












High-redshift Extremely Red Quasars in X-Rays

Andy D. Goulding¹, Nadia L. Zakamska^{2,3} , Rachael M. Alexandroff^{2,4,5} , Roberto J. Assef⁶ , Manda Banerji⁷, Fred Hamann⁸, Dominika Wylezalek^{2,9} , William N. Brandt¹⁰ , Jenny E. Greene¹, George B. Lansbury^{7,11} , Isabelle Pâris¹², Gordon Richards¹³ , Daniel Stern¹⁴ , and Michael A. Strauss¹ 

¹ Department of Astrophysical Sciences, Princeton University, Princeton, NJ 08540, USA; goulding@astro.princeton.edu

² Department of Physics & Astronomy, Johns Hopkins University, Bloomberg Center, 3400 N. Charles St., Baltimore, MD 21218, USA

³ Deborah Lunder and Alan Ezekowitz Founders' Circle Member, Institute for Advanced Study, Einstein Dr., Princeton, NJ 08540, USA

⁴ Canadian Institute for Theoretical Astrophysics, 60 St. George St., University of Toronto, Toronto, ON M5S 3H8, Canada

⁵ Dunlap Institute for Astronomy and Astrophysics, The University of Toronto, Toronto, ON M5S 3H4, Canada

⁶ Núcleo de Astronomía de la Facultad de Ingeniería, Universidad Diego Portales, Av. Ejército Libertador 441, Santiago, Chile

⁷ Institute of Astronomy, University of Cambridge, Madingley Rd., Cambridge CB3 0HA, UK

⁸ Department of Physics & Astronomy, University of California, Riverside, CA 92507, USA

⁹ European Southern Observatory, Karl Schwarzschild Straße 2, D-85748 Garching bei München, Germany

¹⁰ Department of Astronomy & Astrophysics, The Pennsylvania State University, University Park, PA 16802, USA

¹¹ Centre for Extragalactic Astronomy, Department of Physics, Durham University, South Rd., Durham DH1 3LE, UK

¹² Aix-Marseille Université, CNRS, Laboratoire d'Astrophysique de Marseille, UMR 7326, F-13388 Marseille, France

¹³ Department of Physics, Drexel University, 3141 Chestnut St., Philadelphia, PA 19104, USA

¹⁴ Jet Propulsion Laboratory, California Institute of Technology, 4800 Oak Grove Dr., Mail Stop 169-236, Pasadena, CA 91109, USA

Received 2017 November 21; revised 2018 February 13; accepted 2018 February 14; published 2018 March 20

Abstract

Quasars may have played a key role in limiting the stellar mass of massive galaxies. Identifying those quasars in the process of removing star formation fuel from their hosts is an exciting ongoing challenge in extragalactic astronomy. In this paper, we present X-ray observations of 11 extremely red quasars (ERQs) with $L_{\text{bol}} \sim 10^{47} \text{ erg s}^{-1}$ at $z = 1.5\text{--}3.2$ with evidence for high-velocity ($v \geq 1000 \text{ km s}^{-1}$) [O III] $\lambda 5007$ outflows. X-rays allow us to directly probe circumnuclear obscuration and to measure the instantaneous accretion luminosity. We detect 10 out of 11 ERQs available in targeted and archival data. Using a combination of X-ray spectral fitting and hardness ratios, we find that all of the ERQs show signs of absorption in the X-rays with inferred column densities of $N_{\text{H}} \approx 10^{23} \text{ cm}^{-2}$, including four Compton-thick candidates ($N_{\text{H}} \geq 10^{24} \text{ cm}^{-2}$). We stack the X-ray emission of the seven weakly detected sources, measuring an average column density of $N_{\text{H}} \sim 8 \times 10^{23} \text{ cm}^{-2}$. The absorption-corrected (intrinsic) 2–10 keV X-ray luminosity of the stack is $2.7 \times 10^{45} \text{ erg s}^{-1}$, consistent with X-ray luminosities of type 1 quasars of the same infrared luminosity. Thus, we find that ERQs are a highly obscured, borderline Compton-thick population, and based on optical and infrared data we suggest that these objects are partially hidden by their own equatorial outflows. However, unlike some quasars with known outflows, ERQs do not appear to be intrinsically underluminous in X-rays for their bolometric luminosity. Our observations indicate that low X-rays are not necessary to enable some types of radiatively driven winds.

Key words: galaxies: active – quasars: emission lines – quasars: general – X-rays: galaxies

1. Introduction

It has been known for decades that the majority of present-day galaxies host a supermassive black hole and that accretion onto these black holes is the source of quasar emission. Based on the demographics, polarimetry, and X-ray observations of active nuclei, the unification model (Antonucci 1993) posits that obscured (type 2) quasars are seen through a torus of gas and dust but would otherwise appear as classic unobscured type 1 quasars. At the same time, theoretical models of galaxy formation suggest that obscured quasars may represent an early phase of active black hole growth in which galaxy-wide obscuration leads to the observed properties (Sanders & Mirabel 1996; Canalizo & Stockton 2001; Hopkins et al. 2006). In such models, rapidly accreting obscured black holes can drive winds that clear the galaxy of gas, shutting off star formation in the process now commonly referred to as quasar feedback (Silk & Rees 1998; King 2003). Thus, obscuration can be both due to geometric orientation and/or through evolutionary effects, and the role of these factors in quasar demographics and

galaxy evolution is potentially important and remains poorly understood.

Until recently, direct observations of galactic-scale quasar-driven winds have been scarce, but in the last few years this multi-phase phenomenon has been observed via ultraviolet and X-ray absorption (e.g., Hamann et al. 2001; Moe et al. 2009; Nardini et al. 2015), optical line emission (e.g., Greene et al. 2011; Cano-Díaz et al. 2012; Liu et al. 2013a, 2013b; Harrison et al. 2014), and molecular transitions (e.g., Veilleux et al. 2013; Fiore et al. 2017). One conclusion emerging from this work is that in quasars with powerful galaxy-wide outflows, the classical “narrow-line” region is no longer confined by the potential of the host galaxy and exhibits high-velocity dispersions, with full width at half maximum (FWHM) of the forbidden emission lines (e.g., [O III] $\lambda 5007$) of $\geq 1000 \text{ km s}^{-1}$, and blueshifted asymmetries (Liu et al. 2013b). Such criteria on line shapes and widths are widely used for identifying quasars and galaxies with outflows (e.g., Mullaney et al. 2013; Brusa et al. 2015).

The period around $z \sim 2\text{--}3$ is particularly important in galaxy formation because it marks the peak of both star

formation and quasar activity in the universe (Boyle & Terlevich 1998). However, identifying quasars exhibiting feedback at high redshifts is challenging. In unobscured quasars, the emission signatures of winds are difficult to detect in proximity to the bright central source. Furthermore, feedback may be primarily associated with obscured (and thus optically faint) sources which are yet to be identified in large numbers. Over the last few years, our group has developed a range of approaches to identify obscured and reddened luminous quasars at $1.5 < z < 4$ (Alexandroff et al. 2013; Greene et al. 2014; Ross et al. 2015; Hamann et al. 2017) using data from the Sloan Digital Sky Survey (SDSS; Eisenstein et al. 2011; Dawson et al. 2013) and the *Wide-Field Infrared Survey Explorer* (*WISE*; Wright et al. 2010). In particular, extremely red quasars (ERQs), selected on the basis of high infrared-to-optical ratios and high equivalent width of C IV $\lambda 1549$, often show signs of extreme outflow activity in their [O III] $\lambda\lambda 4959, 5007$ emission, unmatched by any other quasar sample (Zakamska et al. 2016; S. Perrotta et al. 2017, in preparation). Because [O III] likely traces relatively low-density gas, these outflows may be occurring on galaxy-wide scales, and therefore these objects could be manifestations of quasars during a strong feedback episode. In order to better constrain the intrinsic power of these quasars, and assess whether they are capable of driving large-scale outflows, we require a direct measurement of their accretion luminosities. Even in the presence of large absorbing column densities, such as those expected in ERQs, hard X-ray emission arising from the quasar provides a robust determination of the bolometric luminosity.

In this paper, we present the first results of our follow-up X-ray programs of this intriguing population of ERQs. In Section 2, we describe the sample and the observations. In Section 3, we present X-ray spectroscopic analysis. We discuss our results in Section 4 and conclude in Section 5. We adopt an $h = 0.7$, $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$ cosmology and identify optical emission lines using their wavelengths in air following long-standing usage. Objects are identified as SDSS Jhhmss.ss+ddmmss.s in Table 1 and as SDSS Jhhmm+ddmm elsewhere. W1, W2, W3, and W4 refer to the $3.6 \mu\text{m}$, $4.5 \mu\text{m}$, $12 \mu\text{m}$, and $22 \mu\text{m}$ Vega-based magnitudes from the AllWISE data release (Wright et al. 2010; Cutri et al. 2013), and u, g, r, i, z to AB-based magnitudes from the SDSS (Eisenstein et al. 2011; Alam et al. 2015).

