The KMOS AGN Survey at High redshift (KASHz): the prevalence and drivers of ionized outflows in the host galaxies of X-ray AGN[★]

C. M. Harrison,[1][†] D. M. Alexander,[1][†] J. R. Mullaney,[2] J. P. Stott,[3,1]
A. M. Swinbank,[4,1] V. Arumugam,[5] F. E. Bauer,[6,7,8] R. G. Bower,[4,1]
A. J. Bunker[3,9] and R. M. Sharples[10,1]

[1]*Centre for Extragalactic Astronomy, Department of Physics, Durham University, South Road, Durham DH1 3LE, UK*
[2]*Department of Physics and Astronomy, University of Sheffield, Sheffield S3 7RH, UK*
[3]*Astrophysics, Department of Physics, University of Oxford, Keble Road, Oxford OX1 3RH, UK*
[4]*Institute for Computational Cosmology, Department of Physics, Durham University, South Road, Durham DH1 3LE, UK*
[5]*European Southern Observatory, Karl-Schwarzschild-Strasse 2, D-85748 Garching, Germany*
[6]*Instituto de Astrofísica, Facultad de Física, Pontifica Universidad Católica de Chile, 306, Santiago 22, Chile*
[7]*Millennium Institute of Astrophysics, Vicuña Mackenna 4860, 7820436 Macul, Santiago, Chile*
[8]*Space Science Institute, 4750 Walnut Street, Suite 205, Boulder, CO 80301, USA*
[9]*Affiliate Member, Kavli Institute for the Physics and Mathematics of the Universe (WPI), Todai Institutes for Advanced Study, The University of Tokyo, 5-1-5 Kashiwanoha, Kashiwa 277-8583, Japan*
[10]*Centre for Advanced Instrumentation, Department of Physics, Durham University, South Road, Durham DH1 3LE, UK*

ABSTRACT

We present the first results from the KMOS (*K*-band Multi-Object Spectrograph) AGN (active galactic nuclei) Survey at High redshift (KASH*z*), a VLT/KMOS integral-field spectroscopic (IFS) survey of $z \gtrsim 0.6$ AGN. We present galaxy-integrated spectra of 89 X-ray AGN ($L_{2-10\,\mathrm{keV}} = 10^{42}$–$10^{45}$ erg s^{-1}), for which we observed [O III] ($z \approx 1.1$–1.7) or Hα emission ($z \approx 0.6$–1.1). The targets have X-ray luminosities representative of the parent AGN population and we explore the emission-line luminosities as a function of X-ray luminosity. For the [O III] targets, $\approx$50 per cent have ionized gas velocities indicative of gas that is dominated by outflows and/or highly turbulent material (i.e. overall line widths $\gtrsim$600 km s^{-1}). The most luminous half (i.e. $L_X > 6 \times 10^{43}$ erg s^{-1}) have a $\gtrsim$2 times higher incidence of such velocities. On the basis of our results, we find no evidence that X-ray obscured AGN are more likely to host extreme kinematics than unobscured AGN. Our KASH*z* sample has a distribution of gas velocities that is consistent with a luminosity-matched sample of $z < 0.4$ AGN. This implies little evolution in the prevalence of ionized outflows, for a fixed AGN luminosity, despite an order-of-magnitude decrease in average star formation rates over this redshift range. Furthermore, we compare our Hα targets to a redshift-matched sample of star-forming galaxies and despite a similar distribution of Hα luminosities and likely star formation rates, we find extreme ionized gas velocities are up to $\approx$10 times more prevalent in the AGN-host galaxies. Our results reveal a high prevalence of extreme ionized gas velocities in high-luminosity X-ray AGN and imply that the most powerful ionized outflows in high-redshift galaxies are driven by AGN activity.

Key words: galaxies: active – galaxies: evolution – galaxies: kinematics and dynamics – quasars: emission lines.

1 INTRODUCTION

Massive galaxies are now known to host supermassive black holes (SMBHs) at their centres. These SMBHs grow through mass accretion events, during which they become visible as active galactic nuclei (AGN). A variety of indirect observational evidence has been used to imply a connection between the growth of SMBHs and the growth of the galaxies that they reside in. For example: (1)

the cosmic evolution of volume-averaged SMBH growth and star formation looks very similar; (2) growing SMBHs may be preferentially located in star-forming galaxies and (3) SMBH mass is tightly correlated with galaxy bulge mass and stellar velocity dispersion (see reviews in e.g. Alexander & Hickox 2012; Kormendy & Ho 2013). Suggestion of a more *direct* connection between SMBH growth and galaxy growth largely comes from theoretical models of galaxy formation. Most successful models propose that AGN are required to regulate the growth of massive galaxies by injecting a fraction of their accretion energy into the surrounding intergalactic medium (IGM) or interstellar medium (ISM; e.g. Silk & Rees 1998; Benson et al. 2003; Granato et al. 2004; Churazov et al. 2005; Bower et al. 2006; Hopkins et al. 2006; Bower, McCarthy & Benson 2008; Somerville et al. 2008; McCarthy et al. 2010; Gaspari et al. 2011; Vogelsberger et al. 2014; Schaye et al. 2015). Without this so-called 'AGN feedback', these models fail to reproduce many key observables of the local Universe, such as the observed SMBH mass–spheroid mass relationship (e.g. Kormendy & Ho 2013); the sharp cut-off in the galaxy mass function (e.g. Baldry et al. 2012) and the X-ray temperature–luminosity relationship observed in galaxy clusters and groups (e.g. Markevitch 1998).

In recent years there has been a large amount of observational work searching for signatures of 'AGN feedback' and to test theoretical predictions (see reviews in e.g. Alexander & Hickox 2012; Fabian 2012; McNamara & Nulsen 2012; Heckman & Best 2014). One of the most promising candidates for a universal feedback mechanism is AGN-driven outflows, which, if they can be driven to galaxy-wide scales, could remove or heat cold gas that would otherwise form stars in the host galaxy. There is now wide-spread observational evidence that galaxy-wide outflows exist in both low- and high-redshift AGN-host galaxies, using tracers of atomic, molecular and ionized gas (e.g. Martin 2005; Rupke, Veilleux & Sanders 2005; Harrison et al. 2012, 2014; Liu et al. 2013; Veilleux et al. 2013; Arribas et al. 2014; Cicone et al. 2014).

Of specific relevance to our study are ionized outflows that have been known for several decades to be identifiable using broad and asymmetric emission-line profiles (e.g. Weedman 1970; Stockton 1976; Heckman et al. 1981). However, it is now possible to use large optical spectroscopic samples, such as the Sloan Digital Sky Survey (SDSS; York et al. 2000), to search for these signatures in hundreds to thousands of $z \lesssim 1$ AGN. By combining these spectroscopic surveys with multi-wavelength data sets, recent studies have provided excellent constraints on the prevalence and drivers of ionized outflows in low-redshift AGN (e.g. Mullaney et al. 2013; Zakamska & Greene 2014; Balmaverde et al. 2015). Follow-up integral-field spectroscopy (IFS) observations of objects drawn from these large samples have made it possible to constrain the prevalence of these outflows on *galaxy-wide* scales to explore their spatially resolved characteristics as a function of AGN and host-galaxy properties and to test theoretical predictions (e.g. Liu et al. 2013; Harrison et al. 2014, 2015; McElroy et al. 2015). These studies have revealed that galaxy-wide ionized outflows are common, perhaps ubiquitous, throughout the most optically luminous low-redshift AGN (i.e. with $L_{AGN} \gtrsim 10^{45}$ erg s^{-1}; although also see Husemann et al. 2013).

Despite the great insight provided by statistical IFS studies of outflows in the low-redshift Universe, these studies do not cover the redshift ranges during the peak epochs of SMBH and galaxy growth (i.e. $z \gtrsim 1$; e.g. Aird et al. 2010; Madau & Dickinson 2014) and consequently the redshift ranges where AGN-driven outflows are predicted to be most prevalent. Spatially resolved spectroscopy of high-redshift AGN has been more limited because the bright optical emission lines, which are excellent traces of ionized gas

kinematics, are redshifted to the near-infrared (NIR), which is much more challenging to observe compared to optical wavelengths. Each study that has searched for galaxy-wide ionized outflows at high redshifts has investigated only a small number of AGN, which were selected in a variety of different ways (e.g. Nesvadba et al. 2006, 2008; Alexander et al. 2010; Cano-Díaz et al. 2012; Harrison et al. 2012; Förster Schreiber et al. 2014; Brusa et al. 2015; Carniani et al. 2015; Cresci et al. 2015; Collet et al. 2015), with the largest, to date, coming from Genzel et al. (2014), which presents 18 confirmed AGN identified using a combination of X-ray, infrared and radio techniques. To make robust conclusions about the AGN population as a whole, and to properly understand the role of ionized outflows in galaxy evolution, it is crucial to place these observations into the context of the parent population of AGN and galaxies. For example, it is particularly important to assess how representative observations are if they have significant implications for our understanding of AGN feedback, such as possible evidence that star formation has been suppressed by ionized outflows in two high-redshift AGN (Cano-Díaz et al. 2012; Cresci et al. 2015). There is clearly a need for IFS observations of large samples of high-redshift AGN that are selected in a uniform way.

It is now possible to efficiently obtain large samples of NIR IFS data thanks to the commissioning of the *K*-band Multi Object Spectrograph (KMOS; Sharples et al. 2004, 2013) on the European Southern Observatory's Very Large Telescope (VLT). This instrument is ideal for systematic studies of the rest-frame optical properties of high-redshift galaxies and AGN, selected in the well-studied extragalactic deep fields (e.g. Sobral et al. 2013; Stott et al. 2014; Stott et al., in preparation; Wisnioski et al. 2015). In this work we present the initial results from our survey of high-redshift AGN: the KMOS AGN Survey at High-*z* (KASH*z*). KASH*z* is an ongoing guaranteed time project, led by Durham University, to observe high-redshift ($z \approx 0.6$–3.6) AGN with KMOS. This survey will provide a huge leap forward for our understanding of AGN outflows and host-galaxy kinematics by measuring the spatially resolved kinematics of an order of magnitude more sources than previous work. Furthermore, this survey has been jointly run with the KMOS Redshift One Spectroscopic Survey (KROSS) of high-redshift star-forming galaxies (Stott et al., in preparation), which makes it possible to place our observations of AGN into the context of the galaxy population.

In this paper we present the first results from KASH*z*. In Section 2 we describe the survey and observations; in Section 3 we describe our data analysis and comparison samples; in Section 4 we present our initial results and discuss their implications; and in Section 5 we give our conclusions. Throughout, we assume a Chabrier IMF (Chabrier 2003) and assume $H_0 = 70$ km s^{-1} Mpc^{-1}, $\Omega_M = 0.30$ and $\Omega_\Lambda = 0.70$; in this cosmology, 1 arcsec corresponds to 7.5 kpc at $z = 0.8$ and 8.4 kpc at $z = 1.4$.

2 SURVEY DESCRIPTION, SAMPLE SELECTION AND OBSERVATIONS

KASH*z* is designed to ultimately obtain spatially resolved emission-line kinematics of $\approx$(100–200) high-redshift ($z \approx 0.6$–3.6) AGN. The overall aim of KASH*z* is to provide insight into the feeding and feedback processes occurring in the host galaxies of high-redshift AGN by using IFS data to measure the ionized gas kinematics traced by the Hα, [O III], Hβ, [N II] and/or [S II] emission lines. The key aspect of KASH*z* is to exploit the unique capabilities of the multiple integral field units (IFUs) in the KMOS instrument to perform such measurements on larger, more uniformly selected

samples of high-redshift AGN than was possible in previous studies that used single IFU instruments. This will make it possible to make conclusions on the overall high-redshift AGN population and to place previous observations, of a few sources, into the context of the parent population of AGN. Until now, this sort of approach has only been possible for *low-redshift* AGN (e.g. Arribas et al. 2014; Harrison et al. 2014; McElroy et al. 2015).

In this paper we present the first results of KASHz, which focuses on the galaxy-integrated emission-line profiles of $z \approx 0.6$–1.7 X-ray detected AGN. In future papers we will present results on the spatially resolved properties of the current sample, as well as expand the sample size when more data are obtained during our ongoing guaranteed-time KMOS observations.

2.1 Parent catalogues

For our target selection we make use of deep X-ray surveys performed in extragalactic fields (COSMOS; CDF-S; UDS and SSA22). These surveys provide an efficient method for AGN selection that is largely free from host-galaxy contamination (e.g. see review in Brandt & Alexander 2015). The chosen X-ray fields are those visible from the VLT with the deepest X-ray data available. These allow us to uniformly select large samples of AGN that can be efficiently observed with KMOS (see Section 2.6.1). The four deep fields are as follows.

(i) Cosmic evolution survey field (COSMOS; see Scoville et al. 2007). We obtain X-ray sources by combining the 1.8 Ms *Chandra* catalogue that covers 0.9 deg² (C-COSMOS; Elvis et al. 2009; Puccetti et al. 2009; Civano et al. 2012) with the wider and shallower 1.5 Ms *XMM–Newton* catalogue that covers 2.13 deg² (*XMM*-COSMOS; Cappelluti et al. 2009; Brusa et al. 2010). For sources detected in both surveys we use the entries in the C-COSMOS catalogue.

(ii) *Chandra* Deep Field South (CDF-S; see Giacconi et al. 2001). We use the X-ray sources from the 4 Ms *Chandra* catalogue that covers 464.5 arcmin² (Xue et al. 2011).

(iii) The *Subaru/XMM–Newton* Deep Survey (SXDS: see Furusawa et al. 2008). We use the X-ray sources from the 400 ks *XMM–Newton* catalogue that covers 1.14 deg² (Ueda et al. 2008). We only observe sources that are positioned inside the central 0.8 deg² that are also covered by the near-IR Ultra Deep Survey (UDS; see Lawrence et al. 2007). We use UDS as the name of this field hereafter.

(iv) The SSA22 $z = 3.09$ protocluster field (see Steidel et al. 1998). We use the X-ray sources from the 400 ks *Chandra* catalogue that covers 330 arcmin² (Lehmer et al. 2009).

To define our sample we primarily make use of the hard-band fluxes ($F_{2–10\,keV}$) for each of the X-ray sources. We also make use of the soft-band fluxes ($F_{0.5–2\,keV}$) when there is no hard-band detection and for calculating the $F_{2–10\,keV}/F_{0.5–2\,keV}$ flux ratios to select X-ray obscured candidates (see Section 2.4). We note that throughout this work, L_X refers to the hard-band (2–10 keV) luminosities. To be consistent across the four fields, we calculate fluxes making similar assumptions to the C-COSMOS catalogue (see Puccetti et al. 2009), i.e. we convert quoted count rates (CR) to fluxes, assuming a photon index of $\Gamma = 1.4$ and a typical galactic absorption of $N_H \approx 2 \times 10^{20}$ cm⁻². For the SSA22 and CDF-S *Chandra* catalogues, we use *CR* to flux conversion factors of 2.87×10^{-11} erg cm⁻² and 8.67×10^{-12} erg cm⁻² for the hard and soft bands, respectively. We note that these conversion factors also take into account the conversion between the quoted 2–8 keV energy-band values to
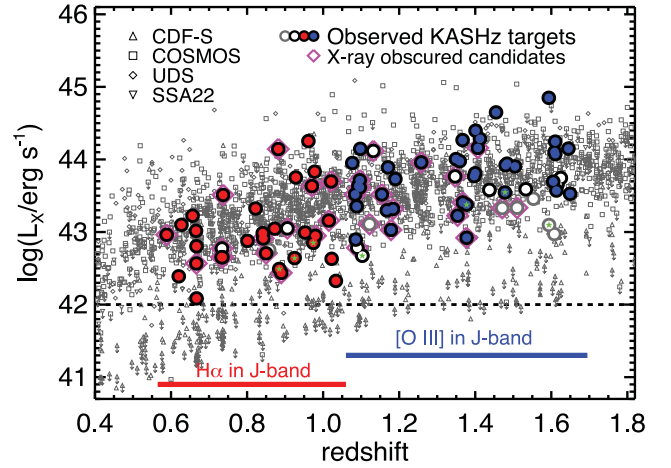
Figure 1. Hard-band (2–10 keV) X-ray luminosity versus redshift for the sources in the four deep fields covered by KASHz (see Section 2.1). The dashed line indicates our luminosity cut for target selection of $L_{2–10\,keV} > 10^{42}$ erg s⁻¹. The large symbols indicate the 89 X-ray AGN that have been observed for KASHz so far. These filled and empty symbols indicate emission-line detections (Hα or [O III]) and no emission-line detections, respectively. Grey circles represent the seven targets that are excluded from the analyses (i.e. resulting in a final sample of 82 targets) because a lack of an emission-line detection may be unphysical (Section 4.1). The stars highlight targets that have $L_{2–10\,keV}$ values estimated from the soft band (see Section 2.1). Overall the targets cover nearly three orders of magnitude in X-ray luminosity (i.e. $L_X \approx 10^{42}$–10^{45} erg s⁻¹).

2–10 keV energy-band values. For the UDS *XMM–Newton* catalogue, we use the conversion factors applicable for $\Gamma = 1.4$ tabulated by Ueda et al. (2008) and for C-COSMOS and *XMM*-COSMOS we use the quoted flux values in Puccetti et al. (2009) and Brusa et al. (2010), respectively.[1] The X-ray fluxes are tabulated in Table A1.

We obtained archival redshifts for the sources in the fields as follows: (1) for C-COSMOS we use the compilation in Civano et al. (2012), using spectroscopic redshifts when available and photometric redshifts otherwise; (2) for CDF-S we only target sources with spectroscopic redshifts as compiled by Xue et al. (2011); (3) for UDS we used the 2010 October spectroscopic redshift compilation provided by the UDS consortium[2] (Smail et al. 2008; Simpson et al. 2012; Akiyama et al., in preparation) and (4) for SSA22 we use the spectroscopic redshifts compiled by Lehmer et al. (2009).[3] For our observed targets the redshifts used throughout this work, z, are those derived from our measured emission lines (i.e. z_L; Section 3.2), in preference to the archival redshifts (i.e. z_A) described above. The archival redshifts and our emission-line redshifts for the targets presented here are tabulated in Table A1.

The sources in our parent catalogues are plotted in the L_X–redshift plane in Fig. 1. We calculate X-ray luminosities following

$$L = 4\pi D_L^2 F (1+z)^{\Gamma-2}, \qquad (1)$$

where D_L is the luminosity distance, F is the X-ray flux in the relevant energy band, z is the redshift (see above) and Γ is the

---

[1] We note that *XMM*-COSMOS catalogue assumes a power-law index of $\Gamma = 1.7$ and $\Gamma = 2.0$ for the hard and soft bands, respectively. However, this results in small differences in the calculated luminosity values from assuming $\Gamma = 1.4$ (i.e. $\lesssim 20$ per cent).
[2] http://www.nottingham.ac.uk/astronomy/UDS/data/data.html
[3] For the target SSA22-39 we used the spectroscopic redshift from Saez et al. (2015).

photon index. As above, we assume $\Gamma = 1.4$ for all sources. Some of the X-ray sources are detected in the soft band (0.5–2.0 keV) but not in the hard band (2–10 keV) and for these sources we make use the catalogued flux upper limits to plot them in Fig. 1. However, for the targets that we observed in this work, that do not have a hard-band detection, i.e. seven targets that are all from the COSMOS field (see Table A1), we estimated hard-band fluxes by extrapolating from the soft band assuming a power law with $\Gamma = 1.4$ (see star symbols in Fig. 1). In all of these cases, the extrapolated hard-band fluxes were consistent with the measured upper limits and we verified that the X-ray emission was AGN-dominated (see below).

2.2 Sample selection

Our KASH*z* targets were selected using the archival redshifts and X-ray luminosities described above. For this paper, we observed targets in the NIR *J* band, which covers a wavelength range of $\lambda \approx 1.03$–1.34 µm. Therefore, we selected targets with redshifts in these two ranges: (1) $z = 1.07$–1.67, for which we could observe the [O III]4959,5007 emission-line doublet, or (2) $z = 0.57$–1.05, for which we could observe the Hα and [N II]6548,6583 emission lines.

To avoid selecting non-AGN X-ray sources (i.e. extreme starburst galaxies), we select sources with a measured hard-band luminosity of $L_{2-10\,\mathrm{keV}} > 10^{42}$ erg s^{-1}. We note that this X-ray luminosity has little impact on the final selection because the X-ray catalogues are not complete down to this luminosity for our redshift ranges of interest (see Fig. 1). Therefore, the majority of the observed targets (i.e. all but four) have X-ray luminosities of $L_{2-10\,\mathrm{keV}} > 3 \times 10^{42}$ erg s^{-1} and will have X-ray emission that is AGN dominated; for X-ray emission of this luminosity to be produced by star formation processes alone, it would require extreme star formation rates of $\gtrsim 1900\,\mathrm{M_\odot\,yr^{-1}}$, based on empirical measurements of star-forming galaxies (following equation 4 of Symeonidis et al. 2011 and Kennicutt 1998 converted to a Chabrier IMF; also see Lehmer et al. 2010). Four of our targets, all of which are in CDF-S, have $L_{2-10\,\mathrm{keV}} = (1$–3$) \times 10^{42}$ erg s^{-1}, which could conceivably be produced by high levels of star formation; however, we verified that the X-ray emission is AGN-dominated in these sources by finding that their star formation rates, taken from Stanley et al. (2015), are too low by a factor of >30 to produce the observed hard-band X-ray luminosities (following the same procedure as above). We note that Xue et al. (2011) also classify all of our CDF-S X-ray targets as AGN, using a more comprehensive X-ray and multi-wavelength identification procedure.

