2.4^{+0.2}_{-0.1} \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$ (0.05–50 keV). The metallicity was fixed to $0.3 Z_{\odot}$ and the redshift was fixed to $z = 0.417$. There was no clear detection of the Fe K line to confirm the redshift. The best-fitting values of temperature and luminosity ($L_X = 1.49^{+0.12}_{-0.06} \times 10^{44} \text{ erg s}^{-1}$ for 0.05–50 keV) place WHL J075052.4+500252 on the L_X – T relation within the errors (Markevitch 1998; Arnaud & Evrard 1999; Pratt et al. 2009).

7 CONCLUSIONS

The new *Chandra* observation of RX J0751.3+5012 has revealed an off-axis merger between two galaxy groups. The NE group hosts a dense, cool core that has been strongly disrupted by ram pressure in the collision and formed a $\sim 100 \text{ kpc}$ -long tail of stripped gas. Sharp surface brightness edges are observed around each group core corresponding to cold fronts where the cooler, dense core gas moves through the warmer ambient medium. The boxy structure and lack of disruption to the SW group suggests the cold front around the core is due to gas sloshing recently triggered by the collision. We observe an additional surface brightness edge $\sim 15 \text{ kpc}$ (30 arcsec) ahead of the NE group core associated with a sharp drop in gas density, temperature and pressure. This is a bow shock front, the first detected in a group merger, with a Mach number $M = 1.9 \pm 0.4$. From the shock velocity of $1100 \pm 300 \text{ km s}^{-1}$ and constituent galaxy redshifts, we estimate that the merger axis lies close to the plane of the sky at an angle of $\sim 10^{\circ}$. Together with the projected distance between the groups of $\sim 90 \text{ kpc}$, we determine that closest approach occurred $\sim 0.1 \text{ Gyr}$ ago. From the global temperature of

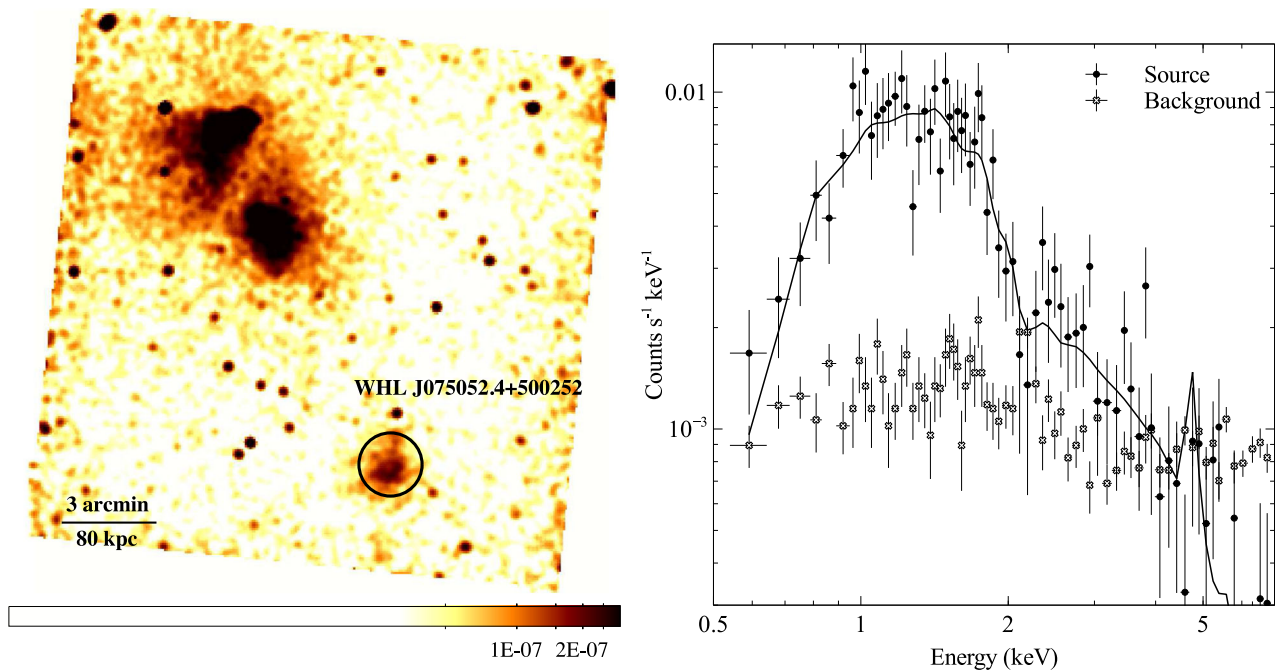


Figure 12. Left: exposure-corrected image showing the full field of the ACIS-I detector in the 0.3–4 keV energy band. The extended emission from WHL J075052.4+500252 is visible on ACIS-I0. The spectral extraction region is shown by the solid circle. Right: background-subtracted spectrum for WHL J075052.4+500252 with best-fitting model overlaid (solid line).

each group, we estimated that the SW group is three times more massive than the NE group and therefore the collision is on the borderline between a major and minor merger.

Although the leading edge of the NE group core is narrow, the sheared sides of the core and stripped tail are broader and appear distorted by Kelvin–Helmholtz instabilities. The gas temperature through the ram pressure stripped tail increases more rapidly along the S edge where the two groups are moving past each other and the shear flow is strongest. In addition, we observe an arc-like surface brightness excess and depression consistent with the development of a large-scale hydrodynamical instability along the S edge of the tail. The growth of these turbulent eddies can be suppressed by magnetic fields and viscosity. Therefore, we show that, for *Spitzer*-like viscosity, the presence of these instabilities is consistent with the critical perturbation length above which instabilities can grow and place an upper limit on the magnetic field strength in the intragroup medium.

The NE group hosts two large elliptical galaxies (UGC 4052) separated by a projected distance of ~ 15 kpc and surrounded by an IR and near-UV ring with a diameter of ~ 40 kpc. The E galaxy is coincident with the X-ray peak and luminous $H\alpha$ line emission located at the centre of the group’s cool core. The W galaxy is positioned ahead of the group core and hosts a ram pressure stripped X-ray corona. The surrounding ring could have been produced by material stripped from a smaller galaxy orbiting around UGC 4052 or it may be a collisional ring generated by a close encounter between the two large galaxies.

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