2. Observations and Data Reduction

2.1. Parent ERQ Samples and Target Selection

ERQs were first identified based on their extremely high infrared-to-optical flux ratios, with $r_{\text{AB}} - W4_{\text{Vega}} > 14$ mag (Ross et al. 2015) within a sample of spectroscopically confirmed, optically selected quasars (Pâris et al. 2014, 2017) in the Baryon Oscillation Spectroscopic Survey (BOSS; Dawson et al. 2013). In addition to their extreme infrared-to-optical colors, some of these objects showed unusual emission line properties. In particular, they display a mix of “type 1” (unobscured) characteristics—such as broad emission lines (FWHM of C IV $\lambda 1550$ is $> 2000 \text{ km s}^{-1}$)—and “type 2” (obscured) characteristics, such as high equivalent widths of emission lines, indicative of continuum suppression.

Hamann et al. (2017) re-examined and formalized the ERQ selection to include both photometric and spectroscopic selection criteria. Using a color cut of $i_{\text{AB}} - W3_{\text{Vega}} > 9.8$ mag and rest

equivalent width of C IV $\lambda 1550 > 100 \text{ \AA}$, Hamann et al. (2017) identified a sample of 97 ERQs at $2 < z < 3$, hereafter referred to as our parent “ERQ-core” sample (shaded region of Figure 1). Based on further analysis of the SDSS DR12 quasar sample, Hamann et al. (2017) defined a second sample of 235 quasars that exhibited any of the following characteristics in the broadband photometry or spectroscopy: $i_{\text{AB}} - W3_{\text{Vega}} > 9.8$ mag, or EW (C IV) $> 100 \text{ \AA}$, or C IV line profiles that are “wingless” or “boxy” (quantified by the line profile’s kurtosis) with $kt_{80} > 0.33$. Specifically, kt_{80} measures the velocity width of the emission line profile at 80% of its peak height divided by the width at 20%. These objects, which display some, but not all, of the properties of the core ERQs, are hereafter referred to as our parent “ERQ-like” sample (Section 5.7 and Table B1 of Hamann et al. 2017). The distribution and selection of the ERQ samples as compared to the SDSS-DR12 quasars are shown in Figure 1. One source in our sample, SDSS J212951.40-001804.3, was determined by Hamann et al. (2017) to be ERQ-like as it was not detected in the version of the *WISE* source catalog used in their previous analysis. In the most recent *WISE* release this source is detected in W3 with signal-to-noise ratio $S/N \sim 4.7$, and here we use its updated *WISE* photometry, $W3_{\text{Vega}} = 11.407$ mag, which would now reclassify the source as an ERQ-core object. However, for ease of comparison to the Hamann et al. (2017) catalog, we show SDSS J212951.40-001804.3 as ERQ-like in Figure 1.

We are conducting an extensive campaign to obtain multi-wavelength follow-up of the ERQ samples, including radio observations (H.-C. Hwang et al. 2017, in preparation), optical spectropolarimetry (Alexandroff et al. 2017), and extensive near-infrared spectroscopy with the VLT, Gemini and Keck (Zakamska et al. 2016; S. Perrotta et al. 2017, in preparation). The near-infrared spectra probe rest-frame optical wavelengths and cover the key diagnostics $H\beta + [\text{O III}] \lambda\lambda 4959, 5007$ and $H\alpha + [\text{N II}] \lambda\lambda 6548, 6563$. Our follow-up near-infrared spectroscopy of ~ 20 sources reveals that ERQs routinely show kinematically disturbed [O III] emission lines (S. Perrotta et al. 2017, in preparation), inconsistent with gas in a galactic potential and passively photo-ionized by the quasar, FWHM reaching an unprecedented $> 5000 \text{ km s}^{-1}$ in some sources (Zakamska et al. 2016; S. Perrotta et al. 2017, in preparation).

As an example, in Figure 2, we show previously unpublished Gemini GNIRS spectra of two objects from our X-ray sample presented in this paper, where we illustrate our identification of [O III] wind signatures. Details of observations and fitting are provided in Alexandroff et al. (2017). We find large velocity widths of [O III] in both objects, with FWHM in the top 1% of the [O III] widths of the low-redshift obscured quasar population (Zakamska & Greene 2014; Yuan et al. 2016). Three of the remaining four near-infrared spectra mentioned in Table 1 are published in Zakamska et al. (2016) and Alexandroff et al. (2017). A complete analysis of our multi-facility near-infrared spectroscopic campaign will be presented by S. Perrotta et al. (2017, in preparation). While it is not yet clear exactly which properties of ERQs—their luminosities, their colors, or their rest-frame UV line shapes—are most strongly associated with the kinematic activity in [O III], ERQs selected on the basis of colors and C IV equivalent widths show near-100% detection rate of [O III] outflows (S. Perrotta et al. 2017, in preparation).

Table 1
High-redshift Extremely Red Quasars with X-Ray Observations, by Optical Selection and Observation Type

| SDSS ID | α (degree) | δ (degree) | z | Class | [O III] spec? | Outflow? | X-ray Telescope ^a | Inst. | t_{exp} (ks) | d_{OAX} (') | $\nu L_{6\ \mu\text{m}}$ (log erg s ⁻¹) |
|--------------------------|----------------------|----------------------|-------|----------|------------------|----------|---------------------------------|----------|--------------------------|-------------------------|--|
| (1) | (2) | (3) | (4) | (5) | (6) | (7) | (8) | (9) | (10) | (11) | (12) |
| SDSS J000610.67+121501.2 | 1.54446 | 12.25033 | 2.309 | ERQ | × | ? | <i>XMM</i> | PN, M1/2 | 23 | on-axis | 47.04 |
| SDSS J011601.43-050503.9 | 19.00596 | -5.08442 | 3.183 | ERQ-like | × | ? | <i>Chandra</i> | ACIS-S | 71 | on-axis | 47.25 |
| SDSS J022052.13+013711.4 | 35.21721 | 1.61983 | 3.138 | ERQ | × | ? | <i>Chandra</i> | ACIS-S | 70 | on-axis | 47.26 |
| SDSS J082653.42+054247.3 | 126.72258 | 5.71314 | 2.578 | ERQ | ✓ | ✓ | <i>Chandra</i> | ACIS-S | 15 | on-axis | 46.78 |
| SDSS J083200.20+161500.3 | 128.00083 | 16.25008 | 2.431 | ERQ | ✓ | ✓ | <i>Chandra</i> | ACIS-S | 15 | on-axis | 46.72 |
| SDSS J083448.48+015921.1 | 128.70200 | 1.98919 | 2.591 | ERQ | ✓ | ✓ | <i>XMM</i> | PN, M1/2 | 35 | on-axis | 47.03 |
| SDSS J091508.45+561316.0 | 138.78521 | 56.22111 | 2.857 | ERQ | × | ? | <i>Chandra</i> | ACIS-S | 23 | 5.0 | 46.69 |
| SDSS J112124.55+570529.6 | 170.35229 | 57.09156 | 2.383 | ERQ-like | × | ? | <i>Chandra</i> | ACIS-I | 5 | 4.6 | 46.80 |
| SDSS J131047.78+322518.3 | 197.69908 | 32.42175 | 3.009 | ERQ | × | ? | <i>XMM</i> | M1/2 | 50 | 11.4 | 47.12 |
| SDSS J153542.40+090341.1 | 233.92667 | 9.06142 | 1.533 | ERQ | ✓ | ✓ | <i>Chandra</i> | ACIS-S | 15 | on-axis | 46.40 |
| SDSS J165202.64+172852.4 | 253.01100 | 17.48122 | 2.942 | ERQ | ✓ | ✓ | <i>Chandra</i> | ACIS-S | 15 | on-axis | 47.19 |
| SDSS J170047.07+400238.7 | 255.19613 | 40.04408 | 2.903 | ERQ-like | × | ? | <i>XMM</i> | PN, M1/2 | 11 | 1.8 | 46.39 |
| SDSS J212951.40-001804.3 | 322.46417 | -0.30119 | 3.206 | ERQ-like | × | ? | <i>XMM</i> | PN, M1/2 | 34 | 9.3 | <46.56 |
| SDSS J232326.17-010033.1 | 350.85904 | -1.00919 | 2.356 | ERQ | ✓ | ✓ | <i>XMM</i> | PN, M1/2 | ... | on-axis | 46.58 |

Note. (1) Full SDSS-DR13 designation; (2)–(3) J2000 coordinates in degrees; (4) redshift; (5) source class based on the ERQ and ERQ-like definitions outlined in Section 2; (6) does the source have available rest-frame optical spectroscopy covering the [O III] emission line at 5007 Å?; (7) If the source has available [O III] spectroscopy, is the [O III] observed to be strongly asymmetric and determined to host an outflow?; (8)–(9) X-ray telescope and associated instruments for the source; (10) exposure time in kiloseconds for the X-ray observation; (11) off-axis distance in arcminutes for the position of the source on the X-ray detector; (12) rest-frame $6_{\mu\text{m}}$ luminosity derived from *WISE* infrared photometry.

