

A MAP OF DUST REDDENING TO 4.5 kpc FROM Pan-STARRS1

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ABSTRACT

We present a map of the dust reddening to 4.5 kpc derived from Pan-STARRS1 stellar photometry. The map covers almost the entire sky north of declination -30° at a resolution of $7''-14''$, and is based on the estimated distances and reddenings to more than 500 million stars. The technique is designed to map dust in the Galactic plane, where many other techniques are stymied by the presence of multiple dust clouds at different distances along each line of sight. This reddening-based dust map agrees closely with the Schlegel et al. (SFD) far-infrared emission-based dust map away from the Galactic plane, and the most prominent differences between the two maps stem from known limitations of SFD in the plane. We also compare the map with *Planck*, finding likewise good agreement in general at high latitudes. The use of optical data from Pan-STARRS1 yields reddening uncertainty as low as 25 mmag $E(B - V)$.

Key words: dust, extinction – ISM: clouds

Online-only material: color figures

1. INTRODUCTION

Dust absorbs and scatters ultraviolet through infrared (IR) light, and re-emits the absorbed energy thermally in the mid-IR through far-IR (FIR). These processes reshape the radiation field of the Galaxy, playing a crucial role in many areas of Galactic astrophysics and additionally obscuring and masking astronomical sources observationally.

Because of this influence, maps of dust and its properties are widely used in astronomy. Dust maps based on neutral hydrogen gas (Burstein & Heiles 1978) and, more recently, on thermal dust emission (Schlegel et al. 1998, SFD) have enjoyed wide use. The advent of wide field digital surveys like the Sloan Digital Sky Survey (SDSS; York et al. 2000) and the Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006) have allowed dust reddening, as opposed to emission, to be mapped accurately on large scales for the first time. Recent SDSS papers on the topic include works using the photometry of stars (Schlafly et al. 2010; Berry et al. 2012) and galaxies (Yasuda et al. 2007; Peek & Graves 2010) as well as spectroscopy (Schlafly & Finkbeiner 2011; Yuan et al. 2013). Studies of dust based on data from 2MASS include photometric studies of reddening (Lombardi & Alves 2001; Rowles & Froebrich 2009; Majewski et al. 2011; Nidever et al. 2012) as well as others using number counts (Dobashi et al. 2005) or combinations of both (Marshall et al. 2006). The work of Lallement et al. (2014) maps the three-dimensional (3D) structure of the dust within about 1 kpc. These works have provided new constraints on the nature of the dust reddening law, the variation in dust properties over the sky, and the distribution of mass in clouds, to mention a few results.

The advent of *Planck* data makes this an especially interesting time for dust mapping (Planck Collaboration et al. 2011, 2013).

The *Planck* data provide full-sky coverage of dust emission from 353–857 GHz. Large-scale reddening maps complement this data, allowing the dust to be studied simultaneously in both emission and extinction.

As part of a larger effort to study the 3D distribution of dust using PS1 photometry, we present here a new map of dust reddening using photometry of stars observed by Pan-STARRS1 (Kaiser et al. 2010). The map is constructed using the technique of Green et al. (2014), in which the distance and reddening to each star is inferred based on its observed photometry, assuming a standard Galactic extinction law with $R_V = 3.1$ (Fitzpatrick 1999). Along each $7'' \times 7''$ line of sight ($14''$ when $|b| > 30^\circ$), we model the reddening as a function of distance to infer the reddening to 4.5 kpc, where the reddening is particularly well constrained. This map has an unprecedented combination of sky coverage, sensitivity, and resolution compared to earlier direct measurements of dust reddening, and is available at our Web site.⁹ The work of Schlafly et al. (2014) has shown the accuracy of the distances estimated by our approach; this work serves to demonstrate the performance of our reddening estimation.

We describe in Section 2 the Pan-STARRS1 survey, which provides the observation on which our map is based. We then present our technique and the resulting reddening map in Sections 3 and 4. We compare it to the Schlegel et al. (1998) dust map and the *Planck* dust map in Section 5. We conclude in Section 6.

2. THE Pan-STARRS1 SURVEY

The Pan-STARRS1 survey, now reaching completion, has observed the entire sky north of declination -30° in five filters

⁹ <http://faun.rc.fas.harvard.edu/eschlafly/2dmap>

covering 400–1000 nm (Stubbs et al. 2010; Tonry et al. 2012). The 1.8 m PS1 telescope has a 7 deg² field of view outfitted with a billion-pixel camera (Hodapp et al. 2004; Tonry & Onaka 2009; Onaka et al. 2008), and reaches a 5 σ single epoch depth of about 22.0, 22.0, 21.9, 21.0, and 19.8 mag in *grizy_{p1}*. The survey pipeline automatically processes images and performs photometry and astrometry on detected sources (Magnier 2006, 2007; Magnier et al. 2008). The photometric calibration of the survey is better than 1% (Schlafly et al. 2012).

We infer the dust reddening to 4.5 kpc from estimates of the distance and reddening to all stars well-observed by Pan-STARRS1. For our purposes, we define “well-observed” to mean that the star has been observed in good conditions in at least the *g_{p1}* filter and three of the other four PS1 filters. We use average PS1 photometry in each band; we ignore the variability information encoded in the multiple PS1 observations of each object. We use point source photometry, and reject galaxies by demanding that the aperture magnitudes of the objects be brighter than the point-spread function (PSF) magnitudes by less than 0.1 mag in at least three bands, a criterion chosen to yield a clean stellar locus at high Galactic latitudes. The resulting catalog contains more than 500 million stars over three-quarters of the sky.

3. METHOD

The observed photometry of a star depends on the light emitted by the star itself, and also on the distance and reddening to the star. Accordingly, we model the photometry of each star individually, with parameters for the distance, reddening, and type of the star, as described in Green et al. (2014). We assume here a fixed $R_V = 3.1$ reddening law from Fitzpatrick (1999), translated into the PS1 bands by Schlafly & Finkbeiner (2011). We use an empirical set of models for the PS1 colors using fits to the shape of the PS1 stellar locus from Green et al. (2014). These colors are associated with absolute magnitudes and metallicities according to the work of Ivezić et al. (2008). We determine the full probability distribution function for each star’s distance, reddening, absolute magnitude, and metallicity. In this analysis, we impose priors on the distribution of stars and their types using the Galactic model of Jurić et al. (2008) and the metallicity study of Ivezić et al. (2008). We adopt a luminosity function assuming a Chabrier (2001) initial mass function and using the PARSEC stellar evolution models (Bressan et al. 2012; Green et al. 2014). By marginalizing out the absolute magnitude and metallicity, we obtain the posterior probability distribution function $p(E, D|\mathbf{m})$ describing the range of possible reddenings E and distances D to each star with PS1 photometry \mathbf{m} . We note that all these priors are spatially smooth and described by only around twenty parameters, and carry no information about the detailed distribution of the dust. This technique is similar to those of Berry et al. (2012), Sale et al. (2009), Sale (2012), Bailer-Jones (2011), and Hanson & Bailer-Jones (2014).

