

The KMOS Deep Survey (KDS) I: dynamical measurements of typical star-forming galaxies at $z \simeq 3.5$ red*

O. J. Turner,^{1,2,red†} M. Cirasuolo,^{2,1} C. M. Harrison,^{2,3} R. J. McLure,¹ J. S. Dunlop,¹
A. M. Swinbank,^{3,4} H. L. Johnson,^{3,4} D. Sobral,^{5,6} J. Matthee,⁶ R. M. Sharples^{3,4}

¹SUPAred†, *Institute for Astronomy, University of Edinburgh, Royal Observatory, Edinburgh EH9 3HJ*

²*European Southern Observatory, Karl-Schwarzschild-Str. 2, 85748 Garching b. München, Germany*

³*Centre for Extragalactic Astronomy, Durham University, South Road, Durham, DH1 3LE, U.K.*

⁴*Institute for Computational Cosmology, Durham University, South Road, Durham, DH1 3LE, U.K.*

⁵*Department of Physics, Lancaster University, Lancaster, LA1 4BY, U.K.*

⁶*Leiden Observatory, Leiden University, PO Box 9513, NL-2300 RA Leiden, the Netherlands*

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ABSTRACT

We present dynamical measurements from the KMOS (*K*-band Multi-Object Spectrograph) Deep Survey (KDS), which is comprised of 77 typical star-forming galaxies at $z \simeq 3.5$ in the mass range $9.0 < \log(M_\star/M_\odot) < 10.5$. These measurements constrain the internal dynamics, the intrinsic velocity dispersions (σ_{int}) and rotation velocities (V_C) of galaxies in the high-redshift Universe. The mean velocity dispersion of the galaxies in our sample is $\sigma_{\text{int}} = 70.8_{-3.1}^{+3.3} \text{ km s}^{-1}$, revealing that the increasing average σ_{int} with increasing redshift, reported for $z \lesssim 2$, continues out to $z \simeq 3.5$. Only $36 \pm 8\%$ of our galaxies are rotation-dominated ($V_C/\sigma_{\text{int}} > 1$), with the sample average V_C/σ_{int} value much smaller than at lower redshift. After carefully selecting comparable star-forming samples at multiple epochs, we find that the rotation-dominated fraction evolves with redshift with a $z^{-0.2}$ dependence. The rotation-dominated KDS galaxies show no clear offset from the local rotation velocity-stellar mass (i.e. $V_C - M_\star$) relation, although a smaller fraction of the galaxies are on the relation due to the increase in the dispersion-dominated fraction. These observations are consistent with a simple equilibrium model picture, in which random motions are boosted in high-redshift galaxies by a combination of the increasing gas fractions, accretion efficiency, specific star-formation rate and stellar feedback and which may provide significant pressure support against gravity on the galactic disk scale.

Key words: galaxies:high-redshift — galaxies:kinematics and dynamics — galaxies:evolution

1 INTRODUCTION

The galaxy population at all redshifts appears to be bimodal in many physical properties (e.g. as described in [Dekel & Birnboim 2006](#)), with a preference for the most massive galaxies to lie on the red sequence, characterised by red optical colours, low star-formation rates (SFRs) and spherical morphologies, and less massive galaxies in the blue sequence, characterised by blue colours, high star-formation rates and disky morphologies. For these blue, star-forming galaxies (SFGs) there is a roughly linear correlation between SFR and stellar mass (M_*) (e.g. [Daddi et al. 2007](#); [Noeske et al. 2007](#); [Elbaz et al. 2007](#)), in the sense that galaxies which have already accumulated a larger stellar population tend to have higher SFRs. This correlation, or ‘main-sequence’, underpins the ‘equilibrium model’, in which the SFR of galaxies is regulated by the availability of gas, with outflows and accretion events sustaining the galaxy gas reservoirs in a rough equilibrium as the galaxy evolves (e.g. [Davé et al. 2012](#); [Lilly et al. 2013](#); [Saintonge et al. 2013](#)). The main-sequence has been studied comprehensively, using multi-wavelength SFR tracers, between $0 < z < 3$ (e.g. [Rodighiero et al. 2011](#); [Karim et al. 2011](#); [Whitaker et al. 2012b](#); [Behroozi et al. 2013](#); [Whitaker et al. 2014](#); [Rodighiero et al. 2014](#); [Speagle et al. 2014](#); [Pannella et al. 2014](#); [Sobral et al. 2014](#); [Sparre et al. 2015](#); [Lee et al. 2015](#); [Schreiber et al. 2015](#); [Renzini & Peng 2015](#); [Nelson et al. 2016](#)), showing evolution of the relation towards higher SFRs at fixed M_* with increasing redshift, reflecting the increase of the cosmic Star-Formation Rate Density (SFRD) in this redshift range (e.g. [Madau & Dickinson 2014](#); [Khostovan et al. 2015](#)). At each redshift slice it has been suggested that galaxies on the main-sequence evolve secularly, regulated by their gas reservoirs, meaning that selecting such populations offers the chance to explore the evolution of the physical properties of typical SFGs across cosmic time. This assumes that high-redshift, main-sequence galaxies are the progenitors of their lower redshift counterparts, which may not be the case (e.g. [Gladders et al. 2013](#); [Kelson 2014](#); [Abramson et al. 2016](#)) and also assumes that we can learn about galaxy evolution (i.e. how individual galaxies develop in physical properties over time) by studying the mean properties of populations at different epochs.

The picture is complicated by the addition of major and minor galaxy mergers which can rapidly change the physical properties of galaxies (e.g. [Toomre 1977](#); [Lotz et al. 2008](#); [Conselice et al. 2011](#); [Conselice 2014](#)) and the relative importance of in-situ, secular stellar mass growth vs. stellar mass aggregation via mergers is the subject of much work involving both observations and simulations (e.g. [Robaina et al. 2009](#); [Kaviraj et al. 2012](#); [Stott et al. 2013](#); [Lofthouse et al. 2017](#); [Qu et al. 2017](#)). To account for the growing number density of quiescent galaxies from $z \simeq 2.5$ to the present day (e.g. [Bell et al. 2004](#); [Faber et al. 2007](#); [Brown et al. 2007](#); [Ilbert et al. 2010](#); [Brammer et al. 2011](#); [Muzzin et al. 2013](#); [Buitrago et al. 2013](#)) there must also be processes which shut-off star-formation within main-sequence galaxies (i.e. quenching) which, to explain observations, must be a function of both mass and environment ([Peng et al. 2010](#); [Darvish et al. 2016](#)).

Recent cosmological volume simulations provide subgrid recipes for the complex interplay of baryonic processes which are at work as galaxies evolve, and can track the development of individual galaxies from early stages, to maturity and through quenching ([Dubois et al. 2014](#); [Vogelsberger et al. 2014](#); [Schaye et al. 2015](#)). Observations can aid the predictive power of such simulations by providing constraints on the evolving physical properties of galaxy populations. The observed dynamical properties of galaxies contain information about the transfer of angular momentum between their dark matter halos and baryons, and the subsequent dissipation of this angular momentum (through gravitational collapse, mergers and outflows e.g. [Fall 1983](#); [Romanowsky & Fall 2012](#); [Fall & Romanowsky 2013](#)), constituting an important set of quantities for simulations to reproduce. Developments in both Integral-Field Spectroscopy (IFS) instrumentation and data analysis tools over the last decade have led to the observation of two-dimensional velocity and velocity dispersion fields for large samples of galaxies of different morphological types, spanning a wide redshift range (e.g. [Sarzi et al. 2006](#); [Flores et al. 2006](#); [Epinat et al. 2008b](#); [Förster Schreiber et al. 2009](#); [Cappellari et al. 2011](#); [Gnerucci et al. 2011](#); [Epinat et al. 2012](#); [Croom et al. 2012](#); [Swinbank et al. 2012a,b](#); [Bundy et al. 2015](#); [Wisnioski et al. 2015](#); [Stott et al. 2016](#); [Harrison et al. 2017](#); [Swinbank et al. 2017](#)). When interpreted in tandem with high-resolution imaging data from the *Hubble Space Telescope (HST)*, these data provide information about the range of physical processes which are driving galaxy evolution. In particular, in recent years the multiplexing capabilities of KMOS ([Sharples et al. 2013](#)) have allowed for IFS kinematic observations for large galaxy samples to be assembled rapidly ([Sobral et al. 2013](#); [Wisnioski et al. 2015](#); [Stott et al. 2016](#); [Mason et al. 2017](#); [Harrison et al. 2017](#)) providing an order-of-magnitude boost in statistical power over previous high-redshift campaigns.

Random motions within the interstellar medium of SFGs appear to increase with increasing redshift between $0 < z < 3$, as traced by their observed velocity dispersions, σ_{obs} ([Genzel et al. 2008](#); [Förster Schreiber et al. 2009](#); [Law et al. 2009](#); [Cresci et al. 2009](#); [Gnerucci et al. 2011](#); [Epinat et al. 2012](#); [Kassin et al. 2012](#); [Green et al. 2014](#); [Wisnioski et al. 2015](#); [Stott et al. 2016](#)). This has been explained in terms of increased ‘activity’ in galaxies during and before the global peak in cosmic SFRD ([Madau & Dickinson 2014](#)), in the form of higher specific star-formation rates (sSFRs) ([Wisnioski et al. 2015](#)), larger gas reservoirs ([Law et al. 2009](#); [Förster Schreiber et al. 2009](#); [Wisnioski et al. 2015](#); [Stott et al. 2016](#)), more efficient accretion ([Law et al. 2009](#)), increased stellar feedback from supernovae ([Kassin et al. 2012](#)) and turbulent disk instabilities ([Law et al. 2009](#); [Bournaud et al. 2007](#); [Bournaud & Frédéric 2016](#)), all of which combine to increase σ_{obs} and complicate its interpretation.

There is also an increasing body of work measuring the relationship between the observed maximum rotation velocity of a galaxy, a tracer for the total dynamical mass, and its stellar mass, known as the stellar mass Tully-Fisher Relation (smTFR) ([Tully & Fisher 1977](#)), with surveys reporting disparate results for the evolution of this relation with redshift (e.g. [Puech et al. 2008](#); [Miller et al. 2011](#); [Gnerucci et al. 2011](#); [Swinbank et al. 2012a](#); [Simons et al. 2016](#); [Tiley et al. 2016](#); [Harrison et al. 2017](#); [Straatman et al. 2017](#); [Ubler et al. 2017](#)). Systematic differences in measurement and modelling techniques at high-redshift, especially with regards to beam-smearing corrections, combine with our poor understanding

of progenitors and descendants to blur the evolutionary picture which these surveys paint. Additionally, there has been increasing focus in recent years on whether the measured velocity dispersions track random motions which provide partial gravitational support for high-redshift galaxy disks (e.g. Burkert et al. 2010; Wuyts et al. 2016; Ubler et al. 2017; Genzel et al. 2017; Lang et al. 2017). These random motions may become an increasingly significant component of the dynamical mass budget with increasing redshift (Wuyts et al. 2016) and pressure gradients across the disk could result in a decrease in the observed rotation velocities (Burkert et al. 2010). Different interpretations of the gaseous velocity dispersions and their role in providing pressure support against gravity also complicate the evolutionary picture.

In this paper we present new results from the KMOS Deep Survey (KDS), which is a guaranteed time programme focusing on the spatially-resolved properties of main-sequence SFGs at $z \simeq 3.5$, a time when the universe was building to peak activity. With this survey, we aim to complement existent studies by providing deep IFS data for the largest number of galaxies at this redshift. By making use of KMOS (with integration times of 7.5-9 hours), we have been able to study [O III] λ 5007 emission in 77 galaxies spanning the mass range $9.0 < \log(M_*/M_\odot) < 10.5$, roughly doubling the number of galaxies observed via IFS at $z > 3$. In order to interpret the evolution of the physical properties of typical star-forming galaxies we have carefully constructed a set of comparison samples spanning $0 < z < 3$. These samples use integral-field spectroscopy to track the ionised gas emission in star-forming galaxies and follow consistent kinematic parameter extraction methods. By doing this we seek to minimise the impact of systematic differences introduced by differing approaches to defining and extracting kinematic parameters.

There are still many open questions which we can begin to answer by studying the emission from regions of ionised gas within individual galaxies at these redshifts:

- (i) What are the dynamical properties of main-sequence galaxies at this early stage in their lifetimes?
- (ii) What are the radial gradients in metal enrichment within these galaxies and what can this tell us about the physical mechanisms responsible for redistributing metals?
- (iii) What is the connection between the gas-phase metallicity and kinematics, particularly in terms of inflows and outflows of material?

This paper focuses on (i) by deriving and interpreting the spatially-resolved kinematics of the KDS galaxies, particularly the rotation velocities and velocity dispersions, using the [O III] λ 5007 emission line, discussing what we can learn about the nature of galaxy formation at $z \simeq 3.5$ and forming evolutionary links with lower redshift work.

The structure of the paper is as follows. In § 2 we present the survey description, sample selection, observation strategy and data reduction, leading to stacked datacubes for each of the KDS galaxies. In § 3 we describe the derivation of morphological and kinematic properties for our galaxies, explaining the kinematic modelling approach and the beam-smearing corrections which lead to intrinsic measurements of the rotation velocities, V_C , and velocity dispersions σ_{int} for each of the galaxies classified as morphologically isolated and spatially-resolved in the [O III] λ 5007 emission line. § 4 presents an analysis of these derived kinematic parameters, comparing with lower redshift work where possible and drawing conclusions about the evolutionary trends and possible underlying physical mechanisms. We discuss these results in § 5 and present our conclusions in § 6. Throughout this work we assume a flat Λ CDM cosmology with $(h, \Omega_m, \Omega_\Lambda) = (0.7, 0.3, 0.7)$.

2 SURVEY DESCRIPTION, SAMPLE SELECTION AND OBSERVATIONS

2.1 The KDS survey description and sample selection

The KDS is a KMOS study of the gas kinematics and metallicity in 77 SFGs with a median redshift of $z \simeq 3.5$, probing a representative section of the galaxy main-sequence. The addition of these data approximately doubles the number of galaxies observed via IFS at this redshift (Cresci et al. 2010; Lemoine-Busserolle et al. 2010; Gnerucci et al. 2011; Troncoso et al. 2014), and will allow for a statistically-significant investigation of the dynamics and metal content of SFGs during a crucial period of galaxy evolution. The key science goals of the KDS are to investigate the resolved kinematic properties of high-redshift galaxies in the peak epoch of galaxy formation (particularly the fraction of rotating disks and the degree of disk turbulence) and also to study the spatial distribution of metals within these galaxies in the context of their observed dynamics. We seek to probe both a ‘field’ environment in which the density of galaxies is typical for this redshift and a ‘cluster’ environment containing a known galaxy over-density, in order to gauge the role of environment in determining the kinematics and metallicities of SFGs during this early stage in their formation history. To achieve this we require very deep exposure times in excess of 7 hours on source to reach the signal-to-noise required to detect line emission in the outskirts of the galaxies where the rotation curves begin to flatten, and to achieve adequate signal-to-noise across several ionised emission lines within individual spatial pixels (spaxels). Consequently, the KDS is one of the deepest spectroscopic datasets available at this redshift.

2.1.1 Sample selection

Target selection for the KDS sample is designed to pick out SFGs at $z \simeq 3.5$, supported by deep, multi-wavelength ancillary data. Within this redshift range the [O III] $\lambda 4959, \lambda 5007$ doublet and the H β emission lines are visible in the *K*-band and the [O II] $\lambda 3727, \lambda 3729$ doublet is visible in the *H*-band, both of which are observable with KMOS. From these lines, [O III] $\lambda 5007$ generally has the highest signal-to-noise and so is well suited to dynamical studies, whereas [O III] $\lambda 4959$, H β and the [O II] $\lambda 3727, \lambda 3729$ doublet complement [O III] $\lambda 5007$ as tracers of the galaxy metallicities. To ensure a high detection rate of the ionised gas emission lines in the KDS we select galaxies in well-studied fields that have a wealth of imaging and spectroscopic data. Most of the galaxies for the KMOS observations had a confirmed spectroscopic redshift (see below). A subset of the selected cluster galaxies in the SSA22 field were blindly-detected in Ly α emission during a narrow-band imaging study of a known overdensity of Lyman Break Galaxies (LBGs) at $z \simeq 3.09$ (Steidel et al. 2000). In each pointing, few sources had no spectroscopic redshift and were selected on the basis of their photometric redshift. We make no further cuts to the sample on the basis of mass and SFR, in order to probe a more representative region of the star-forming main-sequence at this redshift (see Figure 1).

2.1.2 GOODS-S

Two of the three field environment pointings are selected within the GOODS-S region (Guo et al. 2013); accessible from the VLT and with excellent multi-wavelength coverage, including deep *HST WFC3 F160W* imaging with a $0.06''$ pixel scale and $\simeq 0.2''$ PSF, which is well suited for constraining galaxy morphology (Grogin et al. 2011; Koekemoer et al. 2011). We selected targets from the various spectroscopic campaigns which have targeted GOODS-S, including measurements from VIMOS (Balestra et al. 2010; Cassata et al. 2015), FORS2 (Vanzella et al. 2005, 2006, 2008) and both LRIS and FORS2 as outlined in Wuyts et al. (2009). These targets must be within the redshift range $3 < z < 3.8$, have high spectral quality (as quantified by the VIMOS redshift flag equal ‘3’ or ‘4’, and the FORS2 quality flag equal ‘A’) and we carefully excluded those targets for which the [O III] $\lambda 5007$ or H β emission lines, observable in the *K*-band at these redshifts, would be shifted into a spectral region plagued by strong OH emission. The galaxies which remain after imposing these criteria are distributed across the GOODS-S field, and we selected two regions where $\simeq 20$ targets could be allocated to the KMOS IFUs (noting that the IFUs can patrol a $7.2'$ diameter patch of sky during a single pointing). We name these GOODS-S-P1 and GOODS-S-P2, which which we observe 20 and 17 galaxies respectively (see Table 1).¹

2.1.3 SSA22

A single cluster environment pointing was selected from the SSA22 field, (Steidel et al. 1998, 2000, 2003; Shapley et al. 2003), which, as mentioned above, is an overdensity of LBG candidates at $z \simeq 3.09$. Hundreds of spectroscopic redshifts have been confirmed for these LBGs with follow-up observations using LRIS (Shapley et al. 2003; Nestor et al. 2013). A combination of deep *B, V, R* band imaging with the Subaru Suprime-Cam (Matsuda et al. 2004), deep narrow-band imaging at 3640 (Matsuda et al. 2004) and at 4977 (Nestor et al. 2011; Yamada et al. 2012) and archival *HST ACS* and *WFC3* imaging provides ancillary data in excellent support of integral field spectroscopy, albeit over a shorter wavelength baseline and with shallower exposures than in the GOODS-S field. Fortunately at $z \simeq 3.09$ the [O III] $\lambda 5007$ line is shifted into a region of the *K*-band which is free from OH features and so for the cluster environment pointing we filled the KMOS IFUs with galaxies located towards the centre of the SSA22 protocluster (SSA22-P1).

We also added a further field environment pointing to the south of the main SSA22 spatial overdensity where the density of galaxies is typical of the field environment (SSA22-P2). In SSA22-P1 and SSA22-P2 we observe 19 and 21 galaxies respectively. In summary, we have chosen three field environment pointings and a single cluster environment pointing across GOODS-S and SSA22, comprising a total of 77 galaxies, as described in Table 1.²

¹ We note that the number of observed galaxies quoted for GOODS-S-P2 does not include two observed targets which were later found to have $z < 0.5$.

² Additional pointings in the COSMOS and UDS fields were originally scheduled as part of the GTO project, however 50% of the

Table 1. This table summarises the KDS pointing statistics for the full observed sample of 77 galaxies. The columns list the pointing name and galaxy environment probed, the central pointing coordinates, the number of observed, detected, resolved and merging objects as described in § 2.2.2, the waveband observed with KMOS, the exposure time and the PSF measured in the K -band.

Pointing	RA	DEC	N_{Obs}	N_{Det}	% Det.	N_{Res}	% Res	N_{Merg}	% Merg [*]	Band ^a	Exp (ks)	PSF ($''$) ^b
GOODS-S-P1(field)	03:32:25.910	-27:51:58.710	20	17	85	14	70	2	17	K	32.4	0.50
GOODS-S-P2(field)	03:32:32.217	-27:43:08.000	17	14	82	13	76	2	18	K	31.8	0.52
SSA22-P1(cluster)	22:17:11.867	+00:15:44.700	21	15	71	9	46	8	89	HK	38.1	0.62
SSA22-P2(field)	22:17:35.120	+00:09:30.500	19	17	89	12	63	2	18	HK	27.8	0.57

^a We also observed the two GOODS-S pointings in the H -band to cover the $[\text{O II}]\lambda 3727, 3729$ emission lines and a description of these observations will be given in a future work.

^b The PSF values correspond to measurements in the K -band.

^{*} Note that the Merger percentage is computed with respect to the number of resolved galaxies; the other percentages are computed with respect to the total number of galaxies observed in that pointing.

2.2 Observations and data reduction

Our data for the 77 KDS targets were observed using KMOS (Sharples et al. 2013), which is a second generation IFS mounted at the Nasmyth focus of UT1 at the VLT. The instrument has 24 moveable pickoff arms, each with an integrated IFU, which patrol a region $7.2'$ in diameter on the sky, providing considerable flexibility when selecting sources for a single pointing. The light from a set of 8 IFUs is dispersed by a single spectrograph and recorded on a $2\text{k}\times 2\text{k}$ Hawaii-2RG HgCdTe near-IR detector, so that the instrument is comprised of three effectively independent modules. Each IFU has 14×14 spatial pixels which are $0.2''$ in size, and the central wavelength of the K -band grating has a spectral resolution of $R \simeq 4200$ (H -band $R \simeq 4000$, HK -band $R \simeq 2000$).

2.2.1 Observations

To achieve the science goals of the KDS, the target galaxies at $3 < z < 3.8$ were observed in both the K -band, into which the $[\text{O III}]\lambda 5007$ and $\text{H}\beta$ lines are redshifted, and the H -band, into which the $[\text{O II}]\lambda 3727, \lambda 3729$ doublet is redshifted, allowing both dynamical and chemical abundance measurements. The GOODS-S pointings were observed in the H and K -bands separately, however, due to loss of observing time the SSA22 galaxies were observed with the KMOS HK filter, which has the disadvantage of effectively halving the spectral resolution, but allows for coverage of the H -band and K -band regions simultaneously.

We prepared each pointing using the KARMA tool (Wegner & Muschielok 2008), taking care to allocate at least one IFU to observations of a ‘control’ star closeby on the sky to allow for precise monitoring of the evolution of seeing conditions and the shift of the telescope away from the prescribed dither pattern (see § 2.2.2). For the four pointings described above and summarised in Table 1, we adopted the standard object-sky-object (OSO) nod-to-sky observation pattern, with 300s exposures and alternating $0.2''/0.1''$ dither pattern for increased spatial sampling around each of the target galaxies. This procedure allowed for datacube reconstruction with $0.1''$ size spaxels as

observing time was lost to bad weather during these visitor mode observations.

described in § 2.2.2.

The observations were carried out during ESO observing periods P92-P96 using Guaranteed Time Observations (Programme IDs: 092.A-0399(A), 093.A-0122(A,B), 094.A-0214(A,B), 095.A-0680(A,B), 096.A-0315(A,B,C)) with excellent seeing conditions. In GOODS-S-P1 and GOODS-S-P2 the median K -band seeing was $\simeq 0.5''$ and for the SSA22-cluster and SSA22-field pointings the K -band seeing ranged between $\simeq 0.55 - 0.65''$. We observed 17-21 $z \simeq 3.5$ targets in each field (see Table 1), with these numbers less than the available 24 arms for each pointing due to the combination of three broken pickoff arms during the P92/93 observing semesters and our requirement to observe at least one control star throughout an Observing Block (OB).

This paper is concerned with the spatially-resolved kinematics of the KDS galaxies. Consequently, we now focus exclusively on the spatially-resolved $[\text{O III}]\lambda 5007$ measurements in the K -band spectral window. The details of the H -band data reduction and corresponding metallicity analyses will be described in a future study.

2.2.2 Data reduction

The data reduction process primarily made use of the Software Package for Astronomical Reduction with KMOS, (SPARK; Davies et al. 2013), implemented using the ESO Recipe Execution Tool (ESOREX) (Freudling et al. 2013). In addition to the SPARK recipes, custom Python scripts were run at different stages of the pipeline and are described throughout this section.

The SPARK recipes were used to create dark frames and to flatfield, illumination correct and wavelength calibrate the raw data. An additional step, which is not part of the standard reduction process, was carried out at this stage, to address readout channel bias. Differences in the readout process within each 64-pixel wide channel on the detector image lead to varying flux baselines across each of the individual exposures. We corrected back to a uniform flux baseline across the detector image for each object exposure by identifying pixels which are not illuminated in every readout channel and subtracting their median value from the rest of the pixels in the channel. This is a separate issue to non-uniform illumination across the detectors, which is corrected by applying the illumination correction frame to the stacked datacubes.

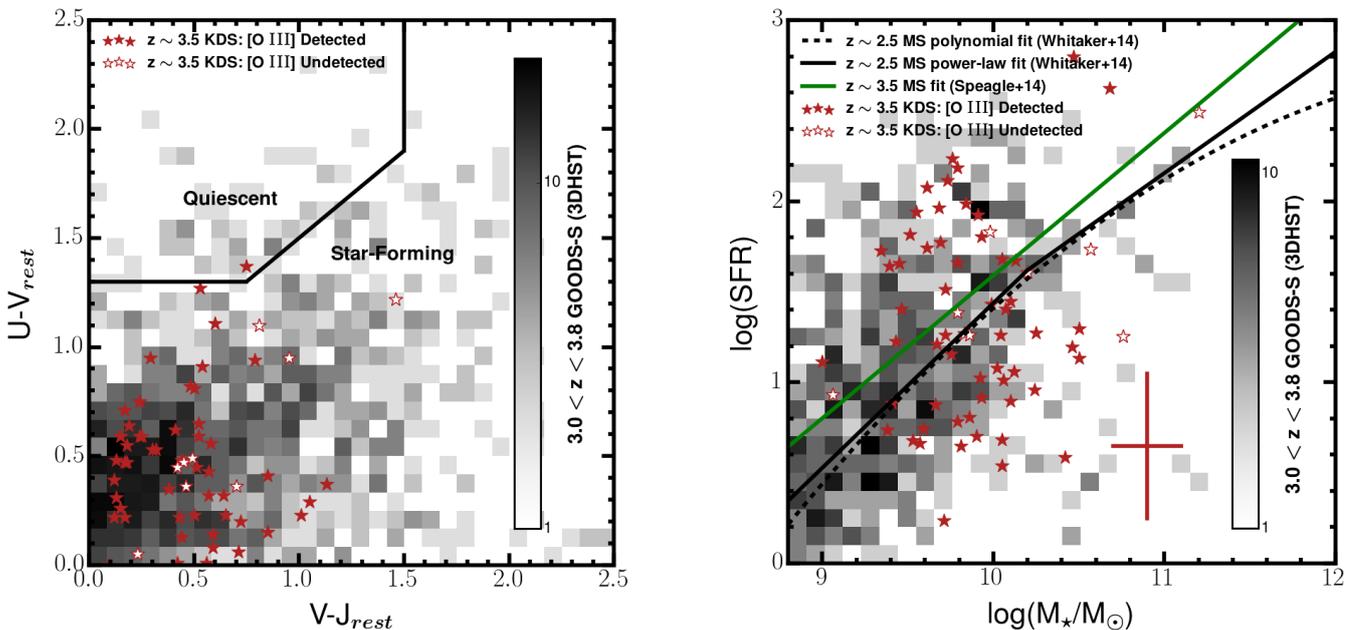


Figure 1. *Left:* The distribution of all 77 KDS galaxies in rest-frame UVJ colour space is plotted, with filled symbols showing galaxies detected in [O III] λ 5007 and open symbols showing those which were not detected. Also plotted in this plane are $\simeq 4000$ galaxies in GOODS-S with $3.0 < z < 3.8$ (mirroring the KDS redshift range) from the 3D-*HST* survey (Brammer et al. 2012; Momcheva et al. 2016), with the filled squares denoting the density of galaxies in that region. We use the galaxy selection criteria defined in Whitaker et al. (2012a) to highlight star-forming and quiescent regions (motivated by the age sequence of quiescent galaxies), finding that all but one of the KDS galaxies are clearly in the star-forming region and overlap with the highest density of 3D-*HST* targets. *Right:* We plot the location of the KDS galaxies in the SFR versus M_* plane, using the same symbol convention. The same GOODS-S galaxies from the 3D-*HST* survey as in the left panel are plotted with the filled squares, as a reference for the typical relationship between SFR and M_* . The black solid line and the dashed line show the $z \simeq 2.5$ broken power-law and quadratic fit to the main-sequence respectively, described in Whitaker et al. (2014). We include the MS relation evaluated at $z = 3.5$ (the median redshift of the KDS sample) described in Speagle et al. (2014) and given in Equation 1 as the green line, with the discrepancy between the two relations representing the expected main-sequence evolution between these redshifts. Within the typical uncertainties (see error bars) the KDS sample is representative of $z \simeq 3.5$ SFGs.