The initial KMOS data of KASH*z* that are presented here were obtained during European Southern Observatory (ESO) periods P92–P95 (details provided in Section 2.6.1). During these observations we obtained KMOS data for 79 X-ray AGN that met our selection criteria. This programme was jointly observed with the KROSS guaranteed time observing (GTO) programme (Stott et al., in preparation) and the choice of these 79 AGN targets, drawn from our parent sample, was dictated by those that could be observed inside the KMOS pointing positions chosen by the KROSS team, with a slight preference to selecting $z \approx 1.1$–1.7 AGN for which we could observe [O III]. In summary, the observed targets are effectively randomly selected from the luminosity and redshift plane of the parent sample (Fig. 1; also see Section 2.3 for more discussion on how representative our targets are).

We supplemented our sample of 79 KMOS targets with archival observations taken with SINFONI (Spectrograph for INtegral Field

Observations in the Near Infrared; Eisenhauer et al. 2003). SIN-FONI contains a single-object NIR IFU, similar to the individual KMOS IFUs (see Section 2.6.2). We queried the ESO archive[4] for *J*-band SINFONI observations at the positions of the X-ray AGN in our parent catalogues. This yielded 10 targets that met our selection criteria to add to the *J*-band sample presented in this paper.

Overall, our current KASH*z* sample contains 89 targets. This sample consists of 54 $z \approx 1.1$–1.7 AGN, for which we targeted the [O III]4959,5007 emission-line doublet and 35 $z \approx 0.6$–1.1 AGN, for which we targeted the Hα and [N II]6548,6583 emission lines (see Fig. 1). The targets are listed in Table A1 and we have used a naming convention that combines their field names and corresponding X-ray ID in the format: 'field'-'X-ray ID'. For C-COSMOS and *XMM*-COSMOS the field names are shortened to COS and XCOS, respectively.

Out of the 89 targets presented here, 83 have a spectroscopic archival redshift and six targets have a photometric redshift. All of the photometric redshift targets are from the C-COSMOS field (Table A1). We note that two of the spectroscopic targets are classified as having 'insecure' redshifts (e.g. based on low signal-to-noise ratios or single-lines; see Table A1) by Xue et al. (2011).[5] For our emission-line detected targets (see Section 4.1), our emission-line redshifts agree with the archival redshifts, such that $|z_L - z_A|/(1 + z_L) \lesssim 0.005$, except for the following three exceptions: COS-44 ($z_A = 1.51$; $z_L = 0.80$) and COS-1199 ($z_A = 0.77$; $z_L = 0.85$), both of which had photometric archival redshifts, and CDFS-561 ($z_A = 0.80$; $z_L = 0.98$), which is a spectroscopic archival redshift quoted by Popesso et al. (2009). We discuss the detection rates of our targets, and how they relate to the archival redshifts, in Section 4.1. For the emission-line detected AGN, throughout this work, we define the redshift as $z = z_L$ and for the undetected targets or non-targeted AGN in the parent sample we use $z = z_A$.

2.3 The KASH*z* sample in context

In Fig. 2 we show histograms of the X-ray luminosities for the X-ray sources that met our selection criteria and of the subset of these that were observed for this work (Table A1). This figure shows that our targets represent the full luminosity range of the parent catalogues from which they were selected. A two-sided Kolmogorov–Smirnov (KS) test yields probability values of 0.53, 0.08 and 0.96 that the two distributions are drawn from the same distribution for COSMOS, CDF-S and UDS, respectively. Hence, there is no evidence that the targets are *not* representative of the parent population from which they were selected. For SSA22 there are only two targets and hence it is not meaningful to perform a KS test. We conclude that our targets are broadly representative of the X-ray AGN population covered by the parent catalogues. It is worth noting that, although X-rays surveys arguably provide the most uniform method for selecting AGN, there will be some unknown fraction of luminous AGN that are detected at other wavelengths, that are very heavily obscured in X-rays and are not detected in these surveys (e.g. Vignali, Alexander & Comastri 2006; Del Moro et al. 2015). We defer a comparison to AGN selected using different observational methods to future work.

---

[4] http://archive.eso.org/eso/eso_archive_main.html

[5] CDFS-549 is a third source with an 'insecure' spectroscopic redshift in Xue et al. (2011) ($z_A = 1.55$); however, Williams et al. (2014) detected Hα at $z = 1.553$ so we re-classify it as secure.
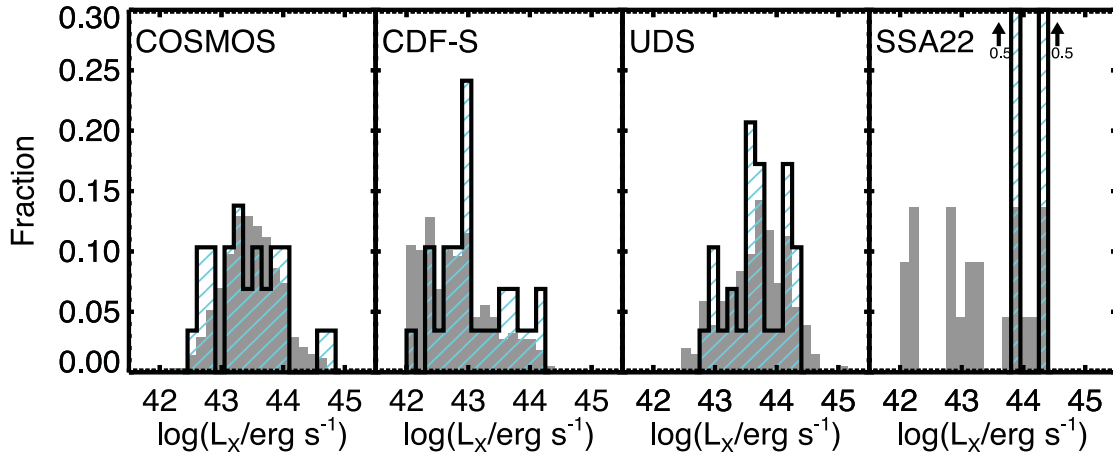
Figure 2. The 2–10 keV X-ray luminosity (L_X) distributions for the sources that met our selection criteria from the four deep fields described in Section 2.1 (solid histograms). Each field is shown in separate panels and we overlay, as hashed histograms, the targets presented in this work that were selected from these respective fields. With the exception of SSA22 (from which there are only two targets presented here), the targets have a range of X-ray luminosities that are representative of the parent catalogues from which they were selected (see Section 2.3).

2.4 Selecting X-ray obscured candidates

For part of this work we compare the emission-line properties of the targets that are most likely to be X-ray obscured (i.e. $N_H \gtrsim 10^{22}$ cm^{-2}) with those that are X-ray unobscured. Due to the use of both *XMM–Newton* and *Chandra* catalogues in our AGN selection and due to the various depths of these catalogues, we opt to use the simplest diagnostic possible to separate X-ray obscured and unobscured sources. We take the ratio of the hard-band to soft-band X-ray fluxes (i.e. $F_{2-10\,keV}/F_{0.5-2\,keV}$; see Table A1) as a proxy for an observed power-law index, Γ_{obs}. For the seven targets where there are no direct hard-band detections, we use the hard-band flux upper limits from the original X-ray catalogues (see Section 2.1). We use a threshold of $\Gamma_{obs} < 1.4$, or equivalently $F_{2-10\,keV}/F_{0.5-2\,keV} > 3.03$, to select our 'obscured candidates'. Assuming a typical *intrinsic* power-law index of $\Gamma = 1.8$ (e.g. Nandra & Pounds 1994; George et al. 2000), this threshold corresponds to an intrinsic column density of $N_H \gtrsim 1 \times 10^{22}$ cm^{-2}, at the median redshift of our Hα targets (i.e. $z = 0.86$), and $N_H \gtrsim 2 \times 10^{22}$ cm^{-2}, at the median redshift of our [O III] targets (i.e. $z = 1.4$). This criterion yields 32 'obscured candidates' and 52 'unobscured candidates' out of the original sample of 89. The other five targets have X-ray upper limit values such that they can not be classified. We identify the classification of each target in Table A1. Reassuringly, we find that only one of the obscured candidates has an identified broad-line region (BLR) component in our data (see Section 3.2), compared to 15 of the unobscured candidates, as expected if X-ray obscured AGN are more likely to also have an obscured BLR. Furthermore, this one exception only just meets our criteria for selecting obscured sources, i.e. it has a flux ratio of $F_{2-10\,keV}/F_{0.5-2\,keV} = 3.1$. This provides extra support that our obscured AGN criterion is reliable.

2.5 Selecting 'radio luminous' AGN

For part of this work we compare KASHz targets that are radio luminous with those that are not (Section 4.3.2). Therefore, we collated the available 1.4 GHz radio catalogues for the COSMOS, CDF-S and UDS fields (a suitable SSA22 catalogue is not available in the literature.). For COSMOS we used the 5σ VLA catalogue of Schinnerer et al. (2010); for CDF-S we use the 5σ VLA catalogue presented in Miller et al. (2013) and for UDS we use the 4σ VLA

catalogue in Arumugam et al. (in preparation). We match the positions of our targets to the radio positions using a 2 arcsec matching radius. For the 87 targets in these three fields (i.e. excluding the two SSA22 targets), we obtain 26 radio matches. We calculate radio luminosities using the aperture-integrated flux densities and assume a spectral index of $\alpha = -0.7$. All three catalogues have typical sensitivities around ≈ 10 μJy beam^{-1} and, therefore, we are complete to 1.4 GHz radio luminosities of $L_{1.4\,GHz} \approx 10^{24}$ W Hz^{-1} at the redshift ranges of interest in this work (see Fig. 1). Therefore, for this work we separate our targets which are 'radio luminous', i.e. with $L_{1.4\,GHz} > 10^{24}$ W Hz^{-1} (11 targets) from those with $L_{1.4\,GHz} < 10^{24}$ W Hz^{-1} (76 targets). Table A1 highlights which targets fall into each category. Above this luminosity threshold, sources are generally thought to have AGN-dominated radio emission and could be predominantly 'radio-loud' sources (e.g. Bonzini et al. 2013; Del Moro et al. 2013; Rawlings et al. 2015). Furthermore, we find that 13 per cent of our targets are in our 'radio luminous' category, which is broadly consistent with the observed radio-loud fraction of ≈ 10 per cent for luminous AGN (e.g. Zakamska et al. 2004; La Franca, Melini & Fiore 2010; Kalfountzou et al. 2014).

2.6 Observations

2.6.1 KMOS observations and data reduction

The majority of the KASHz targets presented here (i.e. 79 out of the 89) were observed using the KMOS instrument on the VLT (Sharples et al. 2004, 2013). KMOS has 24 IFUs that operate simultaneously and can be independently positioned inside a 7.2 arcmin diameter circular field. Each IFU has a field of view of 2.8×2.8 arcsec^2 with a pixel scale of 0.2 arcsec. The 24 IFUs are fed to three spectrographs (eight IFUs per spectrograph). For this initial KASHz paper we present results of sources observed with the YJ grating that covers a wavelength range of 1.03–1.34 μm and has a band-centre spectral resolution of $R \sim 3600$. The KMOS observations were taken during ESO Periods 92–95, sharing IFUs with the KROSS GTO programme, which targeted high-redshift star-forming galaxies in the same fields (see Stott et al., in preparation; also see Section 3.4).

Observations were carried out using ABAABAAB sequences (where A frames are on-source and B frames are on-sky), with 600 s

integrations per position and up to 0.2 arcsec spatial dithering between on-source frames. The targets have total on-source exposure times of 5.4–11.4 ks (i.e. 9–19 on-source exposures). The individual exposure times were a result of the length of observing time with acceptable weather conditions during the various observing runs allocated to the GTO team and the final on-source exposure times are tabulated for individual targets in Table A1. The median *J*-band seeing was 0.7 arcsec with 90 per cent below 1.0 arcsec. The data were reduced using the ESOREX/SPARK pipeline (Davies et al. 2013) which flat-fields, illumination corrects, wavelength calibrates and uses observations of standard stars, taken alongside the science frames, to flux calibrate. Additional sky subtraction was performed using dedicated sky IFUs. Repeated observations of targets were stacked into the final fully reduced data cubes with a clipped mean using the ESOREX pipeline. Finally, we rebinned the cubes on to a 0.1 arcsec pixel scale. For full details of the observations and data reduction see Stott et al. (in preparation).

2.6.2 *SINFONI observations and data reduction*

In addition to the primary observations using KMOS, we supplement the KASH*z* sample with archival observations of 10 targets that met our selection criteria, taken with SINFONI (see Section 2.1). The observations presented here were all observed using SINFONI's 8×8 arcsec^2 field of view, which is divided into 32 slices of width 0.25 arcsec with a pixel scale of 0.125 arcsec along the slices. The observations were carried out using the *J*-band grating which has an approximate resolution of $R \sim 3000$. The observations used a variety of observing strategies, but in all cases we were able to subtract on-sky frames from on-target frames. We reduced the SINFONI data using the standard ESOREX pipeline (Modigliani et al. 2007) that performs flat-fielding, wavelength calibration and cube re-construction. We flux calibrated individual data cubes using standard star observations taken the same night as the science observations at a similar airmass. We stacked individual cubes of the same source by creating white-light (collapsed images), centroiding the cubes based on these images and then performing a median stack with a 3σ clipping threshold, rejecting cubes that could not be well centred. Three of the targets (CDFS-51; CDFS-370 and CDFS-492) were not detected in any of the individual cubes and for these targets, we used the offset pattern in the headers to stack the cubes. The total on-source exposure times range from 2.4 to 25.2 ks and are tabulated for the individual targets in Table A1.

3 ANALYSIS AND COMPARISON SAMPLES

In this section we describe our analyses of the galaxy-integrated spectra for the targets presented in this work (Section 3.1–3.3) and we also describe our comparison samples of low-redshift AGN and high-redshift star-forming galaxies (Section 3.4).

3.1 Galaxy-integrated spectra

We extracted a galaxy-integrated spectrum from the KMOS and SINFONI data cubes using the methods described below. We show example spectra in Figs 3 and 4, and all 89 spectra are shown in Figs A1–A3.

To obtain the galaxy-integrated spectra, we initially define a galaxy 'centroid' by creating wavelength-collapsed images from the data cubes, including both continuum and line emission. For three targets no emission lines or continuum were detected and we
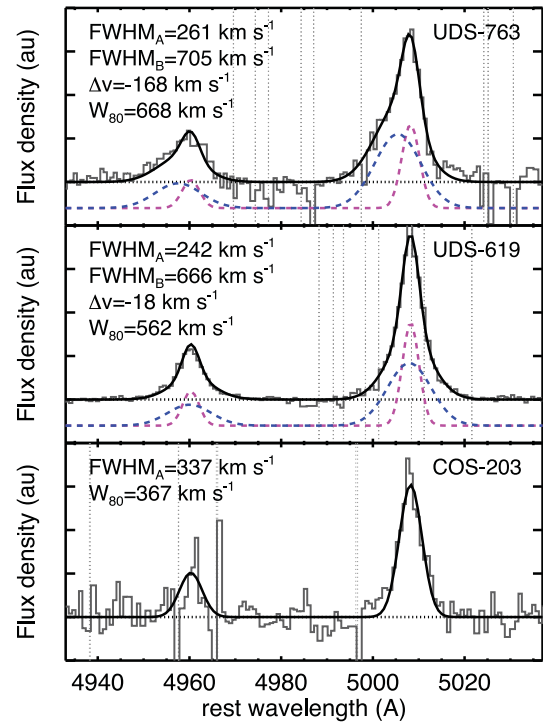
Figure 3. Three examples of our continuum-subtracted, high signal-to-noise ratio, [O III]4959,5007 emission-line profiles. The black curves show our fits to the emission-line profiles and the dashed curves show the individual Gaussian components, where applicable, with an arbitrary offset in the *y*-axis. The vertical dotted lines indicate the wavelengths of the brightest sky lines (Rousselot et al. 2000). These examples demonstrate the diversity of emission-line profiles observed in the sample. Top: a broad, highly asymmetric profile; (middle) a broad, almost symmetric profile; and (bottom) a relatively narrow profile without a strong underlying broad component. Fig. A1 presents the [O III]4959,5007 emission-line profiles for all of the targets.

therefore assumed the centroid was at the centre of the cube for these targets (we later exclude these three targets from our analyses; see Section 4.1). For the targets CDFS-454 and CDFS-606 there are bright continuum sources in the field of view that are spatially offset from the line-emitting regions by $\approx$2 arcsec and $\approx$1.2 arcsec, respectively. For these two targets we centred on the line emission. For each galaxy we summed the spectra from all of the pixels within a circular aperture around the galaxy centroids. We chose diameters to broadly match the physical size scale of the spectra obtained with the SDSS fibres in our low-redshift comparison sample (see Section 3.4). The median redshift of the low-redshift SDSS comparison sample is $z = 0.14$ which means that the spectra are from a median 7.4 kpc diameter aperture due to the 3 arcsec fibres. For our KASH*z* targets, the median redshift of the [O III] sample is $z = 1.4$ (i.e. corresponding to 8.4 kpc arcsec^{-1}) and the median redshift of the Hα sample is $z = 0.86$ (i.e. corresponding to 7.7 kpc arcsec^{-1}). Therefore, for the [O III] targets we used a diameter of 0.9 arcsec and for the Hα targets we used a diameter of 1.0 arcsec. These correspond to physical diameters of 7.6 ± 0.8 kpc and 7.7 ± 0.8 kpc, respectively, where the uncertainties correspond to a 0.1 arcsec pixel scale. To assess the effect of using alternative 'galaxy-wide' apertures, which will cover any losses of flux due to seeing, we also extracted spectra using 2 arcsec diameter apertures for all sources. We discuss the corrections required to the emission-line luminosities in Section 4.2. For the sources significantly detected in both apertures, we find that
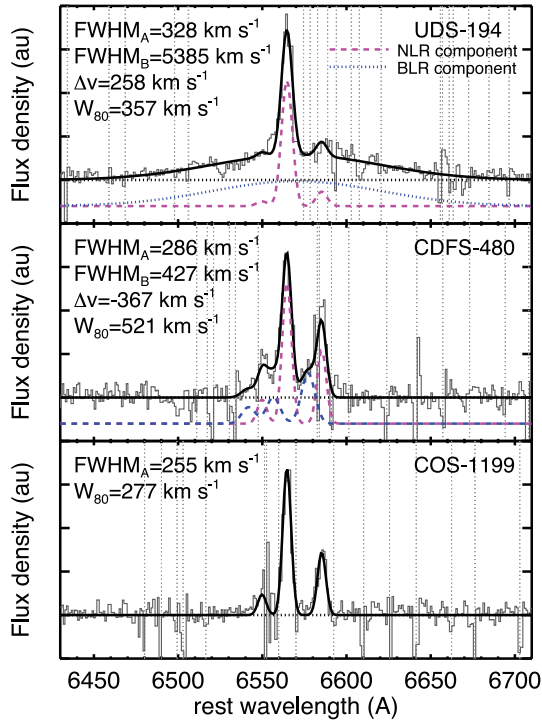
Figure 4. Three examples of our continuum-subtracted Hα+[N II]6548,6583 emission-line profiles. The black curves show our fits to the emission-line profiles and the dashed or dot–dashed curves show the individual Gaussian components, where applicable, with an arbitrary offset. The vertical dotted lines indicate the wavelengths of the brightest sky lines (Rousselot et al. 2000). These examples represent the three different types of profile fits used in this work. Top: a source with a BLR component in addition to the NLR component; (middle) a source with no BLR component but two NLR components seen in Hα and [N II] and (bottom) a source with a single NLR component. Figs A2 and A3 present the Hα+[N II]6548,6583 emission-line profiles for all of the targets.

our emission-line width measurements (Section 3.3) are consistent between both apertures within $1 \times$ the errors for 77 per cent of the targets and within $2 \times$ the errors for 97 per cent of the targets. We defer discussion of how the emission-line widths change as a function of radius to future papers.

3.2 Emission-line fitting

The emission-line profiles were fitted with one or two Gaussian components (with free centroids, line widths and fluxes) and a straight line to define the local continuum (with a free slope and normalization). The continuum regions were defined to be small wavelength regions each side of the emission lines being fitted. The noise in the spectra was also calculated in these regions. The fits were performed using a minimizing-χ^2 method, using the IDL routine MPFIT (Markwardt 2009) and we weighted against the wavelengths of the brightest sky lines (taken from Rousselot et al. 2000; see dotted lines in Figs 3 and 4). Quoted line widths have been corrected for spectral resolution, where we measured the wavelength-dependent spectral resolution using the emission lines in sky spectra. We provide details of how we modelled the emission-line profiles below and tabulate the parameters of the fits in Table A1.

For the [O III]5007,4959 emission-line doublet (see examples in Fig. A1), we simultaneously fit the [O III]5007 and [O III]4959 emission lines using the same velocity-widths and fixing the relative cen-

troids, using the rest-frame wavelengths of 5008.24 Å and 4960.30 Å, respectively (i.e. their vacuum wavelengths). The flux ratio of the doublet was fixed to be 2.99 (Dimitrijević et al. 2007). We initially attempted to fit this emission-line doublet with one Gaussian component (with three free parameters) and then with two Gaussian components (with six free parameters) per emission line. We accepted the two-component fit if there is a significant improvement in the χ^2 values. We required $\Delta\chi^2 > 15$, where this threshold was chosen to provide the best description of the emission-line profiles across the whole sample.[6] We show the fits to the full set of [O III]4959,5007 emission-line profiles in Fig. A1.