We use a HEALPix $N_{\text{side}} = 512$ pixelization of the sky (Gorski et al. 1999) to split the sky into equal-area $7' \times 7'$ lines of sight ($14' \times 14'$ for $|b| > 30^\circ$, $N_{\text{side}} = 256$), and consider the stars along each such line of sight together. We find the reddening profile $E(D)$ that is most consistent with these stars. We parameterize $E(D)$ as an increasing, positive, piecewise-linear function in distance modulus $\mu = 5 \log(D/10\text{pc})$, with parameters α giving the reddening at each anchor point of $E(D)$. The anchors are evenly spaced in μ , with 30 anchors from $\mu = 4.22$ to $\mu = 19.68$. We then find the maximum likelihood α given the observed photometry $\{\mathbf{m}\}$ on each line of sight; that

is, we maximize $p(\alpha | \{\mathbf{m}\})$. As shown in detail by Green et al. (2014) and summarized here, this simplifies to the product over integrals through $p(E, D|\mathbf{m})$ for each star i , because

$$p(\alpha | \{\mathbf{m}\}) \propto p(\alpha) \prod_i p(m_i | \alpha), \quad (1)$$

where we have used Bayes’ rule and assumed the photometry of each star along the line of sight to be independent. Meanwhile

$$p(\mathbf{m} | \alpha) = \int dD d\theta p(\mathbf{m}, D, \theta | \alpha) \quad (2)$$

$$= \int dD d\theta p(\mathbf{m} | D, \theta, \alpha) p(D, \theta) \quad (3)$$

$$= \int dD d\theta p(\mathbf{m} | D, \theta, E(D; \alpha)) p(D, \theta), \quad (4)$$

where θ gives the intrinsic parameters describing each star and $E(D; \alpha)$ gives the reddening profile described by parameters α . The function $p(D, \theta)$ incorporates our prior knowledge about the luminosity function and the spatial distribution of stars and their metallicities. The integral $\int d\theta p(\mathbf{m} | D, \theta, E) p(D, \theta)$ is, up to a normalizing constant, the same as $p(E, D|\mathbf{m})$ when a flat prior on E is adopted. This makes $p(\alpha | \{\mathbf{m}\})$ ultimately the product of line integrals over $p(E, D|\mathbf{m})$.

The resulting $p(\alpha | \{\mathbf{m}\})$ on each line of sight describe the full 3D distribution of the dust, though in the radial direction the resolution is much worse than in the angular direction. Additionally, the radial direction is more vulnerable to systematic errors from mismatch between our model photometry and the observed photometry. In this work, we present an estimate of the cumulative reddening out to $D = 4.5$ kpc from our 3D map. This corresponds to $E(4.5 \text{ kpc}; \alpha)$ for the parameters α that maximize $p(\alpha)$ along each line of sight. The reddening to this distance is particularly well constrained by the PS1 photometry, as it corresponds approximately to the single-epoch *g*-band PS1 completeness limit through 1 mag $E(B - V)$ for main-sequence turn-off stars. We estimate the uncertainty in $E(4.5 \text{ kpc})$ by finding the range of allowed reddening to 4.5 kpc such that $\Delta \log p > -0.5$ (i.e., $\Delta \chi^2 < 1$), while holding the reddening at other distances fixed insofar as possible, subject to the constraint that the reddening must increase with distance. We defer the analysis of the full 3D distribution of the dust to later work.

At Galactic latitudes with $|b| > 30^\circ$, occasionally the fitting process becomes unreliable due to the small number of available stars in each pixel. Accordingly, for the $|b| > 30^\circ$ sky, we use a lower $N_{\text{side}} = 256$ pixelization, with $14'$ pixels. This ensures reliable fits over the entire $\delta > -30$ sky observed by Pan-STARRS1.

4. REDDENING MAP

We present in Figure 1 a map of the reddening to 4.5 kpc of the entire sky north of declination -30° , derived from Pan-STARRS1 photometry. Also shown is the Schlegel et al. (1998, SFD) dust map, which is based entirely on the thermal emission from dust in the FIR, and the difference between the two maps.

Away from the midplane of the inner Galaxy, the two maps agree well. The large discrepancy in the inner Galaxy is expected because the PS1-based dust map measures the reddening out to only 4.5 kpc, while the SFD map includes the entire reddening