Standard star observations were carried out on the same night as the science observations and were processed in an identical manner to the science data. Following this pre-processing, each of the object exposures was reconstructed independently, using the closest sky exposure for subtraction, to give more control over the construction of the final stacks for each target galaxy. Each 300s exposure was reconstructed into a datacube with interpolated $0.1 \times 0.1''$ spaxel size, facilitated by the subpixel dither pattern discussed in § 2.2.1 which boosts the effective pixel scale of the observations.

Sky subtraction was enhanced using the SKYTWEAK option within SPARK (Davies 2007), which counters the varying amplitude of OH lines between exposures by scaling ‘families’ of OH lines independently to match the data. Wavelength miscalibration between exposures due to spectral flexure of the instrument is also accounted for by applying spectral shifts to the OH families during the procedure, and in general the use of the SKYTWEAK option in the *K*-band greatly reduces the sky-line residuals. We monitored the evolution of the atmospheric PSF and the position of the control stars over the OBs, to allow us to reject raw frames where the averaged *K*-band seeing rose above $0.8''$ and to measure the spatial shifts required for

the final stack more precisely. The PSF was determined by fitting the collapsed *K*-band image of the stacked control stars in each pointing with an elliptical gaussian, with the values reported in Table 1. The telescope tends to drift from its acquired position over the course of an OB and the difference between the dither pattern shifts and the measured position of the control stars provides the value by which each exposure must be shifted to create the stack.

We stacked all 300s exposures for each galaxy which pass the seeing criteria using 3-sigma clipping, leaving us with a flux and wavelength-calibrated datacube for every object in the KDS sample. We have found that the thermal background is often under-subtracted across the spatial extent of the cube following a first pass through the pipeline, leading to excess flux towards the long wavelength end of the *K*-band. To account for this, a polynomial function is fit, using the python package LMFIT (Newville et al. 2014), which makes use of the Levenberg-Marquardt algorithm for non-linear curve fitting, to the median stacked spectrum from spaxels in the datacube which contain no object flux and then subtracted from each spaxel in turn.

The central coordinates of each pointing, the number of target galaxies observed, N_{Obs} , the number of galaxies with

[O III] λ 5007 detected as measured by attempting to fit the redshifted line in the integrated galaxy spectrum using the known redshift value, $N_{\text{Det}} = 63/77$ (82%), the number with spatially-resolved [O III] λ 5007 emission, N_{Res} (see § 3.2.1), the on source exposure time and the averaged seeing conditions are listed in Table 1.

2.3 Stellar masses and SFRs

The wealth of ancillary data in both fields allows for a consistent treatment of the SED modelling, providing physical properties which are directly comparable between the cluster and field environments. These derived properties are considered in the context of the galaxy main-sequence, to verify that the KDS sample contains typical SFGs at $z \simeq 3.5$.

2.3.1 SED fitting and main-sequence

In order to constrain their SFRs and stellar masses, the available photometry for the KDS targets was analysed using the SED fitting software described in McLure et al. (2011) and McLeod et al. (2015). The photometry for each target was fit with the same set of solar metallicity BC03 (Bruzual & Charlot 2003) templates adopted by the 3D-HST team Momcheva et al. (2016), and derived stellar masses and SFRs were based on a Chabrier IMF. In addition, the SED fitting software accounts for the presence of strong nebular emission lines according to the line ratios determined by Cullen et al. (2014). During the SED fitting process, dust attenuation was accounted for using the Calzetti et al. (2000) reddening law, with dust attenuation allowed to vary freely within the range $0.0 < A_V < 4.0$. Based on the adopted template set, the median stellar mass for the full observed sample is $\log(M_\star/M_\odot) = 9.8$. Fitting the photometry of the KDS targets with 0.2 Z_\odot templates, rather than solar metallicity templates, typically reduces the derived stellar masses by $\simeq 0.1$ dex, but this change does not affect the conclusions of this work. In GOODS-S, we have compared our derived stellar masses to those in Santini et al. (2015) (which presents an average result from 10 different sets of analyses), finding a median difference between the two sets of values of $\Delta \log(M_\star/M_\odot) = 0.009$. We also note that using the star-forming galaxy templates described in Wuyts et al. (2011) typically leads to stellar masses which are 0.2 dex higher.

In the left panel of Figure 1 we plot the KDS galaxies in the rest-frame $U-V$ vs. $V-J$ colour space. This is a commonly used diagnostic plane for selecting star-forming and quiescent galaxies (e.g. Williams et al. 2009; Brammer et al. 2011; Whitaker et al. 2012a) with the age gradients of the stars within quiescent galaxies placing them in a different region of the plane to those which are actively forming stars. The selection criteria defined in Whitaker et al. (2012a) (which evolve only gently with redshift), separate quiescent and star-forming regions, which are indicated by the black wedge. We also make use of the rest-frame colours of $\simeq 4000$ primarily star-forming galaxies located in GOODS-S between $3.0 < z < 3.8$ (based upon the ‘z.best classification flag’) from the 3D-*HST* survey (Brammer et al. 2012; Momcheva et al. 2016). The filled squares indicate the density of 3D-*HST* targets in colour space and we observe that

the peak density location is consistent with the location of the KDS targets, all but one of which are in the star-forming region.

In the right panel of Figure 1 we plot the M_\star and SFR ‘main-sequence’ for the KDS galaxies with SFR measurements, in combination with the derived physical properties of the same GOODS-S galaxies as in the left panel. We also plot both the linear-break and quadratic $z \simeq 2.5$ main-sequence fits to the 3D-*HST* data described in Whitaker et al. (2014) with the solid and dashed black lines as well as the main-sequence relation described in Speagle et al. (2014) given in Equation 1, evaluated at $z \simeq 3.5$ (where the age of the universe is 1.77 Gyr) with the green line.

$$\log\text{SFR}(M_\star, t) = (0.84 - 0.026 \times t) \log(M_\star/M_\odot) - (6.51 - 0.11 \times t) \quad (1)$$

The difference in position of these relations highlights the main-sequence evolution towards higher SFRs at fixed M_\star between $z \simeq 2.5 - 3.5$. The KDS galaxies scatter, within the errors, consistently above and below the $z \simeq 3$ main-sequence.

When taken together, both panels indicate that the KDS sample is representative of typical star-forming galaxies at $z > 3$.

3 ANALYSIS

3.1 Morphological Measurements

For a robust interpretation of the observed velocity fields, it was necessary to separately determine the morphological properties of the galaxies from high-resolution images. This imaging was used primarily to determine morphological parameters which characterise the size (quantified here through the half-light radius, $R_{1/2}$), morphological position angle, PA_{morph} , and axis ratio, b/a , of the galaxies. In the following sections we describe the approach chosen to recover these parameters, also describing comparisons with matched galaxies in the morphological parameter catalogue of van der Wel et al. (2012). At $3 < z < 4$ and $0.1 < z < 1$ we made use of secure spectroscopic redshifts obtained for SFGs during the ESO public surveys zCOSMOS (Lilly et al. 2007), VUDS (Tasca et al. 2016), GOODS_FORSS2 (Vanzella et al. 2005, 2006, 2008) and GOODS_VIMOS (Balestra et al. 2010) to cross-match with van der Wel et al. (2012). This allowed us to investigate the morphological properties of typical star-forming galaxy populations at two redshift slices in comparison with those determined for the KDS sample. The imaging also helped to distinguish multiple ‘merging’ components with small angular separations from objects which are morphologically isolated, which we discuss in § 3.2.1 and § 3.3 where we refine our sample for dynamical analysis.

3.1.1 Applying GALFIT to the imaging data

We used GALFIT (Peng et al. 2010) to fit 2D analytic functions, convolved with the PSF, to the observed *HST* images of the KDS field galaxies across GOODS-S and SSA22 in

a consistent way. The GOODS-S imaging data used is the latest release of the total field in *WFC3 F160W* band, which traces the rest-frame near-UV at $z \simeq 3.5$, available via the CANDELS (Grogin et al. 2011; Koekemoer et al. 2011) data access portal³. For SSA22 we made use of archival *HST* imaging⁴ data in the *WFC3 F160W* band (P.I. Lehmer: PID 13844; P.I. Mannucci: PID 11735) and the *ACS F814W* band, tracing $\simeq 2500$ light at $z \simeq 3.1$ (P.I. Chapman: PID 10405; P.I. Abraham: PID 9760; P.I. Siana: PID 12527). The *HST* coverage is shallower in SSA22 (exposure times of $\simeq 5$ ks) and the *F160W* coverage is concentrated on the SSA22-cluster and so we resorted to the bluer *ACS F814W* data to derive morphological parameters in SSA22-P2 (also in SSA22-P2, 4 galaxies do not have any *HST* coverage as indicated in Table 2).

We first ran SExtractor (Bertin & Arnouts 1996) on the relevant images to recover initial input parameters and segmentation maps for running GALFIT, and then extracted postage stamp regions around the galaxies in the KDS sample. At this redshift, the galaxies are more compact and generally we could not resolve more complicated morphological features such as spiral arms and bars and so we followed the simple method of fitting Sérsic profiles, with the Sérsic index fixed to the exponential disk value of $n = 1$. In the *F160W* band, the adopted PSF was a hybrid between the Tiny Tim *H*-band model (Krist et al. 2011) in the PSF centre and an empirical stack of stars observed in the *H*-band for the wings (van der Wel et al. 2012) and in the *F814W* band we used the pure Tiny Tim *ACS* high-resolution PSF model. During the fitting process all other morphological parameters, including $R_{1/2}$, the central x and y coordinates, PA_{morph} and inclination, were free to vary.

This method was justified by the recovery of similar mean χ^2 values when fitting floating Sérsic index models and bulge/disk models with both an $n = 1$ and $n = 4$ component (following the procedure described in Bruce et al. 2012) to those recovered from the fixed exponential disk fit. Additionally, 24 GOODS-S objects were also detected in the van der Wel et al. (2012) catalogue, for which the median Sérsic index value is $n = 1.2$. Of these 24 galaxies, 2 are found to have $n > 2.5$ (lbg_32 and lbg_109, see Appendix D), although both show disk-like kinematics with a monotonic velocity gradient. The galaxy lbg_32 can be fit with an $n = 1$ Sérsic profile with small residuals, however lbg_109 has a centrally concentrated light distribution which is not well captured by the $n = 1$ profile. This demonstrates that fitting exponential flux distributions is a valid approach for the sample at large. The use of exponential profiles also facilitates comparison with recent large surveys such as KMOS^{3D} (Wisnioski et al. 2015) and KROSS (Harrison et al. 2017) in which beam-smearing correction factors were applied to the derived kinematic parameters as a function of exponential disk scale length, R_D , defined as $R_{1/2} \simeq 1.68R_D$ (see Appendix A4).

This analysis provided us with three crucial morphological parameters required to support the kinematic anal-

ysis of § 3.2.2, namely the axis ratio b/a , half-light radius $R_{1/2}$ and the morphological position angle PA_{morph} . Example *HST* images and GALFIT outputs are given in Figure 3 for selected galaxies, and maps for the full isolated field sample (see § 3.2.1) are plotted throughout Appendix D.

The errorbars produced by GALFIT are purely statistical and are determined from the covariance matrix used in the fitting, resulting in unrealistically small uncertainties on the derived galaxy parameters (Häußler et al. 2007; Bruce et al. 2012). Throughout the following subsections we discuss the more reasonable adopted errors on each of the relevant morphological parameters in turn. In each of the following sections we also discuss the approach followed to recover the morphological parameters for the galaxies without *HST* coverage. The interpretation of the sample as a whole is not affected by the small number of galaxies without *HST* imaging.

3.1.2 Inclination angles

We used the derived b/a values to determine galaxy inclination angles. As suggested in Holmberg (1958), by modelling the disk galaxies as an oblate spheroid, the inclination angle can be recovered from the observed axis ratio as shown in Equation 2, where $\frac{b}{a}$ is the ratio of minor to major axis of an ellipse fit to the galaxy profile on the sky, i is the inclination angle and q_0 is the axis ratio of an edge-on system.

$$\cos^2 i = \frac{\left(\frac{b}{a}\right)^2 - q_0^2}{1 - q_0^2} \quad (2)$$

To derive the inclination, we selected a value for q_0 , and following the discussion in (Law et al. 2012a) we chose a value appropriate for thick disks, $q_0 = 0.2$ (e.g. Epinat et al. 2012; Harrison et al. 2017). However, as discussed in Harrison et al. (2017), varying the q_0 value by a factor of 2 makes only a small change to the final inclination corrected velocity values, with the difference $< 10\%$ in the case of the KDS galaxies in the isolated field sample (see § 3.3).

The inclination angle calculated for each galaxy is used to correct the observed velocity field, which is the line of sight component of the intrinsic velocity field, with the correction factor increasing with increasing b/a . The median difference between the derived b/a values and those presented in van der Wel et al. (2012) for the 22 matched GOODS-S galaxies is $\Delta b/a = 0.0002$. The b/a distribution shown in Figure 2 (plotted for the 28 galaxies in the isolated field sample with *HST* imaging, § 3.3) is consistent with being uniform between $0.3 < b/a < 0.9$, with a median value of 0.58, corresponding to $i = 57.0^\circ$ using the conversion given in Equation 2, with $q_0 = 0.2$. This compares well with the theoretical mean value of 57.3° computed for a population of galaxies with randomly drawn viewing angles (see e.g. the appendix in Law et al. 2009), which reassured us that we were not biased towards deriving either particularly low or particularly high inclination angles.

As mentioned in § 3.1.1, the error reported by GALFIT on the model b/a value is unrealistically small. Guided by the results of the Monte Carlo approach described in Epinat et al. (2012) we adopted a conservative nominal in-

³ http://candels.ucolick.org/data_access/Latest_Release.html

⁴ <https://archive.stsci.edu/hst/search.php>

Table 2. Physical properties of the resolved and morphologically isolated KDS field galaxies as measured from SED fitting and from applying GALFIT (Peng et al. 2010).

ID	RA	Dec	z	K_{AB}^a	$\log(M_\star/M_\odot)^b$	b/a	i^{oc}	PA_{morph}^o	$R_{1/2}(\text{kpc})^d$
b012141_012208	03:32:23.290	-27:51:57.348	3.471	24.12	9.8	0.36	72	9	1.57
b15573	03:32:27.638	-27:50:59.676	3.583	23.60	9.8	0.28	78	146	0.52
bs006516	03:32:14.791	-27:50:46.500	3.215	23.94	9.8	0.50	61	146	1.91
bs006541	03:32:14.820	-27:52:04.620	3.475	23.44	10.1	0.44	66	168	1.83
bs008543	03:32:17.890	-27:50:50.136	3.474	22.73	10.5	0.50	61	67	1.59
bs009818	03:32:19.810	-27:53:00.852	3.706	24.18	9.7	0.80	37	148	1.24
bs014828	03:32:26.760	-27:52:25.896	3.562	23.58	9.7	0.31	76	63	1.61
bs016759	03:32:29.141	-27:48:52.596	3.602	23.85	9.9	0.65	50	49	0.87
lbg_20	03:32:41.244	-27:52:20.676	3.225	24.97	9.5	0.64	52	1	1.28
lbg_24	03:32:39.754	-27:39:56.628	3.279	24.67	9.6	0.53	60	34	1.27
lbg_25	03:32:29.189	-27:40:22.476	3.322	24.95	9.4	0.30	76	78	1.18
lbg_30	03:32:42.854	-27:42:06.300	3.419	23.85	10.0	0.79	38	66	0.95
lbg_32	03:32:34.399	-27:41:24.324	3.417	23.84	9.9	0.60	54	40	1.88
lbg_38	03:32:22.474	-27:44:38.436	3.488	24.58	9.9	0.58	56	137	0.92
lbg_91	03:32:27.202	-27:41:51.756	3.170	24.65	9.8	0.58	56	79	0.89
lbg_94	03:32:28.949	-27:44:11.688	3.367	24.54	9.9	0.22	84	81	1.16
lbg_105	03:32:24.005	-27:52:16.140	3.092	23.79	9.3	0.56	57	128	1.72
lbg_109	03:32:20.935	-27:43:46.344	3.600	24.63	9.7	0.60	54	119	1.98
lbg_111	03:32:42.497	-27:45:51.696	3.609	24.01	9.7	0.74	42	80	0.64
lbg_112	03:32:17.134	-27:42:17.784	3.617	25.16	9.6	0.79	38	43	0.46
lbg_113	03:32:35.957	-27:41:49.956	3.622	24.01	9.6	0.52	60	15	0.87
lbg_121	03:32:19.606	-27:48:40.032	3.708	25.36	9.1	0.65	50	106	0.42
lbg_124	03:32:33.324	-27:50:07.332	3.794	24.96	9.0	0.56	57	50	0.78
*n3_006	22:17:24.859	+00:11:17.620	3.069	22.98	10.5	0.57	57	156	2.52
n3_009	22:17:28.330	+00:12:11.600	3.069	-	8.7	0.75	42	84	1.06
n_c3	22:17:32.585	+00:10:57.180	3.096	24.96	9.8	0.71	46	94	0.56
lab18	22:17:28.850	+00:07:51.800	3.101	-	8.2	0.85	32	27	0.46
*lab25	22:17:22.603	+00:15:51.330	3.067	-	8.4	0.57	57	87	2.18
s_sa22a-d3	22:17:32.453	+00:11:32.920	3.069	23.46	9.7	0.39	70	125	1.78
s_sa22b-c20	22:17:48.845	+00:10:13.840	3.196	23.91	9.5	0.57	57	76	1.59
*s_sa22b-d5	22:17:35.808	+00:06:10.340	3.175	23.72	10.2	0.57	57	60	3.30
s_sa22b-d9	22:17:22.303	+00:08:04.130	3.084	24.25	10.1	0.65	50	60	0.50
*s_sa22b-md25	22:17:41.690	+00:06:20.460	3.304	24.62	8.6	0.57	57	9	1.39

^a K-band magnitude errors are typically ± 0.05 magnitudes.

^b Representative error of 0.2 dex from SED modelling used throughout this study (see § 2.3.1).

^c Fixed 10 percent inclination error assumed (i.e. $\delta i = i/10$, see § 3.1.2)

^d Fixed 10 percent $R_{1/2}$ error assumed (see § 3.1.4).

* No *HST* coverage: PA_{morph} set to PA_{kin} value; b/a and $R_{1/2}$ estimated as explained throughout the text.

clination angle uncertainty of 10%, i.e. $\delta i = i/10$. For the galaxies without *HST* coverage, the inclination angle was set to the KDS sample median of 57.0° and we used the standard deviation of the KDS inclinations, 13.2° , for the inclination uncertainty. This added an additional factor of $\simeq 1.3$ to the uncertainty in the derived velocities for these galaxies.

3.1.3 Position angles

The second GALFIT parameter was PA_{morph} , which is the direction of the photometric major axis of the galaxy on the sky. We visually inspected each of the GALFIT maps to check that the derived PA_{morph} follows the galaxy light distribution (with PA_{morph} indicated by the orange dashed line in Figure 3) and when this was not the case it was usually an indication of multiple distinct components or tidal streams (see § 3.3 where we remove multiple-component objects from the final sample). We note that discrepancy between PA_{morph} and the kinematic position angle, PA_{kin} , is an indi-

cator of sub-structure in the morphology (e.g. Queyrel et al. 2012; Wisnioski et al. 2015; Rodrigues et al. 2017), and deviations can indicate clumps or mergers which may influence the underlying kinematics or bias the derived PA_{morph} towards a particular direction. We discuss this topic further in § 3.2.6. For the galaxies without *HST* coverage we fixed PA_{morph} equal to PA_{kin} , for the analysis described in § 3.2.2.

3.1.4 Half-light radii

The half-light radius, $R_{1/2}$, provides an indication of the disk sizes, and hence gave us a common fiducial distance from the centre of the galaxy at which to extract rotation velocities. As discussed throughout the introduction and in Appendix A, methods used to derive the intrinsic rotation velocity, V_C , vary widely between studies, and so to compare with previous work it is necessary to extract kinematic parameters consistently from the derived rotation curves. Knowledge of $R_{1/2}$ gives a well defined and consistent point across the extent of the galaxy at which

to extract velocities, which is important when making comparisons to previous IFU studies (see Appendix A and e.g. Förster Schreiber et al. 2009; Epinat et al. 2012; Wisnioski et al. 2015; Stott et al. 2016; Harrison et al. 2017; Swinbank et al. 2017). Throughout our dynamical modelling, and to derive a beam-smearing correction to the rotation velocities and velocity dispersions (see § 3.2.2), we also required the $R_{1/2}$ value for each galaxy. The median difference between the derived $R_{1/2}$ values and those presented in van der Wel et al. (2012) for the 22 matched GOODS-S galaxies is $\Delta R_{1/2} = 0.0054''$ (corresponding to 0.04kpc at $z \simeq 3.5$).

Bruce et al. (2012) present a detailed error analysis of the morphological parameters of $1 < z < 3$ galaxies in the CANDELS fields (which includes *HST* data covering the GOODS-S field which we exploited during this work), reporting that the magnitude of the error on $R_{1/2}$ is typically at the 10% level. We adopted a nominal error of 10% for our recovered $R_{1/2}$ values and note that the errors on V_C and σ_{int} are dominated by measurement errors and uncertainties connected with assumptions about the inclination angle correction and the velocity dispersion distribution in high-redshift galaxies (see Appendix C).

To measure $R_{1/2}$ for the 4 galaxies without *HST* coverage, we fit the [O III] λ 5007 narrow-band images with an exponential profile, and subtracted the appropriate PSF from the recovered $R_{1/2}$ value. We note that the extracted rotation velocities are not sensitive to the precise $R_{1/2}$ values, since we measure these at $2R_{1/2}$ (see § 3.2.4) where the intrinsic model rotation curves have already reached flattening.

3.2 Kinematic Measurements

We proceeded to extract 2D kinematic measurements from the 63 KDS datacubes detected in integrated [O III] λ 5007 emission by fitting the [O III] λ 5007 emission line profiles within individual spaxels. The sample was then refined to galaxies which are spatially-resolved in the [O III] λ 5007 emission line in order to make dynamical measurements. We also describe our dynamical modelling and the approach we have followed to extract V_C and σ_{int} , using the models to derive beam-smearing corrections for each galaxy. We made use of the imaging data described in the previous sections to reduce the number of free parameters in the model, an approach which is necessary when working in this signal-to-noise regime.

3.2.1 Spaxel Fitting

We aimed to extract two-dimensional maps of the [O III] λ 5007 flux, velocity and velocity dispersion from each of the 63 stacked and calibrated KDS datacubes with an [O III] λ 5007 detection. These properties are extracted via modelling emission line profiles within each spaxel, with a set of acceptance criteria imposed to determine whether the fit quality is high enough to allow the inferred properties to pass into the final 2D maps of the flux and kinematics. We concentrated solely on the [O III] λ 5007 emission line which always had signal-to-noise higher than both [O III] λ 4959

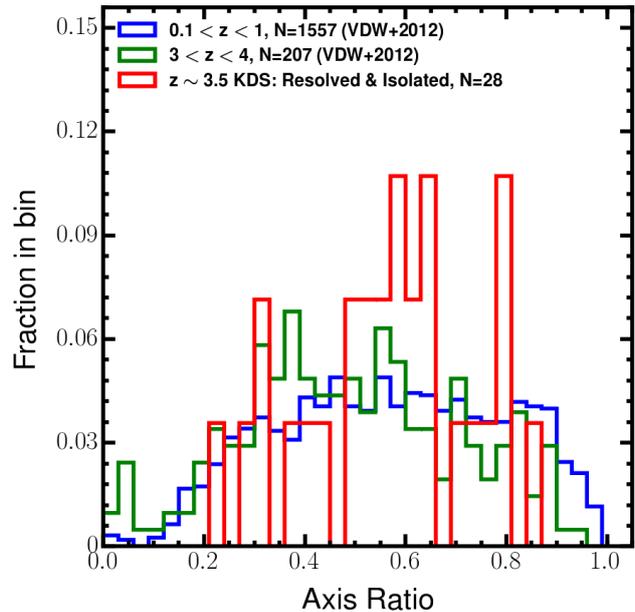


Figure 2. The normalised counts of KDS galaxies in bins of axis ratio is plotted, along with the normalised counts from the morphological catalogue presented in van der Wel et al. (2012) in two redshift ranges. The axis ratio distribution appears to be constant with time as traced by the reference samples, and the KDS values are in good agreement with a relatively uniform distribution spanning $0.3 < b/a < 0.9$. This suggests that we are not biased towards deriving a particular value for the axis ratio.

and $H\beta$ and was detected over a larger area of each IFU.

Each $0.1''$ spaxel across a datacube was considered in turn. The redshift value determined from fitting the [O III] λ 5007 line in the integrated galaxy spectrum, z_{sys} , was used to determine the central wavelength of the redshifted [O III] λ 5007 line, $\lambda_{\text{obs}5007}$. We searched a narrow region around $\lambda_{\text{obs}5007}$ for the peak flux, assumed to correspond to the ionised gas emission, and extracted a 5 region (corresponding to 20 spectral elements) centred on this peak. This truncated spectrum was then used in the fitting procedure for each spaxel. The width of the extraction region was large enough to encompass unphysically large velocity shifts, but not so large as to compromise the fitting by potentially encompassing regions of the spectrum plagued by strong sky-line residuals. A single gaussian model was fit to the extraction region, again using LMFIT (Newville et al. 2014), returning the values of the best-fitting flux, F_g , central wavelength, λ_g , and dispersion, σ_g . These parameters, along with the estimate of the noise in the datacube, were used to assess whether the fitting of the [O III] λ 5007 line is acceptable or not.

The noise was estimated by masking all datacube spaxels containing object flux and computing the standard deviation of the flux values in the spectral extraction region for all remaining spaxels inwards of $0.3''$ from the cube boundary (over the same spectral region as considered to find the peak object flux). This boundary constraint was chosen to mitigate edge effects, where the noise increased sharply due to fewer individual exposures constituting the final stack

in these regions. The final noise value was then taken as the standard deviation of the results from each unmasked spaxel, as shown in Equation 3, where $F_{i,\text{background}}$ are the flux values across the spatial regions containing no object flux in the spectral extraction region, and N_m denotes the noise estimated from masking the datacube.

$$N_m = \text{STD} \left(\sum_i F_{i,\text{background}}^2 \right) \quad (3)$$

The signal in each spaxel was taken as the sum of the flux values across the spectral extraction region, $F_{i,\text{object}}$, as shown in Equation 4, where S denotes the signal.

$$S = \sum_i F_{i,\text{object}} \quad (4)$$

The criteria for accepting the fit within a given spaxel are as follows:

- (i) The uncertainties on any of the parameters of the gaussian-fit as returned by LMFIT must not exceed 30%
- (ii) $\frac{S}{N_m} > 3$

We imposed a further criterion to help remove fits to spectral features unrelated to the galaxy (e.g. fits to skyline residuals), which was important for ‘cleaning up’ accepted spaxels in the 2D maps which were clearly unrelated to the galaxy. For this test we subtracted the galaxy continuum and imposed that $0.7 < F_g/S < 1.3$, where F_g is the flux from the gaussian-fit to the [O III] λ 5007 line and S is the signal in the extraction region from Equation 4. If all three criteria were satisfied, the spaxel was accepted.

From the gaussian-fit, we recovered the centre, λ_g , the width, σ_g , and the area, F_g . The velocity was computed from λ_g using Equation 5 and the velocity dispersion was computed from σ_g using Equation 6, where V_{obs} and σ_{obs} are the observed rotation velocity and velocity dispersion respectively, c is the speed of light, z_{sys} is the redshift determined from the integrated galaxy spectrum, $\lambda_{5007} = 0.500824\mu\text{m}$ is the rest wavelength of [O III] λ 5007 in vacuum and σ_{inst} is a measure of the KMOS instrumental resolution value averaged across gaussian-fits to several skylines (equal to 31.1 km s^{-1} in the K -band and 55.4 km s^{-1} in the HK -band).

$$v_{\text{obs}} = \frac{\lambda_g - (1 + z_{\text{sys}})\lambda_{5007}}{(1 + z_{\text{sys}})\lambda_{5007}} \times c \quad (5)$$

$$\sigma_{\text{obs}} = \sqrt{\left(\frac{\sigma_g}{(1 + z_{\text{sys}})\lambda_{5007}} \times c \right)^2 - \sigma_{\text{inst}}^2} \quad (6)$$

To estimate the errors on the observed quantities, the gaussian-fit was repeated 1000 times, with each flux value in the extraction region perturbed by the noise value, N_m . For each of these 1000 fits the gaussian parameters were used to compute the observed properties. The resultant distributions of observed quantities were gaussian and we measured the median as the final value for V_{obs} , σ_{obs} and F_g within that spaxel, and the standard deviation of the distribution as the observational uncertainty. A map of each

of these observed parameters was constructed by applying this procedure to each spaxel.