For the Hα emission-line profiles, along with the nearby [N II]6548,6583 doublet, we follow a similar approach as for the [O III] emission-line profiles (see example fits in Fig. 4). Based on the atomic transition probabilities, we fix the flux ratio of the [N II] doublet to be 3.06 (Osterbrock & Ferland 2006). To reduce the degeneracy between parameters, we couple the [N II] and Hα emission-line profiles with a fixed wavelength separation between the three emission lines (i.e. by using the rest-frame vacuum wavelengths of 6549.86 Å, 6564.61 Å and 6585.27 Å). We also fix the line widths of the Hα, [N II]6548 and [N II]6583 emission-line components to be the same. The common line width, overall centroid and the individual fluxes of Hα and the [N II] doublet are free to vary. We note that the emission-line coupling is only applied for narrow-line region (NLR) components and is not applicable for BLR components that are not observed in [N II] emission. For these Hα targets, we have two situations to consider: those which exhibit a BLR component and those which do not.

(i) For the sources where a broad component is seen in the Hα emission-line profile but not in the [N II] emission-line profile, these are the BLR or 'Type 1' sources. We note that all seven BLR components that we identify have a full-width at half maximum (FWHM) >2000 km s^{-1} (see Table A1), further indicating that these are true BLRs. For these, we fit one Gaussian component to the [N II]6548,6583 emission-line doublet and two Gaussian components to the Hα emission line. The narrower of the Hα components is coupled to the [N II] emission as described above. The broader of these components, i.e. the BLR component, has a centroid, line width and a flux that are free to vary. Overall, for these sources, there are seven free parameters in the fits and an example is given in the top panel of Fig. 4.

(ii) For sources without a BLR component (i.e. the 'Type 2' sources) we first fit a single Gaussian component to the Hα emission line and [N II] doublet, which are coupled as described above. We then add a second Gaussian component to all three emission lines. This is coupled in the same way, but with the additional constraint that we force the [N II]/Hα flux ratio of the broader Gaussian components to be $\gtrsim 2$ times that of the narrower component. This follows Genzel et al. (2014), who found that this was typical for high-redshift galaxies and AGN. This extra constraint is required to prevent considerable degeneracy in the fits that can lead to unphysical results. As for the [O III] emission-line profiles, we accept the

---

[6] We note that absolute $\Delta\chi^2$ values are commonly used for model selection. For example, the Bayesian Information Criterion (BIC; Schwarz 1978) uses $\Delta\chi^2$ but penalizes against models with more free parameters. This is defined as BIC $= \Delta\chi^2 + k \ln(N)$, where N is the number of data points and k is the number of free parameters. For the fits where we favoured two-component models (including all of the Hα fits) we find a median of ΔBIC $= 34$ between the two- and one-component models, which corresponds to strong evidence in favour of the two-component models (e.g. Mukherjee et al. 1998).

two component Gaussian fit if there is a significant improvement in the χ^2 values, i.e. $\Delta\chi^2 > 15$. Examples of a one- and two-Gaussian component fit are shown in Fig. 4.

The coupling between the [N II] and Hα emission-line profiles is a requirement of the restricted signal-to-noise ratio of the observations to avoid a high level of degeneracy, and is not necessarily physical. The same coupling approach is often followed for high-redshift galaxies and AGN, which are inevitably subject to limited signal-to-noise ratio observations (e.g. Swinbank et al. 2004; Genzel et al. 2011, 2014; Stott et al. in preparation), and makes the underlying assumption that the emission lines are being produced by the same gas, undergoing the same kinematics (see discussion on this in Section 4.3.3). However, our emission-line profile models appear to be a good description of the data and are sufficient for our purposes of measuring the Hα emission-line widths and luminosities and the [N II]/Hα flux ratios. We show the full set of fits to the Hα emission-line profiles in Fig. A2 and Fig. A3 and tabulate the parameters in Table A1.

To calculate the uncertainties on the derived parameters, we follow the same procedure for both the [O III] and Hα+[N II] emission-line profiles. We use our best-fitting models to generate 1000 random spectra by adding random noise to these fits at the level measured in the original spectra, and then re-fit and re-derive the parameters for these random spectra. The quoted uncertainties are from the average of the 16th and 84th percentiles in the distribution of the parameters of these random fits. We add an additional uncertainty of 30 per cent when plotting emission-line luminosities to account for an estimated systematic uncertainty on the absolute flux calibration of the data cubes.

3.3 Measuring the overall emission-line velocity-widths

A key aspect of this work is to characterize the overall velocity-widths of the emission-line profiles. Our spectra have a range in emission-line profile shapes and signal-to-noise ratio values (see Figs A1–A3). Therefore, it is most applicable to characterize the velocity-widths with a single non-parametric measurement that is independent of the number of Gaussian components used in the best-fitting emission-line models (see Section 3.2). Furthermore, many studies at low redshift use non-parametric definitions to describe the very complex emission-line profiles (e.g. Liu et al. 2013; Harrison et al. 2014; McElroy et al. 2015) and it is useful to be able to compare to these studies. Therefore, we use the line width definition, W_{80}, which is the velocity-width that contains 80 per cent of the emission-line flux, i.e. $W_{80} = v_{90} - v_{10}$, where v_{10} and v_{90} are the 10th and 90th percentiles, respectively. For a single Gaussian component $W_{80} = 1.09 \times$ FWHM. These W_{80} values are measured from the models of the emission-line profiles described in Section 3.2 and are tabulated in Table A1. For this work, we are only interested in the kinematics of the host-galaxy gas, and therefore for the seven targets with a BLR component, the W_{80} measurements only refer to the NLR emission (i.e. only the NLR components were used to calculate W_{80}).

3.4 Comparison samples

A key aspect of this work is to compare our high-redshift AGN targets to high-redshift star-forming galaxies and low-redshift AGN. This enables us to assess the evolution in the prevalence of outflows observed in AGN and also to compare high-redshift galaxies that do not host X-ray detected AGN. In the following sub-sections we describe how we constructed our comparison samples.

3.4.1 Low-redshift AGN comparison sample

For a low-redshift AGN comparison sample we make use of the catalogue provided by Mullaney et al. (2013). This sample contains emission-line profile fits to $\approx$24 000 $z < 0.4$ optically selected AGN from the SDSS spectroscopic data base. AGN are identified based on their [O III]/Hβ and Hα/[N II] emission-line ratios (following e.g. Kewley et al. 2006) or the identification of a BLR component. We take the 24 258 AGN in the Mullaney et al. (2013) catalogue, but reject the 37 sources where the [O III] emission-line profile fits have FWHM = 4000 km s^{-1}, which signifies that the fits failed for these sources (Mullaney et al. 2013). This leads to a final sample of 24 221 $z < 0.4$ optically selected AGN.

Mullaney et al. (2013) fit the [O III] emission-line profiles of the low-redshift AGN with one or two Gaussian components, following a very similar procedure to that adopted here (Section 3.2). Following Section 3.3, we measure W_{80} for these sources from the [O III] emission-line profile fits, correcting for the wavelength-dependent SDSS spectral resolution. Due to the nature of the fitting routine in Mullaney et al. (2013), the Hα emission-line profile fits are coupled to the [O III] emission-line profile fits. We are unable to replicate this method for our KASHz Hα targets because we do not have simultaneous [O III] and Hα constraints. Therefore, to avoid a biased comparison of line width measurements, we do not use the Hα emission-line fits provided by Mullaney et al. (2013). Instead, we use the single Gaussian component emission-line fits, as measured by the SDSS team for these targets (Abazajian et al. 2009), to calculate the W_{80} values (see Section 3.3). These SDSS measurements do not separate NLR emission from BLR emission; therefore, we are required to exclude the Type 1 AGN from this comparison sample when comparing the Hα emission-line widths to the KASHz targets (Section 4.3.3). However, we note that the comparison is reasonable for our Type 2 KASHz Hα targets because all but one of the KASHz Type 2 AGN are fitted with a single Gaussian component.

To obtain X-ray luminosities for the $z < 0.4$ AGN sample described above, we match the Mullaney et al. (2013) sample to data release 5 of the *XMM* serendipitous survey (Rosen et al. 2015), using a 1.5 arcsec matching radius. This resulted in 554 matches. We calculate X-ray luminosities using the quoted 2–12 keV fluxes, which we convert to 2–10 keV fluxes using a correction factor of 0.872 and convert to hard-band X-ray luminosities using equation (1).

For part of the analysis in this work (see Sections 4.3.1 and 4.3.3) we are required to construct low-redshift comparison samples that are luminosity-matched to our KASHz samples. To do this, we randomly select sources from the sample described above to construct the following three comparison samples: (1) a low-redshift sample of $\approx$1000 AGN (both Type 1 and Type 2) that has an [O III] luminosity distribution that is the same as our [O III]-detected KASHz targets (see Section 4.1); (2) a low-redshift sample of $\approx$100 AGN (both Type 1 and Type 2) that has an X-ray luminosity distribution that is the same as our [O III]-detected KASHz targets and (3) a low-redshift sample of $\approx$500 Type 2 AGN that has an Hα luminosity distribution that is the same as our Hα-detected KASHz targets (see Section 4.1). It was not possible to construct a low-redshift sample that was X-ray luminosity matched to our KASHz Hα targets due to the lack of X-ray luminous Type 2 AGN in Mullaney et al. (2013).

However, we note that matching by [O III] luminosity compared to matching by X-ray luminosity does not change the conclusions presented in Section 4.3.1.

3.4.2 High-redshift star-forming galaxy comparison sample

To construct a high-redshift galaxy comparison sample we make use of the KROSS survey (Stott et al., in preparation). This is a KMOS GTO survey of $z \approx 0.6$–1.1 star-forming galaxies, which was observed simultaneously with KASHz (see Section 2.6.1). These targets were selected on the basis of their K-band magnitudes and r–z colours to create a sample of star-forming galaxies with stellar masses of $\approx 10^{9-11}$ $M_\odot$ (see Stott et al., in preparation). This data set provides an ideal comparison sample of galaxies for our Hα sample of X-ray detected AGN at the same redshift.

The initial phases of the KROSS survey, i.e. the data obtained during ESO periods P92–P94 and the KMOS commissioning run, contain 514 galaxies (Stott et al., in preparation). For these KROSS galaxies, we extract galaxy-integrated spectra from the KMOS data cubes and perform emission-line profile fits following the same methods as performed on the KASHz sample (see Sections 3.2 and 3.3). We require that Hα emission is detected at $\geq 3\sigma$ and the emission-line profiles are well described using one or two Gaussian components plus a straight-line local continuum (i.e. following our methods described in Section 3.2). Based on these criteria, we end up with W_{80} Hα measurements for 378 of the KROSS galaxies, where we have also removed any X-ray AGN. In our analyses (Section 4.3.3), we also make use of the stellar masses for these galaxies, as described in Stott et al. (in preparation) and note that the average stellar mass of the sample we have constructed is $\log(M_\star/M_\odot) = 10.3$.

3.5 Emission-line profile stacks

As part of our investigation we create *average* emission-line profiles by using spectral stacking analyses on the KASHz targets and comparison samples (see Sections 4.3.1 and 4.3.3). To create the emission-line profile stacks, we de-redshift each continuum-subtracted spectrum to the rest frame and then normalize each spectrum to the peak flux density of the emission-line profile fits. We then construct the stacked average spectra by taking a mean of the flux densities at each spectral pixel, but removing the pixels affected by strong sky-line residuals. An uncertainty on the average at each spectral pixel is obtained by bootstrap resampling, with replacement, the stacks 1000 times and deriving the inner 68 per cent of these stacks. Our analysis is focused on the kinematics in the host galaxies and, therefore, when stacking the Hα emission-line profiles, we do not include any Type 1 sources (i.e. those containing an identified BLR component). We fit the stacked spectra following the procedures described in Section 3.2; however, due to the increased signal-to-noise ratio we include one extra free parameter when fitting the Hα emission-line profiles which allows the flux ratio of the Hα and [N II] emission lines to be free for *both* Gaussian components. This means that it is possible for the [N II] and Hα emission lines to have different line widths (i.e. W_{80} values; see discussion in Section 4.3.3).

4 RESULTS AND DISCUSSION

We present the first results from KASHz, which is an ongoing programme that is utilizing VLT/KMOS GTO to build up a large sample of IFS data of high-redshift AGN (see Section 2). We have obtained new KMOS observations and combined them with archival SINFONI observations, to compile a sample of 89 X-ray detected AGN observed in the J band. These AGN have redshifts of $z = 0.6$–1.7 and hard-band (2–10 keV) X-ray luminosities in the range of $L_X = 10^{42}$–10^{45} erg s^{-1} and are representative of the parent population from which they were selected (see Figs 1 and 2). Of the 89 targets presented here, 54 have $z = 1.1$–1.7 and were targeted to observe the [O III]4959,5007 emission-line doublet and 35 have $z = 0.6$–1.1 and were targeted to observe the Hα+[N II]6548,6583 emission lines. In this paper, we present the galaxy-integrated emission-line profiles for all 89 targets and these are shown, along with their emission-line profile fits, in Figs A1–A3. In the following sub-sections we present: (1) the emission-line detection rates and definition of the final sample of 82 targets used for all further analyses (Section 4.1); (2) the relationship between emission-line luminosities and X-ray luminosities (Section 4.2); and (3) the prevalence and drivers of high-velocity ionized outflows (Section 4.3).

4.1 Detection rates and defining the final sample

We detected continuum and/or emission lines in the IFS data for 86 out of the 89 KASHz targets (i.e. 97 per cent) (see Table A1 for details). Overall, 40 targets were detected in [O III], 32 were detected in Hα and 14 were detected in continuum only. One of the reasons for a lack of an emission-line detection appears to be inaccurate photometric redshifts (photometric redshifts were only used for six COSMOS targets; see Section 2.2). Of the six targets with photometric archival redshifts, only two resulted in emission-line detections and one of these was detected in Hα with $z_L \approx 1.5$, despite an archival redshift of $z_A \approx 0.8$ (see Table A1). To ensure that a target is undetected in line emission because of an intrinsically low emission-line flux, as opposed to an incorrect redshift or position, for all further analyses we only consider non-detections if: (1) they have a secure spectroscopic archival redshift (this excludes five non-detections) and (2) they were detected in continuum so that we have a reliable source position to extract the spectrum from the data cube (this excludes two further non-detections). Based on these exclusions, for the remainder of this work, we only consider 82 of the original 89 targets, which consists of: (a) 40 emission-line detected [O III] targets and eight [O III] targets that were only detected in continuum (i.e. an 83 per cent emission-line detection rate) and (b) 32 emission-line detected Hα targets and two Hα targets that were detected only in continuum (i.e. a 94 per cent emission-line detection rate). Overall, this results in an emission-line detection rate of 88 per cent for our final sample of 82 targets.

We note here that the non-emission-line detected sources are not only associated with the lowest X-ray luminosity sources (see Figs 5 and 6). However, of the 10 targets from our final sample of 82, which were not detected in emission lines, five are classified as X-ray obscured, four are classified as unobscured and one is unclassified, which results in a 50–60 per cent obscured fraction, compared to 33–38 per cent for the 72 emission-line detected targets (see Section 2.4 and Table A1). Although we cannot draw any firm conclusions from this, it is interesting to speculate that this provides evidence that the material obscuring the X-rays may also be responsible for the lack of emission-line luminosity and therefore that the obscuring material may be associated with the host galaxy (i.e. galactic-scale dust; see e.g. Malkan, Gorjian & Tam 1998; Goulding & Alexander 2009; Juneau et al. 2013).
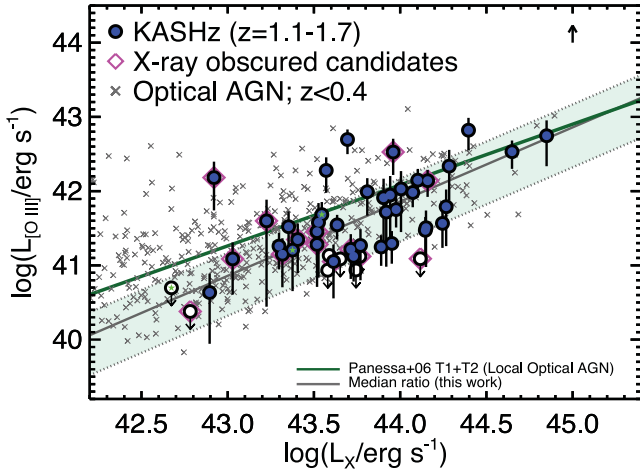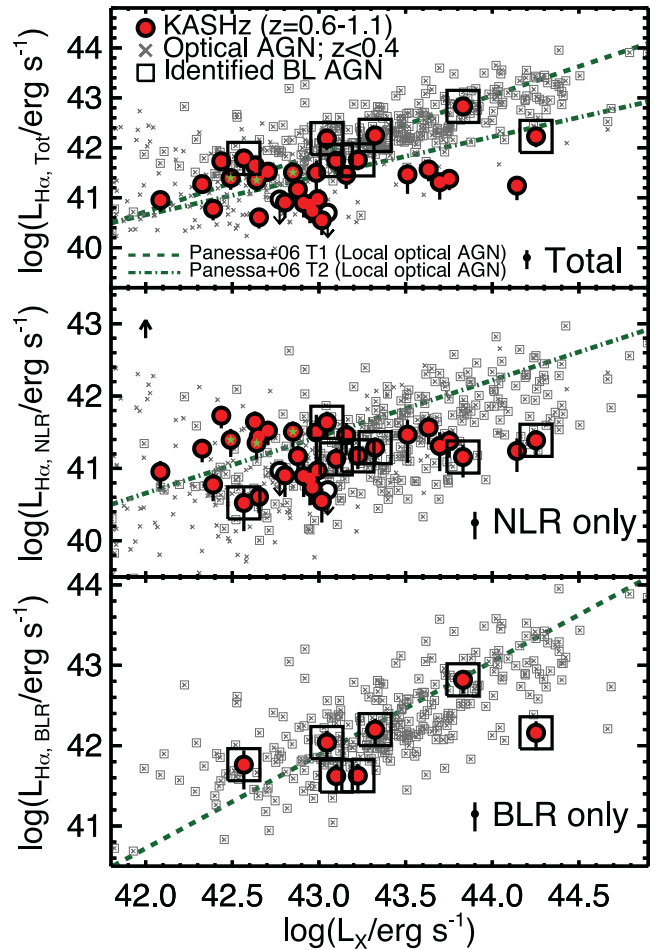
Figure 5. Total [O III] emission-line luminosity ($L_{[O III]}$) versus hard-band (2–10 keV) X-ray luminosity (L_X) for the 48 $z \approx$ 1.1–1.7 KASHz targets. Non-detections are represented as hollow symbols, which signify 3σ upper limits, and stars are the same as in Fig. 1. The solid line shows the median ratio of $\log(L_{[O III]}/L_X) = -2.1^{+0.3}_{-0.5}$ for these 48 KASHz [O III] targets. The shaded region indicates the $\approx 1\sigma$ scatter on this ratio (see Section 4.2). The arrow shows the median aperture correction when using an ≈ 2 times larger aperture (see Section 4.2). The dot–dashed line is the relationship for local Seyferts and QSOs presented in Panessa et al. (2006). We also show our $z <$ 0.4 optical AGN comparison sample (see Section 3.4). The X-ray selected KASHz sample is broadly consistent with the low-redshift optically selected samples, but with the expected tendency towards lower [O III] luminosities (Section 4.2).

4.2 Emission-line luminosities compared to X-ray luminosities

4.2.1 [O III] luminosity versus X-ray luminosity

Both the [O III] emission-line luminosity and X-ray luminosity have been used to estimate total AGN power (i.e. bolometric luminosities) and, therefore, the relationship between these two quantities and any possible redshift evolution is of fundamental importance for interpreting observations (e.g. Mulchaey et al. 1994; Heckman et al. 2005; Netzer et al. 2006; Panessa et al. 2006; LaMassa et al. 2009; Lamastra et al. 2009; Trouille & Barger 2010; Lusso et al. 2012; Berney et al. 2015). Until recently, there has been limited available NIR spectroscopy and therefore limited measurements of the [O III] emission-line luminosities for high-redshift, i.e. $z \gtrsim 1$, X-ray detected AGN. In this sub-section we compare the [O III] and X-ray luminosities for our final KASHz sample of 48 $z \approx 1.1$–1.7 AGN that were targeted to observe [O III] emission (see Section 4.1).

In Fig. 5 we show [O III] luminosity versus X-ray luminosity for the KASHz targets. A correlation is observed between these two quantities, although with a large scatter, in qualitative agreement with studies of low-redshift AGN (e.g. Heckman et al. 2005; Panessa et al. 2006). We note that we observe the same trend when plotting [O III] flux versus X-ray flux and therefore the correlation is not an artefact of flux limits. For the 48 targets, we find a median luminosity ratio of $\log(L_{[O III]}/L_X) = -2.1^{+0.3}_{-0.5}$, where the quoted upper and lower bounds contain the inner 66 per cent of the targets, i.e. roughly the 1σ scatter.[7] To calculate the median ratio of [O III]

[7] We note that we quote the range on the [O III] to X-ray luminosity ratio for the inner 66 per cent, rather than the more standard 68.3 per cent, applicable for 1σ, because eight out of the 48 targets are undetected (i.e. 17 per cent) and therefore we cannot constrain the 15.9 th percentile of the distribution.