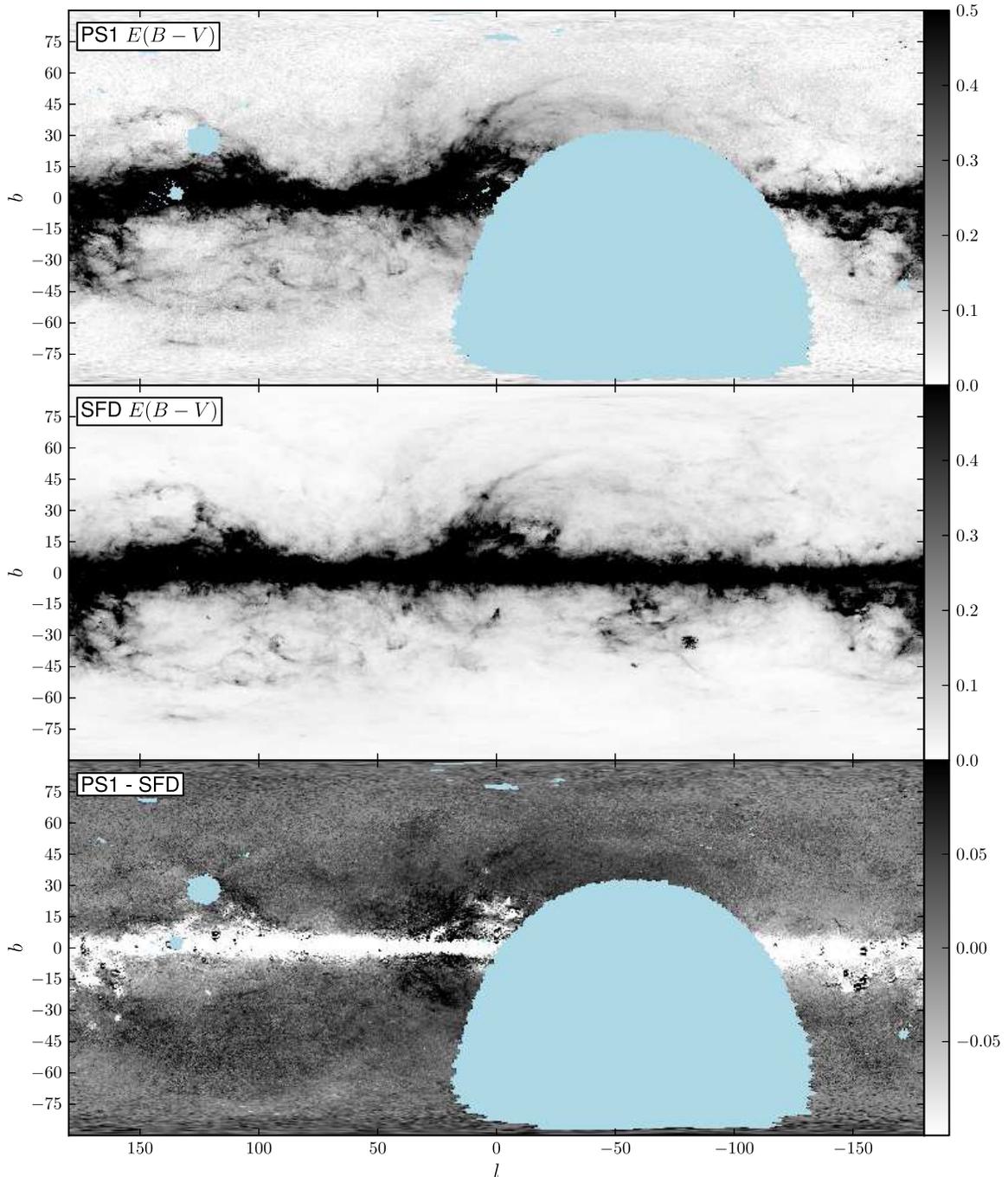


Figure 1. Maximum-likelihood PS1-based map of the dust at 4.5 kpc; the SFD dust map, which estimates the total reddening from its far-infrared emission; and the difference. Outside the Galactic plane, the maps agree closely. Light blue areas denote regions not yet well-observed by PS1; the large blank region is the declination -30° boundary. The full resolution map is available at our Web site.

(A color version of this figure is available in the online journal.)

through the Galaxy. More than five degrees off the Galactic plane, where most dust is within about 1 kpc, both maps essentially estimate the full reddening and agree. This result is expected, but remarkable: though SFD is based on satellite observations of the thermal emission of dust in the FIR, and the PS1 map measures the reddening of optical starlight passing through the dust, the two maps are extremely similar.

The map has $7'$ resolution for $|b| < 30^\circ$ and covers about three quarters of the sky. The noise in the map depends on the local stellar density, the PS1 depth in the region, and the reddening profile. In the best cases, in regions of low reddening but large

stellar density, our map and SFD agree with an rms scatter of only 25 mmag per $7'$ pixel (see Section 5.1).

Our map has a typical formal uncertainty of about ~ 20 mmag $E(B - V)$. However, the PS1 map uncertainty estimates depend significantly on how aggressively we reduce the weight of outliers in the analysis. To make a more empirical internal uncertainty estimate, we can compare the reddening map at 4.5 kpc with the reddening map at 2.8 and 7.4 kpc. When $|b| > 30^\circ$, we expect almost all the dust to be within 1 kpc, and the difference of the 2.8 and 7.4 kpc maps is then a sign of the uncertainty in the map. These two maps have a mean

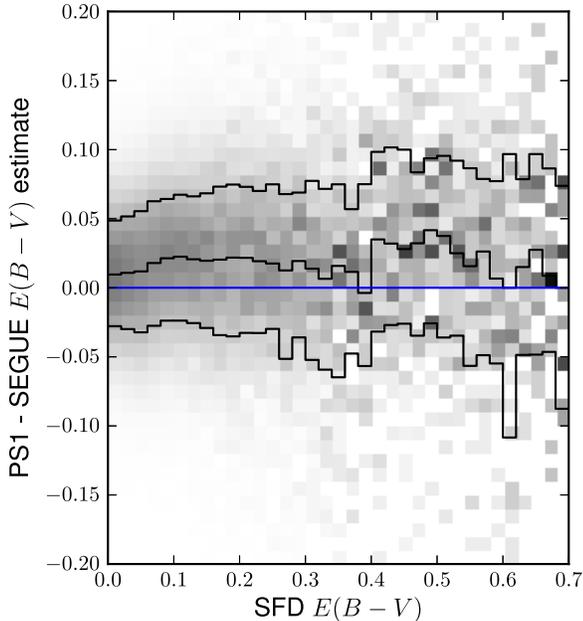


Figure 2. PS1 reddening map compared with reddening estimates toward stars observed spectroscopically by the SDSS having $|b| > 20^\circ$. The distribution of differences between the two reddening estimates are shown as a function of the SFD reddening toward each star. Apart from an ~ 20 mmag overall offset, the agreement is excellent and close to the uncertainty in the SEGUE reddening. (A color version of this figure is available in the online journal.)

difference of about 30 mmag $E(B - V)$, with a scatter similarly of 30 mmag. We therefore budget a systematic error estimate of 30 mmag. This is not far from our formal error estimates, but is more realistic, particularly in light of our comparisons with emission-based maps in Section 5.1. This estimate is based on the high-latitude sky; at low latitudes we expect the systematics in the map to be somewhat worse, as the Galactic and stellar models we use were trained at high latitudes (Green et al. 2014).