If a spaxel did not meet the criteria, we follow the procedure described in Stott et al. (2016) and binned with neighbouring spaxels to create a box of area $0.3 \times 0.3''$. This was carried out by median stacking the extraction region spectra in each of the neighbouring spaxels. The criteria were then re-examined, and if satisfied, the Monte-Carlo procedure was carried out. If the binned spaxels fail the criteria, one final iteration was carried out to bin to a box of area $0.5 \times 0.5''$, roughly equivalent to the size of the seeing disk (see Table 1). If after the widest binning the criteria weren’t satisfied, no value was assigned to that spaxel in the extracted maps.

The process was automated and applied to each of the 63/77 KDS galaxies with an [O III] λ 5007 detection. Following the fitting procedure, the galaxies were classified as spatially-resolved in [O III] λ 5007 emission or not if the spatial extent of the accepted spaxels in the [O III] λ 5007 flux map was equivalent to or greater than the seeing disk. For each pointing, the number of resolved galaxies, N_{Res} , was added to Table 1. We find that 48/63 (76%) of the detected KDS galaxies across both the field (39/48) and cluster (9/48) environments were spatially-resolved and showed clear velocity structure. The 15/63 (24%) galaxies which were not spatially-resolved were dropped at this stage due to the uncertainties associated with deriving kinematic properties and classifications from a single resolution element. The physical properties of all spatially-resolved galaxies in both GOODS-S and SSA22 are listed in Table 2.

The morphological information of § 3.1 was combined with this spatially-resolved condition to define two galaxy subsamples. 6/39 (15%) of the resolved galaxies across the 3 field environment pointings and almost all (8/9) of the cluster environment galaxies in SSA22-P1 have multiple photometric components, which complicate the interpretation of the observed dynamics for these targets. We dropped the cluster environment pointing from the full dynamical analysis, although we show examples of the *HST* imaging, flux, dispersion and velocity maps of the cluster target galaxies in Figure D3. The remaining 6/39 merging candidates in the field environment are referred to as the ‘merger field sample’ throughout the remainder of the paper. The 33/39 morphologically isolated field galaxies which are spatially resolved are referred to as the ‘isolated field sample’ throughout the remainder of the paper.

3.2.2 3D modelling and beam-smearing corrections

As mentioned in the introduction, there is no universal standard for defining V_C and σ_{int} , particularly when a dynamical model is used to provide beam-smearing corrections for σ_{int} and to extrapolate beyond the data in order to measure V_C . Physically motivated models take into account the potential-well of the galaxy through knowledge of the mass distribution (e.g. Genzel et al. 2008; Förster Schreiber et al. 2009; Gnerucci et al. 2011; Wisnioski et al. 2015; Swinbank et al. 2017). Phenomenological models assume a fixed function known to reproduce observed galactic rotation curves

which flatten at large radii (e.g. [Epinat et al. 2010, 2012](#); [Swinbank et al. 2012a](#); [Stott et al. 2016](#); [Harrison et al. 2017](#)). Both of the above examples fit to 1D or 2D fields extracted from the datacubes using line-fitting software. Recently, algorithms such as Galpak^{3D} ([Bouché et al. 2015](#)) and 3D-Barolo ([Di Teodoro & Fraternali 2015](#)) have emerged, fitting directly to the flux in the datacube in order to constrain the kinematic properties.

There are three reasons to construct model velocity fields for the isolated field sample galaxies; the first is to interpolate between data points; the second is to extrapolate the observations smoothly to some fiducial radius from which V_C can be extracted; the third is to estimate the effect of beam-smearing on both the velocity and velocity dispersion fields so that we recover the set of intrinsic galaxy parameters which best describe the observations. As a brief overview, we used a Markov chain Monte Carlo (MCMC) procedure to construct a series of intrinsic model datacubes for each galaxy, convolved these with the atmospheric PSF and then fit to the data. The beam-smearing corrected kinematic parameters were then extracted from the intrinsic galaxy models as described in more detail below. In this analysis, we minimised the number of free parameters by making use of information from *HST* imaging and by assuming a fixed function to describe the rotation curves as described in the following subsection. This is particularly important when dealing with galaxies at $z > 3$, since often there is not enough signal-to-noise away from the centre of the galaxy to observe the flattening of the rotation curve, consequently requiring extrapolation to determine V_C . The following section describes our 3D modelling procedure, and the validation of our results through comparison with the techniques used to determine the dynamical properties presented in [Harrison et al. \(2017\)](#).

3.2.3 Constructing model datacubes

For each galaxy with spatially-resolved [O III] λ 5007 emission we constructed a series of model datacubes with the same spatial dimensions as the observed datacube and populated each spaxel with an [O III] λ 5007 emission line that has central wavelength determined using z_{sys} . This central wavelength was then shifted using the velocity derived at each spaxel from the velocity field model described below. Following the procedure adopted by numerous authors (e.g. [Epinat et al. 2010, 2012](#); [Swinbank et al. 2012a](#); [Stott et al. 2016](#); [Mason et al. 2017](#)) the velocity field of the gas was modelled as a thin disk with the discrete velocity points along PA_{kin} determined by Equation 7, which has been found to reproduce the rotation curves of galaxies in the local universe (e.g. [Courteau 1997](#)):

$$v_r = \frac{2}{\pi} v_{\text{asym}} \arctan\left(\frac{r}{r_t}\right) \quad (7)$$

with ‘ r ’ measured as the distance from the centre of rotation. This model was then projected onto the cube using the inclination angle determined in § 3.1.1 and sampled with 4 times finer pixel scale than the KMOS raw data before binning back to the previous resolution. The pixel scale refinement was necessary to capture the steep velocity gradients across the central regions.

The velocity decreases with a factor of $\cos(\varphi)$ from the kinematic position axis, where φ is the angle measured clockwise from the axis, which points to the direction of the positive side of the velocity field. The 6 free parameters of the model; $(x_{\text{cen}}, y_{\text{cen}})$, i , PA_{kin} , r_t , v_{asym} , were first reduced to 5 by using the inclination from the *HST* axis ratio. We also fixed the $(x_{\text{cen}}, y_{\text{cen}})$ parameters to the location of the stellar continuum peak within the collapsed cubes, a proxy for the centre of the gravitational potential-well, to leave only 3 free parameters to vary in the MCMC sampling. Due to the faintness of the continuum, there are several galaxies for which we could not reliably estimate the rotation centre using this method and instead set $(x_{\text{cen}}, y_{\text{cen}})$ to coincide with the peak [O III] λ 5007 flux location (approach described in [Harrison et al. 2017](#)). Using the continuum to fix the rotation centre was generally more successful for the SSA22 galaxies, since the use of the *HK* filter increases the signal-to-noise of the continuum by collapsing over twice the wavelength range.

The intrinsic flux profile of these simulated [O III] λ 5007 lines in the spatial direction was determined using the GALFIT model derived in § 3.1.1 (which assumes that the [O III] λ 5007 emission profile follows the stellar profile) and the intrinsic velocity dispersion of the line was set to follow a uniform distribution with a default width of 50 km s^{-1} . This default line width was selected following examination of figure 10 of [Wisnioski et al. \(2015\)](#), which places observational constraints on the velocity dispersions of typical star-forming galaxies at $z \simeq 3.0$, and we discuss the impact of this assumption in Appendix C1 by using sensible alternative values. The reported errors on σ_{int} take this assumption into account.

The simulated, intrinsic [O III] λ 5007 cube was then convolved slice by slice using fast Fourier transform libraries with the *K*-band atmospheric seeing profile (Table 1). The velocity and velocity dispersion values were then re-measured in each spaxel to produce ‘beam-smear’d 2D maps of the kinematic parameters. We fit the beam-smear’d velocity field to the observed velocity field, using MCMC sampling with the python package ‘emcee’ ([Foreman-Mackey et al. 2013](#)) to vary the intrinsic model parameters, seeking the combination of parameters which maximises the log-Likelihood function given by Equation 8, which fully accounts for the errors on the observed velocity field:

$$\ln \mathcal{L} = \frac{-0.5 \sum_{i=1}^N (d_i - M_i)^2}{\delta v_i^2} - \ln \left(\frac{1}{\delta v_i^2} \right) \quad (8)$$

where δv_i are the observed velocity errors, d_i are the datapoints and M_i are the convolved model velocity values.

The MCMC sampling provided a distribution of parameter values around those which give the maximum-likelihood. We confirmed also by examining the individual parameter chains that the MCMC run has been properly ‘burned-in’ and that the parameter estimates converged around the 50th percentile values. The 16th and 84th percentiles of these distributions were used to assess the model uncertainties as described in Appendix C1.

Following the construction of a convolved model cube

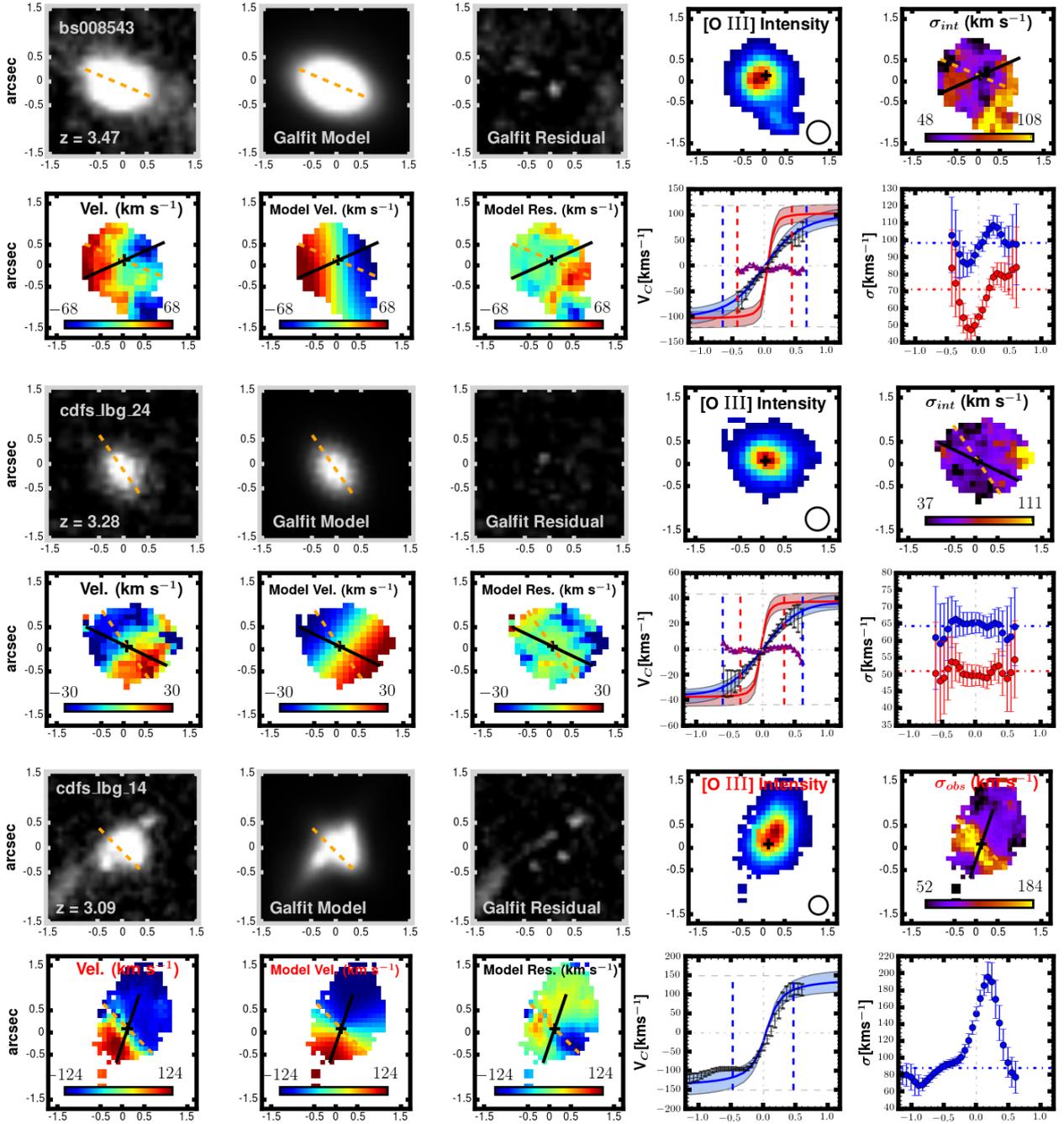


Figure 3. We plot the kinematic grids for one rotation-dominated galaxy (top grid), one dispersion-dominated galaxy (middle grid) and one merger candidate (bottom grid) (see § 3.3) from the KDS. In Appendix D we include these plots for each of the spatially-resolved galaxies in the KDS, divided into rotation-dominated (14/39), dispersion-dominated (19/39) and merger categories (6/39). For each galaxy we plot on the top row the *HST* image, GALFIT model and GALFIT residuals, along with the observed [O III] λ 5007 intensity map and velocity dispersion map. On the bottom row we plot the observed velocity map, the beam-smearred model velocity and data-model velocity residuals with colour bar limits tuned to the velocity model. The solid black line and dashed orange line plotted in various panels show PA_{kin} and PA_{morph} respectively. Also on the bottom row, we plot extractions along PA_{kin} for both the velocity and velocity dispersion. The grey points with errorbars on the velocity extraction plot are from the observed velocity map, the blue line and blue shaded regions represent the beam-smearing model fit to the data and errors respectively, whereas the red line and shaded regions are the intrinsic model from which V_C is extracted. The purple symbols represent the extraction along PA_{kin} from the residual map and the red and blue dashed lines denote the intrinsic and convolved $2R_{1/2}$ values respectively, with V_C extracted at the intrinsic value and V_{BS} extracted at the convolved value of the intrinsic and observed profiles respectively. The intrinsic models flatten at small radii, and so small changes in $R_{1/2}$ have negligible impact on the final, extracted V_C values. The blue points on the velocity dispersion extraction plot show the values extracted from the observed dispersion map and the red are from the beam-smearing corrected map as per Equation 6. Also on the 1D dispersion plot we include with the blue and red dot-dash horizontal lines median values of the observed and intrinsic (i.e. σ_{int}) dispersion maps respectively. The merger candidate in the bottom grid requires several GALFIT components to describe the data, mimics rotation from a purely kinematic perspective and we plot the 1D extractions without beam-smearing corrections, which are not warranted for multiple component objects.

and the measurement of the beam-smearred maps, we extract beam-smearing corrected values for V_C and σ_{int} , as described throughout the following subsections. There have been several different methods used to extract V_C and σ_{int} in previous surveys at different redshifts. We select a set of comparison samples with similar galaxy selection criteria and kinematic parameter extraction methods to those described above in order to make comparisons with the KDS across different redshift slices (see § 4). We have re-evaluated the means and medians of the kinematic properties in these samples in order to make fair and reliable comparisons. An overview of the comparison samples is provided in Appendix A, which also describes the approach used to compute average parameter values.

3.2.4 Beam-smearing corrected rotation velocities

The intrinsic velocity, V_C , was found by reading off from the best-fitting intrinsic 2D model grid at the intrinsic $2R_{1/2}$ value ($\simeq 3.4R_D$, where R_D is the scale length for an exponential disk). This extraction radius is shown with the vertical red dashed line in the panels of Figure 3 showing the 1D velocity extraction. The V_C value extracted at $2R_{1/2}$ is a commonly used measure of the ‘peak’ of the rotation curve (e.g. Miller et al. 2011; Stott et al. 2016; Pelliccia et al. 2017; Harrison et al. 2017; Swinbank et al. 2017) and extracting the velocity at $\geq 2R_{1/2}$ can be crucial for measuring the majority of the total angular momentum (e.g. Obreschkow et al. 2016; Harrison et al. 2017; Swinbank et al. 2017). This is similar to the measurement of the velocity at the radius enclosing 80% of the total light, V_{80} ($\simeq 3R_D$), (Tiley et al. 2016).

The mean and median V_C values for the isolated field sample are $76.7^{+4.9}_{-4.5} \text{ km s}^{-1}$ and $57.0^{+6.6}_{-6.3} \text{ km s}^{-1}$ respectively, which are less than observed for lower redshift typical star-forming galaxies (see § 4.1 for a discussion). In order to determine the beam-smearing correction factor, we also extracted the velocity from the beam-smearred model at the convolved $2R_{1/2}$ radius, represented by the vertical blue dashed line in Figure 3, to recover the beam-smearred velocity, V_{BS} . The magnitude of the beam-smearing correction, (V_C/V_{BS}), has a mean and median value of 1.29 and 1.21 for the isolated field sample and we discuss the observational and model uncertainties which combine to give the V_C errors in Appendix C1.

In 14/39 of the isolated field sample galaxies, the intrinsic $2R_{1/2}$ fiducial radius used to extract V_C from the model was greater than the last observed radius, however only 2 galaxies required extrapolation $> 0.4''$ (i.e. 2 KMOS pixels) with a mean extrapolation of $0.17''$. The extent of the extrapolation for each galaxy is made clear throughout the kinematics plots in Appendix D, with the intrinsic and convolved $2R_{1/2}$ extraction radii marked on each velocity thumbnail.

3.2.5 Beam-smearing corrected intrinsic dispersions

Convolution of the model velocity field with the seeing profile produces a peak in the beam-smearred velocity dispersion map at the centre of rotation. The model therefore provides

a beam-smearing correction for the velocity dispersion in every spaxel. This is dependent primarily on the magnitude of the velocity gradient across the parameter maps, where larger velocity gradients result in larger central velocity dispersions.

By comparing the best-fitting simulated datacube with the assumption for the isotropic velocity dispersion, we recovered the beam-smearing correction for each galaxy. The 2D beam-smearing correction map, σ_{BS} , was computed on a spaxel-by-spaxel basis as the difference between the beam-smearred dispersion produced by the model and the assumed intrinsic velocity dispersion of 50 km s^{-1} i.e. $\sigma_{\text{BS},i} = \sigma_{\text{simulated},i} - \sigma_{\text{assumed},i}$ (where $\sigma_{\text{assumed},i}$ was fixed at 50 km s^{-1} as described in § 3.2.3 and $\sigma_{\text{simulated},i}$ is the simulated velocity dispersion recovered under this assumption). The impact of the value of σ_{assumed} on the magnitude of the $\sigma_{\text{BS},i}$ correction is discussed further in Appendix C1 and the σ_{int} errors take this into account. We note also that not all of the observed σ_{int} profiles peak in the centre, and since we did not attempt to fit the velocity dispersions, this is a feature the model could not reproduce.

To measure the σ_{int} value for each galaxy in our sample, we applied Equation 9 in each spaxel to derive an intrinsic sigma map, where $\sigma_{g,i}$ is the value in the i th spaxel recovered from the gaussian-fit to the [O III] λ 5007 emission line.

$$\sigma_{\text{cor},i} = \sqrt{(\sigma_{g,i} - \sigma_{\text{BS},i})^2 - \sigma_{\text{inst},i}^2} \quad (9)$$

In Equation 9 we have subtracted the beam-smearing correction linearly following the analysis presented in Appendix A of Stott et al. (2016), which shows that linear subtraction recovers input velocity dispersion values more accurately than subtraction in quadrature. However, we note that if quadrature subtraction is applied the sample mean velocity dispersion increases by $\simeq 5\%$, which does not affect the results of this paper. The median of the $\sigma_{\text{cor},i}$ map is taken as the σ_{int} value for each galaxy and the median of the observed map is the σ_{obs} value (see Table 3). The mean and median σ_{int} values in the isolated field sample are $70.8^{+3.3}_{-3.1} \text{ km s}^{-1}$ and $71.0^{+5.0}_{-4.8} \text{ km s}^{-1}$ respectively, higher than for typical star-forming galaxies at lower redshift as has been extensively reported (e.g. Genzel et al. 2006, 2008; Förster Schreiber et al. 2009; Law et al. 2009; Gnerucci et al. 2011; Epinat et al. 2012; Wisnioski et al. 2015). The magnitude of the beam-smearing correction, ($\sigma_{\text{int}}/\sigma_{\text{obs}}$), has a mean and median value of 0.81 and 0.79 respectively for the isolated field sample.

To test the robustness of our beam-smearing correction approach, we also computed V_C and σ_{int} values using the independently derived correction factors, which are a function of the velocity gradient and the ratio $R_{1/2}/R_{\text{psf}}$, detailed in Johnson et al. (in prep) and used to evaluate the final kinematic parameters in Harrison et al. (2017). These correction factors were applied to V_{BS} and σ_{obs} in order to compute V_C and σ_{int} . When following this approach, for the isolated field sample, the derived V_C values agree with our own corrections to within $\simeq 10\%$, and the σ_{int} values to within $\simeq 5\%$, which is within the errors on these quantities. This confirmed that the choice of modelling and kinematic parameter extraction described above is consistent with the $z \simeq 0.9$ KROSS methods (Harrison et al. 2017). We also ob-

served that the shape and width of the observed and intrinsic kinematic parameter distributions are equivalent, suggesting that the corrections do not over-extrapolate from the raw numbers.

We present the observed and model determined kinematic parameters in Table 3 for the isolated field sample.

3.2.6 Kinematic and photometric axes

The dynamical model parameter PA_{Kin} was compared with PA_{morph} computed by GALFIT in § 3.1.1 to indicate the degree of misalignment between the gaseous and stellar components of the galaxies. This follows the analysis described in e.g. Epinat et al. 2008b, 2012; Barrera-Ballesteros et al. 2014, 2015; Wisnioski et al. 2015; Harrison et al. 2017; Swinbank et al. 2017, where rotating galaxies generally have $\Delta PA \equiv |PA_{\text{kin}} - PA_{\text{morph}}| < 30^\circ$, with the discrepancy increasing as a function of axis ratio due to PA_{morph} becoming more difficult to measure as the galaxy becomes more ‘face-on’. Significant misalignment between the two axes can be a signature of galaxy interaction.

The resolved KDS galaxies at $z \simeq 3.5$ generally have larger ΔPA values than those at lower redshift (median $\Delta PA = 37^\circ$) and this observation has several potential explanations. For the majority of galaxies we are not in a regime where prominent morphological features are resolved in the *HST* imaging, so that PA_{morph} is biased by single Sérsic fits as described in Rodrigues et al. (2017). Another factor which could be partially driving increased misalignment in the KDS is the use of the full observed 2D velocity field to determine the kinematic position angle. This generally does not give the same answer as simply connecting the velocity extrema, as we verified by rotating PA_{Kin} in 1° increments and choosing the value which maximises the velocity gradient, calling this PA_{rot} . If the direction of the ‘zero-velocity strip’ across the centre of the velocity maps is not perpendicular to the line connecting the velocity extrema, the PA_{Kin} determined from MCMC will be different to the PA_{Kin} determined from rotation (e.g. GOODS-S galaxy bs006541 in Appendix D2). As a test, we have verified that extracting V_C values along PA_{rot} (which differs from PA_{Kin} typically by only a few degrees) makes only a small difference to the final V_C values, with the errors dominated by measurement uncertainties and model assumptions. Further, the increased prevalence of random motions at $z \simeq 3.5$ (see § 4.2.1), in comparison to at intermediate redshift, could partially explain misalignment of the stellar and gaseous components.

3.3 Morpho-Kinematic classification

We proceeded to further divide the isolated field sample galaxies into two subsamples on the basis of the ratio V_C/σ_{int} . This is a simple empirical diagnostic for whether a galaxy is ‘rotation-dominated’ with $V_C/\sigma_{\text{int}} > 1$, or ‘dispersion-dominated’ with $V_C/\sigma_{\text{int}} < 1$ (i.e. a method to measure the prevalence of rotational and random motions; Epinat et al. 2012; Wisnioski et al. 2015; Stott et al. 2016; Harrison et al. 2017). Previous studies have used various classification schemes of varying complexity to identify rotation-dominated galaxies which also appear to be ‘disk-like’ and we briefly mention some of these in § 3.3.2. We describe our simple classification criteria for the isolated field

sample based on a joint consideration of the high-resolution photometry and derived kinematic parameters in § 3.3.2 and mention the KDS merger candidates in more detail in the following section.

3.3.1 KDS merger candidates

We considered both the *HST* imaging and information from the [O III] $\lambda 5007$ line within the datacubes when classifying galaxies as mergers, using the following three criteria: i) Multiple components were detected in the high-resolution, *HST* imaging, which cannot be fit with a single GALFIT component; ii) The [O III] $\lambda 5007$ flux maps examined in QFitsView (Ott & Thomas 2012) had more than one peak across the spatial extent of the cube; iii) The object spectrum showed peaks at multiple wavelengths, rather than a smoothly varying line profile. When fit using the procedure described in § 3.2.1, this manifests as extreme rotation velocities and velocity dispersion values in the velocity and velocity dispersion maps.

Galaxies which met all three of these criteria were classed as mergers. One benefit of studying the *HST* photometry in tandem with the kinematics was to uncover ‘rotation doppelgängers’, which mimic disk rotation from a purely kinematic perspective (i.e. can be fit with the arctangent model (see § 3.2.2) with small residuals), but are likely two or more potentially interacting components at slightly different redshifts (e.g. the bottom grid in Figure 3).

This has been discussed in Hung et al. (2015), where redshifted local interacting galaxies are misclassified as single rotating disks on the basis of their kinematics. In short, the galaxies with multiple photometric components do not always have disordered kinematics and so the kinematics alone are not sufficient for classification. It was difficult to distinguish in some cases between two *HST* components rotating together in a single disk (i.e. clumpy disk galaxies e.g. Elmegreen et al. 2004; Bournaud et al. 2007; Förster Schreiber et al. 2009) with a receding and approaching side and two [O III] $\lambda 5007$ emitting blobs which are offset along the line of sight. We chose to omit these galaxies from the analysis described in § 4, to minimise the risk of contaminating the kinematics of the sample with mergers.

In 6/39 of the resolved KDS field galaxies (one count per merger) the merging classification criteria were met, which translates into a merger rate of 15%. In 8/9 resolved KDS cluster galaxies the criteria were met, translating into a merger rate of 89% (see Table 1). The higher merger rate observed for the cluster galaxies is perhaps unsurprising since they were drawn from a higher density environment, and these interactions will inevitably have major consequences for the consumption and stripping of gas, as well as the morphological and dynamical evolution within these galaxies. The observation of frequent mergers, which cause major kinematic disturbances, is also in support of the evidence that galaxies living at the peak of the cosmic density field at lower redshift tend to be redder, less active and morphologically early-type, having presumably rapidly exhausted their gas supply throughout these violent interactions at high-redshifts (e.g. Steidel et al. 1998; White et al. 2007; Kodama et al. 2007; Zheng et al. 2009). As mentioned above, we omit galaxies in the cluster pointing (SSA22-P1)

Table 3. The dynamical properties for the KDS isolated field sample galaxies. The observed properties are measured directly from the extracted two-dimensional maps (with σ_{obs} corrected for the KMOS instrumental resolution), whereas V_{C} and σ_{int} have been corrected for the effects of beam-smearing as described throughout the text. Each galaxy is classified as rotation-dominated or dispersion-dominated if the ratio $V_{\text{C}}/\sigma_{\text{int}}$ is greater than or less than 1 respectively, as highlighted in the ‘Class’ column.