Figure 6. Hα emission-line luminosity ($L_{H\alpha}$) versus hard-band (2–10 keV) X-ray luminosity (L_X) for the 34 $z \approx$ 0.6–1.1 KASHz Hα targets. In the top panel we show the total emission-line luminosities; in the middle panel we show the luminosities where the BLR components have been subtracted (i.e. the NLR luminosities) and in the bottom panel we show the BLR luminosities, where applicable. Non-detections are represented as hollow symbols, which signify 3σ upper limits, and the stars are the same as in Fig. 1. For clarity, the median uncertainty on the line luminosities is shown at the bottom of each panel. The arrow in the second panel shows the median aperture correction when using an ≈ 2 times larger aperture (see Section 4.2). The dashed and dot–dashed lines show the luminosity-dependent relationships for Type 1 (T1) and Type 2 (T2) local AGN presented in Panessa et al. (2006). We also show our $z <$ 0.4 optical AGN comparison sample (see Section 3.4).

to X-ray luminosity we have assumed that the non-emission-line detected [O III] targets fall into the bottom 50 per cent of this distribution. Furthermore, to give the quoted range, we have assumed that they intrinsically fall in the bottom 17 per cent of the $L_{[O III]}/L_X$ distribution. These are not unreasonable assumptions given that all but one of these targets have 3σ upper limits that fall very close to, or below, this boundary (see Fig. 5). We note that if we make our [O III] luminosity measurements using a 2 arcsec aperture (see Section 3.1) we find a median aperture correction of 0.28 dex with a scatter of 0.09 dex. Therefore, using 'total' luminosities could increase our quoted $L_{[O III]}/L_X$ ratio by ≈ 0.3 dex (see Fig. 5). Overall, our results indicate that, typically, [O III] luminosities are ≈ 1 per cent of the X-ray luminosities, with a factor of 2–3 scatter, for $z \approx 1.1$–1.7 X-ray detected AGN.

To compare to low-redshift AGN, in Fig. 5, we show the objects from our $z < 0.4$ AGN comparison sample (see Section 3.4) and the relationship found for local Seyfert galaxies and QSOs by Panessa et al. (2006). The KASHz targets appear to broadly cover the same region of the $L_{[O\,III]} - L_X$ parameter space as the $z < 0.4$ AGN, but with a slight tendency towards lower [O III] luminosities. This may be partly due to the different approaches used (or not used) to aperture-correct the luminosities or a lack of correction for reddening to the [O III] emission-line measurements. Furthermore, the $z < 0.4$ AGN are initially drawn from an optically selected SDSS sample, which is [O III] flux limited (see Section 3.4), while, in contrast, the KASHz is initially X-ray selected. Indeed, a plot of [O III] flux versus X-ray flux reveals that the deficit of low $L_{[O\,III]}/L_X$ luminosity ratios in the $z < 0.4$ sample, observed in Fig. 5, is at least partly driven by the [O III] flux limit. Heckman et al. (2005) and Lamastra et al. (2009) present further discussion on the difference between optically selected and X-ray selected samples.

The majority of the KASHz [O III] targets cover a narrow X-ray luminosity range (i.e. $L_X = 2 \times 10^{43}$–2×10^{44} erg s^{-1}; see Fig. 5); therefore, we do not attempt to fit a luminosity-dependent relationship between L_X and $L_{[O\,III]}$. However, we note that the local L_X–$L_{[O\,III]}$ relationship is found to be close to linear in log–log space, when measured over five orders of magnitude in X-ray luminosity [i.e. log $L_{[O\,III]} \propto (0.82 \pm 0.08)$ log L_X; Panessa et al. 2006; also see Lamastra et al. 2009 for an even more linear relationship for a combined sample of optical and X-ray selected AGN]. At the median X-ray luminosity of our [O III] KASHz sample, i.e. log (L_X/erg s^{-1}) $= 43.7$, the local relationship from Panessa et al. (2006) results in a luminosity ratio of log($L_{[O\,III]}/L_X$) $= -1.86$, which is consistent with our KASHz value of $-2.1^{+0.3}_{-0.5}$. Therefore, based on these data, we have no reason to conclude that there is any significant difference between the L_X–$L_{[O\,III]}$ relationship for our $z \approx 1.1$–1.7 sample compared to local AGN.

In Fig. 5 we observe a large scatter of a factor of ≈ 3. This is similar to the scatter seen in local X-ray selected AGN by Heckman et al. (2005). The lack of a correction for X-ray obscuration may account for some of this scatter; however, we note that the obscuration correction to hard-band (2–10 keV) luminosities, at these redshifts, will be small except in the most extreme cases (e.g. Alexander et al. 2008; although see LaMassa et al. 2009). A more significant factor on the amount of scatter will be the effect of various amounts of dust reddening, in the host galaxy, that will affect the [O III] emission-line luminosities (see e.g. Netzer et al. 2006; Panessa et al. 2006). Due to the lack of sufficient constraints on this reddening effect across the sample, we do not attempt to correct for this effect. A more interesting interpretation of the larger scatter in this diagram is a possible lack of uniformity in the amount of gas inside the AGN ionization fields due to variations in opening angles and inclinations, with respect to the host-galaxy gas. These effects will lead to different volumes of gas being photoionized by the central AGN. One further possible interpretation for the large scatter observed in Fig. 5 is the different time-scales of the accretion rates being traced by X-ray versus [O III] luminosity. Whilst the X-ray emission primarily traces nuclear activity associated with the region very close to the accretion disc (e.g. Pringle 1981), the [O III] emission is found in NLRs, which can extend on ≈ 0.1–10 kpc scales (e.g. Boroson, Persson & Oke 1985; Osterbrock & Ferland 2006; Hainline et al. 2013). Therefore, a NLR that is photoionized by an AGN could act as an isotropic tracer of the time-averaged bolometric AGN luminosity over $\approx 10^4$ yr, while in contrast, the X-ray luminosity is an instantaneous measurement of the bolometric AGN luminosity (see e.g. Hickox et al. 2014; Berney et al. 2015; Schawinski et al. 2015

for further discussion on the relative time-scales of different AGN tracers).

4.2.2 Hα luminosity versus X-ray luminosity

The Hα emission-line luminosity is a very common tracer of star formation rates in local and high-redshift galaxies (e.g. see Kennicutt 1998; Calzetti 2013). However, this is complicated in AGN host-galaxies for two main reasons. First, Hα emission is also produced in the sub-parsec scale BLRs around SMBHs, which are most likely to be photoionized by the central AGN and, secondly, the gas in the kpc-scale NLR of AGN can also be photoionized by the AGN (e.g. Osterbrock & Ferland 2006). The relative contribution of AGN versus star formation to producing the total Hα luminosity ($L_{H\alpha}$) will significantly affect how well correlated $L_{H\alpha}$ is with tracers of bolometric AGN luminosity, such as the hard-band X-ray luminosity (L_X). In this sub-section we compare the Hα and X-ray luminosities for our final KASHz sample of 34 $z \approx 0.6$–1.1 AGN that were targeted to observe Hα emission (see Section 4.1).

In Fig. 6 we show $L_{H\alpha}$ versus L_X for our KASHz targets. In the middle and bottom panels we separate the Hα luminosity into that associated with the NLRs (i.e. $L_{H\alpha,NLR}$) and that associated with the BLRs (i.e. $L_{H\alpha,BLR}$), respectively (see Section 3.2 for details of how we separate these components). Following from our [O III] analyses in the previous section, we do not attempt to correct for X-ray obscuration or dust reddening (although see discussion below). In Fig. 6 we also show our $z < 0.4$ AGN comparison sample (see Section 3.4) and the relationships found for local AGN and QSOs by Panessa et al. (2006). We note that if we make our Hα NLR luminosity measurements using a 2 arcsec aperture (see Section 3.1), we find a median aperture correction of 0.25 dex with a scatter of 0.08 dex.

In the bottom panel of Fig. 6, it can be seen that there is a correlation between BLR region luminosity, $L_{H\alpha,BLR}$, and L_X for the $z < 0.4$ AGN. These sources broadly follow the local relationship observed for the Type 1 Seyferts and QSOs by Panessa et al. (2006), i.e. log $L_{H\alpha,T1} \propto (1.16 \pm 0.07)$log L_X. Although we are limited to seven BLR sources for the KASHz sample, our targets are found to be within the scatter of the $z < 0.4$ AGN, and therefore they qualitatively follow the same relationship as low-redshift and local AGN (see Fig. 6). It is not surprising that the BLR luminosities are tightly correlated with the X-ray luminosities (see Fig. 6). The BLR is directly illuminated by the central AGN, which is also responsible for the production of X-rays around the accretion disc (e.g. Pringle 1981; Osterbrock & Ferland 2006). Furthermore, due to the size scales of the BLR that are typically on light-days to light-months (e.g. Peterson et al. 2004), the relative time-scales of the accretion events being probed by the X-ray emission and BLR emission will be much closer compared to the relative time-scales between the X-ray emission and the kpc-scale NLR emission (see the above discussion for the [O III] NLR emission).

In contrast to the BLR luminosities, there is little-to-no evidence for a correlation observed between the NLR Hα luminosity and X-ray luminosity for both the KASHz sample and the $z < 0.4$ comparison sample (see the middle panel of Fig. 6). In qualitative agreement with this for local AGN, the relationship is less steep for Type 2 AGN compared to Type 1 AGN and QSOs by Panessa et al. (2006), i.e. log $L_{H\alpha,T2} \propto (0.78 \pm 0.09)$log L_X. However, some correlation does still exist in local samples, whereas we do not see evidence for this in our current high-redshift sample. In addition to the time-scale arguments already discussed, this may be due to

additional contributions to the NLR Hα luminosities, in addition to photoionization by AGN, such as by star formation processes (e.g. Cresci et al. 2015). Indeed, we observe that a significant fraction of our sample have [N II]/Hα emission-line ratios that could be produced by H II regions (see Section 4.3.3). The NLR region Hα luminosities of the KASHz targets may follow an even shallower trend than the local relationship; however, the deviation is only observed at the highest X-ray luminosities, where we are currently limited to a lower number of sources (see Fig. 6). We also note that the KASHz AGN are likely to be systematically low in $L_{H\alpha, NLR}$ due to the lack of obscuration correction (also see the discussion on aperture effects above). For example, the correction to the NLR luminosities would be ≈0.7 dex, assuming the median $A_{V, H\alpha} = 1.7$ for the star-forming galaxy comparison sample at the same redshift (Stott et al., in preparation; Section 3.4), while the average Balmer decrement of the $z < 0.4$ AGN comparison sample implies a median correction of only ≈0.3 dex for the $z < 0.4$ AGN (following Calzetti et al. 2000). Additionally, our X-ray selection compared to the optically selected comparison samples may also provide a systematic effect towards lower line luminosities for the high-redshift sources (see discussion above for the [O III] targets).

4.3 The prevalence and drivers of ionized outflows

A key aspect of KASHz is to constrain the prevalence of ionized outflow features observed in the emission-line profiles of high-redshift AGN. Additionally, KASHz is designed to assess which AGN and host-galaxy properties are associated with the highest prevalence of high-velocity outflows. One of the most common approaches to search for ionized outflows is to look for very broad and/or asymmetric emission-line profiles in the ionized gas species such as [O III] and non-BLR Hα components (e.g. Heckman et al. 1981; Veilleux 1991; Mullaney et al. 2013; Zakamska & Greene 2014; Genzel et al. 2014). For example, asymmetric emission-line profiles (most commonly a blue wing) that reach high velocities (i.e. ≈1000 km s^{-1}) are very difficult to explain other than through outflowing material (e.g.Veilleux 1991; Zakamska & Greene 2014). Furthermore, extremely broad emission-line profiles (i.e., $W_{80} \gtrsim$ 600 km s^{-1}) are very unlikely to be the result of galaxy kinematics, but instead trace outflows or high levels of turbulence (e.g. Vega Beltrán et al. 2001; Collet et al. 2015; also see discussion in Section 4.3.3) and studies of large samples of low-redshift AGN have shown that the gas that is producing such broad emission-line profiles is not in dynamical equilibrium with their host galaxies (see Liu et al. 2013; Zakamska & Greene 2014).

In the following sub-sections, we assess the prevalence and drivers of ionized outflow features in the galaxy-integrated emission-line profiles of our KASHz AGN sample, following similar methods to Mullaney et al. (2013) who study $z < 0.4$ AGN (see Section 3.4). More specifically, we investigate the distributions of emission-line velocity-widths of individual sources, in combination with emission-line profile stacks. For clarity and ease of comparison to previous studies, we separate the discussion of the $z \approx 1.1$–1.7 [O III] sample from the $z \approx 0.6$–1.1 Hα sample (these are defined in Section 4.1). Furthermore, the [O III] emitting gas is more likely to be dominated by AGN illumination, while the Hα emission may also have a significant contribution from star formation (see discussion in Section 4.3.3). We defer a detailed comparison of these two ionized gas tracers to future papers, which will be based on spatially resolved spectroscopy using both emission lines for the same targets; however, see the works of Cano-Díaz et al. (2012)
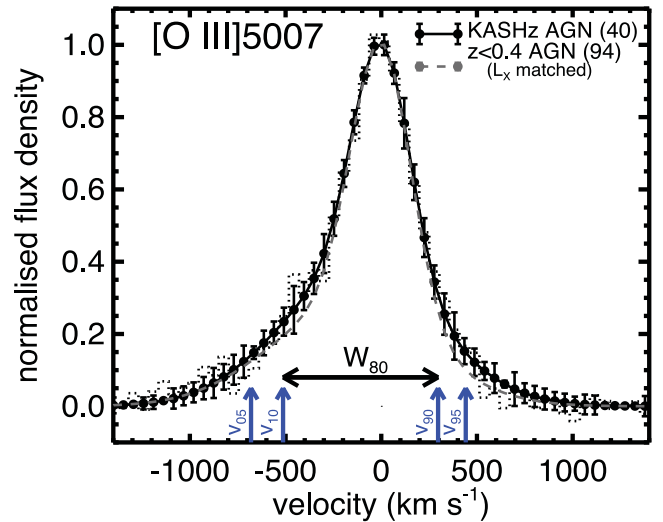
Figure 7. Stacked [O III]5007 emission-line profiles for the 40 [O III] detected KASHz targets and the X-ray luminosity matched $z < 0.4$ AGN comparison sample (see Section 3.4). The dotted curves show the stacked data and the dashed and solid curves are fits to these stacks. The upward arrows show, from left-to-right, the fifth, 10th, 90th and 95th percentile velocities of the KASHz stack. The overall emission-line width of $W_{80} = 810$ km s^{-1} is also illustrated (see Section 3.3). On average, the KASHz AGN show a broad and asymmetric emission-line profile, with velocities reaching ≈1000 km s^{-1}. The low-redshift AGN have a very similar average emission-line profile to the high-redshift AGN for these luminosity-matched samples.

and Cresci et al. (2015) for IFS data covering both Hα and [O III] measurements for two high-redshift AGN.

4.3.1 The distribution of [O III] emission-line velocity-widths

In Fig. A1 we show the [O III] emission-line profiles, and our best-fitting solutions, for all the $z \approx 1.1$–1.7 KASHz targets. The parameters of all of the fits are provided in Table A1. We identify secondary broad components (following the methods described in Section 3.2), with FWHM≈400–1400 km s^{-1}, in the emission-line profiles for 14 out of the 40 [O III] detected targets (i.e. 35 per cent). The velocity offsets of these broad components, with respect to the narrow components, reach up to $|\Delta v| \approx 500$ km s^{-1}. We note that Brusa et al. (2015) find that four out of their eight $z \approx 1.5$ X-ray luminous AGN identify a secondary broad emission-line component at high significance in their [O III] spectra, which is broadly consistent with our fraction given the low numbers involved. While the fraction of broad emission-line components in our KASHz sample already indicates a high prevalence of high-velocity ionized gas in high-redshift X-ray AGN, these measurements do not provide a complete picture across all of the targets. This is because it is very difficult to detect multiple Gaussian components when the signal-to-noise ratio is modest, i.e. the detection of a second Gaussian component is limited to the highest signal-to-noise ratio spectra. Therefore, in the following discussion, we follow two methods to overcome these challenges. First we assess the average emission-line profiles using stacking analysis (see Section 3.5) and, secondly, we use a non-parametric definition to characterize the *overall* line width (i.e. W_{80} which is the width that encloses 80 per cent of the emission-line flux; see Section 3.3).

We show the stacked [O III]5007 emission-line profile for our KASHz targets in Fig. 7. The overall-emission line width of this

average profile is $W_{80} = 810^{+130}_{-220}$ km s^{-1} (see Section 3.3), where the upper and lower bounds indicate the 68 per cent range from bootstrap resampling that stack. This indicates that, on average, the ionized gas in these AGN has high-velocity kinematics that are not associated with the host-galaxy dynamics. Furthermore, the average emission-line profile clearly shows a luminous blueshifted broad wing, which reveals high-velocity outflowing ionized gas out to velocities of ≈ 1000 km s^{-1}. A preference for blueshifted broad wings, compared to redshifted broad wings, has been previously observed for both high- and low-redshift AGN samples (e.g. Heckman et al. 1981; Vrtilek 1985; Veilleux 1991; Harrison et al. 2012; Mullaney et al. 2013; Zakamska & Greene 2014; Balmaverde et al. 2015). This can be explained if the far-side of any outflowing gas, that is moving away from the line of sight, is obscured by dust in the host galaxies (e.g. Heckman et al. 1981; Vrtilek 1985). To observe a redshifted component, the outflow would need to be highly inclined or extended beyond the obscuring material (e.g. Barth, Greene & Ho 2008; Crenshaw et al. 2010; Harrison et al. 2012; Rupke & Veilleux 2013).

While informative, the average emission-line profile shown in Fig. 7 hides critical information on the underlying distribution of ionized gas kinematics in our sample. Therefore, in Fig. 8 we show the distribution of the [O III] velocity-width values, $W_{80,[\mathrm{O\,III}]}$, for the individually [O III] detected targets. In the top panel we show the raw distribution and in the bottom panel we show the cumulative distribution. 1σ uncertainties on the cumulative distribution have been calculated assuming Poisson errors, suitable for small number statistics, following the analytical expressions provided by Gehrels (1986). We use the same uncertainty calculations for all of the percentages presented for the remainder of this section. We find that 50^{+14}_{-11} per cent of the [O III] targets have velocity-widths indicative of tracing ionized outflows or highly turbulent material, i.e. $W_{80,[\mathrm{O\,III}]} > 600$ km s^{-1} (e.g. Liu et al. 2013; Collet et al. 2015). This fraction ranges over (42–58) per cent, when including the eight non-detected targets (see Section 4.1), for which we have no $W_{80,[\mathrm{O\,III}]}$ measurements.

To compare the prevalence of outflow features in our high-redshift AGN sample to low-redshift AGN, we show the distribution of velocity-widths for our $z < 0.4$ luminosity-matched comparison samples in Fig. 8 (see Section 3.4). The luminosity-matching is performed to account for the observed correlation between luminosity and emission-line widths (Mullaney et al. 2013; also see Section 4.3.2). We note that we obtain the same conclusions if we match by either $L_{[\mathrm{O\,III}]}$ or L_X. In Fig. 8, the [O III] line width distributions look indistinguishable between the KASHz sample and the luminosity-matched $z < 0.4$ AGN samples. A two-sided KS test shows no evidence that the two different redshifts have significantly different velocity-width distributions (i.e. an ≈ 60 per cent chance that the two redshift ranges have velocity-width values drawn from the same distribution). In Fig. 7, we further demonstrate that luminosity-matched AGN, from both redshift-ranges, have similar [O III] emission-line profiles by showing that the stacked emission-line profiles look the same. This result implies that the prevalence of ionized outflows in low-redshift and high-redshift AGN is very similar for AGN of the same luminosity. This is despite the fact that the average star formation rates of X-ray AGN are ≈ 10 times higher at $z \approx 1$ compared to $z \approx 0.2$, irrespective of X-ray luminosity (Stanley et al. 2015). Therefore, this result provides indirect evidence that the prevalence of these high-velocity outflows is not greatly influenced by the level of star formation. We draw similar conclusions for the Hα KASHz sample in Section 4.3.3.
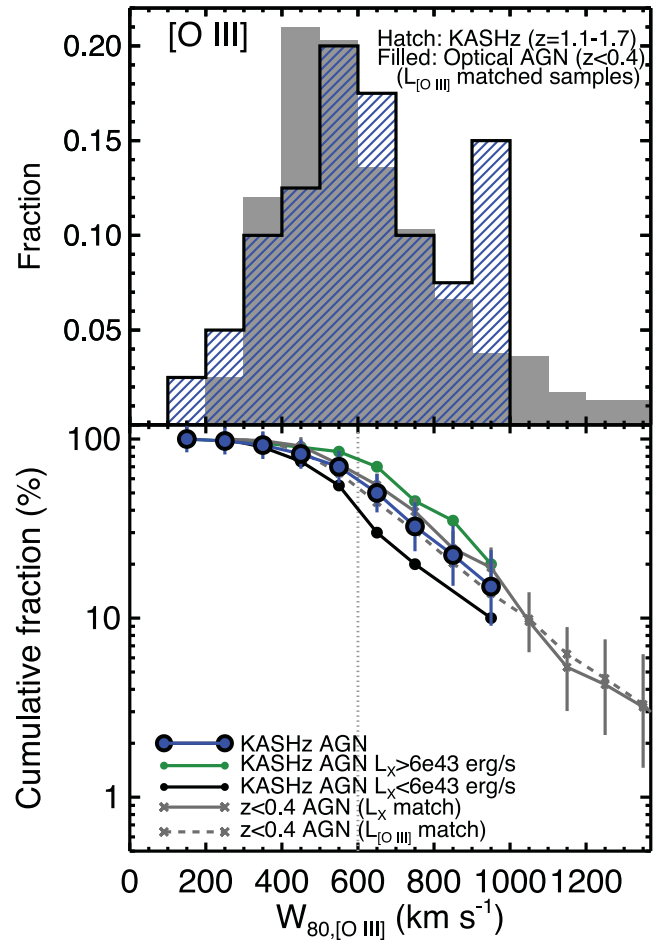
Figure 8. Top: histograms of the overall emission-line velocity-width, $W_{80,[\mathrm{O\,III}]}$, for the 40 $z \approx 1.1$–1.7 [O III] detected KASHz targets (hatched) and a luminosity-matched sample of 1000 $z < 0.4$ optical AGN (filled; see Section 3.4). The KASHz AGN show a very similar distribution of velocities as the luminosity-matched low-redshift AGN sample. This is further demonstrated in the bottom panel which shows the cumulative distributions. We split the KASHz sample in half, separating at $L_X = 6 \times 10^{43}$ erg s^{-1}, and find that high-velocity gas kinematics are more prevalent in higher luminosity AGN; for example, it is $\gtrsim 2$ times more likely that the higher luminosity AGN have $W_{80} \gtrsim 600$ km s^{-1}, compared to the lower-luminosity AGN (see vertical dotted line; Section 4.3.2).