Ideally, we would like to assess our accuracy by means of a comparison with an external set of independent reddening measurements. The work of Schlafly & Finkbeiner (2011) provides color excess measurements for hundreds of thousands of stars observed as part of the SEGUE survey (Yanny et al. 2009). We show in Figure 2 the comparison between our reddening map and the reddening estimates to the SEGUE stars, using their $g-r$ colors; the result is insensitive to the choice of color. Here we use stars only with $|b| > 20^\circ$, to ensure that the stars are behind all the dust. The figure shows the distribution of differences between the SEGUE and PS1 reddening estimates, binned by their SFD-estimated reddening. Solid lines show the 16th, 50th, and 84th percentiles of the distribution. Perfect agreement is indicated by the blue horizontal line.

The agreement between our map and the SEGUE reddening is excellent. There is a small 10–20 mmag overall offset between the measurements, which is within the systematic uncertainties of our method. The scatter between the estimates is only 42 mmag. Since the SEGUE estimates themselves have a typical uncertainty of 33 mmag, the implied uncertainty in our map is only about 25 mmag, comparable to our formal uncertainty estimates. We note that this test is similar to that in Green et al. (2014, Section 7.3 of that work), except that their test compares the SEGUE data with reddening estimates for individual stars, while this test compares the SEGUE data with the map derived from hundreds of individual stars in each pixel.

This accuracy is the best of which we are aware in large-scale reddening maps. For comparison, the 2MASS-based map of Lombardi et al. (2011) near the Galactic anticenter has $3'$ FWHM and $\sigma_{E(B-V)}$ of 90 mmag. Scaling the noise in our map up by a factor of two to account for the worse resolution of our map (half the resolution in the $|b| < 30^\circ$ sky), our map still has significantly better signal-to-noise than the 2MASS map. This reddening sensitivity is a consequence of the deeper PS1 imaging and the use of optical colors. However, because our map relies on optical data, it saturates at an $E(B - V)$ of about 1.5 mag, while the near-infrared map of Lombardi et al. (2011) can reach $E(B - V) > 15$.

5. DISCUSSION

Our map directly measures the reddening from dust over three-quarters of the sky. We compare our map with the FIR-emission-based SFD map and the more recent *Planck* dust map (Planck Collaboration et al. 2013), to get a sense for the trade-offs when estimating extinction directly or indirectly through the dust emission. Emission-based maps provide much higher signal-to-noise, but systematic errors in the conversion from emission to reddening often reduce this benefit.

We note three challenges in the comparison of emission-based maps and the PS1 map. First, our analysis assumes that the reddening is always positive. This biases our reddening estimates in the low $E(B - V)$ sky high by about 20 mmag. Second, the PS1 map tracks dust reddening at 4.5 kpc. Meanwhile, the emission-based maps track the total column in each pixel. At low latitudes, there may be substantial dust beyond 4.5 kpc, present in emission-based maps but not in the PS1 map. At high latitudes, however, nearly all of the dust should be well within 4.5 kpc, allowing comparison between the maps. Third, our map tracks the typical reddening of stars observed behind the dust, while emission-based maps track the PSF-smoothed total column density. In filamentary regions that contain significant structure on scales smaller than one pixel ($7'$ – $14'$ in the PS1 map), the reddening of the typical star can be significantly different from the PSF-smoothed column in a pixel.

5.1. Comparison with SFD

Figure 3 compares the PS1 reddening map with the SFD reddening map in a few regions of interest. Qualitatively, the two maps closely agree in all cases. At high Galactic latitudes, where the dust is well within 4.5 kpc, we expect the best agreement between the maps. This good agreement is verified by the first row of panels in Figure 3. For $60^\circ < l < 80^\circ$, $-30^\circ < b < -15^\circ$, the two maps have an rms difference of only 25 mmag—a value which includes uncertainty in both maps. In the high-latitude $E(B - V) < 0.2$ sky, agreement of 20–30 mmag is typical.

The second row of panels of Figure 3 shows another high Galactic latitude region, this time highlighting a more interesting part of the sky: the Aquila South cloud (Dame et al. 2001). Here again, PS1 and SFD are qualitatively in close agreement. However, the difference map makes clear that SFD is missing a large, filamentary structure that is present in the PS1 map. The reddening in this cloud, absent in SFD, can be quite large: as much as about 0.1 mag $E(B - V)$. A likely explanation for the discrepancy is that the dust in these filaments is substantially colder than the surrounding dust. The reddening $E(B - V)$ was modeled by SFD as