^a All sources observed with *Chandra* are included in the X-ray stacking analysis presented in Section 3.3.

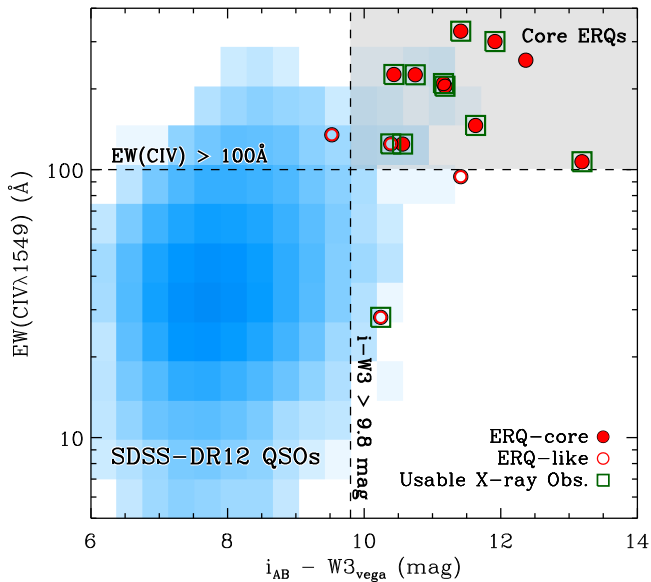


Figure 1. Target selection in the space of optical-to-infrared color and the C IV equivalent width. Shaded contours show the locus of all SDSS-DR12 quasars between $2 < z < 3$ that are detected in the *WISE* W3 band ($S/N > 2$). ERQ-core and ERQ-like objects presented in this paper are shown with filled and open circles, respectively. Those objects covered by X-ray observations that we determined to be of sufficient quality for our purposes are highlighted with green squares. The gray shaded box highlights the region used by Hamann et al. (2017) to describe the ERQ-core sources.

2.2. ERQs with Sensitive X-Ray Observations

In 2014, we started follow-up X-ray observations of $1.5 < z < 3$ ERQs from Ross et al. (2015) and Hamann et al. (2017) with confirmed [O III] outflows. We proposed for three objects to be observed with *XMM-Newton* (PI Alexandroff) in Cycle 14. Two were observed (SDSS J0834+0159; SDSS J2323-0100) by *XMM*. Seven further ERQ targets were proposed for with *Chandra* (PI Zakamska) in Cycle 17. Four were approved and observed. The observed targets are listed in Table 1, and the distribution of targets in the space of infrared-to-optical color and C IV equivalent width—the defining properties of ERQs—is shown in Figure 1. Furthermore, the results of our near-infrared spectroscopy for the six targeted sources are also listed in Table 1. By selection, all six of our observed X-ray targets show [O III] with strong wind signatures, with FWHM of [O III] between $1500\text{--}2800\text{ km}^{-1}$.

In addition to our own targeted observations, we include archival X-ray observations from *XMM* and *Chandra* in our parent ERQ-core and ERQ-like samples. We use NASA’s High Energy Astrophysics Science Archive Research Center (HEASARC) tools to cross-correlate all 97 ERQ-core and 235 ERQ-like quasars against the *xmmaster* catalog (which contains the coordinates of all *XMM* pointings) within $2'$. Outside of this radius, the point-spread function (PSF) becomes highly non-circular, making weak sources difficult to accurately identify. Further, we cross-match the sample against the *XMM* source catalog (*xmmssc*, Rosen et al. 2016) for ERQs at larger off-axis distances. We identified two ERQs (SDSS J0006+1215; SDSS J1700+4002) with on-axis *XMM* observations, and a further two ERQs (SDSS J0220+0137; SDSS J0116-0505) identified within the *xmmssc*. Of the two on-axis *XMM* sources, one was targeted as a comparison object to another sample of reddened quasars (Banerji et al. 2015),

and the other is serendipitously close to the center of an unrelated observation.

To identify any ERQs that overlap with existing *Chandra* observations, instead of the *Chandra* Source Catalog (*cxogsgsrc*, Wang et al. 2016a) whose detection threshold is too high for faint sources such as ours (Goulding et al. 2012), we use the *find_chandra_obsid* tool. Specifically, we include only those observations that cover ERQ positions on Chips I0–3 in ACIS-I mode or Chip S3 in ACIS-S mode, where adverse PSF effects of *Chandra* are reduced. We identified a further four ERQs that met these criteria—two were targeted as hot dust-obscured galaxy candidates (HotDOGs; Section 4.2), and the remaining two are off-axis serendipitous observations.

Our final sample contains 14 ERQ-core or ERQ-like objects with X-ray observations (Table 1), though only 11 ERQs have usable X-ray data (Table 2) as discussed in the next section. Our selection does not include quasars which might be covered by off-axis *XMM* observations if they are not strongly detected in X-rays (i.e., they must be included in the *xmmssc*). Thus, our matching strategy is marginally incomplete and partially biased toward brighter X-ray sources, in that we are not considering the upper limits on X-ray fluxes that might be available for such objects.

2.3. X-Ray Data Reductions

The data for the six ERQ targets observed by *XMM-Newton* were retrieved from the HEASARC database and reduced using standard tools in the *XMM-Newton* SAS software, version 15.0.0. Specifically, we applied the latest calibration files and analyzed the light curves for flaring events for the MOS and PN CCDs from the European Photon Imaging Camera (EPIC). We determined that each of the targets covered by *XMM* observations was subject to some X-ray flaring, typically at the $\sim 10\%$ – 20% level of the total observation exposure time. These periods of flaring were identified and removed from the observations. For SDSS J1700+4002 the entire 11 ks X-ray exposure was subject to high levels of flaring, resulting in no scientifically usable data. We removed this object from our sample and do not discuss it further. Final flare-corrected exposure times are given in Table 2 and are in the range 12–50 ks.

Photon events are extracted from elliptical apertures based upon the size and shape of the PN/MOS PSF at the detector position of the target. As two of the targets were serendipitously detected within the *XMM* observations, they are positioned at large off-axis angles, which significantly distorts the shape of the PSF from that of a circle, which we account for based on *XMM* PSF models. For the two sources with >100 photon counts in the 0.5–10 keV band, we group the counts into bins of 15 counts for χ^2 statistics and construct aspect histograms and response matrices at the positions of the targets in order to derive the final X-ray spectra. Finally, the X-ray emission from SDSS J2323-0100 is heavily contaminated by an unexpectedly bright foreground quasar (angular separation $\sim 15''$). Because of the low spatial resolution of the *XMM* observation, this observation is unusable and we remove this object from further analyses.

Of the eight ERQs observed with *Chandra*, six were observed on-axis with ACIS-S, but one of them (SDSS J0116-0505) is not yet public and is therefore excluded from all analyses. An additional source was serendipitously