Using our representative targets (see Fig. 2 and Section 2.3), it is now possible to place some previous observations of high-redshift X-ray AGN into the context of the overall population. For example, the SINFONI observations of XCOS-2028 were used by Cresci et al. (2015) to present evidence of star formation suppression by the outflow observed in this source (see our spectra for this source in Fig. A1.). This conclusion was based on their observed deficit of NLR Hα emission, a possible tracer of star formation (see Sections 4.2 and 4.3.3), in the location of the outflowing [O III] component. Based on our analyses, this source has an emission-line velocity-width of $W_{80} \approx 730$ km s^{-1}. We find ≈ 30 per cent of the overall X-ray AGN population have these emission-line widths or greater. Therefore, this target does not have an exceptional outflow; however, it is yet to be determined how common the observed deficit of NLR Hα emission in this source is in the overall high-redshift AGN population.
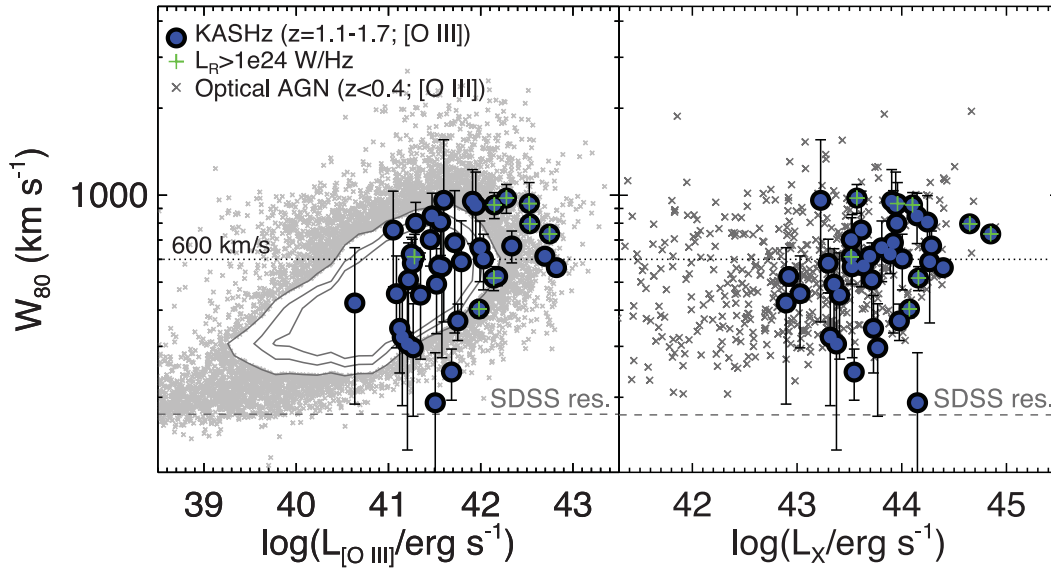
Figure 9. Emission-line velocity-width (W_{80}) versus [O III] luminosity (left) and hard-band (2–10 keV) X-ray luminosity (right) for the $z \approx 1.1$–1.7 KASHz sample (circles) and the $z < 0.4$ AGN comparison sample (contours and crosses). The dashed lines show the W_{80} value for the median spectral resolution for the $z < 0.4$ sources. Although a wide distribution of velocities is observed, the most luminous AGN (either by line luminosity or X-ray luminosity) preferentially host extreme gas velocities. For example, all of the KASHz AGN with $L_{\text{[O III]}} > 10^{42.2}$ erg s^{-1} have velocity-widths of $\gtrsim 600$ km s^{-1} (see dotted line).

*4.3.2 The physical drivers of the high-velocity outflows observed in [O III]*

It is of fundamental importance to constrain how AGN and host-galaxy properties are related to the prevalence and properties of galaxy-wide outflows. For example, in most cosmological models, the accretion rates of the AGN fundamentally determine the velocities and energetics of the outflows (e.g. Schaye et al. 2015), while some models, concentrating on individual sources, have invoked the mechanical output from radio jets as a plausible outflow driving mechanism (e.g. Wagner, Bicknell & Umemura 2012). In this subsection we investigate the role of AGN luminosity, radio luminosity and X-ray obscuration on the prevalence of ionized outflows in our KASHz [O III] sample.

In Fig. 9, we plot the emission-line velocity-width ($W_{80,\text{[O III]}}$) as a function of [O III] luminosity and X-ray luminosity for the [O III] detected KASHz targets. Both [O III] luminosity and X-ray luminosity may serve as a tracer for the bolometric AGN luminosity, potentially on different time-scales (see Section 4.2). We find that the most luminous AGN (both based on [O III] and X-ray) are associated with the highest velocities, although a large spread in velocities is seen at lower luminosities. These same trends are also seen in our low-redshift comparison sample (see crosses in Fig. 9). To quantify the effect of AGN luminosity on the prevalence of ionized outflows, we split the [O III] detected targets into two halves by taking the 20 'lower' and 20 'higher' X-ray luminosity targets, resulting in a luminosity threshold of $L_X = 6 \times 10^{43}$ erg s^{-1} between the two subsets. In Fig. 8 we show the cumulative distributions of line widths for these two subsets. It can be seen that there is higher probability of observing extreme ionized gas velocities in the 'higher' luminosity sub-set. For example, 70^{+24}_{-18} per cent of the 'higher' luminosity targets have line widths of $W_{80} > 600$ km s^{-1}, while only 30^{+18}_{-12} per cent of the 'low' luminosity targets reach these line widths. A two-sided KS test indicates only a ≈2 per cent chance that the two luminosity bins have velocity-width values drawn from the same distribution. In future papers, we will be able to test this result to higher significance as the KASHz sample increases.

To first order, the results described above indicate that the highest outflow velocities are associated with the most powerful AGN. This result has been quoted throughout the literature for low-redshift AGN, for both ionized outflows and for molecular outflows (e.g. Westmoquette et al. 2012; Veilleux et al. 2013; Arribas et al. 2014; Cicone et al. 2014; Hill & Zakamska 2014). However, in their study of ionized outflows of ≈24 000 AGN, Mullaney et al. (2013) found that the highest velocity outflows are more fundamentally driven by the mechanical radio luminosity ($L_{1.4\,\text{GHz}}$) of the AGN, rather than the radiative (i.e. [O III]) luminosity. This result could either be an indication that small-scale radio jets are driving high-velocity outflows, as observed in spatially resolved studies of some low-redshift AGN (e.g. Morganti et al. 2005, 2013; Tadhunter et al. 2014), or that radiatively driven outflows are producing shocks in the ISM which result in the production of radio emission (Zakamska & Greene 2014; Nims, Quataert & Faucher-Giguère 2015; also see Harrison et al. 2015). For our KASHz targets, we are currently limited to only eight [O III] detected targets which we can define as 'radio luminous' (i.e. $L_{1.4\,\text{GHz}} > 10^{24}$ W Hz^{-1}; see Section 2.5). A higher fraction of the radio luminous sample has high velocity line widths of $W_{80} > 600$ km s^{-1} compared to the non-radio-luminous sources, i.e. 6/8 or 75^{+25}_{-30} per cent compared to 13/30 or 43^{+16}_{-12} per cent (see Fig. 9);[8] however, the uncertainties on these fractions are high. Furthermore, the radio luminous sources are preferentially associated with higher X-ray luminosity AGN and we do not currently have sufficient numbers of targets to control for this (see Fig. 9). We note that it *was* possible to control for AGN luminosity, in this case [O III] luminosity, for the low-redshift study of Mullaney et al. (2013). In summary, based on the current sample, we cannot yet determine if the radio luminosity, or the X-ray luminosity, is more fundamental in driving the highest-velocity outflows observed in high-redshift AGN.

---

[8] We note that two of our [O III] detected targets do not have the required radio constraints to classify them as either radio luminous or not (see Section 2.5).
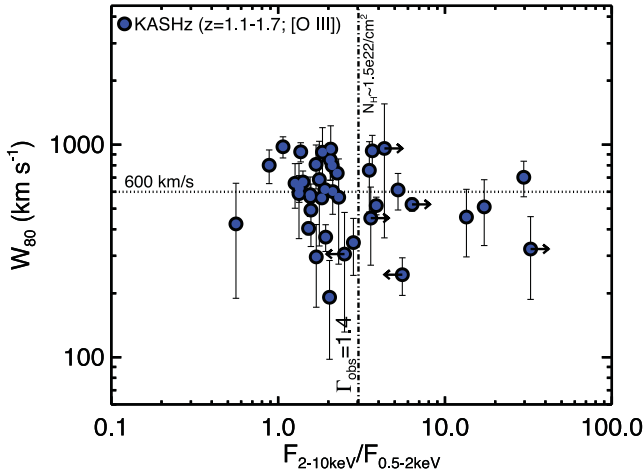
Figure 10. Emission-line width (W_{80}) versus the ratio between 2–10 keV and 0.5–2 keV X-ray fluxes for the 40 [O III] detected KASHz sample (circles). The X-ray axis here serves as a proxy for observed photon index (Γ_{obs}), and also as an estimate for obscuring column density (for an assumed intrinsic photon index and known redshift). We classify the sources with $F_{2-10\,keV}/F_{0.5-2\,keV} > 3.03$ as 'obscured candidates' [corresponding to $N_H \gtrsim (1-2) \times 10^{22}$ cm^{-2}; see Section 2.4]. We find no evidence that high-velocity outflows are preferentially observed in the obscured X-ray AGN.

It has been suggested that massive galaxies may go through an evolutionary sequence where obscured AGN reside in star-forming galaxies during periods of rapid SMBH growth and galaxy growth during which the AGN drive outflows that drive away the enshrouding material to eventually reveal an unobscured AGN (e.g. Sanders et al. 1988; Hopkins et al. 2006). Therefore, it may be expected that outflows are more preferentially associated with X-ray obscured AGN. To test this, we compare the [O III] emission-line velocities of our X-ray 'obscured candidates' to those 'unobscured candidates', which we separate at an obscuring column density of $N_H \approx 2 \times 10^{22}$ cm^{-2} using a simple hard-band to soft-band flux ratio technique (see Section 2.4). In Fig. 10 we show the emission-line velocity-width, $W_{80,[O III]}$, as a function of this flux ratio. We find no significant difference between the two populations; that is, we find that 45^{+20}_{-9} per cent of these sources have $W_{80,[O III]} > 600$ km s^{-1}, compared to a similar fraction of 54^{+15}_{-14} per cent for the unobscured sources. This result implies that high-velocity ionized outflows are certainly not uniquely, and do not appear to be preferentially, associated with X-ray obscured AGN. This may be in disagreement with some evolutionary scenarios for galaxy evolution; however, there may be a population of heavily obscured AGN that are missed from even the deepest X-ray surveys (see Section 2.3) that are not present in the current sample.

4.3.3 The prevalence and drivers of outflows observed in Hα

In Figs A2 and A3 we show the Hα+[N II] emission-line profiles, and our best-fitting solutions, for all targets in the current $z \approx 0.6$–1.1 KASHz sample. The parameters for all of the fits are provided in Table A1. We detected 32 out of the 34 targets (see Section 4.1) and of these 32, we identified a BLR in seven of the targets (see Fig. A3). For this study, we are interested in the kinematics of the host galaxy, or equivalently the NLR, and not the BLR. Therefore, for these seven BLR sources the quoted emission-line velocity-widths (i.e. the $W_{80,H\alpha}$) are for the NLR components only (see Section 3.2).
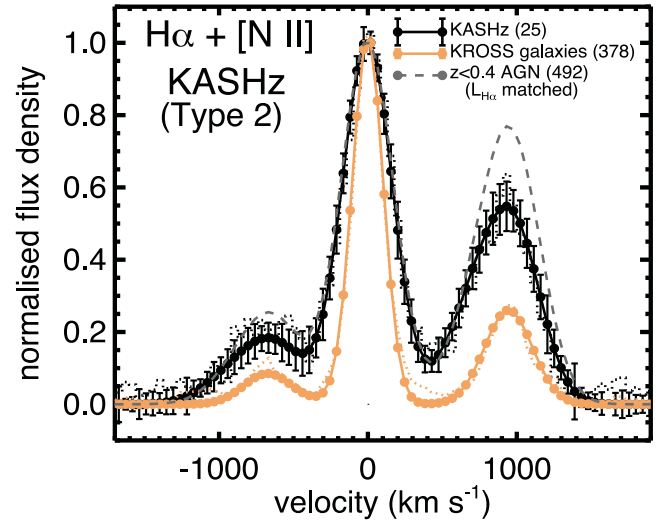
Figure 11. Stacked Hα+[N II] emission-line profiles for the 25 Type 2 Hα detected KASHz targets and the comparison samples of: (1) star-forming galaxies at the same redshift and (2) $z < 0.4$ Type 2 AGN (see Section 3.4). The dotted curves show the stacked data and the dashed and solid and dashed curves are fits to the stacks. The KASHz AGN clearly show much broader emission-line profiles, and higher [N II]/Hα emission-line ratios than the star-forming galaxies.

In Fig. 11, we show the stacked Hα+[N II] emission-line profile for our 25 Hα Type 2 AGN. The overall-emission-line width of this average profile is $W_{80,H\alpha} = 440^{+50}_{-13}$ km s^{-1} (see Section 3.3), where the upper and lower bounds indicate the 68 per cent range from bootstrap resampling that stack. This velocity-width is lower than the $W_{80,[O III]} = 810^{+130}_{-220}$ km s^{-1} observed in the stacked [O III] emission-line profile (Fig. 7). This difference is likely to be due, in part, to the almost order of magnitude difference in the average X-ray luminosity of the Hα sample compared to the [O III] sample (i.e. 2×10^{43} erg s^{-1} compared to 1×10^{44} erg s^{-1}; see Fig. 1), which we have shown be a key driver for the observed prevalence of high-velocity outflows (Section 4.3.2). However, as already briefly mentioned, this may also be due to these two emission lines preferentially tracing different regions of gas in the host galaxy, with Hα likely to have a larger contribution from star-forming regions. While we currently lack systematic IFS studies of AGN covering both [O III] and Hα emission lines, observations of some AGN have already indicated that the NLR Hα emission has a contribution from star-forming regions and is not necessarily dominated by the high-velocity outflows as is the case for [O III] emission (e.g. Cano-Díaz et al. 2012; Cresci et al. 2015). Furthermore, a strong blue wing can be observed in the stacked KASHz emission-line profile for the [N II] emission-line, which has a larger velocity-width than Hα with $W_{80,[N II]} = 590^{+120}_{-50}$ km s^{-1} (Fig. 11). This high-ionization line may be more analogous to the [O III] emission, and preferentially trace outflowing/turbulent material compared to Hα. Indeed, the difficulty in identifying broad-underling outflow components in Hα emission could be that this emission line preferentially traces rotation of the host galaxy and beam-smearing of this emission could dilute centrally located broad components (Förster Schreiber et al. 2014; Genzel et al. 2014). We are unable to de-couple the [N II] from the Hα across the whole Hα sample due to limited signal-to-noise ratios (see Section 3.2) and therefore we may be underestimating the outflow velocities in these sources. However, Hα emission has been used to identify outflows in AGN host-galaxies (e.g. Westmoquette et al. 2012; Arribas et al. 2014; Genzel et al. 2014) and analysis
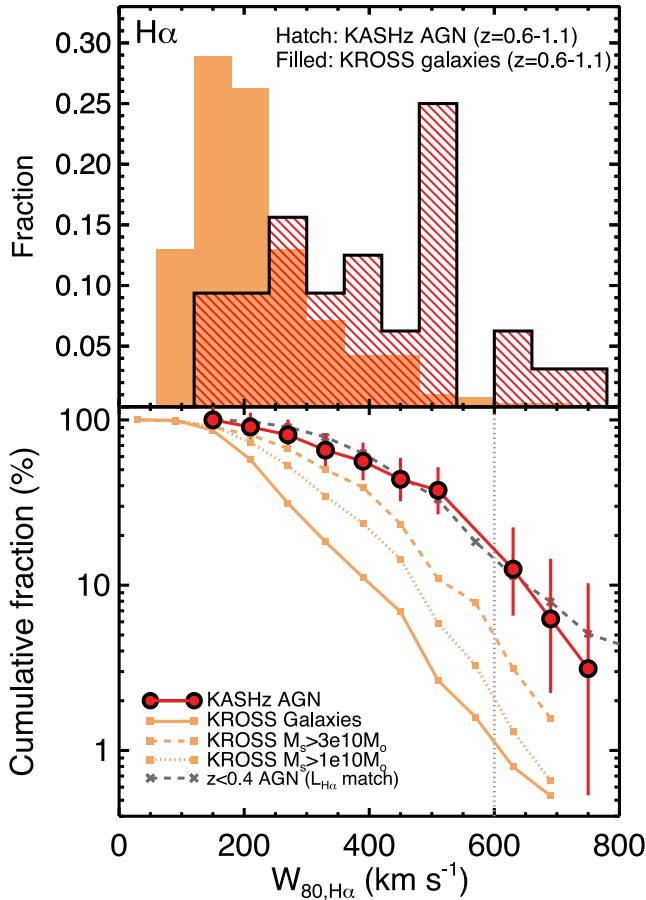
Figure 12. Top: histograms of the overall emission-line velocity-width, $W_{80,H\alpha}$, for the 32 $z \approx 0.6$–1.1, Hα detected KASHz targets (hatched) and the $z \approx 0.6$–1.1 comparison sample of star-forming galaxies from KROSS (filled; see Section 3.4). The AGN preferentially have higher emission-line velocities than the star-forming galaxies. This is further demonstrated in the bottom panel which shows the cumulative distributions. We also show the KROSS sample after applying two stellar mass cuts, where the $M_\star > 3 \times 10^{10} \, \mathrm{M_\odot}$ sub-sample is more comparable to the X-ray AGN host galaxies (Section 4.3.3). We also show the cumulative distribution of the luminosity-matched comparison sample of $z < 0.4$ AGN (see Section 3.4) and find that the velocity distribution is very similar to that of our high-redshift AGN.

of this emission line using these methods provides an informative comparison to other studies.

In Fig. 12 we show the distribution of the Hα line-width values, $W_{80,H\alpha}$, for the 32 detected targets, using the same format as above for the [O III] emission. We find that four of our KASHz targets (i.e. 13^{+9}_{-6} per cent) have $W_{80,H\alpha} > 600$ km s^{-1}, indicative of emission that is dominated by outflowing material (see discussion at the start of this section). We note that the range on this percentage is 12–18 per cent if the two undetected targets are included, for which we have no constraints on $W_{80,H\alpha}$. In agreement with the [O III] results (Section 4.3.1), we find no appreciable difference between the distribution of line widths between the KASHz sample and our low-redshift luminosity-matched comparison sample (see Section 3.4).

In Fig. 12 we compare the Hα emission-line velocity distributions for our KASHz targets with our comparison sample of star-forming galaxies that are at the same redshift from KROSS (see Section 3.4). It can clearly be seen that AGN preferentially have higher emission-line widths than the star-forming galaxies. For example,
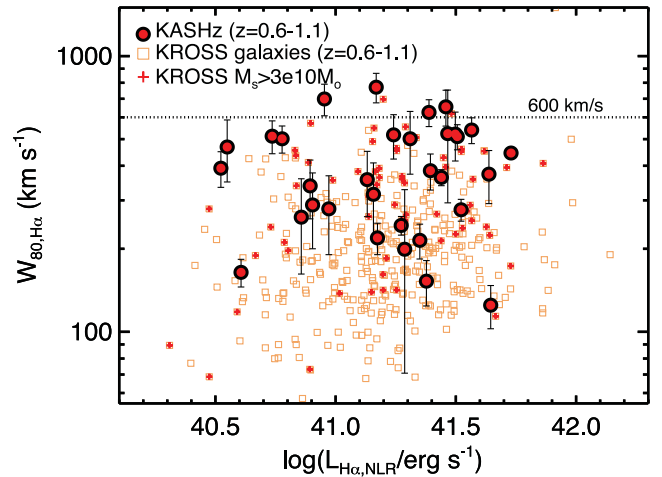
Figure 13. Emission-line velocity-width, $W_{80,H\alpha}$, versus NLR Hα luminosity ($L_{H\alpha,NLR}$) for the KASHz AGN targets (circles) and the KROSS star-forming galaxies at the same redshift (squares). The AGN show a similar distribution of $L_{H\alpha,NLR}$, but preferentially have the higher velocities. $L_{H\alpha,NLR}$ is a tracer of the star formation rates for the galaxies and the AGN (although this will be biased upwards for the AGN; see Section 4.3.3). Therefore, the higher emission-line velocities in the AGN are not the result of higher star formation rates.

only 3/378 (i.e. $0.8^{+0.8}_{-0.4}$ per cent) of the star-forming galaxy sample reach velocity-widths of $W_{80,H\alpha} > 600$ km s^{-1}, compared to 13^{+9}_{-6} per cent found for the AGN. This is also demonstrated in Fig. 11, where we show the stacked emission-line profiles for both samples. These results provide indirect evidence that the high-velocity features we observe are not pre-dominantly driven by star formation. This is because X-ray AGN at these redshifts have average star formation rates that are consistent with the global star-forming galaxy population of the same redshift (e.g. Rosario et al. 2012; Stanley et al. 2015), and are possibly even distributed to lower median star formation rates (Mullaney et al. 2015). We further test this conclusion by plotting the velocity-width as a function of observed NLR Hα luminosity ($L_{H\alpha,NLR}$) for both the KASHz sample and star-forming comparison sample in Fig. 13. The observed NLR Hα luminosity (i.e. excluding any BLR components) is a tracer of the star formation rates in star-forming galaxies; however, for the AGN-host galaxies these will, in general, be relative over-estimates due to the additional contribution of photoionization by the central AGN. We find that the KASHz AGN have a very similar distribution of $L_{H\alpha,NLR}$ to our star-forming galaxy comparison sample, indicating that they have similar, or possibly lower, star formation rates. Despite this, the AGN-host galaxies preferentially have larger velocity-widths, and a higher fraction of sources indicative of hosting high-velocity ionized outflows (Fig. 13).