$$E(B - V) = 0.0184 X_{\text{SFD}} I_{100}, \quad (5)$$

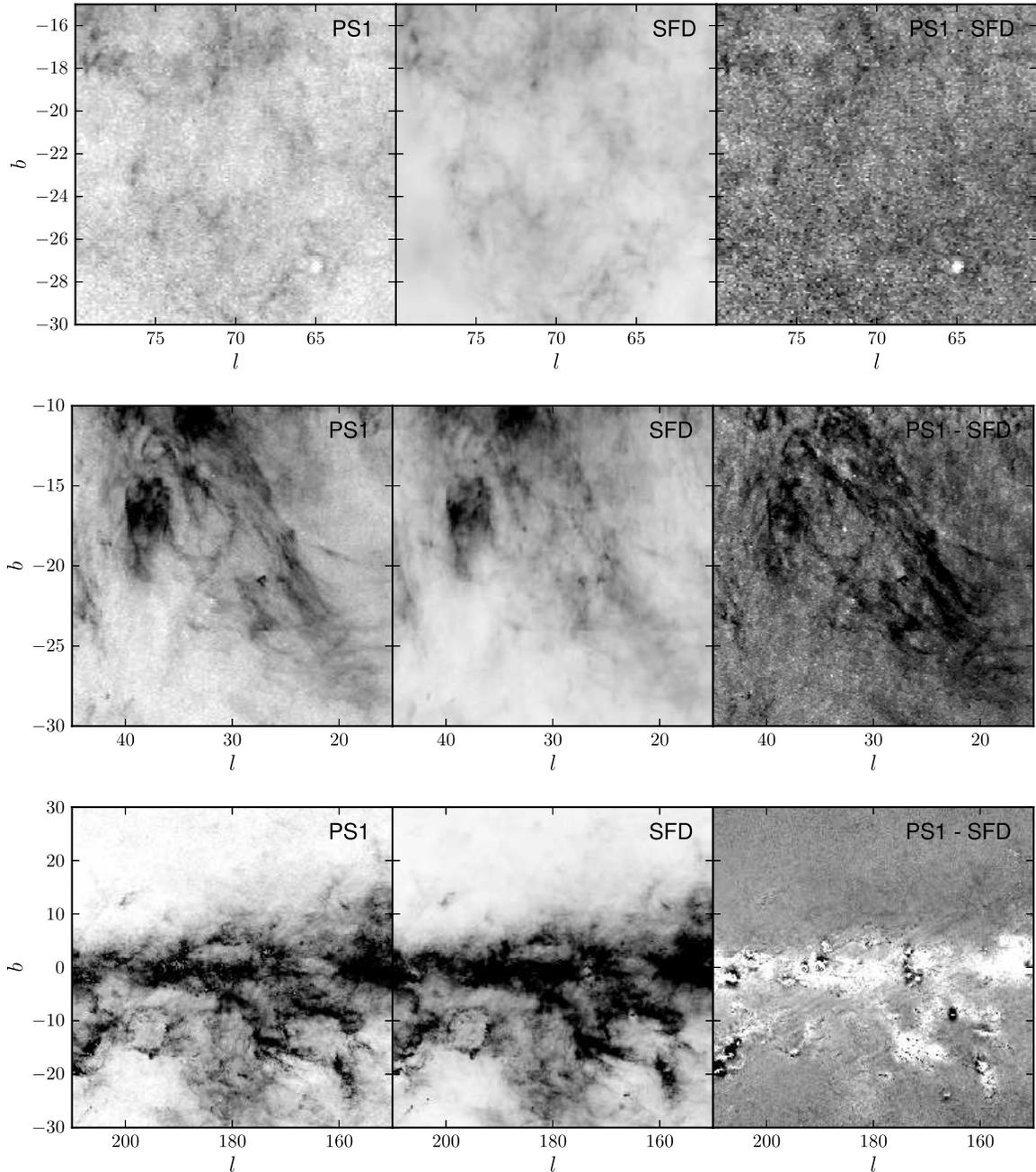


Figure 3. PS1 reddening map (left), SFD dust map (middle), and the difference (right), for three regions of interest: a relatively blank field in the south Galactic cap (top), a region around the Aquila South cloud (middle), and a region in the Galactic anticenter (bottom). The gray scale in the first two rows ranges from 0 to 0.5 mag $E(B - V)$ for the maps and -0.1 to 0.1 mag for the differences, while the last row ranges from 0 to 1 mag $E(B - V)$ for the maps and -0.25 to 0.25 mag for the differences. We have removed a 30 mmag overall offset from the PS1 maps to highlight the structure in the maps, rather than overall offsets.

where I_{100} is the zodiacal-light-subtracted composite IRAS-DIRBE $100 \mu\text{m}$ map, and SFD is a temperature-correction factor derived from the DIRBE $100 \mu\text{m}$ and $240 \mu\text{m}$ maps, normalized to be one at a temperature typical of high-latitude dust. The DIRBE maps have an angular resolution of only about 1° , so these filaments are not resolved in the SFD temperature map. Accordingly, SFD assigns the filaments temperatures typical of the surrounding warmer dust, leading to underestimated reddening. This conclusion is borne out by comparison with the *Planck* maps in the region, which agree better with the PS1 map than SFD does.

The third row of panels shows the maps in the direction of the Galactic anticenter. While the maps agree qualitatively,

large residuals are present in the difference map. Dust beyond 4.5 kpc contributes to these residuals, though other residuals are also prominent. The largest differences take the form of a bright point with a dark halo. This is a signature of bright FIR point sources lighting up a few pixels in the SFD dust map, and then artificially boosting the inferred temperature of the surrounding dust, leading to depressed reddening estimates in the vicinity. This effect is visible directly in SFD alone around $(l, b) = (165^\circ, -9^\circ)$, where the FIR-bright H II region NGC 1579 has drilled a small region out of the map.

Another salient feature of the maps of the Galactic anticenter is that the difference map strikingly resembles the SFD temperature-correction map. Figure 4 shows the difference

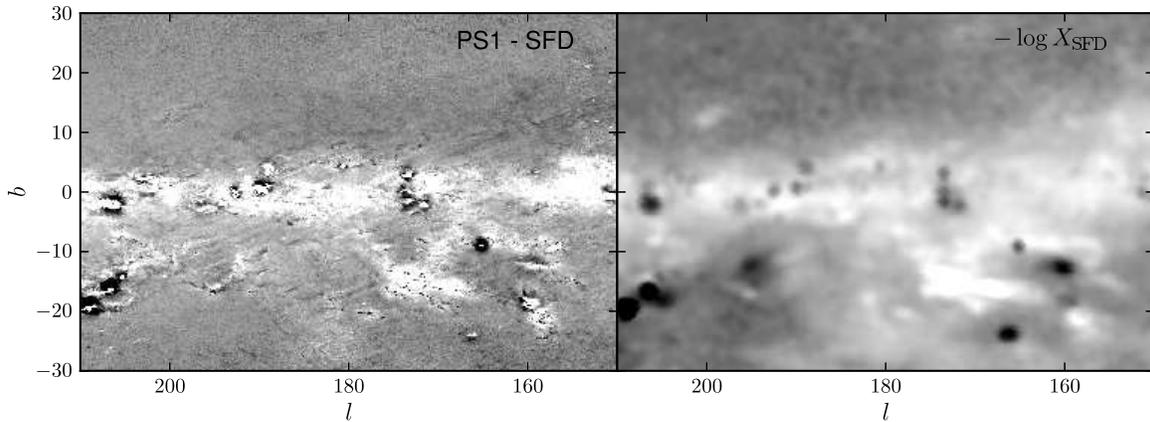


Figure 4. Difference between PS1 and SFD reddening maps (left), compared with the SFD temperature correction factor (right). The strong correlation between the maps shows that the SFD temperature correction factor is problematic.

map toward the Galactic anticenter and the SFD temperature-correction factor (Equation (5)). The strong correlation indicates that SFD could be dramatically improved, as has earlier been suggested (Schlafly et al. 2010; Peek & Graves 2010). Much of the challenge appears to stem from the low resolution of SFD. This is particularly exciting given the recent or forthcoming availability of higher-resolution *WISE*, *AKARI*, and *Planck* maps of dust emission; the combination of these data represent a tremendous advance over earlier *IRAS* and *DIRBE* data. Indeed, comparison with *Planck* (Section 5.2) indicates that substantial progress has been made.