ID	$V_{\text{obs}}(\text{km s}^{-1})$	$V_{\text{C}}(\text{km s}^{-1})$	$\sigma_{\text{obs}}(\text{km s}^{-1})$	$\sigma_{\text{int}}(\text{km s}^{-1})$	$V_{\text{C}}/\sigma_{\text{int}}$	$\text{PA}_{\text{kin}}^{\circ}$	Class ^a
b012141-012208	32±13	93 ⁺⁴⁸ ₋₄₀	90±18	63 ⁺²⁸ ₋₃₀	1.47 ^{+1.04} _{-0.92}	122 ⁺⁹ ₋₉	RD
b15573	61±21	81 ⁺³⁰ ₋₃₃	107±23	84 ⁺²⁸ ₋₂₆	0.97 ^{+0.47} _{-0.52}	118 ⁺¹⁰ ₋₁₈	DD
bs006516	49±4	61 ⁺¹⁰ ₋₉	58±8	45 ⁺¹⁰ ₋₉	1.35 ^{+0.37} _{-0.38}	144 ⁺⁴ ₋₄	RD
bs006541	33±8	65 ⁺²⁴ ₋₂₀	92±7	83 ⁺⁹ ₋₁₆	0.78 ^{+0.34} _{-0.26}	23 ⁺³ ₋₄	DD
bs008543	68±13	111 ⁺²⁰ ₋₁₆	94±7	71 ⁺¹³ ₋₁₀	1.56 ^{+0.37} _{-0.38}	114 ⁺¹ ₋₁	RD
bs009818	39±16	84 ⁺³⁹ ₋₂₃	92±19	79 ⁺²¹ ₋₃₁	1.06 ^{+0.65} _{-0.41}	72 ⁺¹³ ₋₁₅	RD
bs014828	31±15	46 ⁺²³ ₋₂₃	93±15	88 ⁺¹⁵ ₋₁₇	0.53 ^{+0.28} _{-0.28}	44 ⁺¹⁵ ₋₁₁	DD
bs016759	55±9	110 ⁺³⁰ ₋₂₁	77±11	57 ⁺¹⁷ ₋₂₂	1.93 ^{+0.93} _{-0.7}	75 ⁺⁶ ₋₆	RD
lbg_20	20±6	56 ⁺³⁰ ₋₁₉	47±6	40 ⁺⁷ ₋₂₁	1.39 ^{+1.04} _{-0.53}	41 ⁺⁹ ₋₁₀	RD
lbg_24	26±6	42 ⁺⁹ ₋₇	57±5	51 ⁺⁵ ₋₆	0.82 ^{+0.21} _{-0.17}	64 ⁺³ ₋₄	DD
lbg_25	14±5	33 ⁺¹³ ₋₁₃	47±6	41 ⁺⁶ ₋₈	0.8 ^{+0.37} _{-0.34}	87 ⁺⁷ ₋₁₂	DD
lbg_30	20±6	66 ⁺¹⁹ ₋₁₇	89±11	77 ⁺¹⁴ ₋₁₇	0.86 ^{+0.32} _{-0.27}	46 ⁺⁶ ₋₄₂	DD
lbg_32	65±5	121 ⁺¹³ ₋₁₂	100±13	76 ⁺¹⁶ ₋₁₅	1.58 ^{+0.36} _{-0.37}	11 ⁺² ₋₁	RD
lbg_38	26±7	97 ⁺⁷² ₋₃₄	71±10	53 ⁺¹⁶ ₋₅₃	1.82 ^{+2.3} _{-0.86}	47 ⁺² ₋₃	RD
lbg_91	48±10	91 ⁺¹⁵ ₋₁₈	94±8	71 ⁺¹⁶ ₋₁₄	1.28 ^{+0.33} _{-0.39}	40 ⁺⁵ ₋₃	RD
lbg_94	18±8	34 ⁺¹⁹ ₋₁₅	60±6	55 ⁺⁷ ₋₈	0.62 ^{+0.36} _{-0.29}	41 ⁺¹⁸ ₋₉	DD
lbg_105	11±8	32 ⁺¹⁷ ₋₁₆	61±18	59 ⁺¹⁸ ₋₁₈	0.54 ^{+0.34} _{-0.32}	170 ⁺⁴ ₋₇	DD
lbg_109	26±8	53 ⁺²¹ ₋₁₅	96±16	91 ⁺¹⁶ ₋₁₈	0.58 ^{+0.26} _{-0.2}	157 ⁺¹¹ ₋₁₁	DD
lbg_111	36±12	58 ⁺²⁴ ₋₁₇	85±14	81 ⁺¹⁴ ₋₁₆	0.71 ^{+0.33} _{-0.25}	168 ⁺⁴ ₋₅	DD
lbg_112	27±3	63 ⁺¹⁹ ₋₁₁	88±5	78 ⁺⁸ ₋₁₁	0.81 ^{+0.28} _{-0.17}	4 ⁺³⁴ ₋₁₂	DD
lbg_113	50±20	41 ⁺⁴¹ ₋₂₉	113±22	108 ⁺²³ ₋₂₅	0.39 ^{+0.39} _{-0.29}	1 ⁺³⁴ ₋₂₆	DD
lbg_121	33±9	54 ⁺³³ ₋₂₄	67±18	50 ⁺²¹ ₋₂₇	1.06 ^{+0.87} _{-0.65}	55 ⁺³⁴ ₋₂₆	RD
lbg_124	9±8	32 ⁺²⁷ ₋₁₈	52±19	46 ⁺¹⁹ ₋₂₃	0.71 ^{+0.69} _{-0.5}	39 ⁺²⁰ ₋₁₄	DD
*n3_006	36±23	81 ⁺⁴⁷ ₋₂₉	133±26	125 ⁺²⁷ ₋₃₁	0.65 ^{+0.41} _{-0.27}	156 ⁺¹³ ₋₁₂	DD
n3_009	21±9	49 ⁺²⁶ ₋₁₆	56±13	48 ⁺¹⁴ ₋₄₅	1.01 ^{+1.1} _{-0.45}	20 ⁺²⁸ ₋₂₅	RD
n_c3	26±8	51 ⁺¹⁷ ₋₁₈	84±11	72 ⁺¹² ₋₁₃	0.72 ^{+0.28} _{-0.29}	137 ⁺⁴ ₋₁₅	DD
lab18	23±11	41 ⁺¹⁷ ₋₁₄	69±6	62 ⁺⁷ ₋₁₅	0.67 ^{+0.32} _{-0.24}	54 ⁺¹⁸ ₋₂₆	DD
*lab25	37±11	48 ⁺²² ₋₁₆	58±8	49 ⁺¹⁰ ₋₁₉	0.98 ^{+0.59} _{-0.4}	72 ⁺¹⁴ ₋₁₃	DD
s_sa22a-d3	53±17	109 ⁺⁵² ₋₃₂	115±13	93 ⁺²⁰ ₋₃₆	1.16 ^{+0.72} _{-0.43}	30 ⁺² ₋₂	RD
s_sa22b-c20	66±18	139 ⁺³³ ₋₃₀	97±17	48 ⁺³⁴ ₋₅₁	2.87 ^{+3.12} _{-2.14}	37 ⁺²⁶ ₋₁₁	RD
*s_sa22b-d5	27±10	57 ⁺²⁶ ₋₂₃	38±20	23 ⁺²³ ₋₁₈	2.45 ^{+3.44} _{-2.41}	60 ⁺¹⁹ ₋₁₉	RD
s_sa22b-d9	32±14	46 ⁺²² ₋₁₅	83±8	73 ⁺¹⁰ ₋₁₇	0.63 ^{+0.33} _{-0.23}	12 ⁺²³ ₋₇	DD
*s_sa22b-md25	44±15	57 ⁺³⁴ ₋₂₄	87±19	75 ⁺²¹ ₋₃₀	0.76 ^{+0.56} _{-0.39}	9 ⁺¹² ₋₇	DD

^a RD = Rotation-Dominated; DD = Dispersion-Dominated

* No *HST* coverage; see Table 2

from the results presented in § 4 due to the complexity added by galaxy interactions to the interpretation of cluster galaxy velocity fields. We keep the KDS galaxies without *HST* coverage on the basis that we do not identify double components in the [O III]λ5007 maps. In the kinematic grids throughout D3 (as well as the bottom grid in 3) we present the *HST* imaging, kinematic maps and one-dimensional kinematic extractions for the KDS merger candidates in both the field and cluster environments.

3.3.2 Isolated field sample classification

As mentioned above, various methods have been used throughout the literature to quantify the kinematic structure of star-forming galaxies at different redshifts. These include the kinemetry approach which quantifies asymmetries in the 2D moment maps described in Shapiro et al. (2008) and applied throughout in Förster Schreiber et al. (2009) and Cresci et al. (2009). In Wisnioski et al. (2015) a set of 5 criteria (see their section 4) identify rotation-dominated and

disk galaxies by assessing the smoothness of the velocity gradient, the degree of rotational support through the ratio of the velocity to velocity dispersion, V_{rot}/σ , the ΔPA misalignment and the position of the kinematic centre in relation to both the peak in the velocity dispersion map and the centroid of the continuum centre (also see [Rodrigues et al. 2017](#) for a re-analysis of KMOS^{3D} data at $z \simeq 1$ using the same criteria). In [Epinat et al. \(2012\)](#), a detailed description of a morpho-kinematic classification is outlined in their section 4, based on the proximity of counterparts in the imaging, ΔPA and V_C/σ_{int} . In [Gnerucci et al. \(2011\)](#) an inclined plane is fit to the velocity map and the χ^2 of this fit is evaluated in order to test for the presence of smooth velocity gradients, which they find for the 11/33 galaxies described in [Appendix A9](#).

We used a simple approach (dictated by the signal-to-noise of the isolated field sample) and calculated rotation and dispersion-dominated percentages on the basis of V_C/σ_{int} alone with respect to the total number of resolved galaxies in the field environment (39, see [Table 1](#)). The morphologically isolated KDS field galaxies were classified as rotation-dominated if $V_C/\sigma_{\text{int}} > 1$, 14/39 or 36% of the sample, and dispersion-dominated if $V_C/\sigma_{\text{int}} < 1$, 19/39 or 49% of the sample, as indicated in the final column of [Table 3](#). We plot the best-fitting smeared and intrinsic models for the isolated field galaxies in [Appendix D](#), where the galaxies have been divided into rotation and dispersion-dominated categories in [Figures D1 and D2](#) respectively.

3.3.3 Summary of the final sample for further analyses

In summary, we analysed 77 typical $z \simeq 3.5$ SFGs spanning both cluster and field environments, finding 63 with integrated [O III] λ 5007 emission ([§ 2.2.2](#)). This sample was refined to 48 galaxies (62%) with spatially-resolved [O III] λ 5007 emission ([§ 3.2.1](#)). Due to the unusually high merger-candidate rate of 89% in the cluster environment (identified by searching the high-resolution *HST* imaging for photometric counterparts), we omitted the 9 spatially-resolved galaxies in SSA22-P1 from further consideration. In the remaining sample of 39 field environment galaxies we identified 6 merger candidates, giving a field galaxy merger rate of 15%, which constitute the ‘merger field sample’. The remaining 33 galaxies constitute the ‘isolated field sample’, for which we measured morphological parameters using GALFIT to fit exponential disks to the *HST* imaging (see [§ 3.1.1](#) and [Table 2](#)) and derived beam-smearing corrected measurements of the rotation velocities extracted at $2R_{1/2}$ and the intrinsic velocity dispersions (see [§ 3.2.2](#) and [Table 3](#)).

The isolated field sample was further divided into 14/39 rotation-dominated galaxies (36% with respect to full field sample) and 19/39 dispersion-dominated galaxies (49%) on the basis of the ratio V_C/σ_{int} (see [§ 3.3.2](#)), with the remaining 6/39 (15%) classified as mergers.

Throughout the remainder of the paper we discuss the derived kinematic properties of the isolated field sample in further detail.

4 RESULTS

In the previous sections we presented our morphological and kinematic analyses for the 77 galaxies covering both cluster and field environments in the KDS survey. These galaxies are representative of typical star-forming galaxies at $z \simeq 3.5$, as verified by comparing with fits to the galaxy main-sequence at this redshift. Through a combined morpho-kinematic classification described in [§ 3.3](#), we first reduced the sample to 39 spatially resolved field environment galaxies and then to 33 isolated field sample galaxies which have spatially-resolved [O III] λ 5007 emission and are morphologically isolated (i.e. not merger candidates), which is the largest sample of galaxies observed with integral-field spectroscopy at this redshift. For these galaxies we measured morphological properties as listed in [Table 2](#) and extracted intrinsic, beam-smearing corrected measurements of the rotation velocities, V_C , and velocity dispersions, σ_{int} , as listed in [Table 3](#), which we used to classify the galaxies as either rotation or dispersion-dominated. The morphological and kinematic grids for these 33/39 isolated field sample galaxies and the 6/39 merger candidates are plotted in [Appendix D](#).

Throughout the following sections we analyse the derived kinematic parameters of the isolated field sample, placing the galaxies in the context of galaxy evolution using the comparison samples spanning a wide redshift baseline described in [Appendix A](#). In [Appendix A](#) we describe our careful assessment of lower redshift samples from the literature, using tabulated data where possible to re-compute mean and median V_C and σ_{int} values. These fair comparisons help to disentangle true evolution in the kinematic properties of typical star-forming galaxies from systematic effects originating in disparate measurement techniques or galaxy selection criteria. During the following sections we will make statements about dynamical evolution by connecting the dots of these different surveys, assuming that they trace the average properties of typical star-forming galaxies at their respective mean redshifts.

We proceed first to investigate the relationship between rotation velocity and stellar mass for both the rotation and dispersion-dominated isolated field galaxies.

4.1 Rotation velocity and stellar mass

In the ΛCDM paradigm, galaxies form when baryons accumulate and cool in the centres of dark matter halos, subsequently forming a rotating disk of stars and gas. The relationship between stellar mass and rotation velocity describes the connection between the luminous mass and the total dynamical mass in a system. The connection has been studied extensively at low redshift (e.g. [Bell & de Jong 2001](#); [Rhee et al. 2003](#); [Reyes et al. 2011](#)) with fits to late-type, rotating galaxies in the $V_C - M_\star$ plane having a well-defined slope and normalisation. Observing this ‘stellar mass Tully-Fisher relation’ at different epochs thus offers constraints for galaxy formation and evolution models seeking to explain simultaneously the properties of dark matter halos and the dynamical properties of disk galaxies (e.g. [Dubois et al. 2014](#); [Vogelsberger et al. 2014](#); [Schaye et al. 2015](#)). In principle, studying whether the stellar mass Tully-Fisher relation evolves with redshift gives us information about when galaxies accumulate their stel-

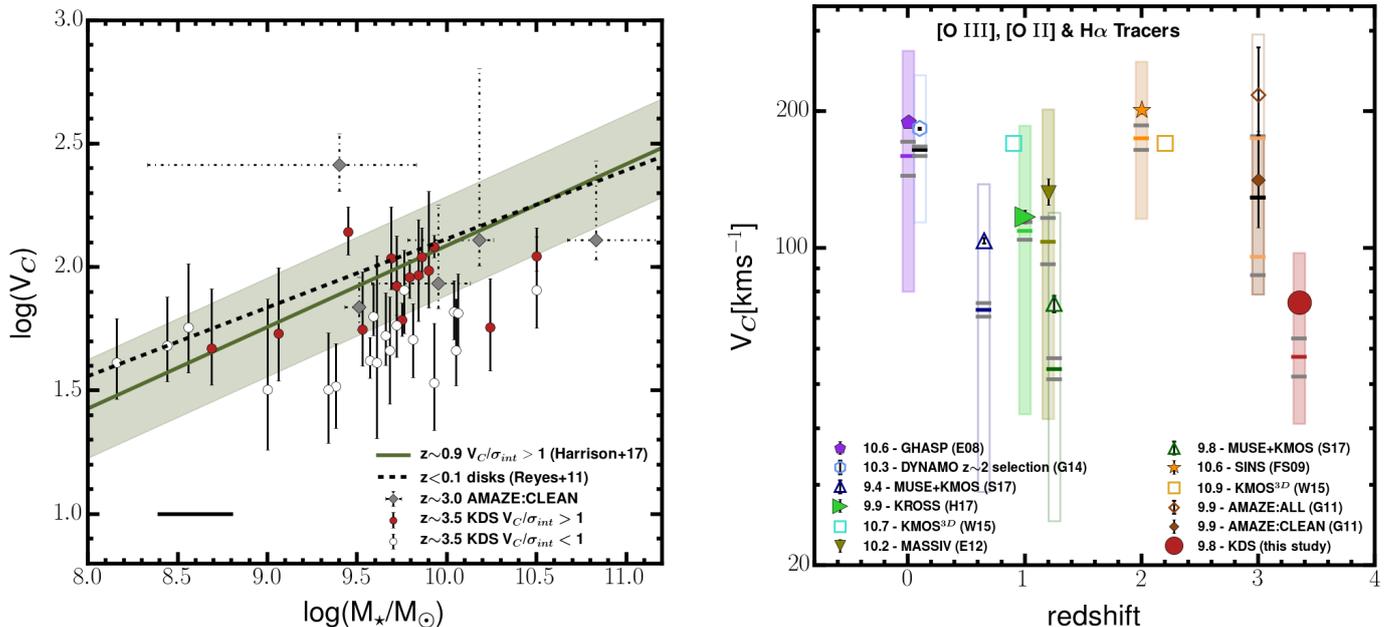


Figure 4. *Left:* We plot $\log(V_C)$ vs. $\log(M_*/M_\odot)$ for the isolated field sample, with the rotation-dominated galaxies in red and dispersion-dominated galaxies with the hollow symbols. We also plot with the solid green line the $z \approx 0.9$ relation recovered from a fit to ≈ 400 rotation-dominated galaxies from the KROSS survey Harrison et al. (2017), with the shaded region showing the associated ≈ 0.2 dex scatter on the fit. This $z \approx 0.9$ relation shows no significant evolution from the fit to a local sample of spiral galaxies with $z < 0.1$, presented in Reyes et al. (2011). The rotation-dominated KDS galaxies, and galaxies classed as reliable from the AMAZE survey (Gnerucci et al. 2011), also show no evolution from this relation within the scatter, with the dispersion-dominated galaxies scattering below the relation. This highlights the need for careful sample selection when constructing this relation. *Right:* The mean, median and distributions of the comparison sample V_C measurements are plotted against redshift. The symbol positions with black errorbars (some not visible), show the mean values and the horizontal filled and grey bars show the median values and associated errors. Note that the median errors for the AMAZE Clean sample are shaded orange to avoid confusion with the full sample. The filled symbols and shaded regions denote surveys deemed directly comparable (as explained in Appendix A), and the mean mass of each survey (in units of $\log(M_*/M_\odot)$) is listed beside the survey name. The V_C distributions in each of the individual studies are wide, reflecting the galaxy diversity within each sample. The low mean value of V_C in the KDS isolated field sample is driven mainly by averaging over a sample which contains a higher fraction of dispersion-dominated galaxies with lower rotation velocities.

lar mass and rotation velocities. Any offset from the local relation towards lower M_* values at fixed rotation velocity suggests that the dynamical mass is in place, but the stellar mass is yet to be formed (e.g. Puech et al. 2008; Cresci et al. 2009; Puech et al. 2010; Gnerucci et al. 2011; Swinbank et al. 2012a; Straatman et al. 2017; Ubler et al. 2017). As has been discussed recently, information about the redshift evolution of gas fractions, $f_g = M_{\text{gas}}/M_{\text{bar}}$ where M_{bar} is the baryonic mass, disk mass fractions, $m_d = M_{\text{bar}}/M_H$ where M_H is the halo mass, dark matter fractions, $f_{\text{DM}}(R_{1/2}) = V_{\text{DM}}^2(R_{1/2})/V_C^2(R_{1/2})$ where V_{DM} is the dark matter circular velocity, may also be encoded in measured offsets from the local stellar mass Tully-Fisher relation (Wuyts et al. 2016; Ubler et al. 2017; Genzel et al. 2017).

We plot the relationship between rotation velocity and stellar mass for the isolated field sample in the left panel of Figure 4. We compare to the relationship from Harrison et al. (2017), in which the inverse stellar mass Tully-Fisher relation is fit to ≈ 400 star-forming galaxies with $9.5 < \log(M_*/M_\odot) < 10.5$ at $z \approx 0.9$, measured to be rotation-dominated with $V_C/\sigma_{\text{int}} > 1$. This is the closest matched sample in terms of analysis and selection criteria, therefore providing a robust comparison. The data

are used to constrain parameters ‘a’ and ‘b’ in the equation $\log V_C = b + a[\log M_* - 10.10]$ (see their section 4.2) finding $a_{z=0.9} = 0.33 \pm 0.11$ and $b_{z=0.9} = 2.12 \pm 0.04$. We plot this best fit relation with the solid green line, with the green shaded region giving the typical scatter of 0.2 dex along the velocity axis.

The rotation-dominated KDS isolated field sample galaxies are broadly consistent within the scatter with the $z \approx 0.9$ relation, but with a tendency to lie below the relation. When considered alone, the mean velocity of the rotation-dominated isolated field sample galaxies ($V_C = 96.7^{+7.3}_{-7.2}$ km s $^{-1}$) at their mean mass of $\log(M_*/M_\odot) = 9.8$ is ≈ 0.05 dex beneath the KROSS relation at the same mass. The dispersion-dominated KDS galaxies sit clearly separated and below the $z \approx 0.9$ relation, bringing significant scatter to lower V_C values at fixed M_* as is also observed in $z \approx 0.9$ dispersion-dominated galaxies and for early-type galaxies in the local universe (e.g. Romanowsky & Fall 2012). We plot also the five galaxies from the AMAZE Clean ($z \approx 3$) sample, which are consistent within the scatter with the relation for typical star-forming galaxies at $z \approx 0.9$. The $z \approx 0.9$ fit parameters are consistent with those measured in the local universe, $a_{z=0} = 0.278 \pm 0.010$ and $b_{z=0} = 2.142 \pm 0.004$, for a sample

of $z < 0.1$ disks from Reyes et al. (2011), which are in a similar mass range to the KROSS galaxies in Harrison et al. (2017) and traced with $H\alpha$ emission. In Reyes et al. (2011), V_C has been extracted at the same radius and we plot this $z \simeq 0.1$ local relation as the dashed black line in the left panel of Figure 4. The agreement between the parameters derived at $z \simeq 0$ and $z \simeq 0.9$ suggests that there is no evolution of the slope/normalisation of the stellar mass Tully-Fisher relation for rotation-dominated star-forming galaxies between these redshifts.

Studies of the redshift evolution of the stellar mass Tully-Fisher relation do not explain how individual, unperturbed galaxies move around the rotation velocity stellar mass plane with redshift. High-redshift galaxy surveys typically do not trace the progenitors of those at lower redshift, as evidenced by the larger observed stellar and baryonic masses (e.g. Cresci et al. 2009; Reyes et al. 2011; Wisnioski et al. 2015; Ubler et al. 2017). Rather, these studies trace the evolution of the relation itself, the position of which is dictated by the mean properties of typical star-forming galaxies at each redshift slice. Some authors have reported that as redshift increases, the relation shifts towards lower masses at fixed rotation velocity (e.g. Puech et al. 2008; Cresci et al. 2009; Puech et al. 2010; Straatman et al. 2017; Ubler et al. 2017). However most find that when selecting rotation-dominated galaxies between $0.5 < z < 2$ (e.g. Flores et al. 2006; Miller et al. 2011; Kassin et al. 2012; Miller et al. 2012; Vergani et al. 2012; Miller et al. 2014; Contini et al. 2016; Di Teodoro et al. 2016; Pelliccia et al. 2017; Molina et al. 2017; Harrison et al. 2017) the results are consistent with zero evolution of the slope and zero-point of the relation. At $z \simeq 3$ observations are limited and in Gnerucci et al. (2011) an offset of -1.29 dex towards lower mass at fixed rotation velocity is claimed, although there is large degree of scatter between the individual galaxies, and the AMAZE Clean sample from Gnerucci et al. (2011) defined in appendix A9 is consistent with the $z \simeq 0$ relation.

Conclusions on the evolution of the stellar mass Tully-Fisher relation are a strong function of sample selection; in Tiley et al. (2016) where a stricter $V_C/\sigma_{\text{int}} > 3$ cut is applied, evolution to lower masses at fixed rotation velocity is observed at $z \simeq 0.9$, which is not the case when using the $V_C/\sigma_{\text{int}} > 1$ cut. Similarly in Cresci et al. (2009) evolution of the normalisation of the stellar mass Tully-Fisher relation is observed at $z \simeq 2$ using only robust rotators from the SINS sample with mean $V_C/\sigma_{\text{int}} = 5$. In both cases the observed evolution is mainly the result of omitting rotation-dominated galaxies with lower V_C values. Conversely, fitting the stellar mass Tully-Fisher relation through a full sample including dispersion-dominated galaxies would shift the zero-point in the opposite sense. This new, robust analyses of typical $z \simeq 3.5$ star-forming galaxies is consistent with the majority of work that shows no evolution for rotation-dominated galaxies across all epochs.

In summary, with the KDS, we have added the largest sample of galaxies observed with integral-field spectroscopy at $z > 3$, and with our simple $V_C/\sigma_{\text{int}} > 1$ cut we conclude that the stellar mass Tully-Fisher relation defined for rotation-dominated star-forming galaxies appears to show

no evolution between $z \simeq 0.1$ and $z \simeq 3.5$. However, as we will show in § 4.3, the $V_C/\sigma_{\text{int}} > 1$ subsamples become increasingly less representative of typical star-forming galaxies with increasing redshift. We also explore this in the context of the mean rotation velocities of the comparison samples throughout the following subsection.

4.1.1 The evolution of maximum circular velocity

In the right panel of Figure 4 we plot the fractional error weighted mean, median and distributions of V_C values reported in the comparison (see Appendix A) and KDS isolated field samples to highlight two observations:

- At each redshift slice the V_C distributions are extremely broad, with mean values sensitive to the average stellar mass of each survey.
- The mean isolated field sample V_C value is less than reported in the lower redshift comparison samples with similar mean M_* values. This is a result of averaging over a sample containing a higher fraction of dispersion-dominated galaxies.

The right panel of Figure 4 indicates that the V_C distributions for the comparison and KDS isolated field samples are wide, reflecting the mass range and the mixture of rotation/dispersion-dominated galaxies which constitute the samples. The mass distribution within each sample drives the location of the mean rotation velocity, with samples that have larger mean stellar mass values showing larger mean rotation velocities. The isolated field sample mean and median V_C values are $76.7^{+4.9}_{-4.5} \text{ km s}^{-1}$ and $57.0^{+6.6}_{-6.3} \text{ km s}^{-1}$. However, the mean and median rotation velocity of the rotation dominated galaxies are $V_C = 96.7^{+7.3}_{-7.2} \text{ km s}^{-1}$ and $V_C = 93.0^{+8.1}_{-7.7} \text{ km s}^{-1}$ respectively, suggesting that averaging over a sample containing a large number of dispersion dominated galaxies brings the averages down.

It is difficult to assess whether the V_C values reported in the comparison samples (the modelling and extraction techniques differ as described in Appendix A) are directly comparable. The GHASP, MASSIV, KROSS and KDS samples should in principle yield comparable V_C values which rely on extracting from the data at a fiducial radius or using a velocity shear value, but the comparison to the dynamical models of SINS and AMAZE is not so clear. Further caveats when making comparisons between the velocity distributions of different samples are variations in disk half-light radii and the spatial resolution of the observations. In Newman et al. (2013) it is shown that it is often not possible to resolve rotational structure in seeing-limited conditions, with this effect more severe for compact galaxies. Consequently, the non-uniform disk-sizes and spatial resolution across the comparison samples may result in additional variations between the inferred rotation velocities. As a result, mainly of the differences in mass distribution between the samples, we do not attempt to interpret any evolutionary trends from the right panel of Figure 4. However, Simons et al. (2017) have demonstrated with a consistently selected and processed sample of star-forming galaxies between $0 < z < 2.5$, that there is a mild drop in rotation velocities with increasing redshift over this interval in both low and high-mass bins.

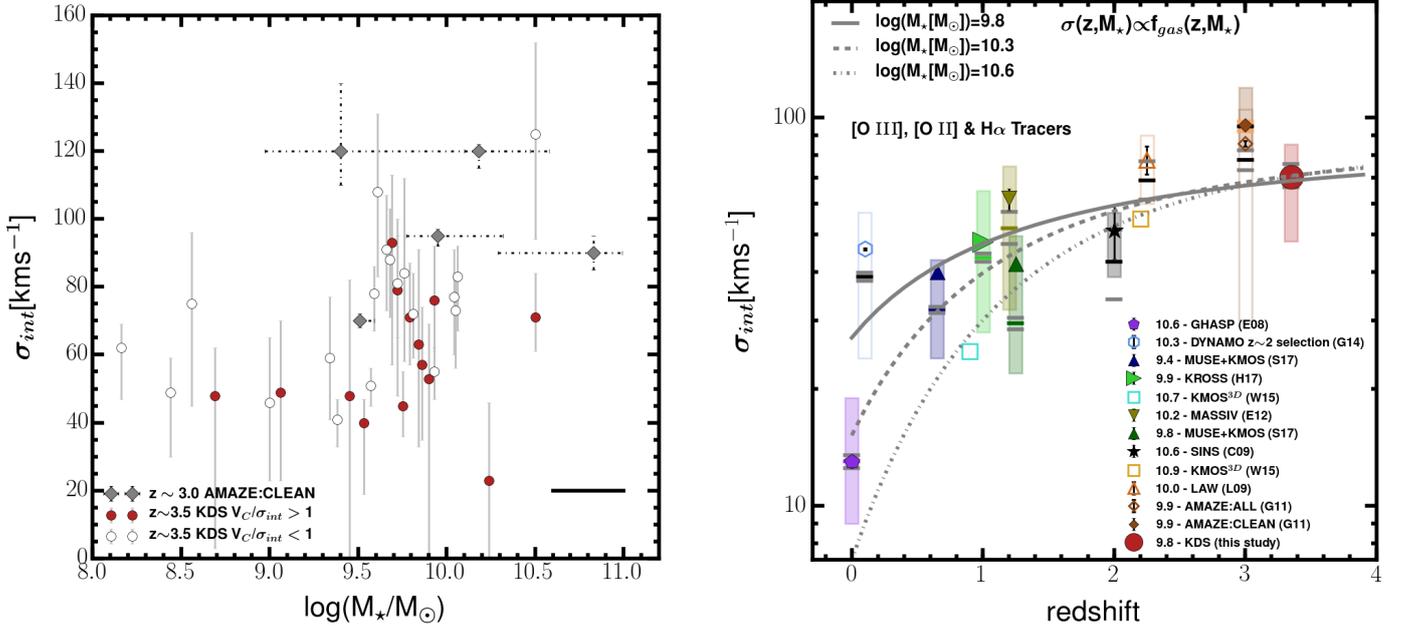


Figure 5. *Left:* We plot the intrinsic velocity dispersion against stellar mass for the rotation and dispersion-dominated isolated field sample galaxies, along with the AMAZE Clean ($z \simeq 3$) sample. The σ_{int} values are typically distributed between 40 – 90 km s⁻¹ and show no clear correlation with stellar mass, highlighting the complicated relationship between these two quantities. *Right:* We present a compilation of literature σ_{int} values plotted against redshift for surveys spanning $0 < z < 4$. The mean σ_{int} values increase from $\simeq 15$ km s⁻¹ in the local universe to > 70 km s⁻¹ at $z > 3$. The symbol convention is the same as in the right panel of Figure 4. The $\sigma_{\text{int}}(z, M_*) \propto f_{\text{gas}}(z, M_*)$ scaling relation from Wisnioski et al. (2015), the form of which is discussed in § 5.1, is plotted for three different M_* values. These tracks indicate the way in which σ_{int} evolves for samples of star-forming galaxies with different mean M_* values, which is a proxy for the dynamical maturity of the galaxies. The KDS datapoint appears to be consistent with the scenario proposed in previous work, whereby the mean σ_{int} increases over cosmic time, with large scatter between individual galaxies, as a result of increased gas fractions (Wisnioski et al. 2015) and the more efficient accretion of cold gas driving disk instabilities (e.g. Law et al. 2009; Genzel et al. 2011).