In addition to investigating the role of the star formation rates, it is also important to consider the effect of host-galaxy masses on the Hα emission-line widths. In the cases where the Hα emission-line profiles are dominated by galaxy kinematics (which is the case for the majority of the star-forming KROSS galaxies; e.g. Stott et al., in preparation; Swinbank et al., in preparation), the line widths will be driven to higher values in galaxies with higher stellar masses due to the increased velocity gradients across these galaxies. To demonstrate this, in Fig. 12, we also show the cumulative distributions of line widths for the star-forming galaxies when we apply increasing mass cuts. As expected, the higher mass galaxies tend to have the broader emission-line widths. We do not attempt to derive stellar
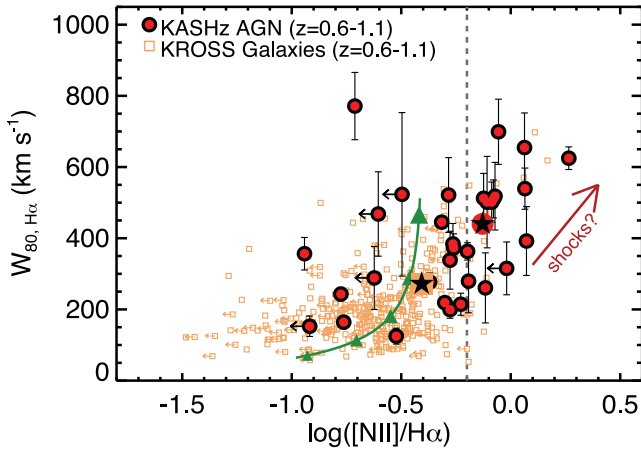
Figure 14. Emission-line velocity-width, $W_{80,H\alpha}$, versus $\log([N\,II]/H\alpha_{NLR})$ emission-line ratio for the KASHz AGN targets (circles) and the KROSS star-forming galaxies (squares). The larger symbols containing the stars are measured from the stacked emission-line profiles (Fig. 11). The vertical dashed line indicates the maximum emission-line ratio expected for photoionized H II regions (e.g. Kewley et al. 2013). The green track shows the predicted trends as a function of mass (following the mass–metallicity and mass–velocity dispersion relations; see Section 4.3.3), where the triangles highlight various $\log(M_\star)$ values in half dex bins starting at 9.0. The positive correlation observed for most of the AGN (i.e. those with $\log([N\,II]/H\alpha) \gtrsim -0.5$ may be indicative of outflows/shocks driving the velocity-widths (see Section 4.3.3). Most of the star-forming galaxies roughly follow the predicted mass-driven trend (Section 3.4) and the stacked average is consistent with the mass-driven prediction at their average mass, i.e. $\log(M_\star/M_\odot) = 10.3$. However, they appear to follow the same trend as the AGN at the largest velocity-widths ($W_{80} \gtrsim 400$ km s^{-1}).

masses for our AGN, due to the variable quality of photometric data sets available for our targets; however, we note that X-ray AGN appear to typically have stellar masses of $\gtrsim 3 \times 10^{10}$ M$_\odot$ (e.g. Bongiorno et al. 2012; Mullaney et al. 2012; Aird et al. 2013; Azadi et al. 2015). Therefore, in Fig. 12, we compare the star-forming galaxy sample, but limiting it to galaxies with stellar masses $>3 \times 10^{10}$ M$_\odot$. We still find that the KASHz AGN have a higher prevalence of the highest velocity-widths, compared to this higher mass subset of the star-forming galaxies, with only 3^{+6}_{-1} per cent exhibiting emission-line velocity-widths of $W_{80,H\alpha} > 600$ km s^{-1}.

To further test the possible role of mass in driving the high emission-line velocity-widths observed in our targets, in Fig. 14 we plot $W_{80,H\alpha}$ as a function of the emission-line ratio $\log([N\,II]/H\alpha)$. This emission-line ratio is a tracer of the metallicity of star-forming galaxies (e.g. Alloin et al. 1979; Denicoló, Terlevich & Terlevich 2002; Kewley & Dopita 2002), as well as an indicator for the source of ionizing radiation (e.g. Kewley et al. 2006; Rich et al. 2014). Furthermore, there is an observed relationship between mass and metallicity and therefore an expected relationship between this emission-line ratio and stellar mass (e.g. Lequeux et al. 1979; Tremonti et al. 2004; Maiolino et al. 2008; Stott et al. 2013). We can also make a crude prediction for the relationship between $W_{80,H\alpha}$ and mass, under the assumption that the line width is a tracer of the stellar velocity dispersion. Therefore, for a given stellar mass, we predict the position galaxies would be located in Fig. 14 by combining: (1) the observed $z = 0.7$ mass–metallicity relation (following Maiolino et al. 2008) and (2) the observed mass–velocity dispersion relationship for massive galaxies (following Bezanson, Franx & van Dokkum 2015). The majority of the star-forming galaxies appear to broadly follow the expected trend. Furthermore, the measure-

ments from the stacked average emission-line profile (Fig. 11) are in agreement with the rough mass-driven prediction for the average mass of these galaxies [i.e. $\log(M_\star/M_\odot) = 10.3$; see Fig. 14]. In contrast, the AGN typically have higher $\log([N\,II]/H\alpha)$ emission-line ratios and higher velocity-widths than the star-forming galaxy sample (also visible in the stacked profiles; Fig. 11) and a positive correlation is observed between these two quantities. Such a positive correlation has been shown to be a tracer of shocks and outflows in the ISM through IFS observations of AGN and star-forming galaxies (e.g. Ho et al. 2014; McElroy et al. 2015). Interestingly, the small number of star-forming galaxies with high line widths (i.e. $W_{80} \gtrsim 400$ km s^{-1}) appear to follow the same relationship as the KASHz AGN, which may indicate that these galaxies also have a contribution from shocks and/or host AGN that were not detected in the X-ray surveys. We clarify that some of the AGN targets have $\log([N\,II]/H\alpha)$ emission-line ratios that could also be photoionized by H II regions (see Fig. 14) and we will explore these ideas further when exploring the spatially resolved outflow kinematics and emission-line flux ratios of the KASHz AGN in future papers.

We conclude that the highest Hα line widths (i.e. those with $W_{80} > 600$ km s^{-1} and possibly those with $W_{80} \gtrsim 400$ km s^{-1}) are at least partially driven by ionized outflow kinematics and/or shocks in the ISM. Furthermore, there is an increased likelihood to find these highest ionized gas velocities in X-ray identified AGN compared to star-forming galaxies at the same redshift that have similar star formation rates and masses. This is in agreement with local IFS studies that have found higher-velocity ionized outflows, traced by Hα emission, in star-forming galaxies that host AGN compared to those which do not (e.g. Westmoquette et al. 2012; Arribas et al. 2014).

5 CONCLUSIONS

We have presented the first results of the ongoing KMOS AGN Survey at High redshift (KASHz). The first 89 targets, which are presented here, are high-redshift ($z \approx 0.6$–1.7) X-ray detected AGN, with hard-band (2–10 keV) luminosities in the range $L_X = 10^{42}$–10^{45} erg s^{-1}. The targets have a distribution of X-ray luminosities that are representative of the parent X-ray AGN population. The majority of the targets (79) were observed with KMOS, supplemented with archival SINFONI observations of 10 targets. All of the observations were carried out in the *J* band and we detected 86 (i.e. 97 per cent) of the targets in continuum and/or emission lines. However, for the analyses in this work, we excluded seven un-reliable non-detections (e.g. due to inaccurate photometric redshifts; Section 4.1). This leaves a final sample of 82 targets, for which we have presented results and discussion based on their galaxy-integrated emission-line profiles.

We detected 72 of the final sample (i.e. 88 per cent) in emission lines, of which 40 out of 48 targets were detected in [O III] (the $z \approx 1.1$–1.7 targets) and 32 out of 34 targets were detected in Hα (the $z \approx 0.6$–1.1 targets). We have explored the emission-line luminosities as a function of X-ray luminosity for our targets and by characterizing the individual emission-line profiles and using stacking analyses, we have investigated the prevalence and drivers of the broad and asymmetric emission-line profiles that are indicative of ionized outflows. Our main conclusions are listed below.

(i) We find a median X-ray luminosity to [O III] luminosity ratio of $\log(L_{[O\,III]}/L_X) = -2.1^{+0.3}_{-0.5}$, where the range is roughly the 1σ scatter. The observed relationship between these two quantities for our high-redshift sample is broadly consistent with that found for

low-redshift AGN and local Seyferts and QSOs (see Fig. 5). Our results indicate that the [O III] luminosities are typically ≈1 per cent of the X-ray luminosities for $z \approx 1.1$–1.7 X-ray AGN. The large scatter of a factor of ≈3 may be due to several observational effects including dust reddening, or due to intrinsic physical effects, such as a higher level of variability in the sub-pc-scale production X-ray emission compared to the kpc-scale production of [O III] emission (Section 4.2).

(ii) Our seven Type 1 Hα targets have BLR luminosities ($L_{\mathrm{H\alpha,BLR}}$) that are broadly consistent with the correlation observed between $L_{\mathrm{H\alpha,BLR}}$ and L_X for low-redshift and local AGN. Although limited by small numbers of very luminous sources, we find no evidence for a correlation between the NLR Hα luminosity and L_X (see Fig. 6; Section 4.2).

(iii) High-velocity emission-line features are common in our [O III] sample, with ≈50 per cent of the targets exhibiting velocities indicative of being dominated by outflowing ionized gas or highly turbulent material (i.e. emission-line velocity-widths of $W_{80,[\mathrm{O\,III}]} > 600$ km s^{-1}; see Figs 7 and 8 and Section 4.3.1). On average, the emission-line profiles have a prominent blueshifted wing, implying outflowing material. Outflowing or highly turbulent material that does not dominate the individual emission-line profiles could be even more common.

(iv) The high-velocity [O III] kinematics are more prevalent for targets with higher AGN luminosities. For example, ≈70 per cent of the $L_X > 6 \times 10^{43}$ erg s^{-1} targets have [O III] line widths of $W_{80,[\mathrm{O\,III}]} > 600$ km s^{-1}, while only ≈30 per cent of the $L_X < 6 \times 10^{43}$ erg s^{-1} targets reach these line widths (see Figs 8 and 9). Using our current sample, we are unable to determine the role of radio luminosity in driving this trend (Section 4.3.2).

(v) Based on our current X-ray detected sample, we find no evidence that the highest ionized gas velocities are preferentially associated with X-ray obscured AGN (i.e. those with $N_{\mathrm{H}} \gtrsim 10^{22}$ cm^{-2}), compared to X-ray unobscured AGN (see Fig. 10), in contrast to the predictions of some evolutionary scenarios (Section 4.3.2).

(vi) We compared the emission-line widths of our Hα detected X-ray AGN targets with a redshift matched sample of star-forming galaxies. Despite a similar distribution of Hα luminosities (excluding the BLR components), and implied similar star formation rates, the AGN-host galaxies exhibit a much higher prevalence of high ionized gas velocities (see Figs 11 and 12). For example, ≈13 per cent of the AGN have Hα emission-line widths of $W_{80} > 600$ km s^{-1}, whilst only ≈1 per cent of the star-forming galaxy sample reach these line widths (Section 4.3.3).

(vii) For both the [O III] and Hα KASHz targets, we find no significant difference between the distribution of velocity-widths for our high-redshift AGN sample and luminosity-matched comparison samples of $z < 0.4$ AGN. Under the assumption that the most extreme ionized gas velocities are associated with outflows, to first order, this implies that it is just as likely to find an ionized outflow of a certain velocity in low-redshift AGN as in high-redshift AGN of the same luminosities (see Figs 8 and 12). This is despite the order-of-magnitude global decrease in average star formation rates towards the lower redshift sample (see Sections 4.3.1 and 4.3.3).

Based on our systematic study of a representative sample of high-redshift X-ray detected AGN, we have evidence that high-velocity ionized outflows are prevalent, in qualitative agreement with theoretical predictions of galaxy formation models. These features appear to be most common in the most powerful AGN, and are equally prevalent in the high-redshift Universe as in the low-redshift Universe for AGN of the same luminosity. Our analyses focused on

searching for emission-line profiles which are dominated by these extreme gas kinematics and lower level outflows or highly turbulent material could be even more common. Due to a higher fraction of galaxies hosting the most luminous AGN at higher redshifts, our results imply that the most extreme ionized outflows are more prevalent in high-redshift galaxies. In future papers, we will present results based on the spatially resolved kinematics of multiple emission lines which will reveal information on the sizes and morphologies of the high-velocity gas, enable us to disentangle outflows from galaxy kinematics, and help us to measure the energetics of the outflows (following e.g. Nesvadba et al. 2008; Cano-Díaz et al. 2012; Harrison et al. 2012, 2014; Cresci et al. 2015). Furthermore, as our sample size grows, we will be able to make more definitive tests on the drivers and impact of ionized outflows in large samples of representative AGN.

REFERENCES

Abazajian K. N. et al., 2009, ApJS, 182, 543
Aird J. et al., 2010, MNRAS, 401, 2531
Aird J. et al., 2013, ApJ, 775, 41
Alexander D. M., Hickox R. C., 2012, New Astron. Rev., 56, 93
Alexander D. M. et al., 2008, ApJ, 687, 835
Alexander D. M., Swinbank A. M., Smail I., McDermid R., Nesvadba N. P. H., 2010, MNRAS, 402, 2211
Alloin D., Collin-Souffrin S., Joly M., Vigroux L., 1979, A&A, 78, 200
Arribas S., Colina L., Bellocchi E., Maiolino R., Villar-Martín M., 2014, A&A, 568, A14
Azadi M. et al., 2015, ApJ, 806, 187
Baldry I. K. et al., 2012, MNRAS, 421, 621
Balmaverde B. et al., 2015, preprint (arXiv:1506.05984)
Barth A. J., Greene J. E., Ho L. C., 2008, AJ, 136, 1179
Benson A. J., Bower R. G., Frenk C. S., Lacey C. G., Baugh C. M., Cole S., 2003, ApJ, 599, 38
Berney S. et al., 2015, MNRAS, 454, 3622
Bezanson R., Franx M., van Dokkum P. G., 2015, ApJ, 799, 148
Bongiorno A. et al., 2012, MNRAS, 427, 3103
Bonzini M., Padovani P., Mainieri V., Kellermann K. I., Miller N., Rosati P., Tozzi P., Vattakunnel S., 2013, MNRAS, 436, 3759
Boroson T. A., Persson S. E., Oke J. B., 1985, ApJ, 293, 120
Bower R. G., Benson A. J., Malbon R., Helly J. C., Frenk C. S., Baugh C. M., Cole S., Lacey C. G., 2006, MNRAS, 370, 645
Bower R. G., McCarthy I. G., Benson A. J., 2008, MNRAS, 390, 1399
Brandt W. N., Alexander D. M., 2015, A&AR, 23, 1
Brusa M. et al., 2010, ApJ, 716, 348

Brusa M. et al., 2015, MNRAS, 446, 2394

Calzetti D., 2013, in Falcón-Barroso J., Knapen J. H., eds, Secular Evolution of Galaxies. Cambridge Univ. Press, Cambridge, UK, p. 419

Calzetti D., Armus L., Bohlin R. C., Kinney A. L., Koornneef J., Storchi-Bergmann T., 2000, ApJ, 533, 682

Cano-Díaz M., Maiolino R., Marconi A., Netzer H., Shemmer O., Cresci G., 2012, A&A, 537, L8

Cappelluti N. et al., 2009, A&A, 497, 635

Carniani S. et al., 2015, A&A, 580, A102

Chabrier G., 2003, PASP, 115, 763

Churazov E., Sazonov S., Sunyaev R., Forman W., Jones C., Böhringer H., 2005, MNRAS, 363, L91

Cicone C. et al., 2014, A&A, 562, A21

Civano F. et al., 2012, ApJS, 201, 30

Collet C. et al., 2015, preprint (arXiv:1510.06631)

Crenshaw D. M., Schmitt H. R., Kraemer S. B., Mushotzky R. F., Dunn J. P., 2010, ApJ, 708, 419

Cresci G. et al., 2015, ApJ, 799, 82

Davies R. I. et al., 2013, A&A, 558, A56

Del Moro A. et al., 2013, A&A, 549, A59

Del Moro A. et al., 2015, preprint (arXiv:1504.03329)

Denicoló G., Terlevich R., Terlevich E., 2002, MNRAS, 330, 69

Dimitrijević M. S., Popović L. Č., Kovačević J., Dačić M., Ilić D., 2007, MNRAS, 374, 1181

Eisenhauer F. et al., 2003, in Iye M., Moorwood A. F. M., eds, SPIE Conf. Ser. Vol. 4841, Instrument Design and Performance for Optical/Infrared Ground-based Telescopes, Society of Photo-Optical Instrumentation Engineers, p. 1548

Elvis M. et al., 2009, ApJS, 184, 158

Fabian A. C., 2012, ARA&A, 50, 455

Förster Schreiber N. M. et al., 2014, ApJ, 787, 38

Furusawa H. et al., 2008, ApJS, 176, 1

Gaspari M., Melioli C., Brighenti F., D'Ercole A., 2011, MNRAS, 411, 349

Gehrels N., 1986, ApJ, 303, 336

Genzel R. et al., 2011, ApJ, 733, 101

Genzel R. et al., 2014, ApJ, 796, 7

George I. M., Turner T. J., Yaqoob T., Netzer H., Laor A., Mushotzky R. F., Nandra K., Takahashi T., 2000, ApJ, 531, 52

Giacconi R. et al., 2001, ApJ, 551, 624

Goulding A. D., Alexander D. M., 2009, MNRAS, 398, 1165

Granato G. L., De Zotti G., Silva L., Bressan A., Danese L., 2004, ApJ, 600, 580

Hainline K. N., Hickox R., Greene J. E., Myers A. D., Zakamska N. L., 2013, ApJ, 774, 145

Harrison C. M. et al., 2012, MNRAS, 426, 1073

Harrison C. M., Alexander D. M., Mullaney J. R., Swinbank A. M., 2014, MNRAS, 441, 3306

Harrison C. M., Thomson A. P., Alexander D. M., Bauer F. E., Edge A. C., Hogan M. T., Mullaney J. R., Swinbank A. M., 2015, ApJ, 800, 45

Heckman T. M., Best P. N., 2014, ARA&A, 52, 589

Heckman T. M., Miley G. K., van Breugel W. J. M., Butcher H. R., 1981, ApJ, 247, 403

Heckman T. M., Ptak A., Hornschemeier A., Kauffmann G., 2005, ApJ, 634, 161

Hickox R. C., Mullaney J. R., Alexander D. M., Chen C.-T. J., Civano F. M., Goulding A. D., Hainline K. N., 2014, ApJ, 782, 9

Hill M. J., Zakamska N. L., 2014, MNRAS, 439, 2701

Ho I.-T. et al., 2014, MNRAS, 444, 3894

Hopkins P. F., Hernquist L., Cox T. J., Di Matteo T., Robertson B., Springel V., 2006, ApJS, 163, 1

Husemann B., Wisotzki L., Sánchez S. F., Jahnke K., 2013, A&A, 549, A43

Juneau S. et al., 2013, ApJ, 764, 176

Kalfountzou E. et al., 2014, MNRAS, 442, 1181

Kennicutt R. C., Jr, 1998, ARA&A, 36, 189

Kewley L. J., Dopita M. A., 2002, ApJS, 142, 35

Kewley L. J., Groves B., Kauffmann G., Heckman T., 2006, MNRAS, 372, 961

Kewley L. J., Dopita M. A., Leitherer C., Davé R., Yuan T., Allen M., Groves B., Sutherland R., 2013, ApJ, 774, 100

Kormendy J., Ho L. C., 2013, ARA&A, 51, 511

La Franca F., Melini G., Fiore F., 2010, ApJ, 718, 368

LaMassa S. M., Heckman T. M., Ptak A., Hornschemeier A., Martins L., Sonnentrucker P., Tremonti C., 2009, ApJ, 705, 568

Lamastra A., Bianchi S., Matt G., Perola G. C., Barcons X., Carrera F. J., 2009, A&A, 504, 73

Lawrence A. et al., 2007, MNRAS, 379, 1599

Lehmer B. D. et al., 2009, ApJ, 691, 687

Lehmer B. D., Alexander D. M., Bauer F. E., Brandt W. N., Goulding A. D., Jenkins L. P., Ptak A., Roberts T. P., 2010, ApJ, 724, 559