5.2. Comparison with *Planck*

The recent Planck Collaboration et al. (2013) dust map combines FIR measurements from Schlegel et al. (1998) with *Planck* microwave data to more completely trace the spectral energy distribution of the Galaxy’s dust. The improved wavelength coverage and resolution of the *Planck* data relative to the *COBE/DIRBE* data used by Schlegel et al. (1998) for studying dust temperature should allow for dramatic improvements in dust mapping.

The Planck Collaboration et al. (2013) provide two estimates of dust column density: one based on the total dust radiance, \mathcal{R} , and another based on the optical depth of the dust at 353 GHz, τ_{353} . We have adopted the calibration of Planck Collaboration et al. (2013) to convert each of these maps into $E(B - V)$ for comparison. Throughout, when referring to the \mathcal{R} or τ_{353} maps, we are referring to these maps after calibration into units of $E(B - V)$. The dust radiance \mathcal{R} is defined as $\mathcal{R} = \int I_\nu d\nu$, the integral of the dust emission spectrum over frequency. This quantity is not obviously appropriate as a reddening template, given its dependence on the interstellar radiation field (ISRF) and the corresponding temperature of the dust. The primary motivations for adopting \mathcal{R} as a probe of the dust column instead of τ_{353} are that it is less sensitive to degeneracies in the fit parameters used in the *Planck* dust fits, and that it is less affected by the cosmic IR background anisotropy. Due to the ISRF dependence of the \mathcal{R} map, Planck Collaboration et al. (2013) recommends using τ_{353} instead of \mathcal{R} to trace dust column in translucent and dense clouds where the ISRF is significantly modified.

The accuracy of the Planck Collaboration et al. (2013) dust maps is largely tested by comparison to N_H , as N_H is the only all-sky map which is a good proxy for dust column. Calibration of the *Planck* map to reddening $E(B - V)$ was

made using a sample of quasars probing the very low extinction sky ($E(B - V) \lesssim 0.1$). Our current reddening map covers three-quarters of the sky and $E(B - V)$ up to about 1.5 mag before saturating, allowing a much broader test.

5.3. The Overall $E(B - V)$ Scale

A basic test is verifying that the various maps share the same $E(B - V)$ scaling. The PS1 map measures reddening in the PS1 bands, rather than $E(B - V)$ directly. It is tied to $E(B - V)_{\text{SFD}}$ according to the reddening vector of Schlafly & Finkbeiner (2011). Provided that Schlafly & Finkbeiner (2011) is correct, this places our map and SFD on the same overall scale. However, Schlafly & Finkbeiner (2011) finds that the $E(B - V)$ scale of SFD is off by 14%. Our map inherits that offset, though see below for further discussion.

We show the distribution of differences between our map, SFD, and the *Planck* τ_{353} and \mathcal{R} maps as a function of SFD in Figure 5. As in Figure 2, the solid black lines indicate the 16th, 50th, and 84th percentiles of the distribution in each bin, and the solid blue line indicates perfect agreement. We include here only the $b > 30^\circ$ sky to ensure that essentially all of the dust is within 4.5 kpc, and to focus on the region of sky where the agreement between the PS1 map and the emission-based maps is best. Figure 5 shows that the PS1 map and the emission maps systematically disagree as a function of $E(B - V)$.

There is a clear residual trend in $\text{SFD} - \text{PS1}$ and $\tau_{353} - \text{PS1}$ with the sense that, at reddenings less than 0.15 mag, PS1 finds more reddening per unit SFD than the emission-based maps, while at larger $E(B - V)$, PS1 finds less reddening per unit SFD. That is, the slope of the residuals in Figure 5 is negative for $E(B - V) < 0.1$, and positive for $E(B - V) > 0.2$. This trend is not expected; a priori we would have expected PS1 to be biased high at low reddening, leading to the opposite trend at low $E(B - V)$. That effect is however very small, and causes only the diagonal envelope of points with $E(B - V)_{\text{SFD}} < 0.03$ and positive residuals in each panel of Figure 5. The *Planck* \mathcal{R} map meanwhile has a different residual trend, with a constant negative slope with no change in behavior around $E(B - V) = 0.15$ mag.

The fact that PS1 finds consistently less reddening than SFD at $E(B - V) > 0.3$ is a surprise. The PS1 map was intended to inherit the overall reddening scale of Schlegel et al. (1998), through its adoption of the reddening vector of Schlafly & Finkbeiner (2011). Nevertheless, the slope $\Delta\text{SFD}/\Delta\text{PS1}$ and $\Delta\tau_{353}/\Delta\text{PS1}$ is about 1.15 for $E(B - V) > 0.3$. This result is

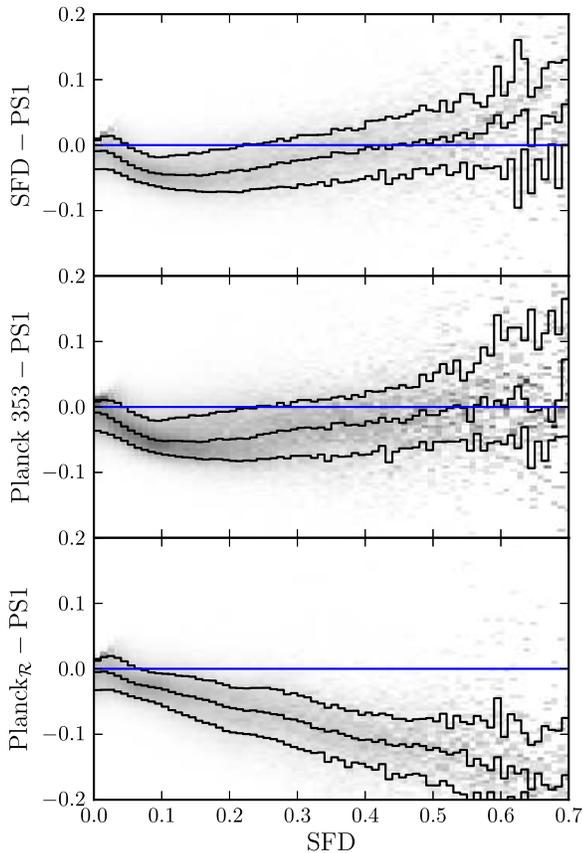


Figure 5. Distribution of differences between SFD, *Planck* τ_{353} , and *Planck* \mathcal{R} as a function of SFD, for the high-latitude sky ($|b| > 30^\circ$). Systematic differences between the PS1 map and emission-based maps exist for all maps.