In a rotationally supported system, the circular velocity is set by the requirement to support the total mass in the system, i.e. the combination of gas mass, stellar mass and dark matter mass, against gravitational collapse. V_C increases with increasing total mass and so we would expect the systems with larger average stellar mass, assuming that these are indicative of systems with greater total mass, to have larger V_C , which appears generally to be the case in the right panel of Figure 4 where the KMOS^{3D}, GHASP and SINS galaxies have the largest inferred V_C values. Using the average stellar mass values and velocity extraction radii in the isolated field sample, $\log(M_*/M_\odot) = 9.8$ and $2R_{1/2} = 3.2$ kpc respectively, and assuming simple, circular Keplerian orbits where the rotational motions alone support the mass of the system, we can calculate a rough lower limit of $\langle V_C \rangle > 95$ km s⁻¹ for the baryonic material in the galaxies (since the inclusion of the gas mass and dark matter mass would raise this estimate). This value is greater than the mean and median values of the isolated field sample, but comparable to the mean and median of the rotation dominated galaxies, suggesting that rotation in these systems could support, at most, the stellar mass of the galaxy.

As mentioned above, the observation of a low mean V_C value in the full isolated field sample is primarily explainable through averaging over a sample with a higher fraction of dispersion-dominated galaxies with lower ro-

tation velocities. This is seen in the left panel of Figure 4 where the dispersion-dominated galaxies sit clearly below and separated from the stellar mass Tully-Fisher relation. However, it has also been discussed that pressure support, provided by turbulence and measured through the galaxy velocity dispersions, may partially compensate the gravitational force within galaxy disks and contribute to lowering the rotation velocities through an ‘asymmetric drift’ correction (e.g. Burkert et al. 2010; Newman et al. 2013; Genzel et al. 2017). This topic is discussed further throughout § 5.

We proceed to consider the cause of the high number of dispersion-dominated galaxies in our sample and the obvious candidate is the increase of velocity dispersion with redshift (e.g. Genzel et al. 2006, 2008; Förster Schreiber et al. 2009; Law et al. 2009; Gnerucci et al. 2011; Epinat et al. 2012; Wisnioski et al. 2015). There are also some galaxies in our sample where the observed velocity field has clear rotational structure and reaches flattening (e.g. see Figure D2 lbg_24 with $\log(M_*/M_\odot) = 9.75$), with inferred V_C smaller than σ_{int} . This suggests again that the increase in random motions could influence the rotation velocities and that the accumulation of a massive stellar population is necessary to stabilise a galaxy (e.g. Law et al. 2009, 2012c,b;

Newman et al. 2013; Wisnioski et al. 2015). These ideas are explored in the following sections.

4.2 Velocity dispersions

In the left panel of Figure 5 we plot the isolated field galaxies, split into dispersion and rotation-dominated classes, along with the AMAZE Clean ($z \simeq 3$) sample, in the σ_{int} vs. M_{\star} plane. This is an interesting relationship to explore as it has been suggested that rotation velocity alone may not be a good tracer of total dynamical mass at high-redshift and that random motions traced by σ_{int} may contribute to supporting some fraction of this total dynamical mass (e.g. Kassin et al. 2007; Burkert et al. 2010; Kassin et al. 2012; Law et al. 2012c,b; Ubler et al. 2017; Genzel et al. 2017). However, the extent to which the gaseous σ_{int} values reported here trace stellar mass is unclear, with physical processes such as disk turbulence, gas accretion and subsequent disk instabilities contributing to the magnitude of the velocity dispersion. All of these processes are also more prevalent at high-redshift (Genzel et al. 2006; Förster Schreiber et al. 2009; Law et al. 2009; Genzel et al. 2011; Wisnioski et al. 2015; Wuyts et al. 2016), leaving σ_{int} as a challenging property to interpret.

We observe no clear correlation between the gaseous velocity dispersion and stellar mass in any of the subsamples, highlighting the complicated relationship between these quantities at high-redshift. The beam-smearing corrected σ_{int} values, with distribution median, 16th percentile and 84th percentile equal to 71.0 km s^{-1} , 48.0 km s^{-1} and 85.4 km s^{-1} respectively, are larger than in the local universe where typically $\sigma_{\text{int}} = 10 - 20 \text{ km s}^{-1}$ (Epinat et al. 2008a) (see § 4.2.1 and § 5 for a discussion). Generally, the dispersion-dominated galaxies have higher σ_{int} values in the left panel of Figure 5, however the rotation-dominated galaxies also span the full distribution width.

4.2.1 The evolution of velocity dispersion

In the right panel of Figure 5 we plot the fractional error weighted mean, median and distributions of σ_{int} from the comparison and KDS isolated field samples. The values increase sharply from $\sigma_{\text{int}} \simeq 10 - 20 \text{ km s}^{-1}$ as traced by $z \simeq 0$ measurements of typical SFGs to $\sigma_{\text{int}} \simeq 30 - 60 \text{ km s}^{-1}$ at $z \simeq 1$ to $\sigma_{\text{int}} \simeq 40 - 90 \text{ km s}^{-1}$ at $z \simeq 3$, but with a wide range of values observed at each redshift. There appears to be a trend for the surveys with lower average stellar mass (MASSIV, DYNAMO, LAW 2009, KROSS, AMAZE) to have higher velocity dispersions than the surveys with larger average stellar mass (GHASP, KMOS^{3D}, SINS).

In agreement with the evolution of gas fractions described in § 5.1 (and see the tracks in the right panel of Figure 5), we find that the mean velocity dispersion measurements from samples of typical star-forming galaxies are decreasing with redshift, with a dependence on mass such that samples with higher stellar mass have lower velocity dispersions at each epoch. The redshift and mass dependence of these tracks is discussed in Wisnioski et al. (2015) and discussed in more detail throughout § 5.1, where we also discuss possible physical explanations for the

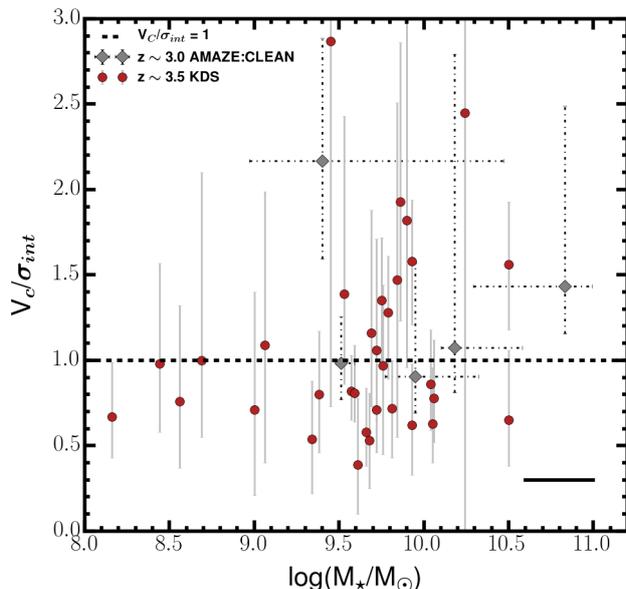


Figure 6. The distribution of galaxies in the V_C/σ_{int} vs. M_{\star} plane is shown for the isolated field sample, along with the galaxies in the AMAZE Clean ($z \simeq 3$) sample. We mark the $V_C/\sigma_{\text{int}} = 1$ line dividing rotation and dispersion-dominated galaxies in this panel, finding that more massive galaxies (e.g. $\log(M_{\star}/M_{\odot}) > 9.5$) are more likely to be rotation-dominated, with dispersion-dominated galaxies spread across the stellar mass range.

increase in velocity dispersions with redshift.

These new data from the KDS sustain the observation that velocity dispersion is increasing with redshift and provide strong constraints on the internal dynamics of typical star-forming galaxies between $z = 3 - 3.8$. When considered in conjunction with data from the literature in the right panel of Figure 5, it appears that the stellar mass of galaxies may play a role in mediating the velocity dispersion values (see § 5.1).

We proceed now to consider both the rotation velocities and velocity dispersions in tandem by considering V_C/σ_{int} and the rotation-dominated fraction of galaxies in the isolated field sample.

4.3 V_C/σ_{int} and the rotation-dominated fraction

In Figure 6 we plot the ratio of V_C/σ_{int} against M_{\star} for the isolated field sample, also including galaxies from the AMAZE Clean ($z \simeq 3$) sample for reference. In the local universe, disk, SFGs are observed to have V_C/σ_{int} values in excess of 10 (i.e. the disks are ‘well settled’), due to the decline of velocity dispersion across cosmic time, as gas fractions and disk turbulence decrease (e.g. Epinat et al. 2008a,b). Clearly the isolated field sample galaxies are in a different physical regime, with mean and median V_C/σ_{int} values of $1.08^{+0.18}_{-0.15}$ and $0.97^{+0.14}_{-0.11}$ respectively. This is despite many of the dispersion-dominated galaxies showing clear velocity gradients across the observed velocity maps (Appendix D), with the arctangent model providing a good

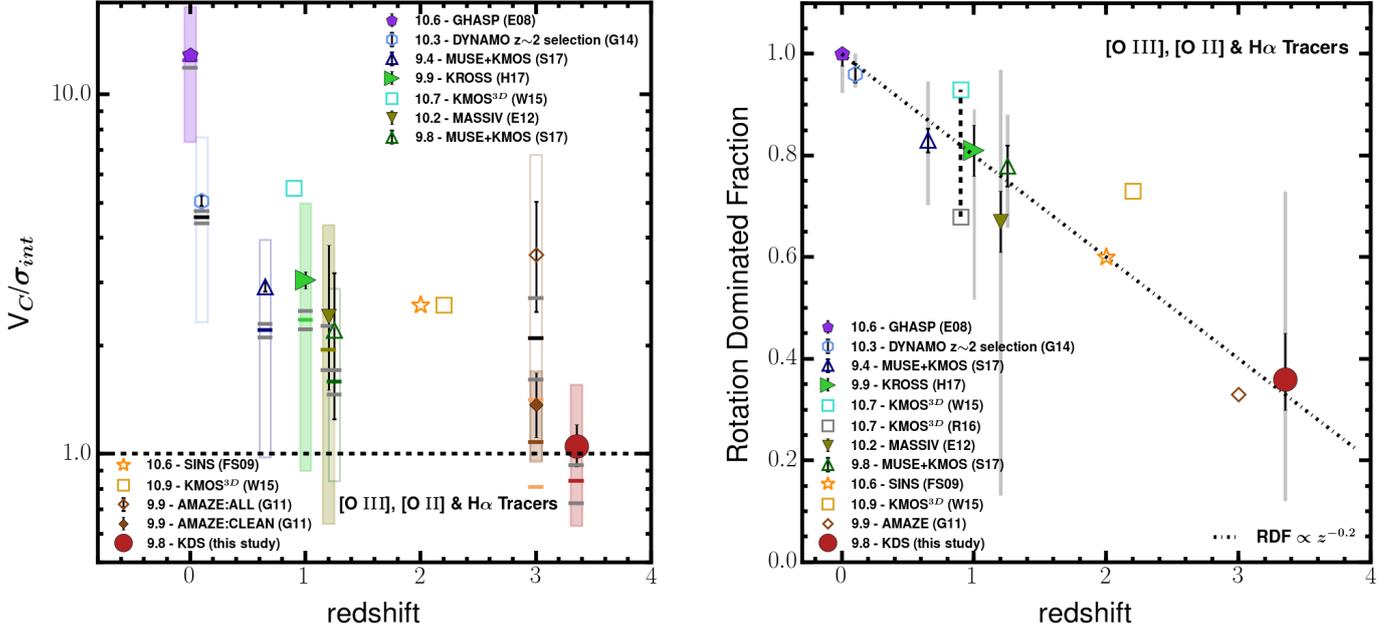


Figure 7. *Left:* V_C/σ_{int} is plotted against redshift for the comparison samples spanning $0 < z < 4$. We see a clear decline in the mean V_C/σ_{int} values with redshift, but with wide ranges in the individual measurements. There is a connection between V_C/σ_{int} and the stellar mass of the surveys, with larger mean M_* surveys reporting higher mean V_C/σ_{int} . *Right:* We plot the associated rotation-dominated fraction, defined as the fraction of galaxies in each sample with $V_C/\sigma_{\text{int}} > 1$. The narrow shaded regions represent the maximum and minimum RDF, computed as described in Appendix A. It appears that the rotation-dominated fraction drops from $\simeq 100\%$ in the local universe to 60% at intermediate redshifts and then to $\simeq 35\%$ at $z \simeq 3.5$, again with the trend for high V_C/σ_{int} values in the larger M_* surveys reflected by higher rotation-dominated fractions. The dot-dash line is the $\text{RDF} \propto z^{-0.2}$ cosmic decline suggested in Stott et al. (2016), which appears to follow the redshift decline of the rotation dominated fraction. In both panels, the symbol convention is the same as in the right panel of Figure 4.

fit to the rotation curves. This suggests that although rotational motions are present within the disks, they are small or comparable with the elevated σ_{int} values observed throughout the $z \simeq 3.5$ sample.

The $V_C/\sigma_{\text{int}} = 1$ line dividing rotation and dispersion-dominated galaxies is marked on Figure 6, separating the dispersion and rotation-dominated galaxies by definition. There is a spread of rotation and dispersion-dominated galaxies across the stellar mass range, with rotation-dominated galaxies more likely to be found above $\log(M_*/M_\odot) > 9.5$. The Rotation-Dominated Fraction (RDF), calculated with respect to the full field sample of 39 galaxies, is 0.36 ± 0.08 , which is substantially lower than in the local and intermediate redshift universe (e.g. Epinat et al. 2008a,b; Green et al. 2014). We interpret these results in the context of the cosmic evolution of the rotation dominated fraction using the comparison samples in § 4.3.1.

4.3.1 The evolution of the rotation-dominated fraction

In the left and right panels of Figure 7 we plot V_C/σ_{int} and the closely connected rotation-dominated fraction respectively, as a function of redshift, for the KDS isolated field sample and the comparison samples described in Appendix A. In each sample, the rotation dominated fraction

is computed as the fraction of galaxies with $V_C/\sigma_{\text{int}} > 1$, which, given the varying data quality across the comparison samples, is a simple and fair way to study the redshift evolution of this quantity. The filled and hollow symbols with black errorbars in the left panel of Figure 7 represent the mean and error on the mean, and the horizontal lines show the median and error on the median. The shaded regions give the 16th and 84th percentiles of the distribution of individual V_C/σ_{int} measurements, plotted here to indicate the range of measurements at each redshift slice. In the right panel of Figure 7, we have also computed the maximum and minimum RDF for these surveys by adding and subtracting the V_C/σ_{int} error values respectively and recomputing the rotation-dominated fraction. These are shown with the thin grey shaded regions and are intended to give an indication of the limits of the RDF at different redshift slices given the size of the errors on the individual points.

The rotation-dominated fraction, traced simply by the V_C/σ_{int} ratio inferred from ionised gas emission lines, drops from $\simeq 100\%$ of galaxies in the local universe to $\simeq 2/3$ of galaxies in the $z \simeq 1-2$ universe and to $\simeq 1/3$ above $z \simeq 3$, albeit with individual surveys reporting wide V_C/σ_{int} distributions at each redshift slice. This is highlighted in the right panel of Figure 7 by overplotting the $\text{RDF} \propto z^{-0.2}$ line which appears to roughly follow the decline of the rotation-dominated with redshift as described in Stott et al. (2016).

The width of the V_C/σ_{int} distributions indicate the galaxy diversity at each redshift slice, driven by the collection of galaxy masses, sSFRs, gas fractions, morphologies and evolutionary states which comprise each sample. Scatter above and below the $\text{RDF} \propto z^{-0.2}$ evolution line may be attributed in part to the different mean stellar masses of the surveys, with those that have larger stellar masses tending to scatter above this line. This draws again from the idea that higher stellar mass values suggest that the galaxies are more evolved, with more gas having been converted into stars for the bulk of the sample, which provides stability for a rotating disk and pushes the rotation-dominated fraction higher (see § 5). This is shown explicitly in [Simons et al. \(2017\)](#), where galaxies in higher stellar mass bins have larger V_C/σ_{int} at every epoch. It is important to note that the results in the right panel of [Figure 7](#) are subject to the same half-light radius and spatial resolution caveats discussed in section § 4.1.1, which may lead to an increase in velocity dispersions and a decrease in rotation velocities when galaxies are compact and the spatial resolution of the observations is poorer. Specifically for the KDS sample, the disk sizes are smaller than at intermediate redshift, and so it may be more challenging to fully resolve the rotational structure of the galaxies (e.g. [Newman et al. 2013](#)).

In summary, the rotation-dominated galaxies in the isolated field sample appear to conform with a redshift invariant $V_C - M_\star$ relation. Due to increased velocity dispersions, these rotation-dominated galaxies account for only 36% of the field sample and so are less representative of typical star-forming galaxies at $z \simeq 3.5$ than at intermediate and low redshifts. Explanations for the increase of σ_{int} with redshift (and hence the decline of V_C/σ_{int}) with cosmic time as well as the nature of dispersion-dominated galaxies have been discussed at length throughout other studies (e.g. [Förster Schreiber et al. 2009](#); [Law et al. 2009](#); [Burkert et al. 2010](#); [Newman et al. 2013](#); [Wisnioski et al. 2015](#)) and we revisit these in the following discussion.

5 DISCUSSION

Throughout § 4 we have investigated the derived kinematic parameters of the isolated field sample, showing first that, within the scatter, the rotation-dominated galaxies sit on a rotation velocity - stellar mass relationship consistent with what is observed in the local universe. The mean rotation velocity of the isolated field sample is low as a result of averaging over a sample containing a large number of dispersion-dominated galaxies (66% of the sample) and probable pressure support from turbulent motions (e.g. [Burkert et al. 2010](#)). This increased number of dispersion-dominated galaxies in comparison to low and intermediate redshifts appears to be driven by high velocity dispersions at $z \simeq 3.5$. In this section we discuss possible physical origins for the elevated random motions traced by σ_{int} and their impact on the galactic dynamics.

5.1 The origin and impact of increased velocity dispersions at high-redshift

Previous studies have claimed that the elevated velocity dispersions observed at high-redshift are a consequence of both internal and external processes. [Law et al. \(2009\)](#) observe galaxies with elevated velocity dispersions at $z \simeq 2.2$ (see [Figure 5](#)) and attribute this to gravitational instabilities following efficient accretion, which can occur when the cold gas becomes dynamically dominant.

[Genzel et al. \(2011\)](#) show that for dispersion-dominated galaxies, the scales susceptible to gravitational collapse are of the order of the galactic half-light radius. Systems such as these form highly unstable and roughly spherical stellar distributions with large internal motions at early times, before more gradual accretion from the galactic halo forms an extended, stable, gaseous disk which instigates kinematic order ([Law et al. 2009](#); [Genzel et al. 2011](#)). The physical mechanisms responsible for the increase in σ_{int} may also result in a reduction in the observed rotation velocities by partially compensating the gravitational force throughout the galactic disk (e.g. [Burkert et al. 2010](#); [Wuyts et al. 2016](#); [Ubler et al. 2017](#); [Genzel et al. 2017](#); [Lang et al. 2017](#)). Following this, σ_{int} decreases over time as ordered rotation begins to dominate and pressure support becomes increasingly insignificant (e.g. [Burkert et al. 2010](#); [Ubler et al. 2017](#); [Genzel et al. 2017](#); [Lang et al. 2017](#)). This is in line with the evolution of gas fractions observed at intermediate and high-redshifts (e.g. [Saintonge et al. 2013](#); [Tacconi et al. 2013, 2017](#)), with the gas reservoirs of galaxies fed by the accretion of cold gas.

[Burkert et al. \(2010\)](#) show, starting from the hydrostatic equation, that negative pressure gradients sourced by the disk turbulence generated through gravitational instabilities result in a decrease in observed rotation velocities. This has prompted others (e.g. [Wuyts et al. 2016](#); [Ubler et al. 2017](#); [Lang et al. 2017](#)) to incorporate a velocity dispersion term in the derivation of maximum circular velocities, which contributes to (but is not entirely responsible for) the observed shift in the zero-point of the stellar mass Tully-Fisher relation between $z \simeq 0 - 0.9$. [Ubler et al. \(2017\)](#) also claim that without the inclusion of this term incorrect conclusions surrounding the evolution of the Tully-Fisher relation will be drawn. We have not incorporated such a term when deriving the rotation velocities of the isolated field sample, but note that as a consequence of our galaxies showing some of the most extreme intrinsic dispersions observed, this would undoubtedly lead to a different distribution in the $V_C - M_\star$ plane. The contribution of pressure to rotation velocity for the KDS sample will be explored in more detail in a future work.

In [Wisnioski et al. \(2015\)](#) the authors present a scaling relation describing the increase of σ_{int} with redshift in terms of the key observable properties: gas fraction, f_{gas} , gas depletion timescale, t_{dep} , and specific star-formation rate, sSFR. Here we explore how well this model describes the evolution in velocity dispersion, incorporating our new measurement at $z \simeq 3.5$ from the KDS. The function describing the evolution of the velocity dispersion with redshift is shown below in [Equation 10](#), where Q_{crit} is the critical Toomre parameter for stability against gravitational collapse ([Toomre 1964](#)), and a is a dimensionless parameter which depends on

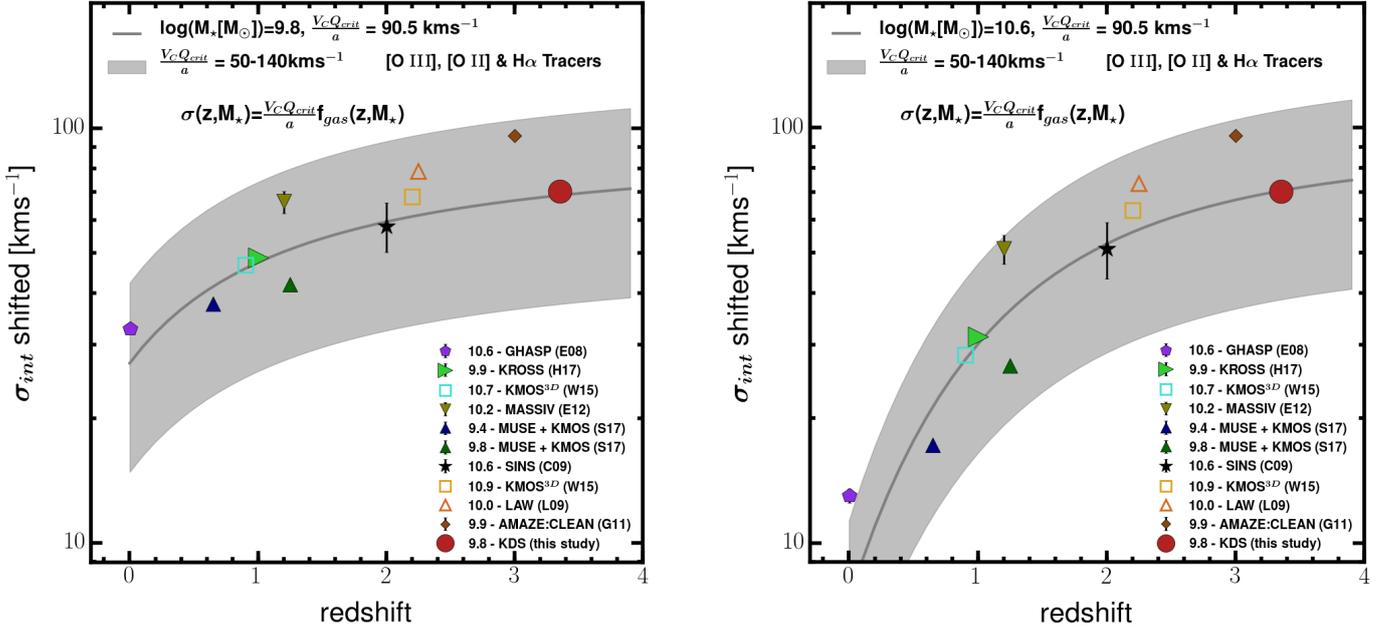


Figure 8. *Left:* We plot the σ_{int} values for the comparison samples as well as the isolated field sample, shifted as described in the text to a reference mass of $\log(M_*/M_\odot) = 9.8$ (the median mass of the isolated field sample). The grey line shows the model prediction with $V_C Q_{\text{crit}}/a = 90.5 \text{ km s}^{-1}$, as dictated by our data. The grey shaded region encompasses the model predictions with lower and upper bounds using $V_C Q_{\text{crit}}/a = 50 - 140 \text{ km s}^{-1}$ respectively (corresponding roughly to the velocity range $V_C = 70 - 200 \text{ km s}^{-1}$, spanned by galaxies in the comparison samples, under the assumption that $Q_{\text{crit}} = 1.0$ and $a = \sqrt{2}$ as discussed in the text). The shifted points are in line with the scenario whereby the sample averaged velocity dispersions increase with redshift as a result of increasing average gas fractions. *Right:* The same as in the left panel for a reference mass of $\log(M_*/M_\odot) = 10.6$. The steeper slope beyond $z \simeq 2$ highlights the model decrease in gas fraction, and hence velocity dispersion, for galaxies which have accumulated a larger stellar population.

the assumed distribution of gas and gravitational potential Forster Schreiber et al. (2006).

$$\sigma(z, M_*) = \frac{V_C Q_{\text{crit}}}{a} f_{\text{gas}}(z, M_*) \quad (10)$$

The gas fraction is a function of gas depletion timescale, with an assumed redshift dependence, and specific star-formation rate, with an assumed mass and redshift dependence (see Wisnioski et al. 2015 Equations 3-6, summarised below by Equations 11 and 12).

$$f_{\text{gas}} = \frac{1}{1 + (t_{\text{dep}} \text{SFR})^{-1}} \quad (11)$$

$$= \frac{1}{1 + (1.5 \times 10^a (M_*) (1+z)^{b(M_*) + \alpha})^{-1}}$$

$$a(M_*) = -10.73 + \frac{1.26}{1 + e^{(10.49 - \log(M_*/M_\odot))/(-0.25)}} \quad (12)$$

$$b(M_*) = 1.85 + \frac{1.57}{1 + e^{(10.35 - \log(M_*/M_\odot))/(0.19)}}$$

This relation is an expression of the belief that high-redshift galaxies experience more intense and efficient gas inflow, with the accumulation of a more massive stellar population providing stability against perturbations. In this equilibrium model framework (e.g. Davé et al. 2012; Lilly et al. 2013; Saintonge et al. 2013), populations of disk galaxies

with higher average stellar mass have lower average velocity dispersions due to this stability condition.

In the right panel of Figure 5 we plot Equation 10 for three different input mass values ($\log(M_*/M_\odot) = 9.8, 10.3, 10.6$, roughly spanning the mass range of the comparison samples), using the updated depletion time scaling relation from Tacconi et al. (2017) (i.e. $t_{\text{dep}} \propto (1+z)^{-0.57}$, meaning $\alpha = 0.57$ in Equation 11). The Equation $\sigma(z, M_*) \propto f_{\text{gas}}(z, M_*)$ is scaled so that the lowest mass track coincides with the KDS datapoint, finding a proportionality constant $V_C Q_{\text{crit}}/a = 90.5 \text{ km s}^{-1}$, which we assume when plotting all three tracks in Figure 5.

In this model, the velocity dispersions increase more steeply with increasing stellar mass, as shown by the solid and dashed lines in Figure 5. At higher redshift, these lines converge, suggesting that the predictions for the gas fraction become independent of stellar mass in this regime. In general, the surveys with higher stellar mass have lower velocity dispersions, in support of the model hypothesis.

Since stellar mass appears to be a crucial factor in regulating the velocity dispersions within galaxies, and the comparison samples have disparate stellar mass ranges, we use the model to shift the σ_{int} values according to their stellar mass to a reference mass value as per Wisnioski et al. (2015). This is done by computing the difference between the model predicted σ_{int} value at the stellar mass of each comparison survey, and at the reference mass for the appropriate redshift, and shifting the reported σ_{int} by this value.

We do this for two reference mass values; $\log(M_*/M_\odot) = 9.8$ (the median mass of the isolated field sample, left panel of Figure 8) and $\log(M_*/M_\odot) = 10.6$ (right panel of Figure 8).