Lequeux J., Peimbert M., Rayo J. F., Serrano A., Torres-Peimbert S., 1979, A&A, 80, 155

Liu G., Zakamska N. L., Greene J. E., Nesvadba N. P. H., Liu X., 2013, MNRAS, 436, 2576

Lusso E. et al., 2012, MNRAS, 425, 623

McCarthy I. G. et al., 2010, MNRAS, 406, 822

McElroy R., Croom S. M., Pracy M., Sharp R., Ho I.-T., Medling A. M., 2015, MNRAS, 446, 2186

McNamara B. R., Nulsen P. E. J., 2012, New J. Phys., 14, 055023

Madau P., Dickinson M., 2014, ARA&A, 52, 415

Maiolino R. et al., 2008, A&A, 488, 463

Malkan M. A., Gorjian V., Tam R., 1998, ApJS, 117, 25

Markevitch M., 1998, ApJ, 504, 27

Markwardt C. B., 2009, in Bohlender D. A., Durand D., Dowler P., eds, ASP Conf. Ser. Vol. 411, Astronomical Data Analysis Software and Systems XVIII, Astron. Soc. Pac., San Francisco, p. 251

Martin C. L., 2005, ApJ, 621, 227

Miller N. A. et al., 2013, ApJS, 205, 13

Modigliani A. et al., 2007, preprint (arXiv:0701297)

Morganti R., Oosterloo T. A., Tadhunter C. N., van Moorsel G., Emonts B., 2005, A&A, 439, 521

Morganti R., Frieswijk W., Oonk R. J. B., Oosterloo T., Tadhunter C., 2013, A&A, 552, L4

Mukherjee S., Feigelson E. D., Jogesh Babu G., Murtagh F., Fraley C., Raftery A., 1998, ApJ, 508, 314

Mulchaey J. S., Koratkar A., Ward M. J., Wilson A. S., Whittle M., Antonucci R. R. J., Kinney A. L., Hurt T., 1994, ApJ, 436, 586

Mullaney J. R. et al., 2012, MNRAS, 419, 95

Mullaney J. R., Alexander D. M., Fine S., Goulding A. D., Harrison C. M., Hickox R. C., 2013, MNRAS, 433, 622

Mullaney J. R. et al., 2015, MNRAS, 453, L83

Nandra K., Pounds K. A., 1994, MNRAS, 268, 405

Nesvadba N. P. H., Lehnert M. D., Eisenhauer F., Gilbert A., Tecza M., Abuter R., 2006, ApJ, 650, 693

Nesvadba N. P. H., Lehnert M. D., De Breuck C., Gilbert A. M., van Breugel W., 2008, A&A, 491, 407

Netzer H., Mainieri V., Rosati P., Trakhtenbrot B., 2006, A&A, 453, 525

Nims J., Quataert E., Faucher-Giguère C.-A., 2015, MNRAS, 447, 3612

Osterbrock D. E., Ferland G. J., 2006, Astrophysics of Gaseous Nebulae and Active Galactic Nuclei. Univ. Science Books, Mill Valley, CA

Panessa F., Bassani L., Cappi M., Dadina M., Barcons X., Carrera F. J., Ho L. C., Iwasawa K., 2006, A&A, 455, 173

Peterson B. M. et al., 2004, ApJ, 613, 682

Popesso P. et al., 2009, A&A, 494, 443

Pringle J. E., 1981, ARA&A, 19, 137

Puccetti S. et al., 2009, ApJS, 185, 586

Rawlings J. I. et al., 2015, MNRAS, 452, 4111

Rich J. A., Kewley L. J., Dopita M. A., 2014, ApJ, 781, L12

Rosario D. J. et al., 2012, A&A, 545, A45

Rosen S. R. et al., 2015, preprint (arXiv:1504.07051)

Rousselot P., Lidman C., Cuby J.-G., Moreels G., Monnet G., 2000, A&A, 354, 1134

Rupke D. S. N., Veilleux S., 2013, ApJ, 768, 75

Rupke D. S., Veilleux S., Sanders D. B., 2005, ApJ, 632, 751

Saez C. et al., 2015, MNRAS, 450, 2615

Sanders D. B., Soifer B. T., Elias J. H., Neugebauer G., Matthews K., 1988, ApJ, 328, L35

Schawinski K., Koss M., Berney S., Sartori L. F., 2015, MNRAS, 451, 2517

Schaye J. et al., 2015, MNRAS, 446, 521

Schinnerer E. et al., 2010, ApJS, 188, 384

Schwarz G., 1978, Ann. Stat., 6, 461

Scoville N. et al., 2007, ApJS, 172, 1

Sharples R. M. et al., 2004, in Moorwood A. F. M., Iye M., eds, SPIE Conf. Ser. Vol. 5492, Ground-based Instrumentation for Astronomy. SPIE, p. 1179

Sharples R. et al., 2013, The Messenger, 151, 21

Silk J., Rees M. J., 1998, A&A, 331, L1

Simpson C. et al., 2012, MNRAS, 421, 3060

Smail I., Sharp R., Swinbank A. M., Akiyama M., Ueda Y., Foucaud S., Almaini O., Croom S., 2008, MNRAS, 389, 407

Sobral D. et al., 2013, ApJ, 779, 139

Somerville R. S., Hopkins P. F., Cox T. J., Robertson B. E., Hernquist L., 2008, MNRAS, 391, 481

Stanley F., Harrison C. M., Alexander D. M., Swinbank A. M., Aird J. A., Del Moro A., Hickox R. C., Mullaney J. R., 2015, MNRAS, 453, 591

Steidel C. C., Adelberger K. L., Dickinson M., Giavalisco M., Pettini M., Kellogg M., 1998, ApJ, 492, 428

Stockton A., 1976, ApJ, 205, L113

Stott J. P. et al., 2013, MNRAS, 436, 1130

Stott J. P. et al., 2014, MNRAS, 443, 2695

Swinbank A. M., Smail I., Chapman S. C., Blain A. W., Ivison R. J., Keel W. C., 2004, ApJ, 617, 64

Symeonidis M. et al., 2011, MNRAS, 417, 2239

Tadhunter C., Morganti R., Rose M., Oonk J. B. R., Oosterloo T., 2014, Nature, 511, 440

Tremonti C. A. et al., 2004, ApJ, 613, 898

Trouille L., Barger A. J., 2010, ApJ, 722, 212

Ueda Y. et al., 2008, ApJS, 179, 124

Vega Beltrán J. C., Pizzella A., Corsini E. M., Funes J. G., Zeilinger W. W., Beckman J. E., Bertola F., 2001, A&A, 374, 394

Veilleux S., 1991, ApJS, 75, 383

Veilleux S. et al., 2013, ApJ, 776, 27

Vignali C., Alexander D. M., Comastri A., 2006, MNRAS, 373, 321

Vogelsberger M. et al., 2014, MNRAS, 444, 1518

Vrtilek J. M., 1985, ApJ, 294, 121

Wagner A. Y., Bicknell G. V., Umemura M., 2012, ApJ, 757, 136

Weedman D. W., 1970, ApJ, 159, 405

Westmoquette M. S., Clements D. L., Bendo G. J., Khan S. A., 2012, MNRAS, 424, 416

Williams R. J., Maiolino R., Santini P., Marconi A., Cresci G., Mannucci F., Lutz D., 2014, MNRAS, 443, 3780

Wisnioski E. et al., 2015, ApJ, 799, 209

Xue Y. Q. et al., 2011, ApJS, 195, 10

York D. G. et al., 2000, AJ, 120, 1579

Zakamska N. L., Greene J. E., 2014, MNRAS, 442, 784

Zakamska N. L., Strauss M. A., Heckman T. M., Ivezić Ž., Krolik J. H., 2004, AJ, 128, 1002

APPENDIX A: EMISSION-LINE PROFILES AND TABULATED DATA

In this Appendix we provide the galaxy-integrated spectra, emission-line profile fits and tabulated data for all 89 of the KASH*z* targets presented in this work.
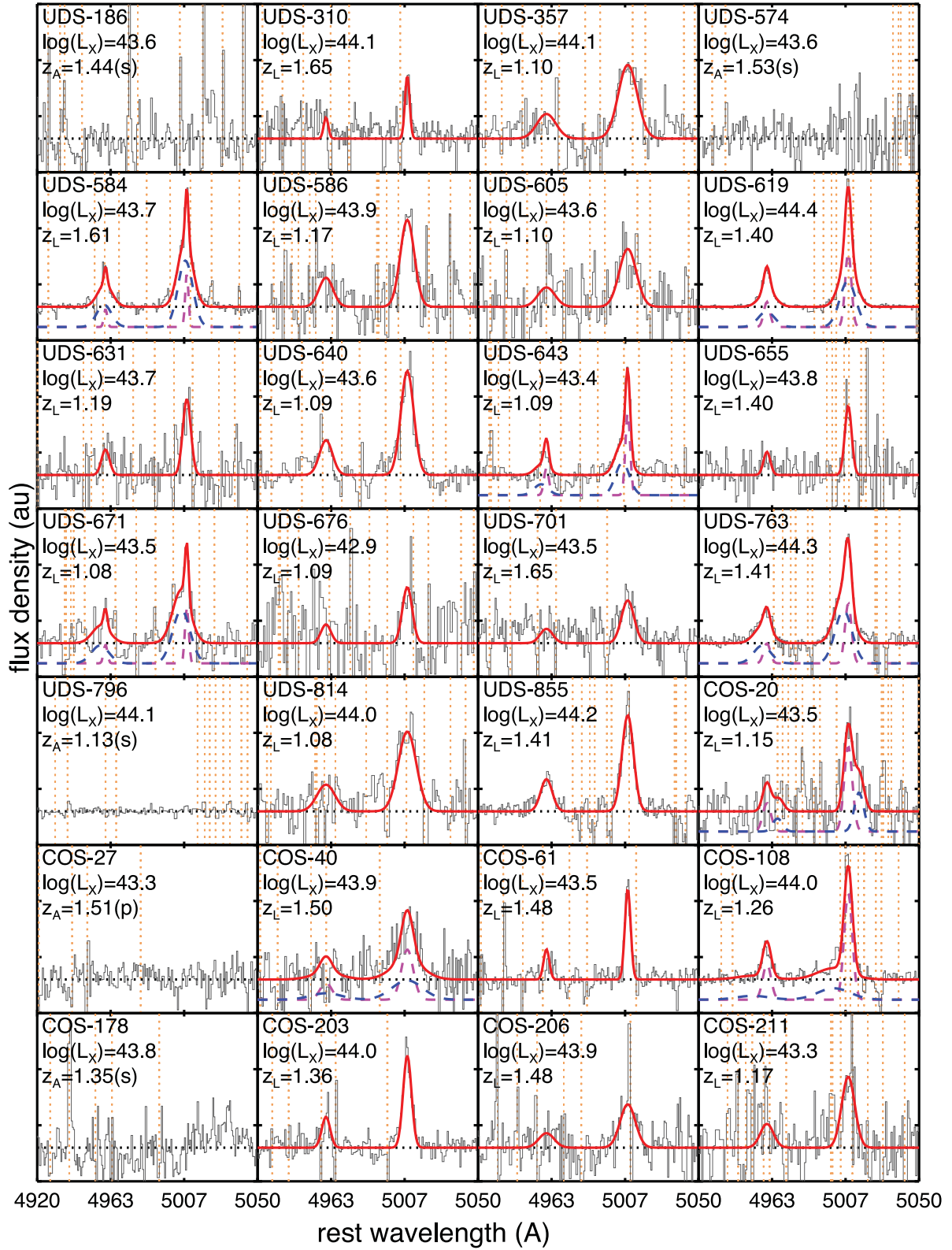
Figure A1. Galaxy-integrated spectra, shifted to the rest frame, around the [O III]4959,5007 emission-line doublet, in arbitrary flux density units, for the $z \approx 1.1$–1.7 targets. For descriptions of the different curves, see Fig. 3. The vertical dotted lines indicate the wavelengths of the brightest sky lines (Rousselot et al. 2000). In each panel we also identify the target name, the logarithm of the hard-band X-ray luminosities (erg s^{-1}) and the redshifts (Table A1).
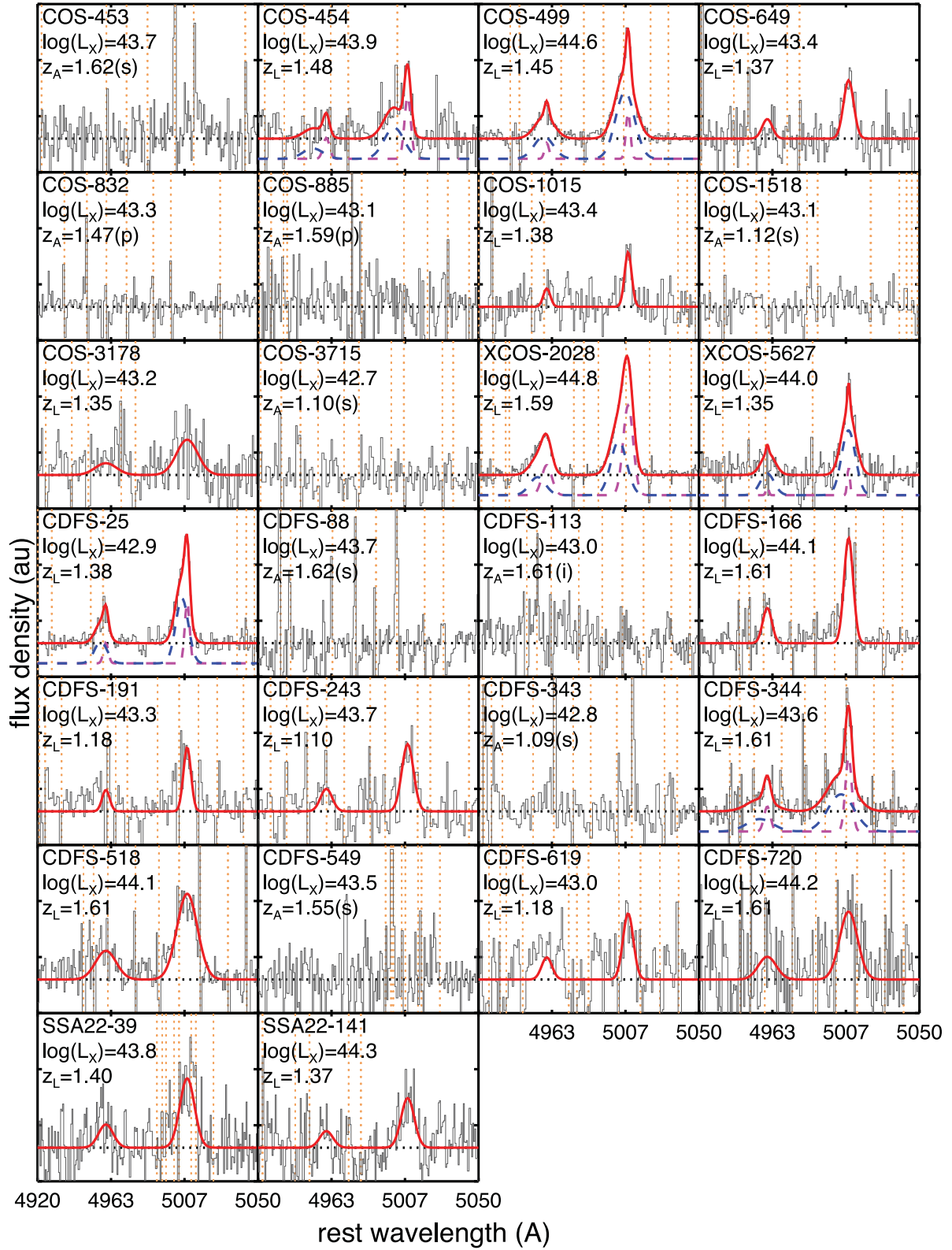
Figure A1 – *continued*
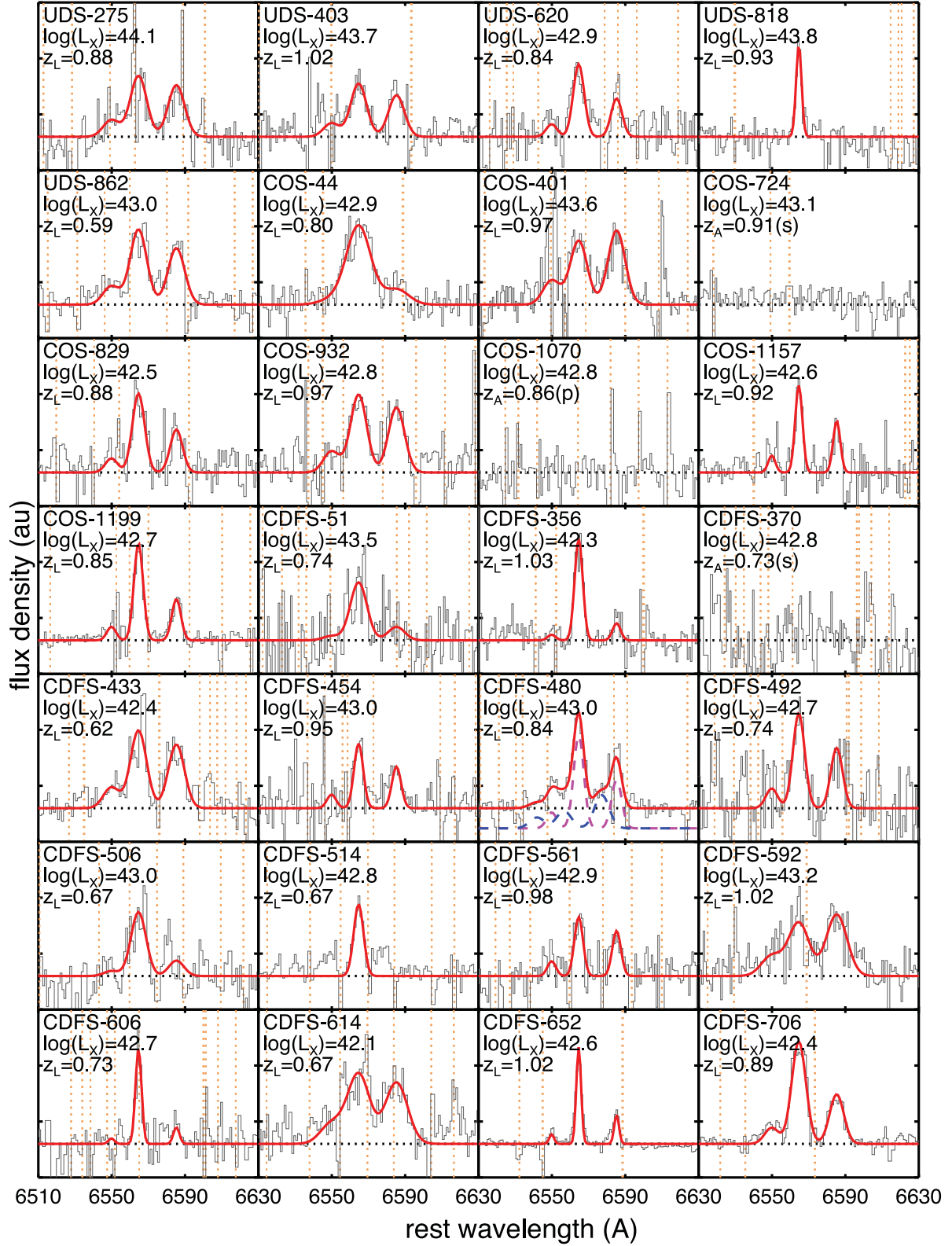
Figure A2. Galaxy-integrated spectra, shifted to the rest frame, around the Hα and [N II]6548,6583 emission lines, in arbitrary flux density units, for the $z \approx 0.6$–1.1 targets that do not show a BLR component (i.e. only the Type 2 sources). For descriptions of the different curves, see Fig. 4. The vertical dotted lines indicate the wavelengths of the brightest sky lines (Rousselot et al. 2000). In each panel we identify the target name, the logarithm of the X-ray luminosity (in units of erg s^{-1}) and the redshift (see Table A1).
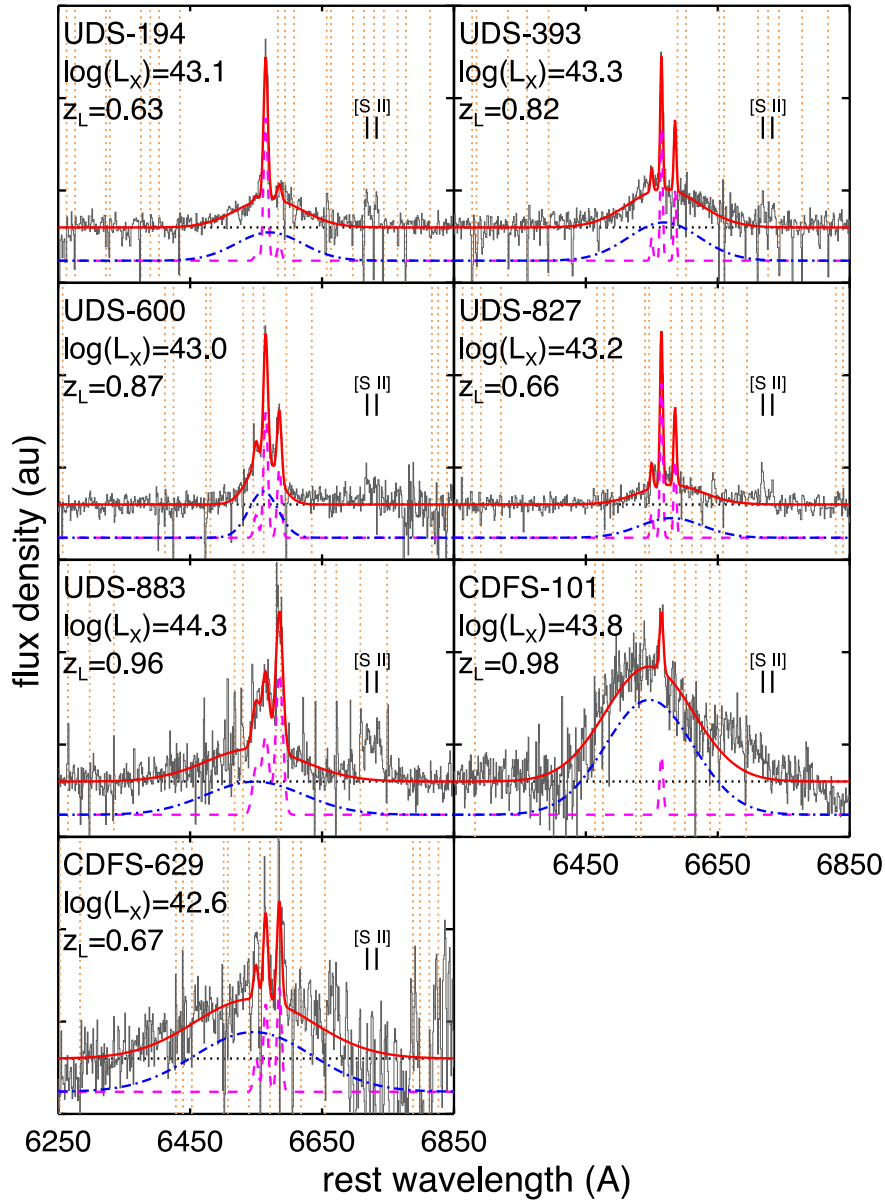
Figure A3. Same as Fig. A2 but for the targets with an Hα BLR component (i.e. the Type 1 sources). In each panel, the small thick lines show the wavelengths of the [S II]6716,6731 emission-line doublet.