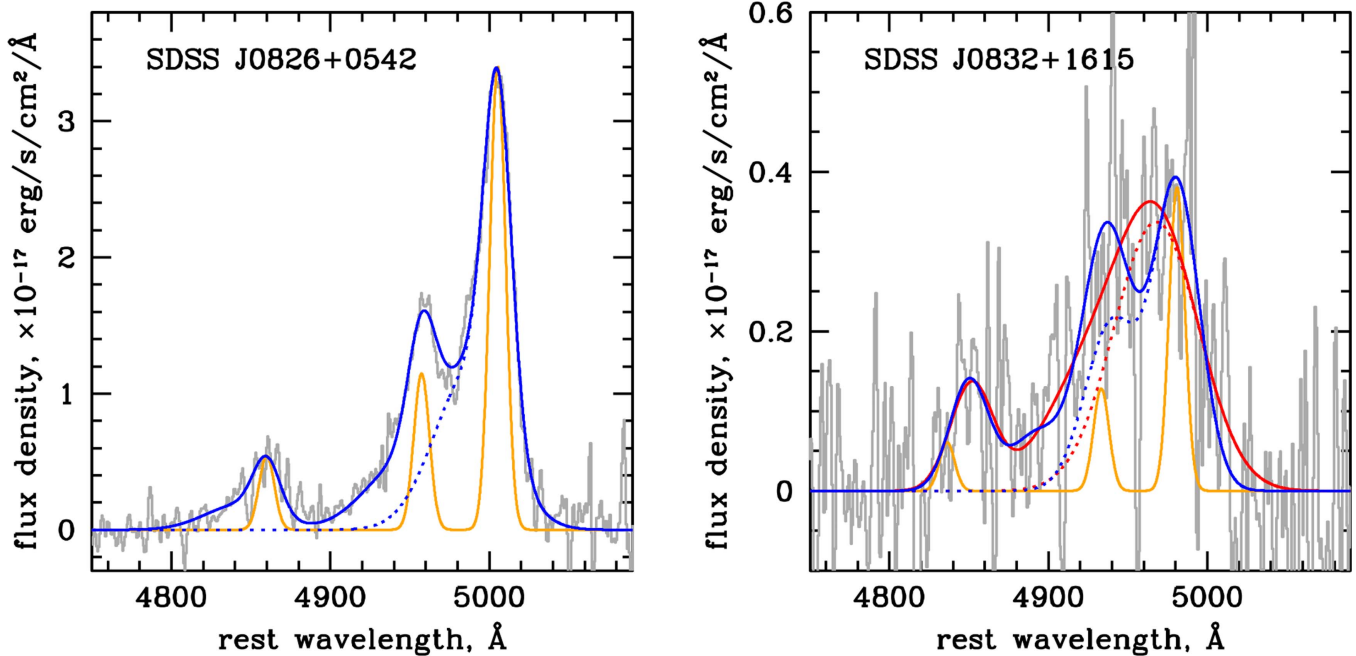


Figure 2. Gemini GNIRS [O III] $\lambda\lambda 4959, 5007 + H\beta$ spectroscopy (PI Alexandroff) of two of the objects presented here. Solid lines show the overall fit (blue for two-Gaussian fits, red for one-Gaussian fits), and dotted lines show the [O III] $\lambda 5007$ contribution separately. In SDSS J0826 + 0542, outflow signatures include a blueshifted broad wing of [O III] and the large velocity width of the entire line ($\text{FWHM}[\text{O III}] \simeq 1600 \text{ km s}^{-1}$). SDSS J0832 + 1615 is much fainter, so the spectrum has a lower S/N, but [O III] $\lambda 4959$ and [O III] $\lambda 5007$ are strongly blended, suggesting velocity $\text{FWHM} \simeq 3800 \text{ km s}^{-1}$ in both the single-Gaussian (red) and the double-Gaussian (blue) fitting attempts. For comparison, the orange line shows [O III] + $H\beta$ profiles with kinematically undisturbed $\text{FWHM} = 700 \text{ km s}^{-1}$.

identified in an ACIS-S observation, ~ 5 arcmin off-axis on the ACIS-S3 chip, and the last *Chandra* source was observed ~ 4.6 arcmin off-axis with ACIS-I. The individual observation identification numbers (ObsIDs) and the X-ray fluxes and counts are provided in Table 2. For each of the *Chandra* observations, we carry out data processing using the *Chandra* X-ray Center software packages available in CIAO v4.8 in conjunction with the latest calibration files (CALDB 4.7.3) applied using `chandra_repro`. Streak events, bad pixels, pixel randomization and cosmic rays are removed with `STATUS = 0` and screened with the typical grad set during the implementation of `acis_process_events`. Flares greater than 3σ above the background are identified and removed using `lc_clean` to create Level-2 events files. Flare-corrected exposure times for the *Chandra* observations are in the range ~ 5 –30 ks. Point source photometry is performed using circular apertures based on the size of a region required to enclose 90% of the PSF (r_{90}) at 0.5, 2 and 5 keV, with background regions defined by annuli with outer radii equal to $5r_{90}$ and inner radii equal to $1.3r_{90}$. Counts are extracted and exposure maps are constructed through the implementation of the `srcflux` tool in CIAO for energy bands 0.3–1, 1–4 and 4–7 keV (see Section 3.1).

3. X-Ray Spectral Analysis

We identify X-ray detections and measure X-ray fluxes and spectral shapes in Section 3.1 and summarize these results in Table 2. In Section 3.2 we discuss some of the sources individually. In Section 3.3 we stack the *Chandra* sources—which are too weakly detected for individual spectral fits—and conduct a spectral analysis of the stack.

3.1. X-Ray Flux and Hardness Measurements

We divide the observed energy range into three bands: soft (0.3–1 keV), middle (1–4 keV), and hard (4–7 keV). At $z \sim 2.5$, this provides us with rest-frame energy ranges of ~ 1 –3.5, ~ 3.5 –14, and ~ 14 –24 keV. All seven sources observed with *Chandra* were detected in at least one energy band and, similarly, three of four sources were detected with *XMM*. Only four of the 11 objects were detected in the soft band, two of each in *Chandra* and *XMM*. We detect all seven of the *Chandra* objects in the middle band, with four of these also detected in the hard band. We detected three (two) out of four sources in the middle (hard) bands with *XMM*. The majority of the detected sources are observed with $S/N > 5$ above the background in the middle band with either *Chandra* or *XMM*. Given the relatively short *Chandra* exposures, the expected background counts in the source apertures in the middle band are typically < 1 count. The object most weakly detected with *Chandra* (SDSS J1535+0903) has three counts at $E \sim 1$ –7 keV. More typically, our sources have 5–15 counts. The individual source counts and fluxes (or 3σ upper limits) are listed in Table 2. For consistency between the telescopes, we have chosen to limit both the *Chandra* and *XMM* data to 7 keV for all of our energy-band analyses. Even though the effective area of *XMM* extends beyond 7 keV, we find that limiting the energy range in this manner does not adversely affect our conclusions given that the two sources observed by *XMM* that do not have the required counts to construct a spectrum are already not detected in the 4–7 keV range, and hence, we are not excluding potential photons at 7–10 keV from our analyses.

It is conventional to quote X-ray luminosities over the rest-frame 2–10 keV range ($L_{2-10 \text{ keV}}$ hereafter). For a power-law X-ray spectrum with a typical slope $\Gamma = 1.8$ ($dN/dE \propto E^{-\Gamma}$; Nandra & Pounds 1994; Reeves & Turner 2000; Page et al.

Table 2
X-Ray Properties of ERQs, by Instrument

| <i>Chandra</i> | | | 0.3–1 keV | | 1–4 keV | | 4–7 keV | | HR ₁ ^a | HR ₂ ^a | $N_{\text{H,HR1}}^{\text{b}}$ | $N_{\text{H,HR2}}^{\text{b}}$ | $L_{\text{X,abs}}$ | $L_{\text{X,int}}$ | N_{H}^{c} |
|----------------|-------|------------------|------------------|--|------------------|--|------------------|--|------------------------------|------------------------------|-------------------------------|-------------------------------|-------------------------|-------------------------|---------------------------|
| ID | ObsID | t_{GTI} | C_{net} | F_{net} | C_{net} | F_{net} | C_{net} | F_{net} | | | (log) | (log) | (10^{44}) | (10^{44}) | (10^{22}) |
| | | (ks) | | (10^{-15}) | | (10^{-15}) | | (10^{-15}) | | | (cm^{-2}) | (cm^{-2}) | (erg s^{-1}) | (erg s^{-1}) | |
| | | | | ($\text{erg s}^{-1} \text{cm}^{-2}$) | | ($\text{erg s}^{-1} \text{cm}^{-2}$) | | ($\text{erg s}^{-1} \text{cm}^{-2}$) | | | | | | | |
| J0826+0542 | 18206 | 14.86 | <6.9 | <6.51 | 14.0 ± 3.7 | $8.76_{-3.34}^{+4.44}$ | 4.9 ± 2.2 | $13.0_{-7.5}^{+12.5}$ | $0.88_{-0.08}^{+0.12}$ | $-0.62_{-0.29}^{+0.18}$ | $23.3_{-0.2}^{+1.7}$ | $23.5_{-4.5}^{+0.3}$ | 5.0 | $9.7_{-4.8}^{+\infty}$ | HR1 |
| J0832+1615 | 18207 | 14.86 | <6.9 | <7.40 | 3.9 ± 2.0 | $1.73_{-1.10}^{+1.93}$ | <5.4 | <10.3 | $0.80_{-0.10}^{+0.16}$ | $-0.26_{-0.23}^{+0.22}$ | $23.1_{-0.4}^{+0.3}$ | $24.0_{-0.3}^{+0.2}$ | 0.9 | $4.6_{-3.7}^{+12.5}$ | HR2 |
| J1535+0903 | 18208 | 14.86 | <5.4 | <5.75 | 1.8 ± 1.4 | $1.08_{-0.91}^{+2.04}$ | <6.9 | <13.2 | $0.47_{-0.25}^{+0.39}$ | $-0.62_{-0.38}^{+0.20}$ | $21.8_{-2.8}^{+1.3}$ | $23.3_{-4.3}^{+0.4}$ | 0.2 | $2.4_{-1.2}^{+9.9}$ | HR1 |
| J1652+1728 | 18205 | 15.31 | <5.4 | <6.25 | 9.9 ± 3.1 | $5.27_{-2.32}^{+3.31}$ | <5.4 | <10.0 | $0.57_{-0.23}^{+0.43}$ | $-0.02_{-0.46}^{+0.50}$ | $22.1_{-3.1}^{+2.9}$ | $23.9_{-0.6}^{+0.4}$ | 3.9 | 15.0_{-11}^{+43} | HR2 |
| J0220+0137 | 18708 | 31.32 | <6.9 | <3.05 | 6.8 ± 2.6 | $1.96_{-1.02}^{+1.57}$ | 2.9 ± 1.7 | $3.27_{-2.36}^{+4.52}$ | $0.65_{-0.17}^{+0.29}$ | $-0.16_{-0.32}^{+0.14}$ | $22.8_{-0.7}^{+0.7}$ | $24.2_{-0.4}^{+0.2}$ | 1.7 | $10.0_{-7.0}^{+17}$ | HR2 |
| J0915+5613 | 04821 | 23.00 | 1.6 ± 1.4 | $1.73_{-1.52}^{+3.76}$ | 35.9 ± 6.0 | $14.1_{-3.5}^{+4.4}$ | 12.5 ± 3.6 | $21.7_{-9.0}^{+12.0}$ | $0.88_{-0.06}^{+0.09}$ | $-0.29_{-0.15}^{+0.14}$ | $23.3_{-0.1}^{+0.3}$ | $24.0_{-0.3}^{+0.1}$ | 10.2 | 45.1_{-25}^{+24} | HR2 |
| J1121+5705 | 06958 | 4.72 | 6.0 ± 2.4 | $10.3_{-5.6}^{+8.7}$ | 12.9 ± 3.6 | $24.9_{-9.8}^{+13.4}$ | 1.9 ± 1.4 | $17.6_{-14.1}^{+31.7}$ | $0.52_{-0.19}^{+0.22}$ | $-0.54_{-0.26}^{+0.19}$ | $22.3_{-3.3}^{+0.8}$ | $23.5_{-4.5}^{+0.3}$ | 11.7 | $16.2_{-7.0}^{+16}$ | HR1 |