(A color version of this figure is available in the online journal.)

unexpected given that Schlafly & Finkbeiner (2011) find good agreement in scale with SFD, and we find good agreement in scale with Schlafly & Finkbeiner (2011) (Figure 2). The disagreement stems from the different regions of sky used when comparing with SEGUE in Figure 2 and between reddening maps in Figure 5. The SEGUE comparison is limited to stars observed by SEGUE, within the SDSS footprint, while the reddening map comparison covers essentially the entire $|b| > 30^\circ$, $\delta > -30^\circ$ sky. Restricted to the locations on which SEGUE stars were observed, the agreement between SEGUE and SFD is very similar to the agreement between PS1 and SFD. This is demonstrated in Figure 6, which shows the distribution of residuals between SFD, SEGUE, and PS1 over different regions of sky. The solid black lines give the 16th, 50th, and 84th percentiles of the distribution, while the solid blue line indicates perfect agreement. The top two panels of Figure 6 show the differences between SFD and the SEGUE reddening estimates, followed by the differences between SFD and the PS1 reddening estimates, both over the SDSS footprint with $|b| > 20^\circ$. The second panel shows the good agreement in overall scale between SFD and PS1 in this region. However, the third panel shows that in the $|b| > 20^\circ$ sky outside the SDSS footprint, there is a clear residual slope between SFD and PS1 at $E(B - V) > 0.3$. We conclude only that while the calibration of Schlafly & Finkbeiner (2011) is appropriate over the SDSS footprint (albeit with 10% local variations), extrapolation outside of the SDSS footprint is risky as cloud properties can vary substantially over the sky.

The *Planck* τ_{353} map and SFD have very similar residual behavior as a function of $E(B - V)$. This is expected as both

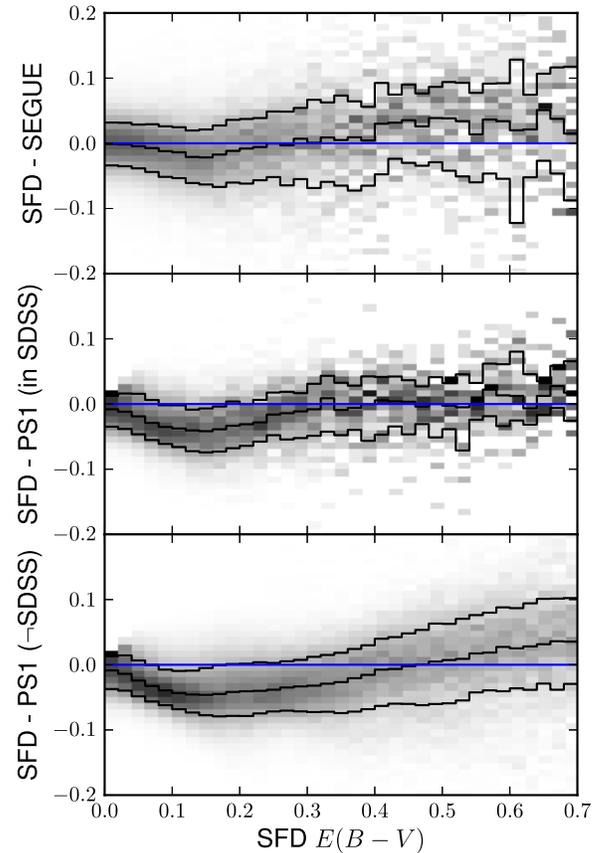


Figure 6. Comparison between the SEGUE and SFD reddening measurements on the $|b| > 20^\circ$ sky sampled by SEGUE (top panel); between PS1 and SFD over the same sky footprint (middle panel); and between SFD and PS1 over the $|b| > 20^\circ$ sky outside the SEGUE footprint (bottom panel). While there is good agreement in overall scale between SFD and PS1 within the SDSS footprint, outside this region, at $E(B - V)_{\text{SFD}} > 0.3$, the PS1 and SFD reddening estimates have systematically different slopes.

(A color version of this figure is available in the online journal.)

maps are essentially thermal fits to $100\ \mu\text{m}$ dust emission and longer wavelength emission. The *Planck* map, however, includes a correction for the variation in the emissivity of the dust, omitted by SFD. This correction seems to have had only a minor effect on the overall scaling of the *Planck* τ_{353} map when $E(B - V) < 1$. The cause of the change in residual slope between the PS1 map and both the SFD and τ_{353} maps around $E(B - V) = 0.15$ is not understood.

The *Planck* \mathcal{R} map residuals behave differently. The map predicts less reddening than the PS1 map everywhere, by a factor of about 1.3. Half of this discrepancy may be due to our map's intended tie to SFD rather than $E(B - V)$. This leaves a discrepancy of about 15% which we are unable to explain. However, it is ultimately a disagreement between the reddening calibration of Planck Collaboration et al. (2013) and the reddening measurements of Schlafly & Finkbeiner (2011), rather than a problem in the technique we have adopted to map dust using PS1. The \mathcal{R} residuals show no change in slope at $E(B - V) = 0.15$, unlike the SFD and τ_{353} residuals. This suggests that the FIR modeling of the dust may be causing the change in slope in the SFD and τ_{353} residuals.

5.4. Map Comparisons

Figure 7 shows our PS1 map, followed by the difference between our PS1 map and the two *Planck* reddening maps.