Table A1. Properties for the KASHz targets observed so far. Notes: (1) source name ('field' – 'X-ray ID'), those targets followed by an asterisk (*) are excluded from the analyses presented in this paper (see Section 4.1); (2) archival redshift; (3) archival redshift type (i.e. photometric [p], secure spectroscopic redshift [s] or insecure spectroscopic redshift [i]); (4) and (5) soft-band (0.5–2 keV) and hard-band (2–10 keV) X-ray fluxes ($\times 10^{-15}$ erg s^{-1} cm^{-2}), where the hard-band values followed by a dagger (†) were estimated from the soft-band fluxes (see Section 2.1); (6) flag to indicate X-ray obscured candidates (O) and unobscured candidates (U), those without sufficient constraints are labelled with '?'; (7) flag to indicate the 'radio luminous' targets with $L_{1.4\mathrm{GHz}} > 10^{24}$ W Hz^{-1} (Section 2.5); (8) instrument used for the observations (K = KMOS and S = SINFONI); (9) primary targeted emission line (O = [O III] and H = Hα); (10) on-source exposure time (kilo-seconds); (11) note on data (L = line detected; C = continuum only detected; N = no detection); (12) note for the identification of a BLR component (α for Hα and β for Hβ); (13) redshift derived from the narrowest Gaussian component of the emission-line profile fit; (14) and (15) flux ($\times 10^{-17}$ erg s^{-1} cm^{-2}) and FWHM (km s^{-1}) of the narrower Gaussian component; (16) and (17) flux ($\times 10^{-17}$ erg s^{-1} cm^{-2}) and FWHM (km s^{-1}) of the broader Gaussian component (for sources with 'α' in column 10, this is the BLR component); (18) velocity offset between the two Gaussian components (km s^{-1}); (19) flux of [N II]λ6583 ($\times 10^{-17}$ erg s^{-1} cm^{-2}) where applicable; (20) overall emission-line width (km s^{-1}). All of the quoted uncertainties are the random errors on the fits due to the noise in the spectra; however, we note that throughout this work we add an extra 30 per cent systematic error to the emission-line fluxes to account for the uncertainty in the flux calibration (see Section 3.2).

Name (1)	z_A (2)	Type (3)	$F_{0.5-2}$ (4)	F_{2-10} (5)	Obs. (6)	RL (7)	Inst. (8)	Line (9)	t_e (10)	Note (11)	BLR (12)	z_L (13)	S_A (14)	FW_A (15)	S_B (16)	FW_B (17)	Δv (18)	$S_{[\mathrm{N\,II}]}$ (19)	W_{80} (20)
UDS-186	1.437	s	4.07	5.06	U	✗	K	O	9.0	C	–	–	<0.7	–	–	–	–	–	–
UDS-310	1.645	s	6.96	14.15	U	✗	K	O	5.4	L	–	1.645	1.8 ± 0.7	176 ± 86	–	–	–	–	192 ± 94
UDS-357	1.100	s	16.16	33.18	U	✗	K	O	5.4	L	β	1.098	4.5 ± 1.0	779 ± 154	–	–	–	–	848 ± 168
UDS-574	1.533	s	4.55	4.51	U	✗	K	O	6.0	C	–	–	<0.9	–	–	–	–	–	–
UDS-584	1.602	s	2.74	5.22	U	✗	K	O	9.0	L	–	1.605	4.5 ± 0.8	107 ± 26	25.0 ± 1.1	611 ± 28	−52 ± 13	–	615 ± 26
UDS-586	1.171	s	11.97	16.01	U	✗	K	O	6.0	L	–	1.171	2.3 ± 0.3	575 ± 83	–	–	–	–	625 ± 90
UDS-605	1.096	s	2.77	9.76	O	✗	K	O	6.0	L	–	1.097	1.7 ± 0.7	696 ± 253	–	–	–	–	758 ± 276
UDS-619	1.410	s	19.18	35.01	U	✗	K	O	6.0	L	β	1.401	21.2 ± 4.9	242 ± 29	33.8 ± 4.3	666 ± 68	−18 ± 15	–	562 ± 33
UDS-631	1.192	s	3.79	10.70	U	✗	K	O	8.4	L	–	1.190	1.6 ± 0.4	318 ± 95	–	–	–	–	346 ± 104
UDS-640	1.092	s	6.61	10.32	U	✗	K	O	6.0	L	–	1.093	5.4 ± 0.3	523 ± 36	–	–	–	–	570 ± 39
UDS-643	1.087	s	3.46	5.45	U	✗	K	O	6.0	L	β	1.088	2.4 ± 0.7	146 ± 66	2.8 ± 0.9	565 ± 149	−189 ± 121	–	492 ± 161
UDS-655	1.397	s	4.92	8.32	O	✗	K	O	6.0	L	–	1.397	1.5 ± 0.6	272 ± 114	–	–	–	–	297 ± 125
UDS-671	1.083	s	0.27	8.02	U	✗	K	O	6.0	L	–	1.084	0.8 ± 0.5	77 ± 66	3.6 ± 0.8	712 ± 176	−191 ± 90	–	702 ± 134
UDS-676	1.086	s	3.41	1.91	O	✗	K	O	9.0	L	–	1.086	0.7 ± 0.3	389 ± 215	–	–	–	–	424 ± 235
UDS-701	1.653	s	1.45	3.35	U	✗	K	O	9.0	L	–	1.649	2.1 ± 1.1	520 ± 268	–	–	–	–	566 ± 291
UDS-763	1.413	s	18.88	26.59	U	✗	K	O	9.0	L	β	1.412	5.4 ± 3.0	261 ± 78	12.3 ± 3.3	705 ± 142	−168 ± 90	–	668 ± 84
UDS-796	1.132	s	3.80	29.03	O	✗	K	O	9.0	C	–	–	<1.7	–	–	–	–	–	–
UDS-814	1.074	s	25.01	22.10	U	✗	K	O	9.0	L	–	1.077	3.1 ± 0.5	735 ± 133	–	–	–	–	800 ± 144
UDS-855	1.407	s	5.19	20.10	U	✗	K	O	8.4	L	–	1.407	11.4 ± 1.1	475 ± 45	–	–	–	–	517 ± 49
COS-20	1.156	s	1.34	7.01	U	✓	K	O	9.6	L	–	1.154	1.7 ± 0.5	324 ± 130	0.9 ± 0.6	380 ± 214	381 ± 25	–	612 ± 120
COS-27*	1.510	p	<0.73	2.59	O	✓	K	O	5.4	C	–	–	–	–	–	–	–	–	–
COS-40	1.510	s	4.69	9.69	U	✗	K	O	9.0	L	–	1.504	2.8 ± 1.3	485 ± 171	2.8 ± 1.6	1277 ± 375	0 ± 72	–	953 ± 273
COS-61	1.478	s	1.45	4.41†	U	?	K	O	8.4	L	–	1.478	3.5 ± 0.6	225 ± 45	–	–	–	–	245 ± 49
COS-108	1.253	s	4.39	16.10	O	✓	S	O	12.0	L	–	1.258	25.6 ± 3.0	331 ± 26	10.9 ± 4.6	1360 ± 409	−506 ± 220	–	936 ± 170
COS-178	1.347	s	1.47	8.92	O	✗	S	O	12.0	C	–	–	<1.2	–	–	–	–	–	–
COS-203	1.360	s	7.37	14.20	U	✗	K	O	7.2	L	–	1.359	5.1 ± 0.6	337 ± 49	–	–	–	–	367 ± 53
COS-206	1.483	s	5.89	10.40	U	✗	K	O	7.2	L	β	1.480	3.8 ± 2.0	630 ± 323	–	–	–	–	685 ± 352
COS-211	1.166	s	2.65	4.12	U	✓	K	O	7.2	L	–	1.167	2.4 ± 0.5	535 ± 113	–	–	–	–	582 ± 123
COS-453	1.625	s	3.01	5.70	U	✗	K	O	9.0	C	–	–	<0.5	–	–	–	–	–	–
COS-454	1.478	s	5.91	10.90	U	✓	K	O	9.0	L	β	1.484	2.5 ± 1.3	250 ± 110	3.9 ± 1.9	827 ± 415	−480 ± 198	–	922 ± 280
COS-499	1.459	s	27.40	57.80	U	✓	K	O	9.0	L	β	1.455	3.6 ± 1.2	165 ± 49	22.1 ± 1.7	786 ± 58	−109 ± 28	–	796 ± 52
COS-649	1.369	s	<1.05	3.77	O	✓	K	O	5.4	L	–	1.367	2.0 ± 0.7	415 ± 166	–	–	–	–	451 ± 180
COS-832*	1.471	p	<0.86	2.71	O	✓	K	O	8.4	C	–	–	–	–	–	–	–	–	–
COS-885*	1.594	p	0.44	1.34†	?	✗	K	O	9.0	C	–	–	–	–	–	–	–	–	–
COS-1015	1.379	s	1.14	3.47†	U	✗	K	O	7.2	L	–	1.377	1.4 ± 0.6	281 ± 160	–	–	–	–	306 ± 174
COS-1518*	1.122	s	<0.24	2.90	O	✗	K	O	9.6	N	–	–	–	–	–	–	–	–	–
COS-3178	1.355	s	<0.59	2.54	O	✗	K	O	7.2	L	–	1.354	3.6 ± 2.3	882 ± 547	–	–	–	–	960 ± 596
COS-3715	1.103	s	0.36	1.11†	?	✗	K	O	5.4	C	–	–	<0.7	–	–	–	–	–	–
XCOS-2028	1.592	s	33.50	75.70	U	✗	S	O	25.2	L	β	1.593	17.3 ± 5.1	410 ± 36	16.4 ± 5.4	657 ± 88	−332 ± 76	–	734 ± 17
XCOS-5627	1.337	s	7.25	15.40	U	✗	S	O	12.6	L	β	1.349	0.7 ± 0.7	62 ± 62	9.0 ± 2.1	573 ± 191	−5 ± 14	–	602 ± 131
CDFS-25	1.374	s	<0.19	1.21	O	✗	K	O	6.0	L	–	1.377	3.5 ± 2.1	184 ± 51	9.7 ± 2.4	518 ± 94	−183 ± 64	–	524 ± 37

Table A1 – *continued*

Name (1)	z_A (2)	Type (3)	$F_{0.5-2}$ (4)	F_{2-10} (5)	Obs. (6)	RL (7)	Inst. (8)	Line (9)	t_e (10)	Note (11)	BLR (12)	z_L (13)	S_A (14)	FW_A (15)	S_B (16)	FW_B (17)	Δv (18)	$S_{[NII]}$ (19)	W_{80} (20)
CDFS-88	1.616	s	3.07	4.67	U	✗	K	O	9.0	C	–	–	<0.7	–	–	–	–	–	–
CDFS-113*	1.608	i	0.57	0.98	U	✗	K	O	9.0	N	–	–	–	–	–	–	–	–	–
CDFS-166	1.605	s	8.14	12.39	U	✓	K	O	9.0	L	–	1.611	5.7 ± 0.4	371 ± 25	–	–	–	–	404 ± 27
CDFS-191	1.185	s	<0.13	4.18	O	✗	K	O	6.0	L	–	1.184	1.8 ± 0.6	297 ± 125	–	–	–	–	323 ± 136
CDFS-243	1.097	s	0.71	12.29	O	✗	K	O	6.0	L	–	1.096	2.5 ± 0.8	468 ± 160	–	–	–	–	510 ± 174
CDFS-343	1.090	s	<0.08	1.46	O	✗	K	O	9.6	C	–	–	<0.4	–	–	–	–	–	–
CDFS-344	1.617	s	3.64	3.89	U	✓	K	O	7.8	L	–	1.613	3.7 ± 0.9	258 ± 40	7.4 ± 1.4	1089 ± 149	−270 ± 83	–	977 ± 112
CDFS-518	1.603	s	9.74	13.30	U	✓	K	O	9.6	L	–	1.609	8.3 ± 1.0	850 ± 88	–	–	–	–	926 ± 96
CDFS-549*	1.553	s	1.36	3.25	U	✗	K	O	11.4	N	–	–	–	–	–	–	–	–	–
CDFS-619	1.178	s	0.16	2.17	O	✗	K	O	7.8	L	–	1.179	1.6 ± 0.6	419 ± 147	–	–	–	–	457 ± 160
CDFS-720	1.609	s	10.93	18.44	U	✗	K	O	9.0	L	–	1.610	2.2 ± 0.5	742 ± 173	–	–	–	–	808 ± 189
SSA22-39	1.397	s	7.19	9.08	U	?	S	O	3.0	L	–	1.400	8.2 ± 1.8	605 ± 143	–	–	–	–	658 ± 156
SSA22-141	1.370	s	20.53	27.38	U	?	S	O	2.4	L	–	1.367	5.4 ± 2.2	541 ± 209	–	–	–	–	589 ± 227
UDS-194	0.627	s	5.80	10.13	U	✗	K	H	9.0	L	α	0.628	8.1 ± 0.9	328 ± 42	25.2 ± 2.3	5385 ± 604	258 ± 190	0.93	357 ± 46
UDS-275	0.883	s	7.27	52.96	O	✗	K	H	5.4	L	–	0.882	4.5 ± 0.9	476 ± 87	–	–	–	3.85	518 ± 95
UDS-393	0.822	s	9.70	9.38	O	✗	K	H	9.0	L	α	0.822	6.0 ± 0.5	183 ± 22	49.1 ± 3.4	5811 ± 457	176 ± 164	3.18	199 ± 24
UDS-403	1.021	s	2.30	13.81	O	✗	K	H	9.0	L	–	1.021	3.7 ± 0.9	461 ± 118	–	–	–	2.89	502 ± 129
UDS-600	0.873	s	4.41	4.34	O	✗	K	H	6.0	L	α	0.873	11.6 ± 1.3	343 ± 36	29.2 ± 3.1	2401 ± 216	−158 ± 108	6.39	373 ± 39
UDS-620	0.842	s	1.08	3.44	O	✗	K	H	9.0	L	–	0.843	2.3 ± 0.4	311 ± 75	–	–	–	1.21	338 ± 81
UDS-818	0.928	s	14.55	19.28	U	✗	K	H	8.4	L	–	0.928	5.5 ± 0.6	140 ± 26	–	–	–	<0.66	153 ± 29
UDS-827	0.658	s	4.65	12.22	U	✗	K	H	8.4	L	α	0.657	8.0 ± 0.4	201 ± 15	22.7 ± 2.1	5129 ± 548	682 ± 196	4.02	219 ± 17
UDS-862	0.589	s	1.35	8.50	O	✗	K	H	8.4	L	–	0.589	3.8 ± 0.5	471 ± 64	–	–	–	2.88	512 ± 70
UDS-883	0.961	s	24.39	56.58	U	✗	K	H	8.4	L	α	0.961	5.2 ± 0.4	574 ± 29	30.5 ± 2.9	7970 ± 795	−748 ± 271	9.50	625 ± 32
COS-44	1.513	p	1.23	3.56	O	✓	S	H	9.0	L	–	0.801	4.9 ± 0.4	709 ± 87	–	–	–	0.95	771 ± 95
COS-401	0.969	s	<1.14	13.30	O	✗	K	H	2.4	L	–	0.971	7.6 ± 0.9	495 ± 53	–	–	–	8.80	539 ± 58
COS-724	0.906	s	<0.51	4.04	O	✓	K	H	6.0	C	–	–	<1.2	–	–	–	–	–	–
COS-829	0.885	s	0.39	1.17†	?	✓	K	H	8.4	L	–	0.885	6.4 ± 0.8	353 ± 53	–	–	–	3.50	384 ± 57
COS-932	0.975	s	0.71	2.17†	U	✗	K	H	8.4	L	–	0.974	6.5 ± 0.6	469 ± 49	–	–	–	5.46	511 ± 53
COS-1070*	0.858	p	1.06	2.30	U	✗	K	H	8.4	C	–	–	–	–	–	–	–	–	–
COS-1157	0.915	s	0.49	1.50†	?	✗	K	H	9.6	L	–	0.925	5.2 ± 0.6	197 ± 29	–	–	–	3.07	215 ± 31
COS-1199	0.771	p	0.67	2.10	O	✗	K	H	9.0	L	–	0.850	9.5 ± 0.7	255 ± 23	–	–	–	4.05	277 ± 25
CDFS-51	0.737	s	1.81	18.33	O	✗	S	H	2.7	L	–	0.737	11.9 ± 3.5	481 ± 211	–	–	–	<3.79	523 ± 230
CDFS-101	0.977	s	11.57	20.66	U	✗	K	H	9.0	L	α	0.978	2.9 ± 0.5	290 ± 68	133.5 ± 3.7	7150 ± 202	−843 ± 77	<2.78	315 ± 74
CDFS-356	1.034	s	0.29	0.57	O	✗	K	H	6.0	L	–	1.034	3.3 ± 0.2	223 ± 17	–	–	–	0.55	243 ± 19
CDFS-370	0.734	s	1.07	3.38	O	✓	S	H	2.7	C	–	–	<3.8	–	–	–	–	–	–
CDFS-433	0.617	s	1.04	2.03	U	✗	K	H	9.6	L	–	0.620	3.7 ± 0.4	461 ± 52	–	–	–	3.03	501 ± 57
CDFS-454	0.952	i	1.42	3.17	U	✗	K	H	11.4	L	–	0.953	2.0 ± 0.5	257 ± 82	–	–	–	1.30	279 ± 89
CDFS-480	0.839	s	2.48	4.08	U	✗	K	H	6.0	L	–	0.841	7.3 ± 2.1	286 ± 41	2.0 ± 2.1	427 ± 280	−367 ± 169	4.80	521 ± 105
CDFS-492	0.735	s	0.13	2.55	O	✗	S	H	5.4	L	–	0.735	11.2 ± 0.6	334 ± 22	–	–	–	7.13	363 ± 24
CDFS-506	0.665	s	4.30	7.35	O	✗	K	H	7.8	L	–	0.666	1.8 ± 0.4	430 ± 109	–	–	–	<0.46	468 ± 119
CDFS-514	0.664	s	3.29	4.49	O	✗	K	H	7.8	L	–	0.667	4.2 ± 0.8	265 ± 81	–	–	–	<0.99	289 ± 88
CDFS-561	0.798	s	0.08	2.69	U	✗	K	H	7.8	L	–	0.980	1.4 ± 0.4	239 ± 91	–	–	–	1.11	261 ± 99
CDFS-592	1.016	s	0.69	4.03	O	✗	K	H	9.0	L	–	1.015	5.3 ± 0.9	601 ± 89	–	–	–	6.13	655 ± 97
CDFS-606	0.733	s	0.34	2.58	O	✗	K	H	7.8	L	–	0.733	1.7 ± 0.2	151 ± 17	–	–	–	0.29	164 ± 19
CDFS-614	0.668	s	0.38	0.86	U	✗	K	H	9.0	L	–	0.667	4.7 ± 0.5	642 ± 84	–	–	–	4.09	699 ± 92
CDFS-629	0.667	s	0.83	2.60	O	✗	K	H	9.0	L	α	0.667	1.7 ± 0.5	360 ± 89	30.1 ± 6.3	9647 ± 1706	−846 ± 549	2.04	392 ± 97
CDFS-652	1.020	s	0.66	1.18	U	✗	K	H	7.8	L	–	1.022	8.0 ± 0.7	115 ± 21	–	–	–	2.40	125 ± 22
CDFS-706	0.891	s	<0.15	1.02	O	✗	K	H	9.0	L	–	0.890	13.7 ± 0.4	409 ± 16	–	–	–	6.63	445 ± 18

This paper has been typeset from a TeX/LaTeX file prepared by the author.