| <i>XMM PN</i> | | | 0.3–1 keV | | 1–4 keV | | 4–7 keV | | HR ₁ | HR ₂ | $N_{\text{H,HR1}}$ | $N_{\text{H,HR2}}$ | $L_{\text{X,abs}}$ | $L_{\text{X,int}}$ | $N_{\text{H}}^{\text{?}}$ |
|---------------|-----------|------------------|------------------|--|------------------|--|------------------|--|------------------------|-------------------------|----------------------|----------------------|-------------------------|-------------------------|---------------------------|
| ID | ObsID | t_{GTI} | C_{net} | F_{net} | C_{net} | F_{net} | C_{net} | F_{net} | | | (log) | (log) | (10^{44}) | (10^{44}) | |
| | | (ks) | | (10^{-15}) | | (10^{-15}) | | (10^{-15}) | | | (cm^{-2}) | (cm^{-2}) | (erg s^{-1}) | (erg s^{-1}) | |
| | | | | ($\text{erg s}^{-1} \text{cm}^{-2}$) | | ($\text{erg s}^{-1} \text{cm}^{-2}$) | | ($\text{erg s}^{-1} \text{cm}^{-2}$) | | | | | | | |
| J0834+0159 | 762260101 | 22.8 | <22.4 | <1.08 | <31.5 | <5.35 | <22.7 | <10.4 | $0.06_{-0.51}^{+0.72}$ | $0.02_{-0.59}^{+0.65}$ | $22.1_{-3.1}^{+1.0}$ | $24.1_{-0.5}^{+0.5}$ | <3.1 | <120 | HR2 |
| J1310+3225 | 020540401 | chip-gap | ... | ... | ... | ... | ... | ... | $0.61_{-0.20}^{+0.28}$ | $-0.80_{-0.20}^{+0.10}$ | $23.0_{-0.3}^{+0.4}$ | $20.0_{-1.0}^{+1.9}$ | 2.6 | $4.0_{-1.6}^{+2.7}$ | HR1 |
| J2129-0018 | 729160501 | 31.4 | 29.3 ± 7.6 | $1.02_{-0.27}^{+0.26}$ | 56.3 ± 11.1 | $6.94_{-1.37}^{+1.37}$ | <19.9 | <6.59 | $0.26_{-0.20}^{+0.22}$ | $-0.78_{-0.22}^{+0.10}$ | $22.6_{-0.4}^{+0.3}$ | $20.0_{-1.0}^{+2.1}$ | 7.1 | $9.1_{-2.7}^{+3.6}$ | Spec |
| J0006+1215 | 763780701 | 12.7 | 37.5 ± 7.9 | $3.25_{-0.69}^{+0.68}$ | 136.0 ± 13.6 | $41.4_{-4.1}^{+4.2}$ | 33.9 ± 7.5 | $27.8_{-6.2}^{+6.1}$ | $0.55_{-0.07}^{+0.08}$ | $-0.49_{-0.09}^{+0.09}$ | $22.7_{-0.1}^{+0.1}$ | $23.5_{-0.4}^{+0.1}$ | 17.0 | $27.1_{-6.2}^{+4.5}$ | Spec |

| <i>XMM MOS1-2</i> | | | M1 0.3–1 keV | | M1 1–4 keV | | M1 4–7 keV | | M2 0.3–1 keV | | M2 1–4 keV | | M2 4–7 keV | |
|-------------------|-----------|------------------|------------------|--|------------------|--|------------------|--|------------------|--|------------------|--|------------------|--|
| ID | ObsID | t_{GTI} | C_{net} | F_{net} | C_{net} | F_{net} | C_{net} | F_{net} | C_{net} | F_{net} | C_{net} | F_{net} | C_{net} | F_{net} |
| | | (ks) | | (10^{-15}) | | (10^{-15}) | | (10^{-15}) | | (10^{-15}) | | (10^{-15}) | | (10^{-15}) |
| | | | | ($\text{erg s}^{-1} \text{cm}^{-2}$) | | ($\text{erg s}^{-1} \text{cm}^{-2}$) | | ($\text{erg s}^{-1} \text{cm}^{-2}$) | | ($\text{erg s}^{-1} \text{cm}^{-2}$) | | ($\text{erg s}^{-1} \text{cm}^{-2}$) | | ($\text{erg s}^{-1} \text{cm}^{-2}$) |
| J0834+0159 | 762260101 | 27.8 | <7.4 | <1.38 | <8.0 | <2.91 | <4.3 | <4.90 | <6.1 | <1.13 | <6.84 | <2.47 | <3.1 | <3.53 |
| J1310+3225 | 020540401 | 47.9 | <9.8 | <1.06 | 15.2 ± 5.4 | $3.21_{-1.14}^{+1.13}$ | <7.6 | <5.02 | <8.7 | <0.94 | 14.9 ± 5.4 | $3.14_{-1.14}^{+1.14}$ | <8.1 | <5.36 |
| J2129-0018 | 729160501 | 34.3 | 10.9 ± 4.1 | $1.64_{-0.62}^{+0.60}$ | 26.2 ± 6.4 | $7.71_{-1.90}^{+1.86}$ | <9.5 | <8.78 | 11.2 ± 4.1 | $1.69_{-0.64}^{+0.60}$ | 30.2 ± 6.4 | $8.89_{-1.92}^{+1.80}$ | 8.2 ± 3.6 | $7.58_{-3.33}^{+3.22}$ |
| J0006+1215 | 763780701 | 17.2 | 11.4 ± 4.0 | $3.42_{-1.22}^{+1.20}$ | 62.4 ± 8.6 | $38.6_{-7.0}^{+8.1}$ | 10.2 ± 3.9 | $18.8_{-7.2}^{+7.2}$ | 9.6 ± 3.5 | $2.89_{-1.10}^{+1.02}$ | 72.4 ± 8.9 | $42.5_{-5.3}^{+5.2}$ | 8.3 ± 3.3 | $15.3_{-6.1}^{+6.1}$ |

Notes.^a Hardness ratios as defined in Section 3.^b Column density inferred from hardness ratio diagnostics.^c The diagnostic (HR₁; HR₂; X-ray spectra) used to determine the final value of N_{H} adopted in order to calculate the intrinsic X-ray luminosity ($L_{\text{X, int}}$).

2005; Piconcelli et al. 2005), and in the absence of obscuration, the ratio of the 3.5–14 keV to the 2–10 keV luminosities is close to unity, making our middle band well-suited for the determination of X-ray luminosity. As the cross-section for photo-electric absorption steeply declines with energy, soft X-ray photons are absorbed by the intervening material a lot more effectively than the hard X-ray photons. As a result, for a Compton-thick absorber with $N_{\text{H}} \sim 3 \times 10^{24} \text{ cm}^{-2}$ the ratio of the apparent 3.5–14 keV to the 2–10 keV luminosities is ~ 10 . We are interested in the intrinsic absorption-corrected X-ray luminosities, so the redshift of the sample works in our favor in moving the relatively less absorbed parts of the spectrum into the observable energy range. Nonetheless, to determine the intrinsic luminosities we need measurements of absorption and calculations of absorption corrections.