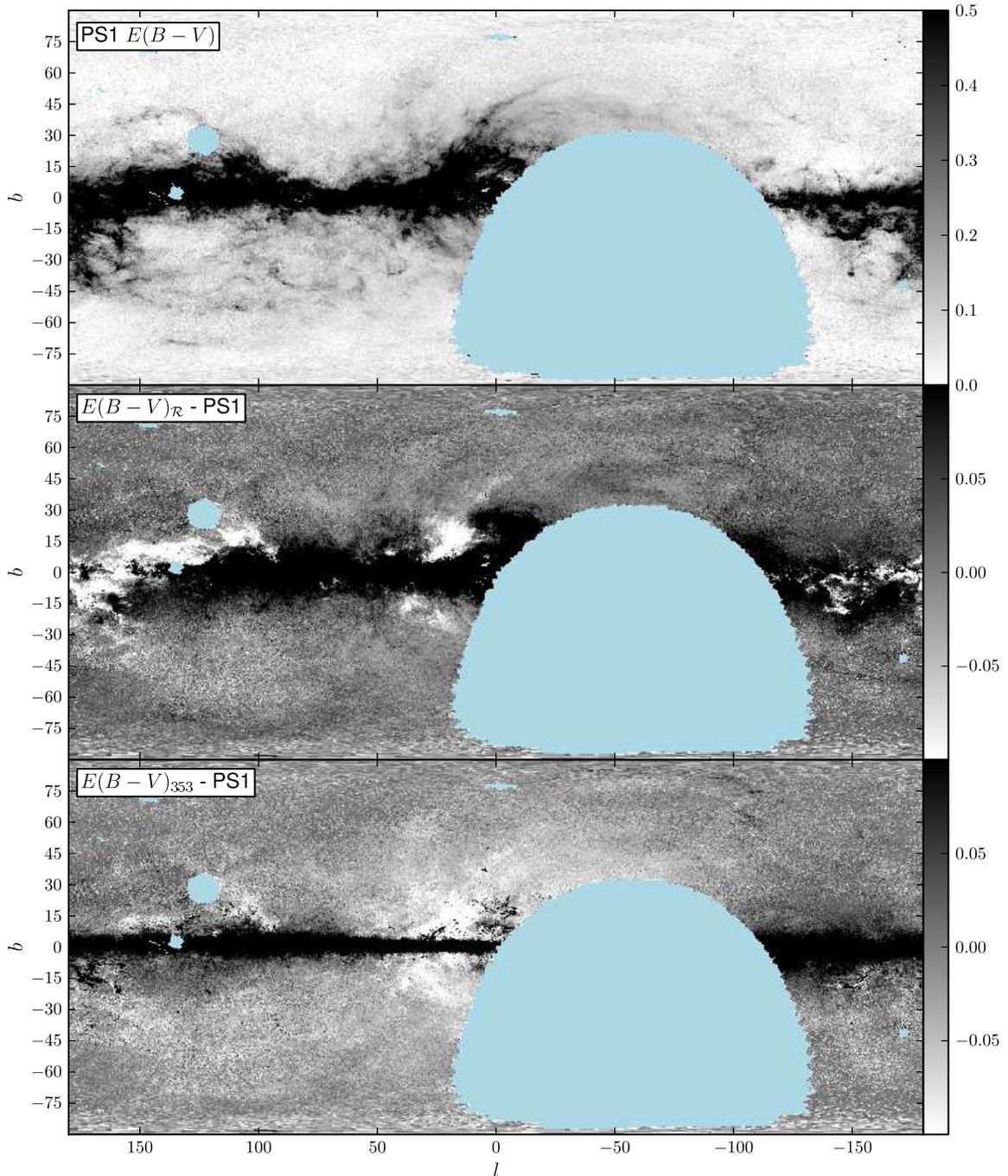


Figure 7. Comparison between PS1 and *Planck* reddening maps. The top panel shows our PS1 reddening map, and the following two panels show the difference between it and the *Planck* dust-brightness-based and dust-column-based $E(B - V)$ maps. At high latitudes all maps are qualitatively in good agreement, though in detail we find that the *Planck* brightness-based map is somewhat more consistent. Below latitudes of about 20° – 30° , however, the column-based map is much more accurate.

(A color version of this figure is available in the online journal.)

Consistent with Section 5.3, we have rescaled the *Planck* τ_{353} map by 30% to provide a better match to the PS1 map. The high-latitude sky is in close agreement and subtracts well in the difference maps. However, the lower latitude and higher $E(B - V)$ sky shows large differences between the maps. Both *Planck* maps substantially underpredict the amount of extinction outside the plane east of the Galactic center, as well as in the Cepheus Flare and at high latitudes toward the anticenter. Meanwhile at low latitudes the *Planck* $E(B - V)$ are generally larger than the PS1 $E(B - V)$, which is expected as the PS1

maps saturate at about 1.5 mag and there is often significant dust beyond 4.5 kpc in this area.

In general at low latitudes ($|b| \lesssim 30^\circ$), but outside the plane, the PS1 map agrees more closely with the *Planck* τ_{353} map than with \mathcal{R} . This was anticipated by Planck Collaboration et al. (2013) given the varying ISRF in these environments, but our map makes clear that this effect is generic at moderately low latitudes. Outside of $|b| < 7^\circ$ our map and *Planck* τ_{353} agree within about 0.05 mag outside the Galactic center, while the region $7^\circ < |b| < 15^\circ$ is unreliable in \mathcal{R} . The presence of

regions of both overestimated and underestimated reddening shows that even in the high-latitude sky no simple rescaling of the Planck maps can bring them into agreement with the reddening estimates.

6. CONCLUSION

We present a reddening map based on photometry from Pan-STARRS1 covering almost three-quarters of the sky. The map has a unique combination of high angular resolution ($7''$ – $14''$ pixels), low noise (~ 25 mmag $E(B - V)$), and wide sky coverage that makes it ideal for studying the Galactic dust over a wide range of environments.

Comparison with the widely used SFD dust map highlights some known shortcomings of that map. The combination of low-resolution DIRBE data with high-resolution *IRAS* data leads to peculiar features in filamentary dust clouds, and can lead to SFD either overestimating or underestimating the true reddening. Comparison with the Planck Collaboration et al. (2013) maps shows improvement, but the maps have similar problems despite the wider frequency coverage and better resolution. The *Planck* \mathcal{R} map traces the dust well for $|b| > 15^\circ$, but is unreliable outside this region due to variation in the ISRF.

These techniques can be used in any multiband optical survey. Application to the SkyMapper survey would allow the dust in the remaining quarter of the sky to be mapped. The Dark Energy Survey will also map much of the remaining sky, and could efficiently map the $\delta < -30^\circ$ Galactic plane ($|b| < 5^\circ$). As foreseen by Bailer-Jones (2011), even the photometric component of *Gaia* alone is extremely interesting from the perspective of this technique—but the addition of parallax information will make that mission truly revolutionary for maps of dust.

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