In both panels, the solid grey line denotes the model prediction using the parameter combination determined above; $V_C Q_{\text{crit}}/a = 90.5 \text{ km s}^{-1}$. The grey shaded regions encompass model predictions with lower and upper bounds defined by $V_C Q_{\text{crit}}/a = 50 - 140 \text{ km s}^{-1}$ respectively. Assuming that $Q_{\text{crit}} = 1.0$ (as for a quasi-stable thin disk e.g. Forster Schreiber et al. 2006; Burkert et al. 2010) and $a = \sqrt{2}$ (as for a disk with constant rotation velocity Wisnioski et al. 2015), this range corresponds to $V_C = 70 - 200 \text{ km s}^{-1}$, which is roughly the velocity range spanned by our comparison samples, and the solid grey line corresponds to $V_C = 128 \text{ km s}^{-1}$. Whilst the KDS mean rotation velocity sits at the bottom of this range, possible explanations are: (1) The rotation velocities of the KDS galaxies are underestimated as a result of the velocity dispersion contribution discussed above; (2) $Q_{\text{crit}} \neq 1$ or $a \neq \sqrt{2}$ as a consequence of changing physical conditions at high-redshift.

Although the details of the coefficients of Equation 10 are unclear (i.e. V_C is not fixed at one value for the comparison samples and is itself linked to stellar mass, and both Q_{crit} and a could vary with redshift), the shape of the velocity dispersion evolution appears to follow that of the gas fraction, as suggested by the model. The difference between the slopes in the left and right panels highlights the role of stellar mass in the kinematic settling of galaxies between $z \simeq 3$ and the local universe. The isolated field sample data, in combination with comparison sample results spanning $z = 0 - 3$, thus suggest that the time evolution of the velocity dispersion measured from ionised gas emission lines is intimately linked to the consumption of gas and the accumulation of stars in a stable disk (e.g. Law et al. 2009, 2012c,b; Wisnioski et al. 2015). Rotation-dominated galaxies become less representative of the typical star-forming population at these redshifts as a result of the increased velocity dispersions (shown in Figure 7), constituting only $\simeq 1/3$ of the isolated field sample at $z \simeq 3.5$. This suggests that we are probing more chaotic and unstable systems, subject to gravitational instabilities and collapse on the disk scale length (Burkert et al. 2010; Genzel et al. 2011), residing in a much more active period of the universe in terms of both accretion and star-formation. Recent results indicate that these systems may become strongly baryon-dominated on the galaxy disk scale at high-redshift (Ubler et al. 2017; Lang et al. 2017; Genzel et al. 2017), which could imply that the increasing contribution of dark matter to galaxy disks with decreasing redshift plays a role in the evolution of velocity dispersions.

The precise interpretation of the velocity dispersion remains unclear, although there is mounting evidence that random motions have an increasingly significant contribution to the mass budget of high-redshift galaxies (e.g. Kassin et al. 2007; Law et al. 2009; Burkert et al. 2010; Kassin et al. 2012; Wuyts et al. 2016; Lang et al. 2017; Ubler et al. 2017; Genzel et al. 2017). The nature of dispersion-dominated galaxies at high-redshift has been discussed in Newman et al. (2013) using Adaptive-Optics (AO) and seeing limited observations of the same galaxies. The

AO observations show that, in some cases, galaxies which are classified as dispersion-dominated with seeing limited observations reveal much larger velocity gradients at higher spatial resolution. This highlights that beam-smearing effects can wash-out observed rotational motions, particularly for small galaxies. We have attempted to account for beam-smearing using half-light radius and velocity gradient dependent dynamical modelling, but we acknowledge that these observations are challenging and cannot rule out the presence of unresolved velocity gradients in our seeing-limited dataset. Smaller galaxies do have intrinsically smaller rotation velocities, whilst there is no observed dependence of velocity dispersion on galaxy size (e.g. Newman et al. 2013). Dispersion-dominated galaxies are therefore likely to be at an earlier evolutionary stage and unstable to gravitational instabilities which generate turbulence which is then maintained by the potential energy of the gas in the disk (Burkert et al. 2010; Newman et al. 2013).

Bringing the above points together, turbulent motions consume an increasingly significant fraction of the energy budget at high-redshift when typical star-forming galaxies are at earlier evolutionary stages. The KDS isolated field sample presents an extension of these trends to higher redshift, where galaxies are even more likely to be dispersion-dominated, and pressure support may be more significant than explored at $z \simeq 2.3$ (Wuyts et al. 2016; Ubler et al. 2017; Genzel et al. 2017; Lang et al. 2017). However, the precise interpretation of σ_{int} measured from ionised gas emission lines remains unclear, and future observations of both stellar and gaseous velocity dispersions in tandem with *JWST* will allow for unambiguous investigations into the nature of gaseous velocity dispersions at high-redshift.

In summary, the isolated field sample galaxies are diverse and kinematically immature, with $\simeq 2/3$ of the sample being dominated by random motions which may contribute to supporting a portion of the dynamical mass (e.g. Kassin et al. 2007; Burkert et al. 2010; Kassin et al. 2012; Newman et al. 2013; Straatman et al. 2017; Ubler et al. 2017; Lang et al. 2017). In the context of the equilibrium model explored here, these galaxies will evolve gradually with redshift along the velocity-mass relationship into stable disk galaxies which have converted their gas reservoirs into stellar mass (e.g. Lilly et al. 2013; Tacconi et al. 2013; Wisnioski et al. 2015; Tacconi et al. 2017). Future work at higher resolution and redshift will trace the extent to which gas-derived velocity dispersions trace the stellar dispersions and will investigate even earlier and more chaotic stages in the lifetimes of late-type galaxies.

6 CONCLUSIONS

We have presented new dynamical measurements of 33 typical star-forming galaxies spanning both cluster and field environments at $z \simeq 3.5$ as part of the KMOS Deep Survey, based on IFU data observed with KMOS. These measurements push back the frontier of IFU observations in the early universe and provide more robust constraints on the internal and rotational dynamics of $9.0 < \log(M_*/M_\odot) < 10.5$ typical star-forming galaxies at these redshifts. By using a com-

binned morpho-kinematic classification based on broad-band *HST* imaging and our IFU data, we have separated interacting galaxies from the sample, finding merger rates consistent between the field pointings ($\simeq 15\%$) and a very high merger rate in the cluster pointing ($\simeq 89\%$). We have made beam-smearing corrected measurements of V_C and σ_{int} for the remaining isolated field galaxies, and we interpret these in the context of previous dynamical studies using IFU data (Appendix A and Table A1). The main conclusions of this work are summarised as follows:

- We use a fractional error weighted mean to derive the mean kinematic parameters of the KDS isolated field sample of $V_C = 76.7^{+4.9}_{-4.5} \text{ km s}^{-1}$ and $\sigma_{\text{int}} = 70.8^{+3.3}_{-3.1} \text{ km s}^{-1}$ (right panels of Figures 4 and 5 respectively). Rotation (dispersion) dominated galaxies are defined as those with $V_C/\sigma_{\text{int}} > (<)1$.
- We plot the mean V_C values of the comparison samples and the KDS isolated field sample, finding significant diversity in V_C measurements in each survey (right panel of Figure 4). The spread in measurements is a result of the samples containing a mixture of rotation and dispersion-dominated galaxies with varying M_* values, and the low mean KDS value is primarily the result of averaging over a sample with a high fraction of dispersion-dominated galaxies and indicating the impact pressure support may have on reducing rotation velocities. At higher M_* , the accumulation of a stellar population provides stability, leading to an interplay between σ_{int} and V_C mediated across redshift by the cosmic decline of gas fractions, sSFRs and accretion (e.g. Law et al. 2012c; Wisnioski et al. 2015).
- When the V_C values are viewed as a function of mass in the inverse stellar mass Tully-Fisher Relation, the rotation-dominated galaxies are correlated with mass and lie within the errors on the same relation as derived for rotation-dominated galaxies with $9.0 < \log(M_*/M_\odot) < 11.0$ at $z \simeq 0.9$ in the KROSS survey (Figure 4). This relation is consistent with star-forming galaxies at $z = 0$ and so we report no significant evolution in the slope or zero-point of the smTFR between $z = 0-3.5$. Consistency of sample selection and measurement techniques are crucial factors when determining evolution in the V_C vs. M_* plane by comparing samples at different redshifts. The dispersion-dominated galaxies in the isolated field sample scatter below the trend.
- When considering the rotation-dominated galaxies alone, the mean and median rotation velocities are $V_C = 96.7^{+7.3}_{-7.2} \text{ km s}^{-1}$ and $V_C = 93.0^{+12.0}_{-12.5} \text{ km s}^{-1}$ respectively, which are roughly equivalent to a simple lower limit virial theorem calculation using the mean mass and half-light radius of the sample. This suggests that pressure forces, which are more significant at high-redshift, may play a role in supporting the total mass in all of the galaxies in the isolated field sample (e.g. Burkert et al. 2010).
- 14/33 isolated field sample galaxies are rotation-dominated and 19/33 are dispersion-dominated, with a tendency for galaxies with higher M_* to have larger values of V_C/σ_{int} (Figure 6). This gives a rotation-dominated fraction of $36 \pm 8\%$ in the isolated field sample, substantially less than surveys at lower redshift, although there is significant diversity amongst the individual V_C/σ_{int} measurements at each redshift slice (left and right panels of Figure 7). When viewed

as a function of redshift, the rotation dominated fraction appears to follow the scaling relation $\text{RDF} \propto z^{-0.2}$.

- We plot mean σ_{int} values computed in a consistent way for SFG samples spanning a wide redshift baseline (the comparison samples described in Appendix A), finding a sharp increase in σ_{int} values between $z = 0-1$ and a fairly shallow increase thereafter, mediated by the mean M_* of the galaxy samples (right panel of Figure 5). This is in line with a simple equilibrium model prescription in which the gas fractions and the impact of several physical mechanisms such as accretion of gas from the IGM, stellar feedback and turbulence increase with redshift and combine to increase random motions within high-redshift galaxies.

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APPENDIX A: COMPARISON SAMPLES

In § 4 we compare the isolated field sample results with the results of surveys tracing dynamics with ionised gas emission across a wide redshift baseline, to determine the evolving physical state of SFGs as the age of the universe increases. The galaxy selection criteria in these surveys, with the exception of DYNAMO (Green et al. 2014) as discussed below, consistently picks out representative star-forming galaxies at each epoch and the dynamical properties are traced by observing either the [O III] λ 5007, [O II] λ 3727,3729 or H α ionised gas emission lines.

Given the range of modelling and kinematic parameter extraction methods it is important to verify the extent to which the results from these surveys can be directly compared and treated as forming an evolutionary sequence, which we test in the following subsections by considering each survey in turn.

We make use of tabulated data from the surveys, where available, to compute sample averages using the following method. The fractional error weighted mean of V_C/σ_{int} , σ_{int} and V_C in each of the samples is computed (i.e. we do not want the derived values with extremely large errors to dominate the averages). The errors on these mean values are computed in a statistical sense, by generating bootstrapped samples, with replacement, with size equivalent to the original survey sample size and with values perturbed by a random number drawn from a gaussian distribution with width given by the error on the original point. The same process is applied to compute the errors on the sample medians and we

Table A1. The mean and median kinematic properties used throughout Figures 4, 5 and 7 for the different surveys. The errors on the mean and median represent the statistical errors from bootstrap resampling. For each of the comparison samples, the cross or tick after each of the mean kinematic properties indicates whether we have defined this as a fair value to compare against, with further details provided throughout Appendices A1 - A9.

Survey	$\langle z \rangle$	$\left\langle \log \left(\frac{M_\star}{M_\odot} \right) \right\rangle$	$\langle V_C \rangle$	med(V_C)	$\langle \sigma_{\text{int}} \rangle$	med(σ_{int})	$\left\langle \frac{V_C}{\sigma_{\text{int}}} \right\rangle$	med($\frac{V_C}{\sigma_{\text{int}}}$)	RDF
GHASP	0.001	10.6	189.0 $^{+3.5}_{-3.0}$ ✓	159.4 $^{+12.0}_{-15.0}$	13.0 $^{+0.5}_{-0.5}$ ✓	13.0 $^{+0.5}_{-0.5}$	12.9 $^{+0.5}_{-0.4}$ ✓	12.5 $^{+0.6}_{-0.6}$	1.00 $^{+0.00}_{-0.02}$ ✓
DYNAMO	0.1	10.3	183.0 $^{+1.0}_{-1.0}$ ✗	164.0 $^{+3.0}_{-5.0}$	45.9 $^{+0.3}_{-0.3}$ ✗	39.0 $^{+0.9}_{-1.0}$	5.1 $^{+0.2}_{-0.2}$ ✗	4.6 $^{+0.2}_{-0.2}$	0.96 $^{+0.02}_{-0.02}$ ✗
MUSE+KMOS	0.65	9.4	103.8 $^{+1.5}_{-1.6}$ ✗	73.0 $^{+2.5}_{-2.5}$	40.0 $^{+0.3}_{-0.3}$ ✓	32.0 $^{+0.6}_{-0.5}$	2.9 $^{+0.1}_{-0.1}$ ✗	2.21 $^{+0.1}_{-0.1}$	0.83 $^{+0.02}_{-0.02}$ ✗
*KROSS	0.9	9.9	117.0 $^{+4.0}_{-4.0}$ ✓	109.0 $^{+5.0}_{-5.0}$	- ✓	-	3.1 $^{+0.2}_{-0.2}$ ✓	2.4 $^{+0.1}_{-0.1}$	0.81 $^{+0.05}_{-0.05}$ ✓
KMOS ^{3D}	1.0	10.7	170.0	-	25.0 ✗	-	5.5 ✗	-	0.93 ✗
MASSIV	1.2	10.2	132.1 $^{+10.4}_{-8.2}$ ✓	103.0 $^{+13.4}_{-11.0}$	61.8 $^{+3.8}_{-4.2}$ ✓	52.0 $^{+5.2}_{-4.7}$	2.4 $^{+1.4}_{-0.9}$ ✓	2.0 $^{+0.3}_{-0.2}$	0.67 $^{+0.06}_{-0.06}$ ✓
MUSE+KMOS	1.25	9.8	75.5 $^{+3.0}_{-3.3}$ ✗	54.0 $^{+3.0}_{-2.8}$	42.0 $^{+0.5}_{-0.5}$ ✓	29.5 $^{+1.0}_{-1.0}$	2.2 $^{+1.0}_{-1.0}$ ✗	1.59 $^{+0.1}_{-0.1}$	0.78 $^{+0.04}_{-0.04}$ ✗
SINS(C09)	2.0	10.6	232.0 $^{+12.8}_{-12.7}$ ✓	240.0 $^{+18.0}_{-17.2}$	51.2 $^{+8.0}_{-7.9}$ ✓	42.5 $^{+9.1}_{-8.5}$	5.0 $^{+0.9}_{-1.0}$ ✗	4.7 $^{+1.0}_{-1.0}$	1.0 ✓
SINS(F09)	2.0	10.6	201.3 $^{+4.3}_{-4.0}$ ✓	174.0 $^{+12.1}_{-10.3}$	- ✗	-	2.6 ✗	-	0.60 ✗
KMOS ^{3D}	2.2	10.9	170.0 ✗	-	55.0 ✗	-	2.6 ✗	-	0.73 ✗
LAW 09	2.3	10.0	- ✗	-	78.0 $^{+6.5}_{-6.5}$ ✗	69.0 $^{+8.3}_{-7.5}$	- ✗	-	0.73 ✗
AMAZE (Full)	3.0	10.0	217.0 $^{+59.1}_{-40.2}$ ✗	129.0 $^{+47.5}_{-42.8}$	85.9 $^{+1.5}_{-1.4}$ ✗	78.0 $^{+4.5}_{-4.8}$	3.6 $^{+1.5}_{-1.1}$ ✗	2.1 $^{+0.6}_{-0.5}$	0.33 ✗
AMAZE (Clean)	3.0	10.0	140.8 $^{+40.1}_{-30.2}$ ✓	129.0 $^{+45.0}_{-33.4}$	95.7 $^{+2.0}_{-2.0}$ ✓	95.0 $^{+2.6}_{-2.6}$	1.4 $^{+0.3}_{-0.3}$ ✓	1.1 $^{+0.3}_{-0.3}$	- ✗
KDS	3.5	9.8	75.7 $^{+4.4}_{-4.4}$ ✓	57.5 $^{+5.5}_{-5.5}$	70.3 $^{+3.3}_{-3.1}$ ✓	71.0 $^{+5.0}_{-4.8}$	1.1 $^{+0.2}_{-0.1}$ ✓	0.84 $^{+0.1}_{-0.1}$	0.36 $^{+0.08}_{-0.08}$ ✓

* The intrinsic velocity dispersion values for the KROSS sample will be presented in Johnson et al. *in prep.*

report the 16th and 84th percentiles of the distributions of each of the above quantities as an indicator of the distribution width, and hence galaxy diversity, at each redshift slice. These results are listed in Table A1 and discussed in detail throughout § 4, where we will make statements about dynamical evolution by connecting the dots of these different surveys, assuming that on average they are tracing a population of SFGs which evolve across cosmic time. Throughout the plots in § 4 we highlight which values can be directly compared, due to consistency of sample selection and measurement methods, using filled and hollow symbols. This is also highlighted with the ticks and crosses beside the mean values in Table A1.

A1 GHASP ($z \simeq 0$)

The GHASP survey (Epinat et al. 2008a,b, E08) makes use of Fabry-Perot observations of 203 spiral and irregular galaxies with median $\log(M_\star/M_\odot) = 10.6$ in the local universe ($z \simeq 0$) to produce 3D $H\alpha$ datacubes. The data extend out to several half-light radii and in each case tilted ring models are fit to determine V_C at large radii where the rotation curves have flattened, with σ_{int} given as the dispersion in the rotation velocity across the whole field. In this survey, the mean and median kinematic properties computed from the tabulated literature values are directly comparable to the KDS and we used filled symbols in the plots throughout § 4 to highlight this.

A2 DYNAMO ($z \simeq 0.1$, but selected with physical properties resembling $z \simeq 2$)

In the DYNAMO survey (Green et al. 2014, G14), $H\alpha$ IFU data of 67 galaxies with median $\log(M_\star/M_\odot) = 10.3$ at

$z \simeq 0.1$ are presented. The selection criteria is such that half of the sample have high sSFRs representative of SFGs at $z \simeq 2$, and are also found to have higher f_{gas} values than locally, mimicking the physical conditions of $z \simeq 2$ SFGs. V_C values are extracted at a fixed radius from an arctangent model fit to the velocity fields ($1.6R_{1/2}$ as measured in the r-band) and σ_{int} is taken as the luminosity weighted average of the beam-smearing corrected dispersion field. We make use of tabulated data from this survey, but use hollow symbols throughout the plots in § 4 to highlight the difference in selection criteria.

A3 MUSE and KMOS ($z \simeq 0.65$ and $z \simeq 1.25$)

In Swinbank et al. (2017) a collection of $\simeq 400$ star-forming galaxies observed in $[\text{O II}]\lambda 3727, 3729$ with MUSE and $H\alpha$ emission with KMOS is described. Of this parent sample, inclination corrected velocities and intrinsic velocity dispersions are derived for 179 galaxies, to which we make comparisons. The observations span the redshift range $0.2 < z < 1.6$ and we split the data into two subsets with $z < 1.0$ (129 galaxies with mean redshift of $z = 0.65$, mean stellar mass of $\log(M_\star/M_\odot) = 9.80$) and $z > 1.0$ (50 galaxies with mean redshift of $z = 1.25$, mean stellar mass $\log(M_\star/M_\odot) = 9.80$). The reported velocities are extracted at $3R_D$, comparable to the extraction radius for the KDS sample, but these have not been beam-smearing corrected. The intrinsic velocity dispersions have been corrected for beam-smearing effects. For this reason, the velocity dispersions are deemed directly comparable to the KDS, reflected by the filled symbols throughout the plots in § 4, but the points involving velocity measurements are represented by hollow symbols due to the lack of a beam-smearing correction.

A4 KROSS ($z \simeq 0.9$)

We make comparisons to the most recent KROSS results (Harrison et al. 2017, H17), which presents KMOS H α observations of $\simeq 600$ SFGs with median $\log(M_*/M_\odot) = 9.9$ and $z \simeq 0.9$. V_C values are extracted from exponential disk fits to the data at $2R_{1/2}$ (differing from arctangent fits by a median of 0.5%) with the σ_{int} values also extracted at $2R_{1/2}$ when the data extend to this radius and when they do not σ_{int} is taken to be the median of the σ_{obs} map. These quantities are then beam-smearing corrected as described in H17 section 3.3.3. Due to the consistency of modelling and measurement techniques, the $z \simeq 0.9$ galaxies in this sample are directly comparable to the KDS and are plotted throughout § 4 with filled symbols.

A5 KMOS^{3D} ($z \simeq 1.0$ and $z \simeq 2.3$)

The KMOS^{3D} results described in (Wisnioski et al. 2015, W15) describe KMOS H α observations of $\simeq 191$ massive SFGs across two redshift slices. At $z \simeq 1.0$ the galaxies have $\log(M_*/M_\odot) = 10.7$ and at $z \simeq 2.3$ the median is $\log(M_*/M_\odot) = 10.9$, over an order-of-magnitude above the KDS median of $\log(M_*/M_\odot) = 9.8$. The rotation velocity is taken as $V_C = \frac{1}{2\sin(i)}(v_{\text{obs,max}} - v_{\text{obs,min}})$ and σ_{int} is extracted from the σ_{obs} map far from the kinematic centre, where the effects of beam-smearing are negligible. No tabulated values are provided for this survey, and so we cannot re-compute the mean and median values. We make use of values quoted in W15 at both redshift intervals throughout § 4, plotting these with hollow symbols due to the inconsistency of measurement method.

A6 MASSIV ($z \simeq 1.2$)

The MASSIV sample (Epinat et al. 2012, E12) uses SINFONI H α observations of 50 SFGs with median $\log(M_*/M_\odot) = 10.2$ at $z \simeq 1.2$. An arctangent function is fit to the data and V_C is extracted at $\simeq 1.7R_{1/2}$ and the σ_{int} value is derived from the σ_{obs} map by subtracting in quadrature the beam-smearing correction value as per § 3.2.5. We make use of tabulated values presented in this survey to compute the mean and median kinematic properties, plotted with filled symbols throughout § 4.

A7 SINS ($z \simeq 2.0$)

The SINS survey (Förster Schreiber et al. 2009, FS09) presents SINFONI H α observations of 80 massive galaxies with median $\log(M_*/M_\odot) = 10.6$ at $z \simeq 2.0$. Galaxies in this sample are classified as rotation or dispersion-dominated following FS09 section 9.5.1 which classifies rotation-dominated galaxies using the observed velocity and integrated velocity dispersion criteria $V_{\text{obs}}/(2\sigma_{\text{tot}}) > 0.4$.

V_C values for rotation-dominated galaxies are computed using a combined velocity gradient + width approach (Forster Schreiber et al. 2006) and for dispersion-dominated galaxies using the velocity width. Only integrated velocity dispersions are presented for the full sample. In addition, (Cresci et al. 2009, C09) model robust rotators in the SINS

sample using the IDL code DYSMAL, which derives rotation curves given an input radial mass distribution. In this approach, the V_C value comes from the best-fitting model parameter and σ_{int} is calculated using Equation A1, with the best-fitting σ_{01} used (reflecting thin and thick disks; see their text for more detail) and with σ_{02} , an additional component of isotropic velocity dispersion throughout the disk, left as a free parameter in the fitting.

$$\sigma_{01} = \sqrt{\frac{v^2(R)h_z}{R}} \quad \text{OR} \quad \sigma_{01} = \frac{v(R)h_z}{R}$$

$$\sigma_{\text{int}} = \sqrt{\sigma_{01}^2 + \sigma_{02}^2} \quad (\text{A1})$$

It is unclear whether the V_C values for the full sample from FS09 are directly comparable, but we make use of the tabulated values for completeness and plot these with filled symbols throughout § 4. The velocity dispersion values, V_C/σ_{int} and the rotation dominated fraction from FS09 are plotted with hollow symbols throughout § 4 (using the results listed in W15 for the point locations) due to the difference in measurement method for the velocity dispersion values.

The V_C and σ_{int} values in C09 are likely biased towards rotation-dominated galaxies with a well-settled disk. Nonetheless, we plot these values with filled symbols in § 4 as a directly comparable dataset, omitting the V_C/σ_{int} values and the rotation dominated fraction, which are listed for the full FS09 sample.

A8 LAW 09 ($z \simeq 2.3$)

In (Law et al. 2009, L09) OSIRIS [O III] $\lambda 5007$ and H α observations are collected for 13 galaxies with median $\log(M_*/M_\odot) = 10.0$ at $z \simeq 2.3$. The ‘velocity shear’ is computed as $v_{\text{shear}} = \frac{1}{2}(v_{\text{max}} - v_{\text{min}})$ without inclination correction and σ_{int} is the flux-weighted mean of σ_{obs} . We note that the σ_{mean} values and errors tabulated in the Law et al. (2009) paper are the flux-weighted mean and standard deviation of the measurements in individual spaxels in each galaxy, not corrected for beam-smearing in any way. These measurements are expected to be larger than the σ_{int} measurements at similar redshift which have been beam-smearing corrected, or extracted at the galaxy outskirts where the effects of beam-smearing are much smaller. We use tabulated data σ_{int} data in L09 to compare this distribution with other results, but plot with a hollow symbol due to the lack of a beam-smearing correction.

A9 AMAZE ($z \simeq 3.0$)

The AMAZE sample (Gnerucci et al. 2011, G11) presents SINFONI [O III] $\lambda 5007$ measurements for 33 galaxies with median $\log(M_*/M_\odot) = 9.9$ at $z \simeq 3$, closest in redshift to the KDS galaxies. In this study, rotation curves and intrinsic velocity dispersions are derived from a modelled exponential mass distribution, with the extracted V_C value taken as the large radius limit of the rotation curve and the σ_{int} as the maximum of the difference in quadrature between the σ_{obs} map and the σ_{model} map (which also takes

into account instrumental resolution and beam-smearing; see their Equation 8). A caveat when comparing with the G11 values is that dynamical properties are computed for 11 rotation-dominated galaxies in their sample, with the remaining 22 galaxies not analysed.

Also, for 5 galaxies in G11, V_C is not well constrained and no errorbar is given; for these objects we take the fractional error on V_C equal to 1 when computing the sample averages. The σ_{int} errors for individual galaxies are generally very small, and for 3 galaxies the σ_{int} value is consistent with 0 km s^{-1} . For the mean value computation for the full sample we set these equal to the SINFONI resolution limit of 30 km s^{-1} . A total of 6/11 galaxies either have no errors on V_C or have σ_{int} values consistent with 0 km s^{-1} , which we refer to as the AMAZE ‘Unconstrained’ sample and represent these with hollow symbols throughout the plots in § 4.

This Unconstrained sample also contains the galaxy s_sa22a-M38, common between the KDS and G11, classified here as a merger due to double *HST* components and twin peaks in the object spectrum, but classified in G11 as a rotating galaxy with $V_C = 346 \text{ km s}^{-1}$ (although the authors discuss the possibility that the galaxy could be either a close pair or two clumps embedded within a rotating disk). We use tabulated values from the full analysed sample of 11 galaxies, plotting these with hollow symbols throughout § 4 as a consequence of the caveats listed above. The 5/11 galaxies which are not in the Unconstrained sample are referred to as the AMAZE ‘Clean sample’ and are directly comparable to the KDS. These are plotted with filled symbols throughout § 4. We plot the quoted rotation-dominated fraction of 33% as a hollow symbol, noting that this could be higher if any of the 22/33 galaxies not analysed are rotation-dominated.

This comparison sample summary shows that there are many different approaches for computing the same intrinsic kinematic parameters, dictated by data quality and model preference, however, despite this diversity, over the past decade the studies appear to be converging on ‘modelling out’ the effects of beam-smearing using similar approaches (e.g. as described in Davies et al. 2011, where removing the effects of beam-smearing with spectrally and spatially convolved models is also the least biased approach). This trend towards consistency is encouraging, and, as described in § 3.2.2, we have also shaped our approach towards extracting intrinsic kinematic parameters after correcting for beam-smearing effects quantified in the modelling. We stress that, keeping in mind sample selection, mass ranges and kinematic parameter extraction methods are crucial when comparing results between different surveys.