To gauge whether absorption is important in our sample we first look at the flux ratios. For a power-law spectrum with index $\Gamma \neq 2$, the flux ratio between two bands from E_1 – E_2 and from E_2 – E_3 is $F_{E_2-E_3}/F_{E_1-E_2} = (E_3^{2-\Gamma} - E_2^{2-\Gamma}) / (E_2^{2-\Gamma} - E_1^{2-\Gamma})$. For standard values of $\Gamma = 1.65$ – 1.95 the expected hard-to-middle band flux ratio is $F_{4-7 \text{ keV}}/F_{1-4 \text{ keV}} = 0.42$ – 0.56 , where a higher ratio would indicate lower values of Γ and harder spectra. The six detected objects have much harder spectra, with a median $F_{4-7 \text{ keV}}/F_{1-4 \text{ keV}} \simeq 1.2$, and the four remaining objects with hard-band upper limits consistent with this value. Only one source—SDSS J0006+1215—with the hard-to-middle flux ratio of 0.36 is marginally consistent with a power-law spectrum with $\Gamma \geq 1.95$.

The most natural explanation for the hardness of the observed spectra is photoelectric absorption, strong enough to suppress rest-frame 3.5–14 keV flux. Measurements of the photoelectric absorption due to intervening gas typically require detailed modeling of the observed X-ray spectrum. However, the majority of the ERQs are only weakly detected in the X-rays and lack the necessary counts to perform such an analysis. In the absence of detailed X-ray spectroscopy, we can estimate the obscuring column density (N_{H}) through the use of X-ray hardness ratios (HRs). These compare the observed photon counts between two bands, with the harder (higher-energy) photons being less affected by photoelectric absorption than the softer photons. Here we use two sets of hardness ratios derived from the 0.3–1 keV and 1–4 keV bands and the 1–4 keV and 4–7 keV bands, defined as:

$$\text{HR}_1 = \frac{C_{1-4 \text{ keV}} - C_{0.3-1 \text{ keV}}}{C_{1-4 \text{ keV}} + C_{0.3-1 \text{ keV}}}, \quad (1)$$

$$\text{HR}_2 = \frac{C_{4-7 \text{ keV}} - C_{1-4 \text{ keV}}}{C_{4-7 \text{ keV}} + C_{1-4 \text{ keV}}}. \quad (2)$$

Defined in this way for $z = 2.5$ targets, these values are roughly equivalent to the standard HR bands used for analysis of local systems with *Chandra*/*XMM* in the case of HR_1 or *NuSTAR*/*Suzaku* for HR_2 .

We use the Bayesian Estimation of Hardness Ratios (BEHR) software package (Park et al. 2006) to measure the HRs and estimate their uncertainties for all of the ERQs in our sample. BEHR computes posterior draws for the HRs assuming each detected photon (source or background) is an independent Poisson random variable. It is particularly useful in the low-count Poisson regime as it is capable of producing posterior draws irrespective of whether the source is formally detected in both energy bands. As is recommended with BEHR when

computing HRs, we assume uniform priors and employ a Gaussian quadrature algorithm to compute the Markov chain Monte Carlo integrals with the default 1000 bins. In Figure 3 we provide the HRs measured from the median of the posterior draws with the quoted uncertainties derived from the 17th and 83rd percentiles of the posteriors.

To approximate the absorbing gas column density we use the Portable, Interactive Multi-Mission Simulator (PIMMS) to simulate a power-law spectrum with $\Gamma = 1.8$ subject to a set of column densities $\log N_{\text{H}} = \{20, 22, 23, 23.5, 24, 24.5\}$ at the rest-frame of the target; these spectra are then folded through the response curves of *Chandra* ACIS-S and *XMM*-PN. The minimum column of $N_{\text{H}} = 10^{20} \text{ cm}^{-2}$ is used to represent foreground galactic absorption, which we include for all of the simulations. The typically low number of detected counts for the ERQ sample considered here necessitate the use of a simple absorbed power-law model over more complex models such as those that can be constructed using MYTORUS (Yaqoob 2012). Using PIMMS we track the expected count-rate as a function of redshift, and use this to predict HR_1 and HR_2 for *Chandra* and *XMM* observations. A particular advantage to harnessing both HR_1 and HR_2 is that they are sensitive to different levels of incident absorption. HR_1 is more sensitive to columns below a few $\times 10^{23} \text{ cm}^{-2}$, while HR_2 is more reliable toward the identification of heavily obscured and Compton-thick columns in the range $(0.3\text{--}5) \times 10^{24} \text{ cm}^{-2}$.

In Figure 3, we provide our BEHR-produced HR_1 and HR_2 measurements for our ERQ sample as a function of redshift and theoretical tracks of predicted constant N_{H} as a function of redshift produced from our modeling with PIMMS for both the *XMM* EPIC-PN and *Chandra* ACIS-S detectors. The significant offset between the predicted HR_1 values for PN and ACIS-S at low column densities is due to the lower quantum efficiency at low energies of the *Chandra* ACIS detectors associated with the deposition of contaminating materials onto the detectors. This contamination is a function of mission time, and given that the majority of the *Chandra* observations were performed in Cycle 16, these observations are relatively insensitive to photons with $E < 1 \text{ keV}$, significantly reducing the dynamic range of HR_1 for the *Chandra* observations. For the one source which fell on a PN chip-gap, we used the counts observed in MOS1/2, and for consistency with the other ERQs observed in *XMM*-PN, we applied a model-dependent correction of -0.25 to HR_1 for this source to convert from MOS to PN. No correction is required to convert HR_2 between MOS and PN for the observed low HR_2 .

When comparing the measured HR_1 values to those predicted from the models, we show that all of the detected ERQs are consistent with being moderately to heavily obscured, with $N_{\text{H}} \gtrsim 3 \times 10^{22} \text{ cm}^{-2}$. Even at $z \sim 3$, HR_1 is insensitive to changes in the observed spectral energy distribution (SEDs) caused by high column densities. Hence, for the five ERQs with HR_1 measurements consistent with $N_{\text{H}} \gtrsim 3 \times 10^{23} \text{ cm}^{-2}$, we can harness additional information from the HR_2 , which is capable of observing steep rises in the SED above rest-frame $\sim 10 \text{ keV}$ which is typical of Compton-thick column densities. Indeed, from the combined N_{H} predictions from HR_1 and HR_2 , we find evidence that four of these five ERQs are potentially Compton-thick with column densities of $N_{\text{H}} \approx 10^{24} \text{ cm}^{-2}$. Hence, based on the population of ERQs studied here, we find evidence that ERQs presenting strong outflows are also accompanied by significant gas

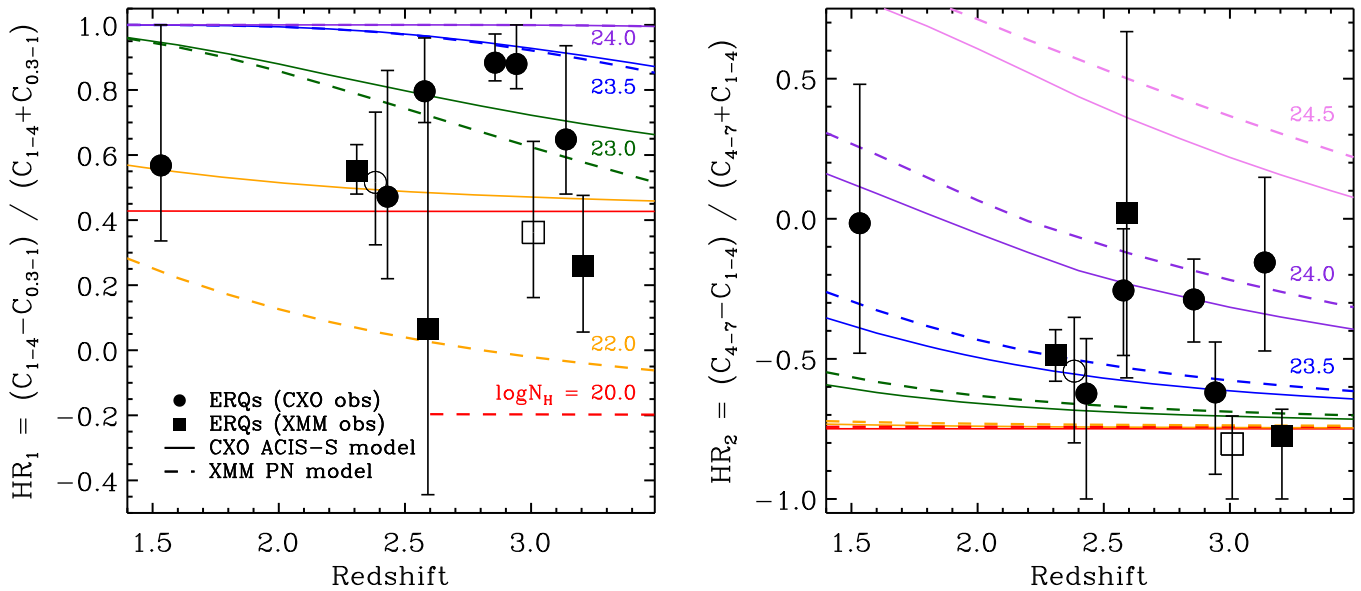


Figure 3. Observed photon hardness ratios vs. redshift for our sample of ERQ sources with $HR_1 = (C_{1-4\text{ keV}} - C_{0.3-1\text{ keV}}) / (C_{1-4\text{ keV}} + C_{0.3-1\text{ keV}})$ (left) and $HR_2 = (C_{4-7\text{ keV}} - C_{1-4\text{ keV}}) / (C_{4-7\text{ keV}} + C_{1-4\text{ keV}})$ (right). HRs are computed using the Bayesian Estimation of Hardness Ratios (BEHR) software package. Uncertainties are derived from 67% of the posterior samples. Filled (open) circles represent sources observed with *Chandra* ACIS-S (ACIS-I) and filled (open) squares are those ERQs observed with *XMM*-PN (MOS1/2). Model HRs obtained for an absorbed power-law spectrum with slope $\Gamma = 1.8$ blocked by a range of redshifted column densities (N_H) are shown for *XMM*-PN (dashed lines) and *Chandra* ACIS-S (solid lines) detectors.

column densities obscuring their central nuclei. Owing mainly to the relatively large 1σ uncertainties (typically ~ 0.8 dex), the N_H values predicted by HR_1 and HR_2 tend to be consistent. However, we note that there are four sources that are estimated to be strongly obscured with $N_H \sim 10^{24} \text{ cm}^{-2}$ based on HR_2 , but are predicted to have significantly lower N_H based on HR_1 , and are each inconsistent at the $1-2\sigma$ level. As we show in the following sections, based on more sophisticated stacking analyses, this could in part be due to an optically thin scattered component that is contributing to the emission at soft-energies (rest-frame $E < 4 \text{ keV}$) that is not currently being included in our simple absorbed power-law model for converting HR to N_H .

Given the uncertainty in the intrinsic power-law spectrum and the measurement uncertainty, clearly the column density estimates from the HRs, while highly suggestive of strong obscuration, are uncertain by a factor of a few. We directly demonstrate this uncertainty in the N_H estimates in the next section through comparison of the HR predictions and the N_H measurements from spectral fits for two of our ERQs. This uncertainty in N_H further implies that the intrinsic (obscuration-corrected) luminosities of our sources are very uncertain, as the 2–10 keV rest-frame luminosity correction depends very sensitively on the value of column density in this regime. For example, as N_H ranges from 0.8 to $1.6 \times 10^{24} \text{ cm}^{-2}$, the luminosity correction factor ranges between 14 and 87. With the exception of SDSS J0006+1215 and SDSS J2129–0018, which have sufficient X-ray counts to perform a spectral fit (Section 3.2), we use the measured HRs to provide an estimate of the absorbing column in each ERQ considered here in order to predict the intrinsic X-ray luminosity at 2–10 keV. As outlined above, the HR_2 is not sensitive to N_H values below $\approx 10^{23} \text{ cm}^{-2}$. Hence, for those ERQs that have an HR_2 measurement that is consistent (at the 1σ level) with no absorption (i.e., $HR_2 \lesssim -0.75$), we choose to use HR_1 to estimate N_H as these sources likely have lower N_H values, otherwise, we use HR_2 . In Table 2 we provide the predicted N_H values for the ERQ sample along with the adopted indicator

(i.e., X-ray spectra; HR_1 ; HR_2). We use these N_H estimates to predict intrinsic L_X with their associated associates computed by combining in quadrature the 1σ X-ray flux and HR uncertainties.

3.2. Comments on Individual Sources

Before looking into the average X-ray properties of our sources in Section 3.3, here we discuss sources that are apparent outliers within our sample. Specifically, SDSS J0006+1215 and SDSS J2129–0018 are the two most strongly detected objects with sufficient counts to enable spectral analysis. SDSS J0834+0159 is the only non-detected source. Finally, SDSS J1535+0903 is a redshift outlier with a unique optical spectrum.

In Figure 4 we show the X-ray spectra of two sources, SDSS J0006+1215 and SDSS J2129–0018, which have >100 detected counts in their *XMM* data, sufficient to perform a spectral fit. For both sources, we fit a power-law model combined with foreground Galactic absorption determined from Stark et al. (1992) and a rest-frame absorber which is intrinsic to the source. The unbinned data from the PN and MOS detectors are fit simultaneously using Cash (1979) statistics, which allows spectral parameter estimation using the maximum likelihood method, especially well suited for low count rates. For SDSS J0006+1215 and SDSS J2129–0018, we find that an absorbed power-law produces a reasonable fit with C-statistic ~ 181.5 and 134.9 for 208 and 159 degrees of freedom, respectively. Hence, based on the available data there is no statistical evidence for requiring a more complex model. The best-fit model parameters for SDSS J0006+1215 are $\Gamma \sim 1.74 \pm 0.07$ and $N_H \sim (5.8 \pm 0.7) \times 10^{22} \text{ cm}^{-2}$. This is consistent with the value of N_H of $\sim (6 \pm 2) \times 10^{22} \text{ cm}^{-2}$ predicted from using HR_1 . However, the HR_2 measurement would suggest a value of N_H that is a factor $\sim 3-5$ greater than that measured from the spectroscopy. For SDSS J2129–0018, the best-fit spectral model yields a similar absorbing column of

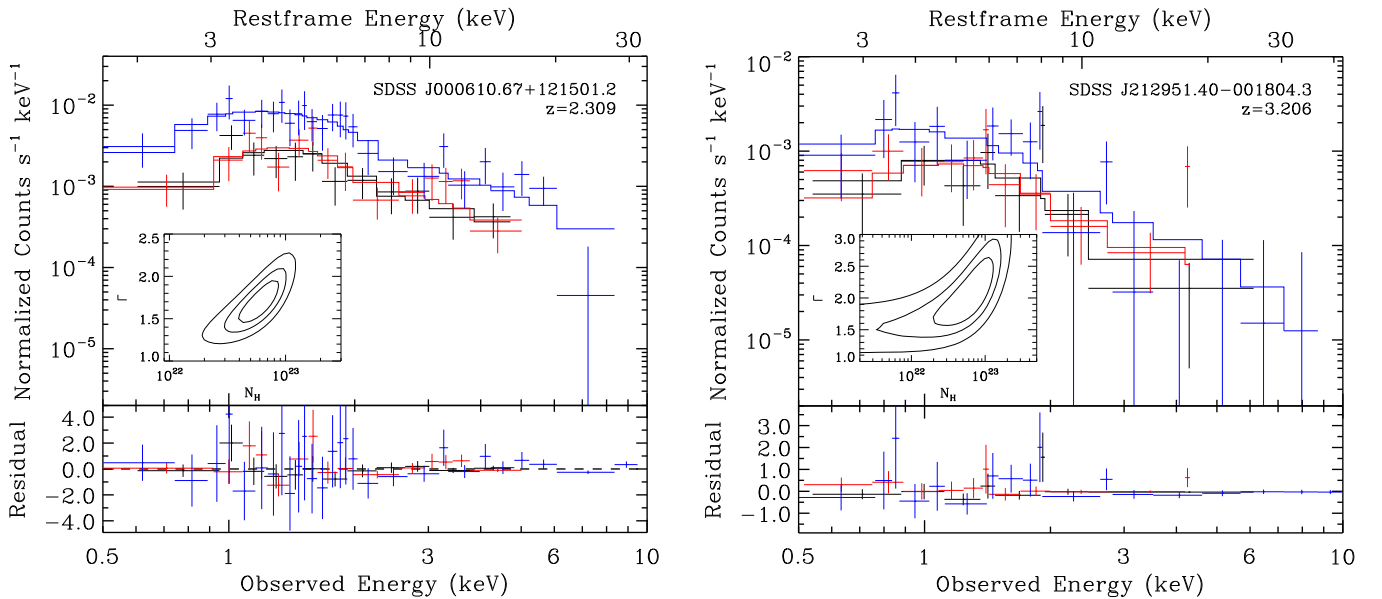


Figure 4. Top panels: *XMM-Newton* spectra for SDSS J0006+1215 (left) and SDSS J2129-0018 (right). Spectra are binned to have at least 15 counts in each bin. EPIC-PN and MOS1/2 are shown with blue, black, and red symbols, respectively. Insets provide the 1, 2, and 3 σ contours for the N_{H} and spectral slope (Γ) parameters produced from a simple rest-frame absorbed power-law model that has been fit to the data. Lower panels provide the residual between the best-fit model and the data.

$N_{\text{H}} \sim (6.1 \pm 0.6) \times 10^{22} \text{ cm}^{-2}$, again with a marginally flatter spectral slope of $\Gamma \sim 1.7 \pm 0.1$ than the one we had assumed for the HRs. In this case, the spectroscopic N_{H} value is consistent with the one resulting from our HR analysis. Based on these two sources, we find that the uncertainty in the HR method is large, and N_{H} estimates from the HRs should only be taken as suggestive, not conclusive (especially given the simplicity of the spectral model used in hardness ratio simulations).

One source SDSS J0834+0159 (with a high-velocity [O III] outflow discussed by Zakamska et al. 2016) is not detected in any of the X-ray bands. Non-detections can arise because the source is intrinsically X-ray weak, because it is obscured, or both. However, we cannot distinguish between these scenarios on the basis of the existing data.