APPENDIX B: INTRINSIC PARAMETER DISTRIBUTIONS

In Figure B1 we plot the distribution of the isolated field sample in the σ_{int} vs. $\log(V_C)$ plane. The $V_C = \sigma_{\text{int}}$ line roughly bisects the KDS sample, and separates the dispersion-dominated and rotation-dominated galaxies by definition. The isolated field sample galaxies are clustered

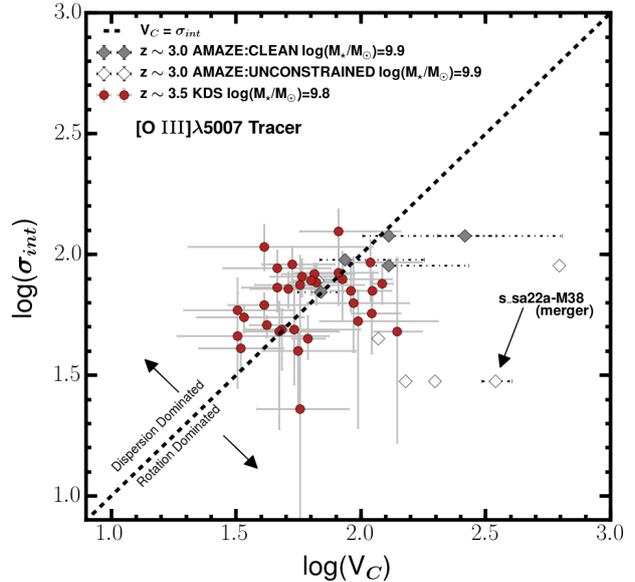


Figure B1. The distribution of the isolated field sample in the σ_{int} vs. V_C plane is plotted with the red symbols. The AMAZE Clean ($z \simeq 3$) sample galaxies (see appendix A9) are plotted with grey symbols and the AMAZE Unconstrained sample with hollow symbols. There is no strong correlation between the two parameters in either of these surveys and we note that the KDS and AMAZE Clean galaxies are generally in agreement. The $V_C/\sigma_{\text{int}} = 1$ line divides the galaxies classified as rotation-dominated and dispersion-dominated. The galaxy s_sa22a-M38, classified here as a merger but classified in Gnerucci et al. (2011) as a rotating galaxy with $V_C = 346 \text{ km s}^{-1}$ is marked on the plot.

around a relatively tight region in both V_C and σ_{int} and there is good agreement between the KDS and AMAZE Clean ($z \simeq 3$) samples. The galaxies in the AMAZE Unconstrained sample generally have much larger V_C values despite occupying a similar stellar mass range to our isolated field sample, with this being a consequence of large model extrapolations beyond the observed data. We follow the bootstrapping procedure described in Appendix A to compute error weighted averages and statistical errors.

APPENDIX C: KINEMATIC PARAMETERS ERROR ESTIMATES

C1 3D modelling kinematic parameter error estimates

As mentioned in § 3.2.2, the MCMC sampling provides distributions for each of the model parameters, from which we can extract the 84th and 16th percentile values as the $\pm 1\text{-}\sigma$ errors. During the following procedure, PA_{kin} is fixed to the maximum-likelihood value for all model evaluations. We proceed to reconstruct the beam-smearred and intrinsic dynamical models for each of the isolated field galaxies using the 16th and 84th percentile parameters, with the lower velocity 16th percentile model constructed using the upper limit on the inclination, and the 84th percentile model using the lower limit. To clearly distinguish the $1\text{-}\sigma$ error regions for both the beam-smearred and intrinsic models, the

region between the 16th and 84th percentile evaluations is shaded blue and red respectively in the velocity extraction panels throughout Figures D1 and D2. The +/- 1- σ error values for both V_C are then calculated using the equations below, which also take into account the measurement errors, with $\bar{\sigma}_{v_{\text{obs}}}$ equal to the average observational uncertainty extracted along PA_{kin} . The subscripts ‘16th’ and ‘84th’ denote the parameters used to construct the model from which the velocity has been extracted.

$$\delta V_{C+} = \sqrt{(V_{C-84} - V_C)^2 + (\bar{\delta}V_{\text{obs}})^2} \quad (\text{C1})$$

$$\delta V_{C-} = \sqrt{(V_C - V_{C-16})^2 + (\bar{\delta}V_{\text{obs}})^2} \quad (\text{C2})$$

The upper and lower errors on the σ_{int} are calculated using a similar approach. To incorporate the uncertainty introduced in the modelling by assuming a fixed value of $\sigma_{\text{int}} = 50\text{km s}^{-1}$, we make two further model evaluations using both the 16th and 84th percentile parameters with $\sigma_{\text{int}} = 40\text{km s}^{-1}$ and $\sigma_{\text{int}} = 80\text{km s}^{-1}$, with these limits roughly encompassing the range of observed and predicted velocity dispersion values at $z \simeq 3.5$ as shown in Wisnioski et al. (2015) figure 10. When assuming the broader $\sigma_{\text{int}} = 80\text{km s}^{-1}$ value in each spaxel the beam-smearing correction decreases, as the impact of convolution with spectral lines in adjacent spaxels with shifted velocity centres is less severe. Conversely, assuming $\sigma_{\text{int}} = 40\text{km s}^{-1}$ increases the beam-smearing. We calculate the minimum and maximum intrinsic velocity dispersion, $\sigma_{\text{int-84-40}}$ and $\sigma_{\text{int-16-80}}$, using equations C3 and C4 respectively:

$$\sigma_{\text{int-84-40}} = \sqrt{(\sigma_{\text{obs}} - \sigma_{\text{bs-84-40}})^2 - \sigma_{\text{sky}}^2} \quad (\text{C3})$$

$$\sigma_{\text{int-16-80}} = \sqrt{(\sigma_{\text{obs}} - \sigma_{\text{bs-16-80}})^2 - \sigma_{\text{sky}}^2} \quad (\text{C4})$$

where $\sigma_{\text{bs-84-40}} = \sigma_{\text{model-84-40}} - 40$ and $\sigma_{\text{bs-16-80}} = \sigma_{\text{model-16-80}} - 80$ are the beam-smearred maps. The lower and upper errors on σ_{int} are then given by the following:

$$\delta \sigma_{\text{int}+} = \sqrt{(\sigma_{\text{int}} - \sigma_{\text{int-84-40}})^2 + (\bar{\sigma}_{\sigma_{\text{obs}}})^2} \quad (\text{C5})$$

$$\delta \sigma_{\text{int}-} = \sqrt{(\sigma_{\text{int-16-80}} - \sigma_{\text{int}})^2 + (\bar{\sigma}_{\sigma_{\text{obs}}})^2} \quad (\text{C6})$$

with $\bar{\sigma}_{\sigma_{\text{obs}}}$ equal to the average of the velocity dispersion measurement errors. Once these quantities have been measured, the upper and lower errors on the ratio V_C/σ_{int} can be computed using equations C7 and C8.

$$\delta \frac{V_C}{\sigma_{\text{int}}} + = \frac{V_C}{\sigma_{\text{int}}} \sqrt{\left(\frac{\delta V_{C+}}{V_C}\right)^2 + \left(\frac{\delta \sigma_{\text{int}+}}{\sigma_{\text{int}}}\right)^2} \quad (\text{C7})$$

$$\delta \frac{V_C}{\sigma_{\text{int}}} - = \frac{V_C}{\sigma_{\text{int}}} \sqrt{\left(\frac{\delta V_{C-}}{V_C}\right)^2 + \left(\frac{\delta \sigma_{\text{int}-}}{\sigma_{\text{int}}}\right)^2} \quad (\text{C8})$$

APPENDIX D: KINEMATICS PLOTS

In Figures D1, D2 and D3 we plot the kinematic grids for the KDS galaxies classified as rotation-dominated, dispersion-dominated and merger candidates respectively. The figure captions provide more information on each of the panels of these grids.

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.

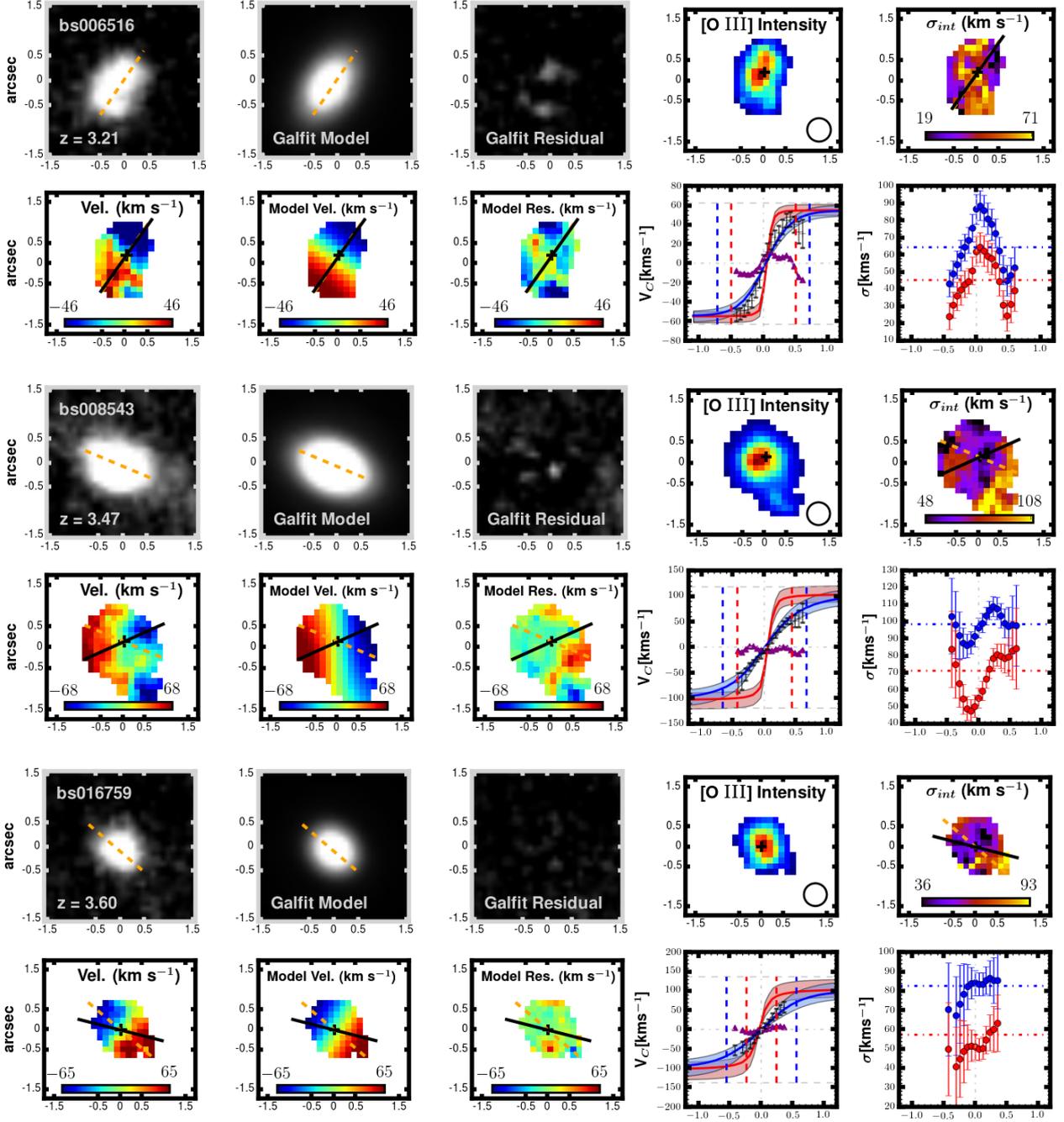


Figure D1. We plot the kinematic grids for the 14/33 isolated field sample galaxies classified as rotation-dominated. For each galaxy we plot the *HST* image, GALFIT model and GALFIT residuals, along with the observed [O III] λ 5007 intensity map, dispersion map and velocity map extracted from the datacubes with gaussian-fits to the individual spaxels as described in § 3.2.1. The solid black line and dashed orange line plotted in various panels show PA_{kin} and PA_{morph} respectively. On the bottom row we plot the beam-smearing model velocity and data-model velocity residuals with colour bar limits tuned to the velocity model. Also on the bottom row, we plot extractions along PA_{kin} for both the velocity and velocity dispersion. The grey points with errorbars on the velocity extraction plot are from the observed velocity map, the blue line and blue shaded regions represent the beam-smearing model fit to the data and errors respectively, whereas the red line and shaded regions are the intrinsic model from which V_C is extracted. The purple symbols represent the extraction along PA_{kin} from the residual map, and the two vertical dashed lines denote the intrinsic (inner) and convolved (outer) $2R_{1/2}$ values, with V_C extracted at the intrinsic value and V_{BS} extracted at the convolved value of the intrinsic and observed profiles respectively. The intrinsic models flatten at small radii, and so small changes in $R_{1/2}$ have negligible impact on the final, extracted V_C values. The blue points on the velocity dispersion extraction plot show the values extracted from the observed dispersion map and the red are from the beam-smearing corrected map as per Equation 6, and the vertical dashed lines have the same meaning as for the velocity extraction plot. Also plotted with the blue and red dot-dash horizontal lines are the median values of the observed and intrinsic (i.e. σ_{int}) dispersion maps respectively.

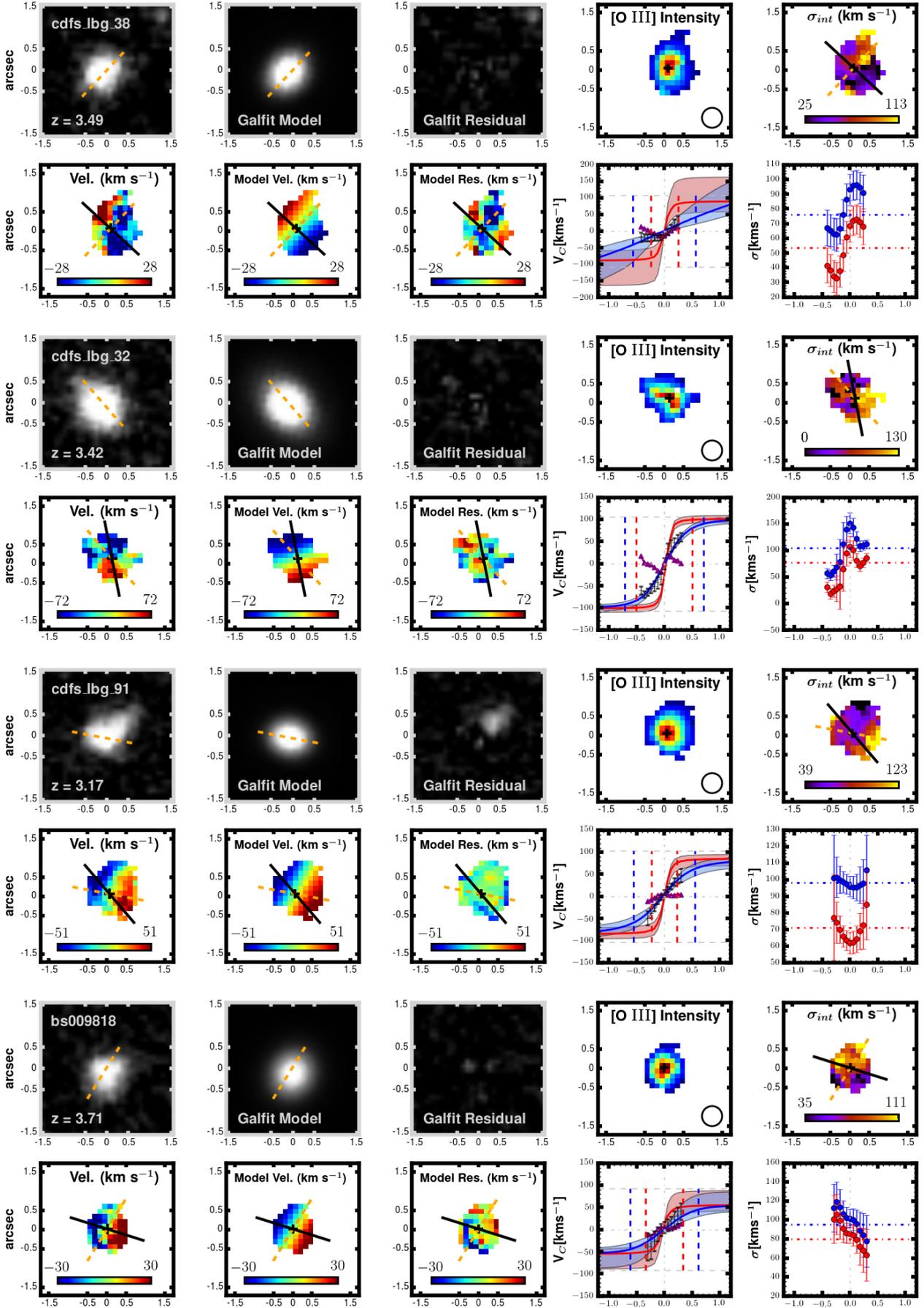


Figure D1. Continued.

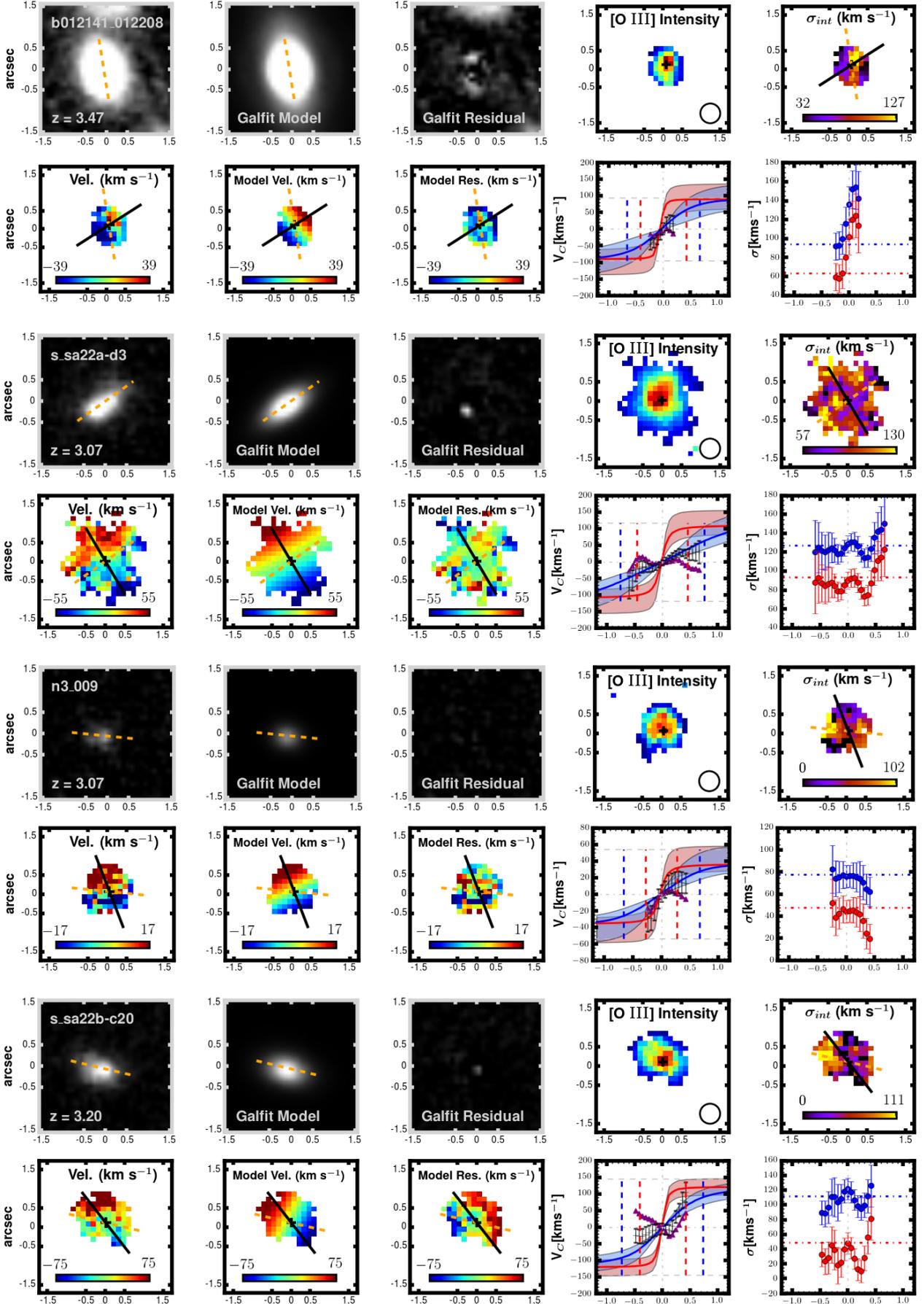


Figure D1. Continued.

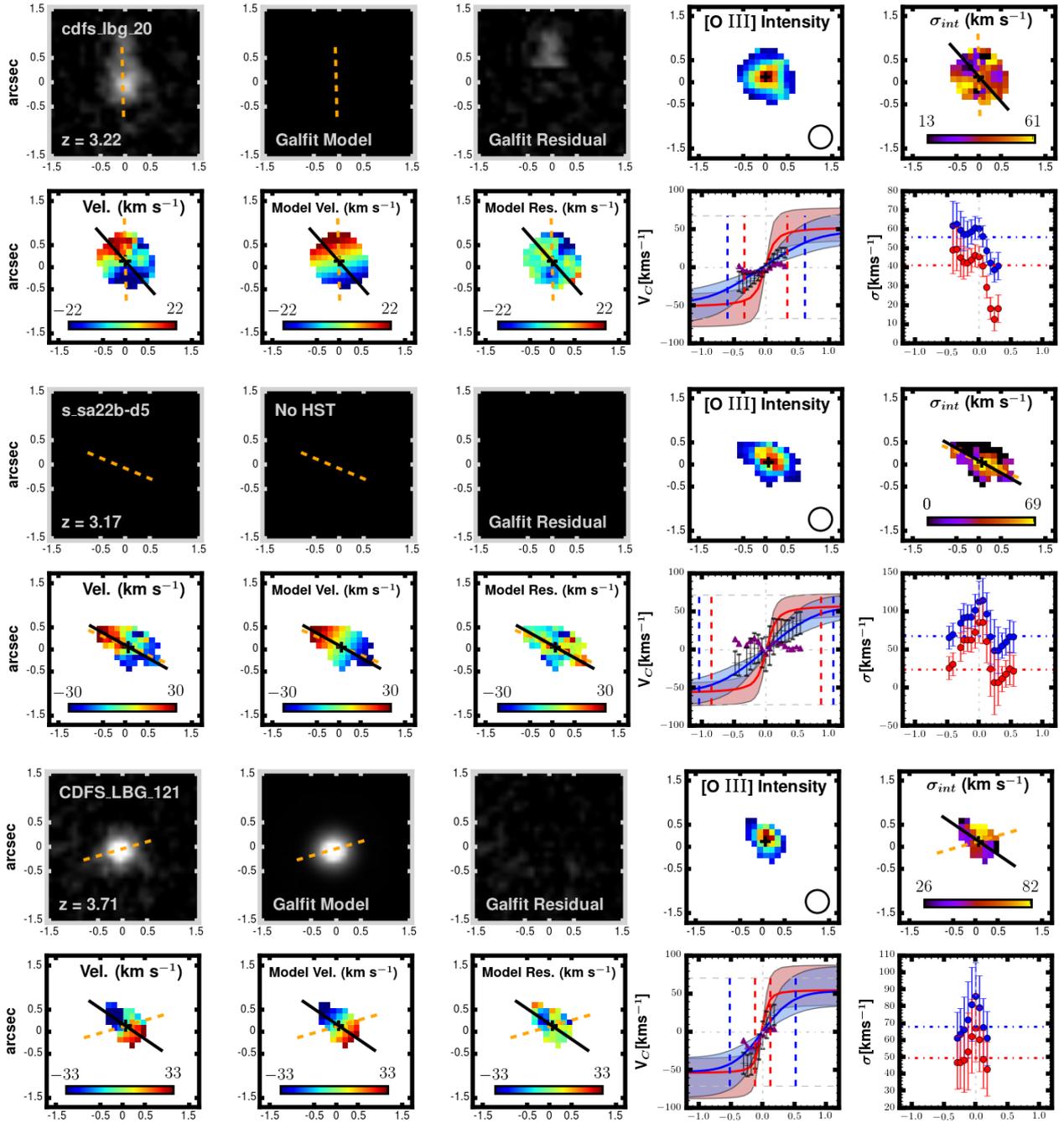


Figure D1. Continued.

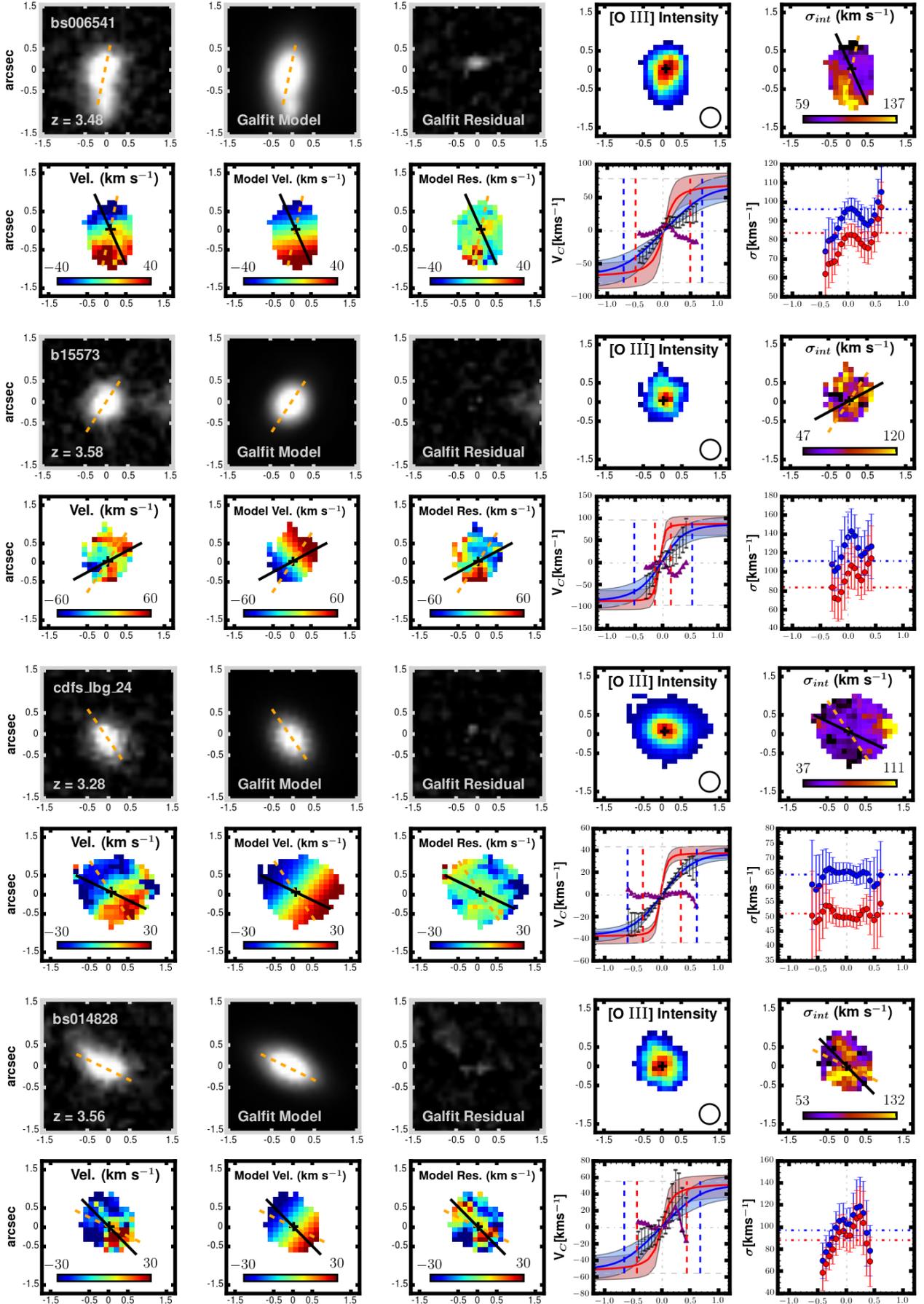


Figure D2. The same as in figure D1 but for the dispersion-dominated galaxies.

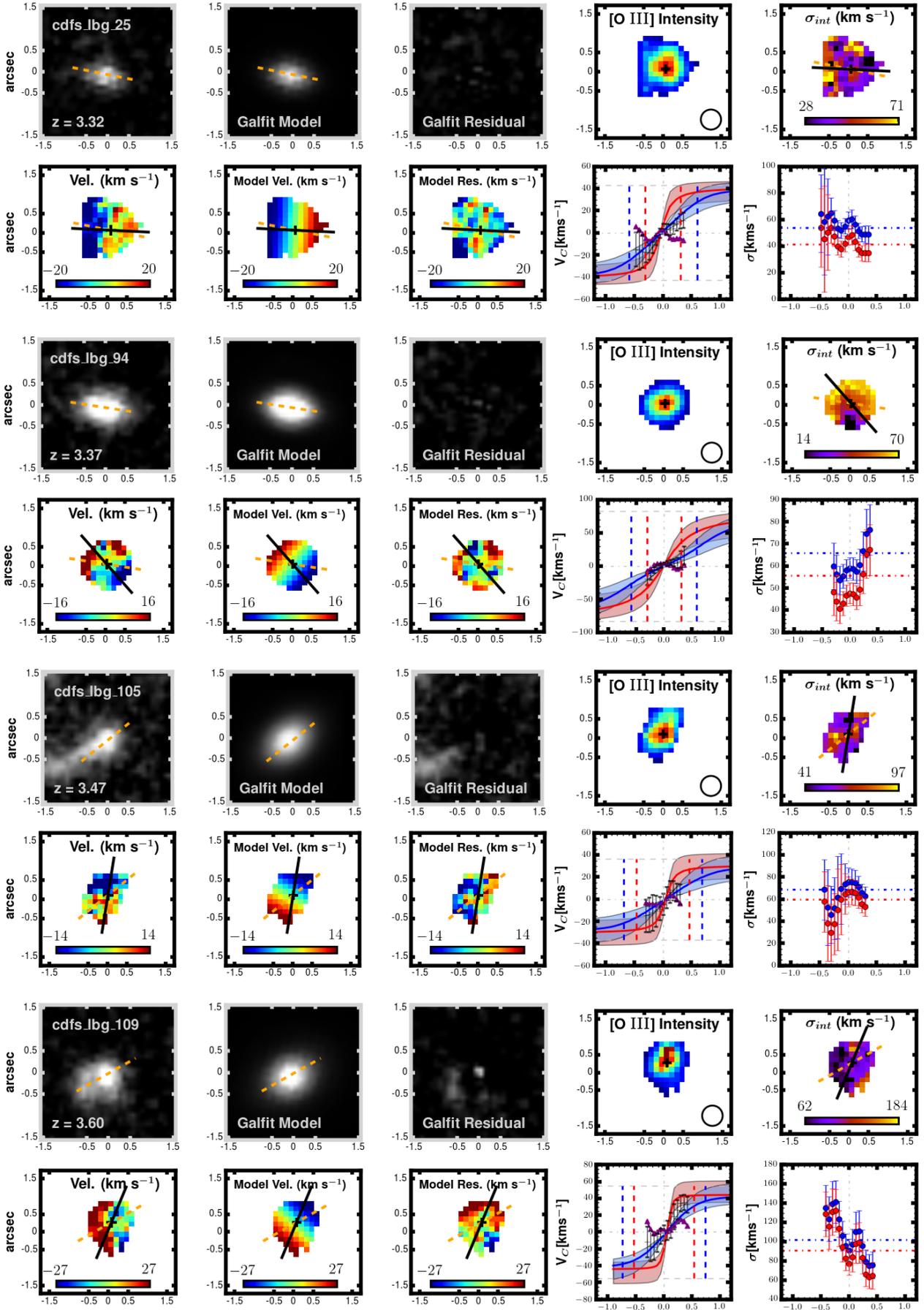


Figure D2. Continued.

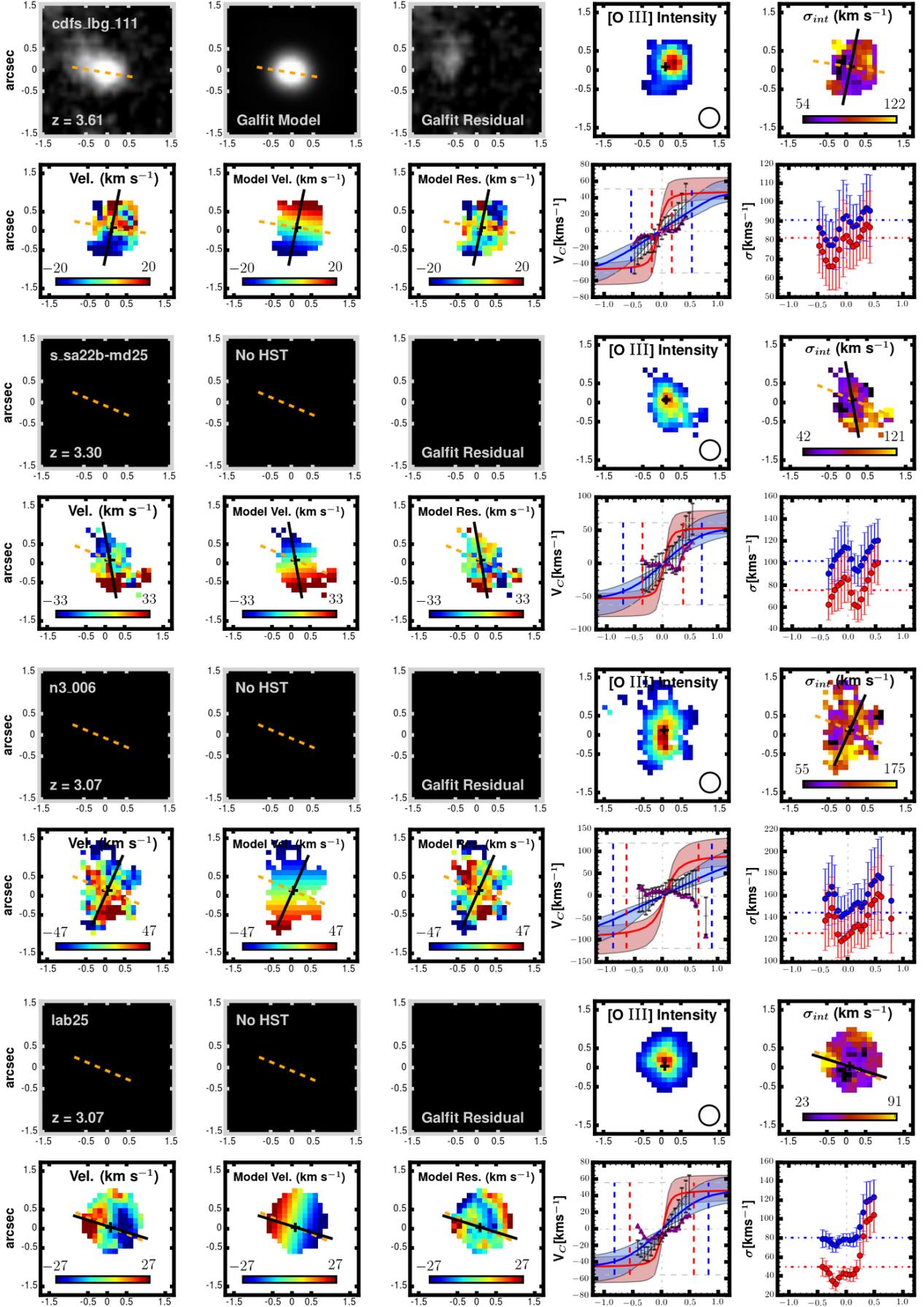


Figure D2. Continued.

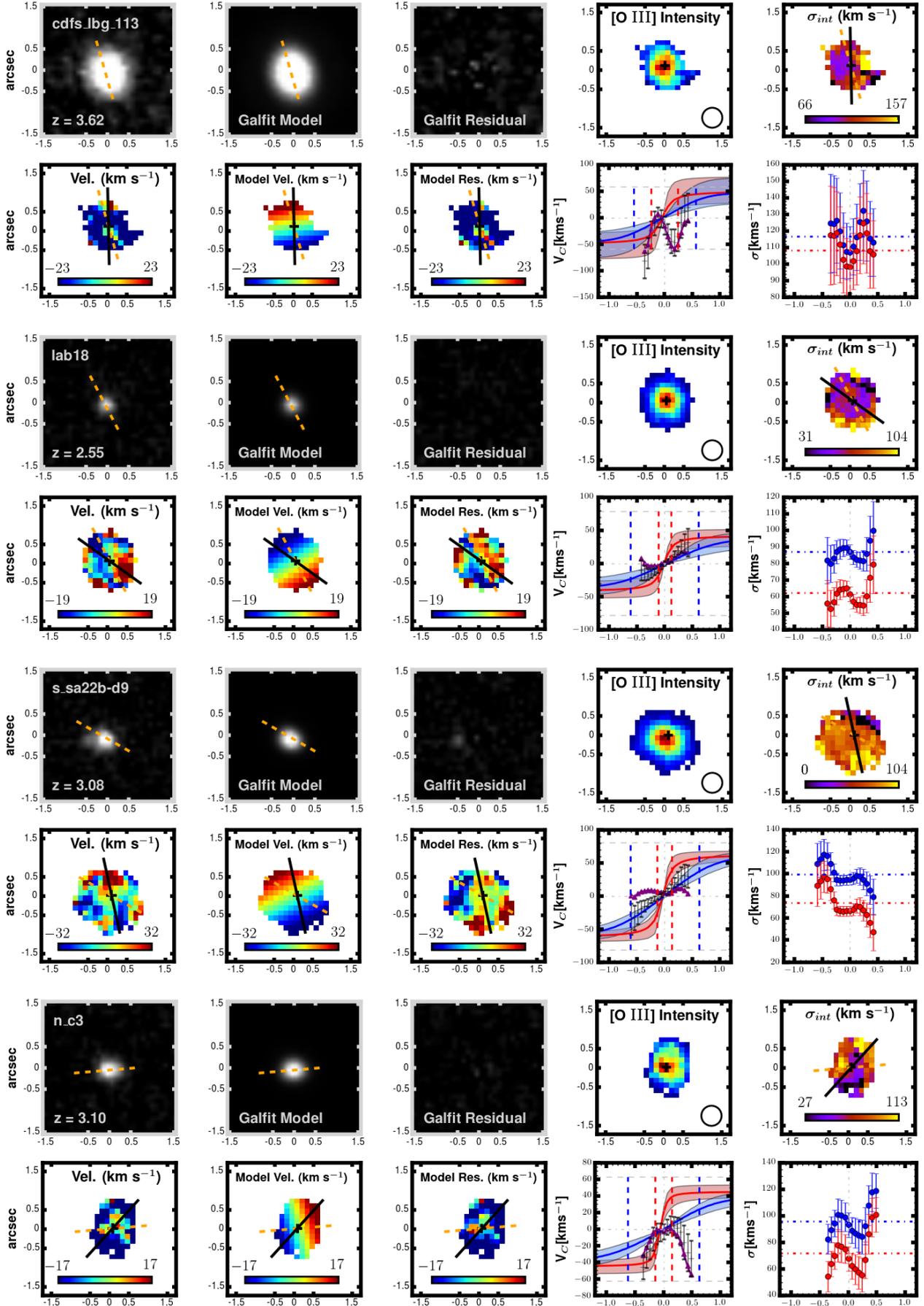


Figure D2. Continued.

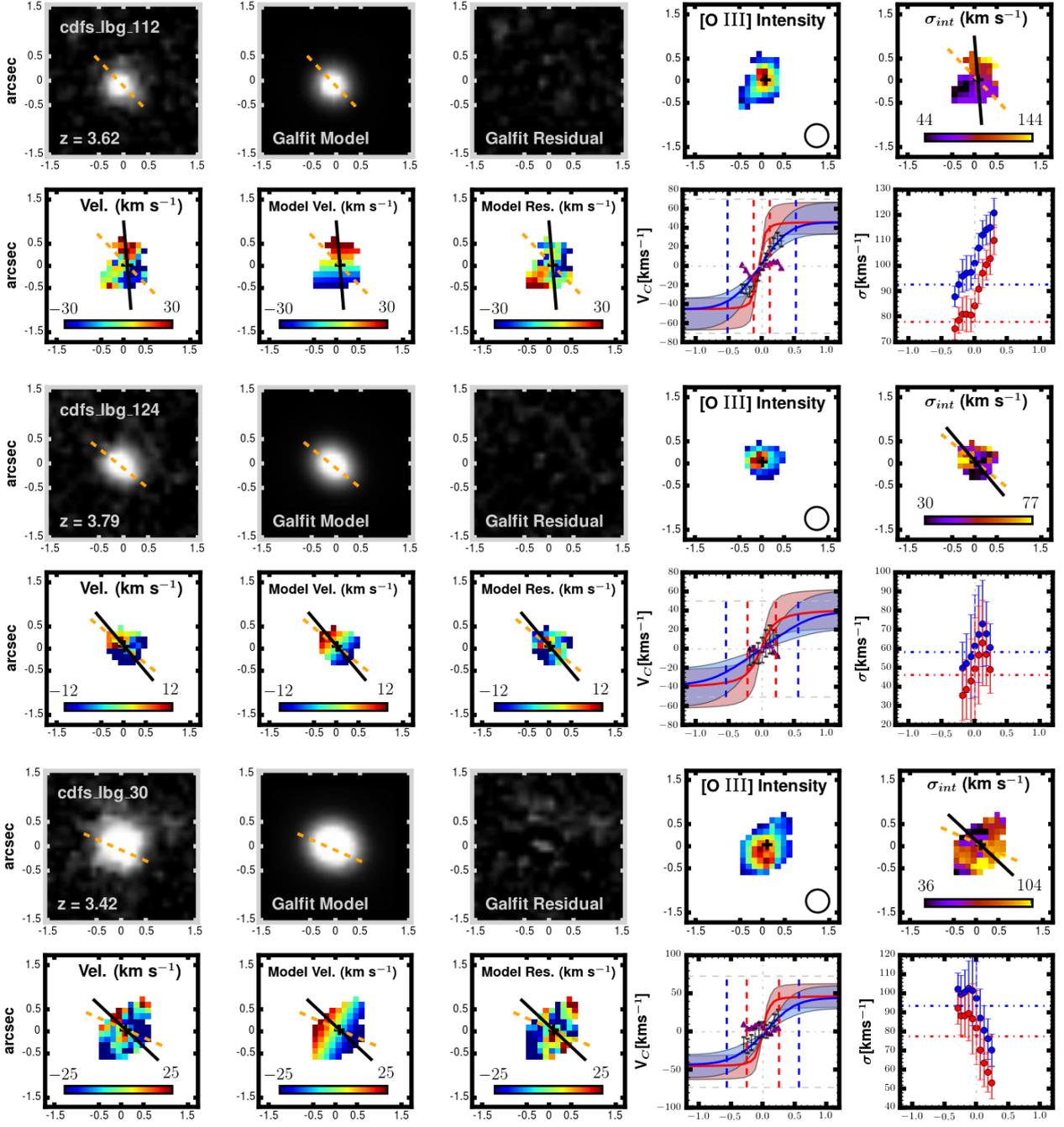


Figure D2. Continued.

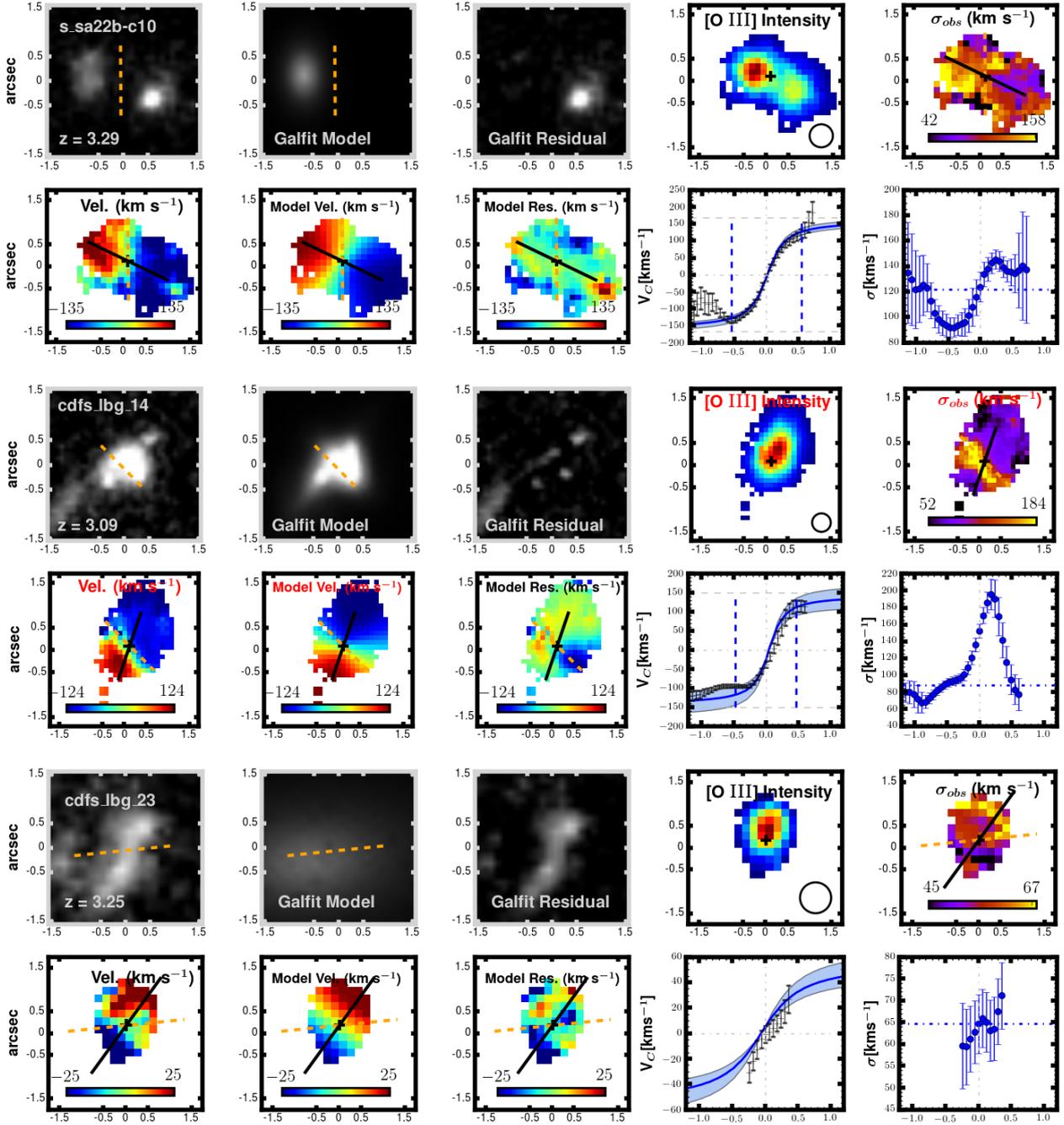


Figure D3. The same as for figure D1 but for the merger candidates. In this case we plot only fits to the data in the velocity extraction plot with the blue line, rather than attempting the full beam-smearing analysis. Several of the galaxies here mimic rotation from a purely kinematic perspective, but have two or more *HST* components and an accompanying double peak in the object spectrum at the object centre, leading to large velocity dispersions when single gaussian-fits are attempted.

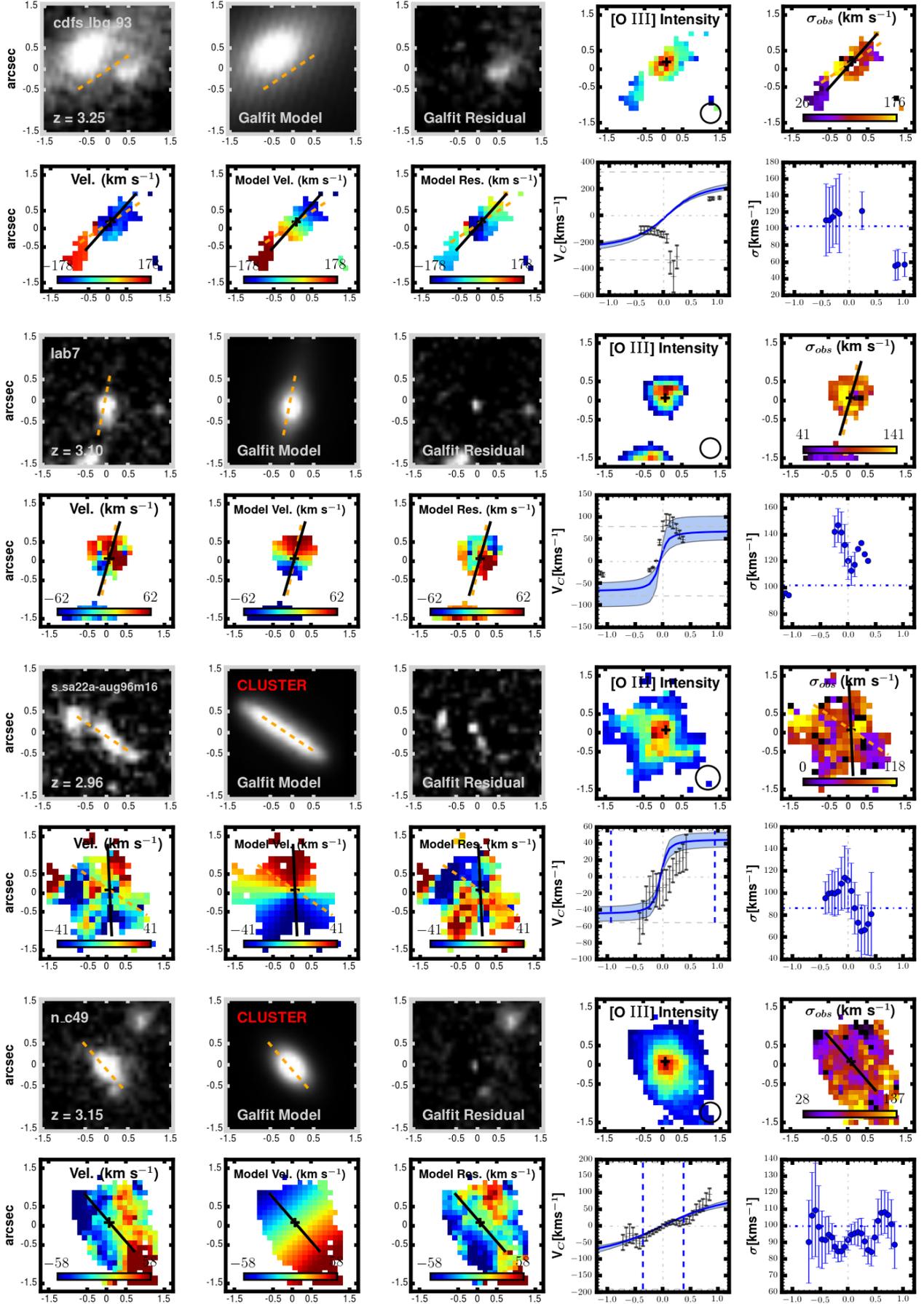


Figure D3. Continued.

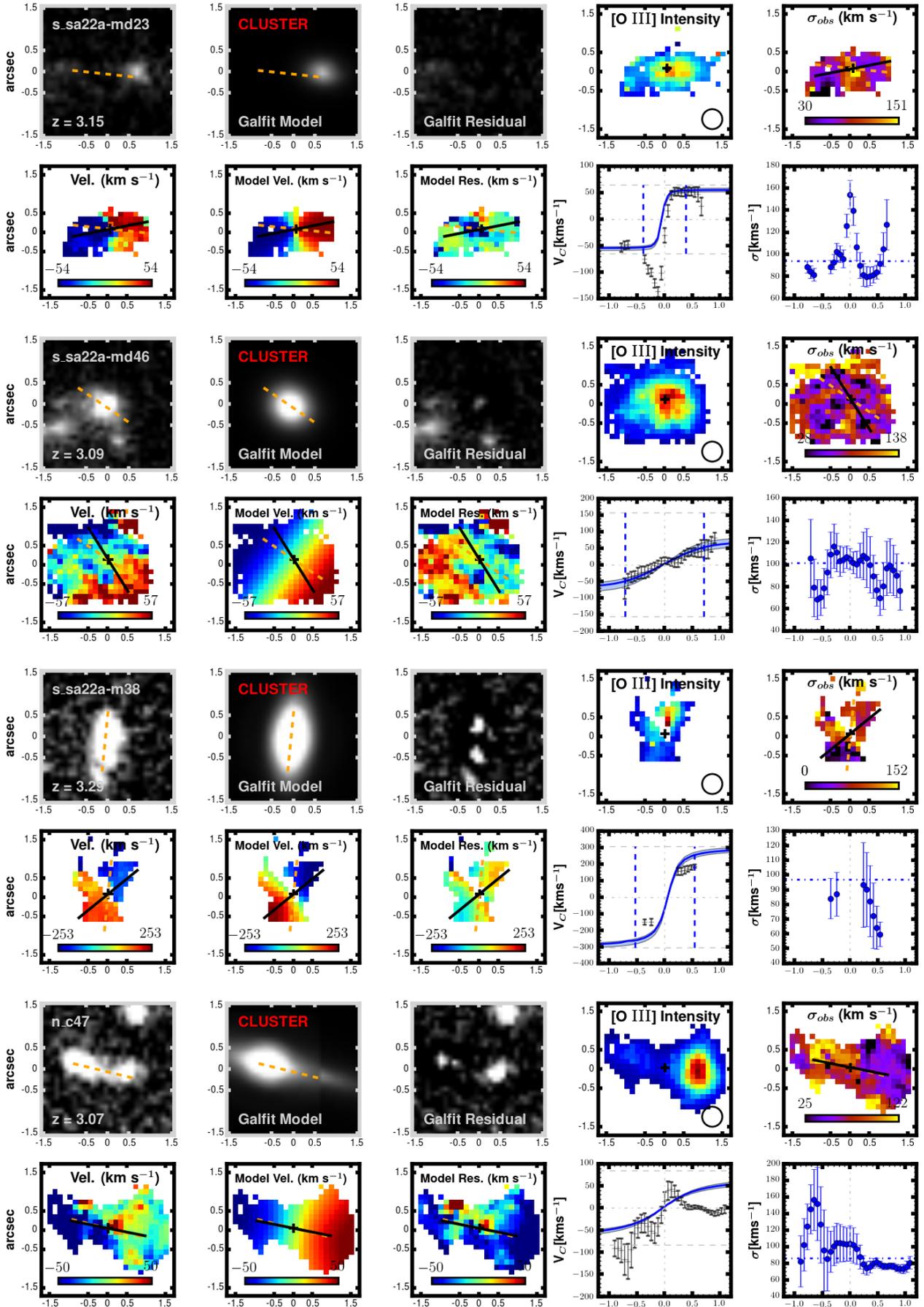


Figure D3. Continued.

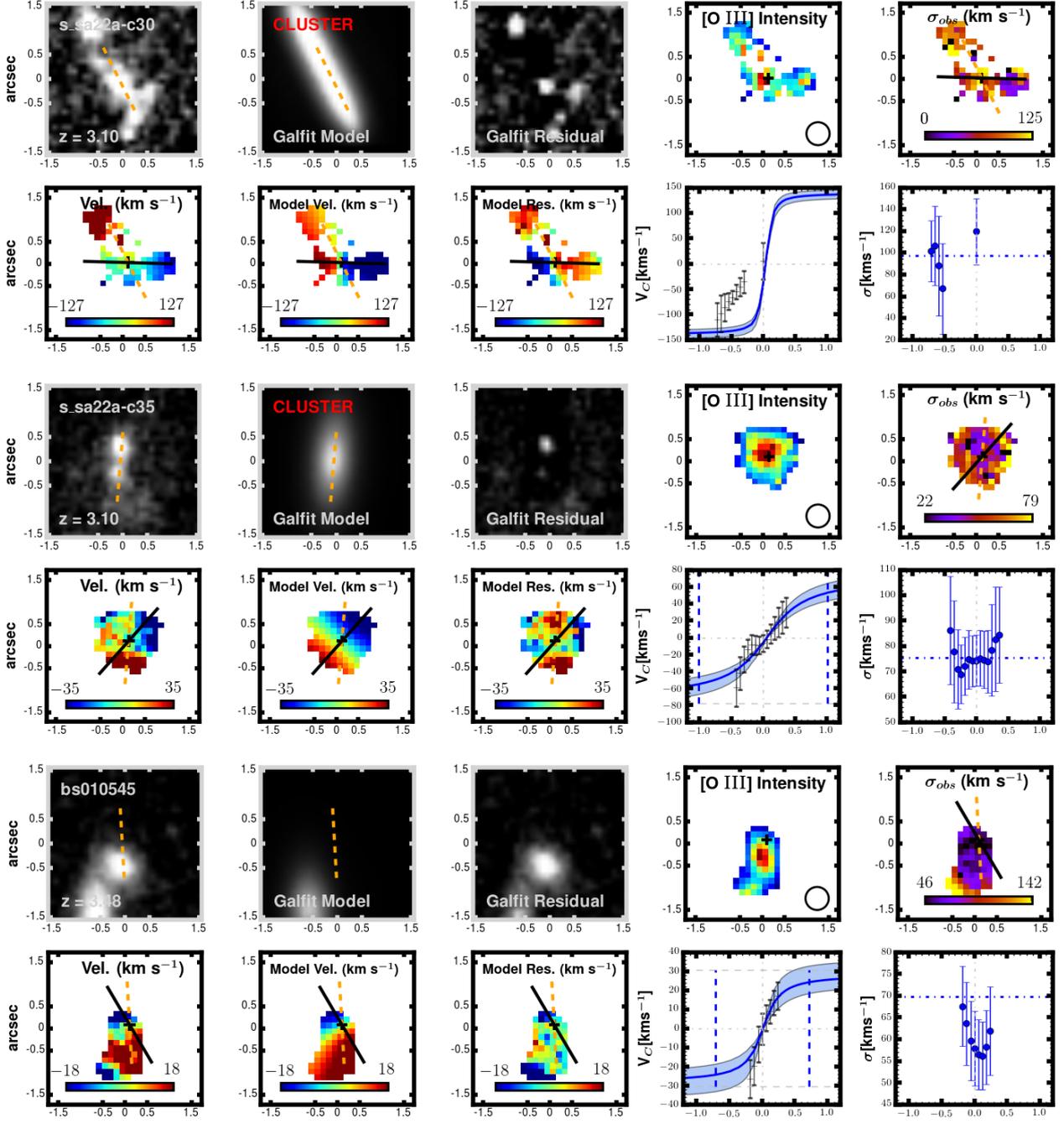


Figure D3. Continued.