SDSS J1535+0903 is at a substantially lower redshift ($z \sim 1.5$) compared with $z \sim 2.3\text{--}3.2$ characteristic of the remainder of the ERQs considered here. This object was recognized as having an extremely unusual optical spectrum (Ross et al. 2015) which was analyzed by Wang et al. (2016b). The optical spectrum is dominated by high equivalent width Fe II emission which Wang et al. (2016b) interpret as being resonantly scattered by extended outflows that are viewed nearly edge-on. The [O III] profile (S. Perrotta et al. 2017, in preparation) shows a strong blueshifted outflow component. This source has weak apparent X-ray emission and is only marginally detected ($S/N \sim 2$) at $E \sim 1\text{--}7$ keV which, as in the case of SDSS J0834+0159, is suggestive of intrinsic X-ray weakness and/or heavy obscuration.

3.3. Stacking the Weakly Detected Sources

All seven ERQs that were observed with *Chandra* are only weakly detected in their *Chandra* X-ray observations, and their low count rates preclude us from analyzing their spectra individually. Instead, we stack these seven sources to determine their average X-ray properties. We use the SDSS spectroscopic redshifts for each of the ERQs to reproject the source aperture

photons to the rest-frame energies of the individual sources. We further reproject the individual energy redistribution matrices to the source rest-frames and interpolate onto a common binning. The source aperture photons are then summed across the observations before being combined with the response matrices within `xspec`. We do not include the ERQs covered by *XMM* observations into the stacked spectrum given the large background differences between *XMM* and *Chandra* and the strongly differing responses between the instruments. Furthermore, the inclusion of the two ERQs that are strongly detected in *XMM* would dominate the stacked spectrum, whereas the purpose of stacking is to determine the average properties of the weakly detected population. The stacked X-ray spectrum contains 138 counts, and is shown in Figure 5.

We begin by modeling the stacked X-ray spectrum using a simple power-law subject to local galactic extinction (assumed to be $N_{\text{H}} = 10^{20} \text{ cm}^{-2}$), finding an extremely flat spectral slope of $\Gamma \sim 0.7 \pm 0.2$ with $C_{\text{stat}} = 78.8$ and 118 d.o.f. Such flat apparent slopes indicate that the observed spectrum is much harder than the intrinsic power-law spectra of unobscured active nuclei (Nandra & Pounds 1994) and are commonly taken as indicators of absorption (Alexander et al. 2001). Fixing the slope to a more typical value of $\Gamma = 1.9$, we use the stacked spectrum to find evidence for intrinsic X-ray absorption of $N_{\text{H}} \sim 1.1 \pm 0.2 \times 10^{23} \text{ cm}^{-2}$, albeit with a slightly reduced significance of $C_{\text{stat}} = 88.1$ and 117 degrees of freedom. The stacked spectrum in Figure 5 displays a clear inversion of the X-ray spectrum around rest-frame 4 keV, with a rising continuum out to $E \sim 20$ keV, which is also a signature of absorption. Taken together these spectral measurements and features provide evidence that the stacked sources are drawn from a moderately to heavily obscured quasar population.

As the data appear to be consistent with significant absorption toward the quasars in our sample, we fit the stacked spectrum with a more physically motivated model using the pre-calculated `MYTorus` tables (Yaqoob 2012), which

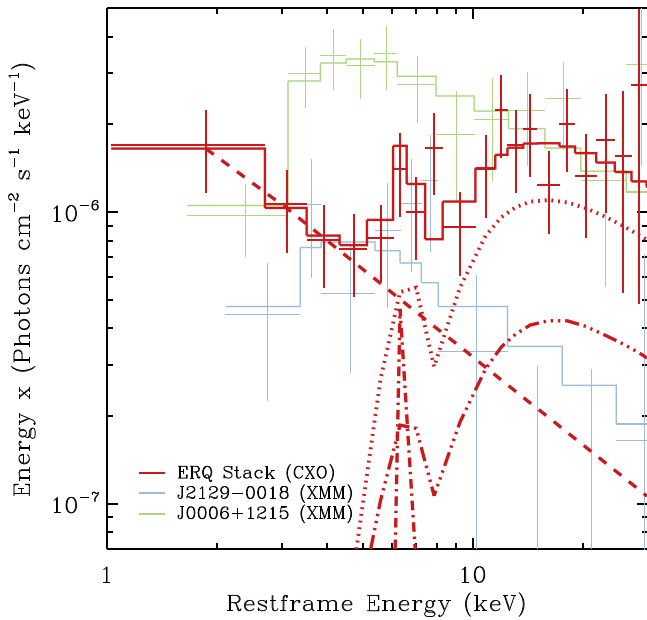


Figure 5. Stacked X-ray spectrum of the ERQ and ERQ-like objects that were observed with *Chandra* (red crosses). The solid red histogram shows the best fit MYTORUS model composed of a zeroth-order power-law continuum (dotted red line) that is additionally subject to Compton-scattering (dot-dash-dot red line), fluorescent line emission (dash-dot red), and an optically thin scattered continuum (dashed red line). For comparison purposes, the rest-frame XMM-PN X-ray spectra of the two sources shown in Figure 4 are provided in light-blue and light-green.

simulate the reprocessed X-ray emission from a toroidal structure in the heavily obscured and Compton-thick regime. Specifically, we use a MYTORUS model consisting of a transmitted zeroth-order power-law continuum, a Compton-scattered continuum, and emission-line fluorescence due to FeK at $E \sim 6.4$ keV. The advantage of this model is that it self-consistently includes the Compton-scattered component for heavily absorbed sources (Lansbury et al. 2014). In the model, we assume an inclination angle of the torus of 85° , a power-law slope and its normalization, and the column densities for all three components are tied together. The relative normalizations of the Compton-scattered and line-emission components are additionally tied together. Based on our HR- N_H analysis in the previous section, there are several weakly detected sources that have apparently contradicting N_H estimates from HR1 and HR2 due to significant excess emission at soft energies. This soft excess is also observed in the stacked spectrum at $E < 4$ keV, and could arise due to a scattered continuum, which has been previously observed in similar populations of red quasars (LaMassa et al. 2016; Glikman et al. 2017). To account for the observed soft X-ray emission at $E < 4$ keV, we additionally include an optically thin scattered quasar continuum into our spectral model.

The stacked spectrum is well characterized by our physically motivated model with a best-fit column density for the stacked spectrum of $N_H \sim (9.1 \pm 3.8) \times 10^{23} \text{ cm}^{-2}$ and $C_{\text{stat}} = 78.5$ for 113 degrees of freedom. Therefore, we find that on average the weakly detected ERQs are significantly obscured in the X-ray, and approaching the Compton-thick threshold ($N_H \geq 10^{24} \text{ cm}^{-2}$). The intrinsic power-law is not particularly well-constrained with $\Gamma \simeq 1.55 \pm 0.49$, which is within the standard range of power-law slopes measured in active nuclei. However, for this particular value of Γ , we find an unusually high and unconstrained

scattering fraction of $\sim 19\%$, which we suggest is unlikely. Fixing the spectral slope to $\Gamma = 1.9$, we find a consistent column density to before with $N_H \sim (1.1 \pm 0.3) \times 10^{24} \text{ cm}^{-2}$ but with a more realistic, though still relatively high, scattering fraction of $5.9 \pm 2.5\%$. The absorbed flux (at rest-frame 2–10 keV) is $8 \times 10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2}$, and the absorption-corrected flux is $4.1 \times 10^{-14} \text{ erg s}^{-1} \text{ cm}^{-2}$. At the median redshift of the ERQs observed with *Chandra* ($z \sim 2.7$), this corresponds to an absorption-corrected rest-frame luminosity of $L_{2-10 \text{ keV}} \sim 2.6 \times 10^{45} \text{ erg s}^{-1}$.

To validate our stacking analysis, we conduct a simultaneous spectral fit directly to all (*Chandra* and XMM) data using Cash (1979) statistics. We extract the photons for each ERQ using the same point-source apertures, and treat them in XSPEC as though they were individual source spectra. We then fit the spectra simultaneously with the same combined spectral model, redshifted to the rest-frame of each individual object. We start with a simple absorbed power-law model, finding an extremely flat slope $\Gamma = 0.6 \pm 0.2$ with an unconstrained column density $N_H \sim (0.05 \pm 1.3) \times 10^{22} \text{ cm}^{-2}$. This is consistent with that found for a simple absorbed power-law fit to our stacked spectrum. As found previously, such a best-fit model is non-physical, and suggestive of a need for a more complex model involving reflection components. We used the same physically motivated MYTORUS model as above, finding best-fit values for the spectral slope and column density of $\Gamma \sim 1.45 \pm 0.51$ and $N_H = (5.5 \pm 2.5) \times 10^{23} \text{ cm}^{-2}$, which are fully consistent with the fit to the stacked X-ray spectrum presented in Figure 5.

4. Discussion

4.1. X-Ray Properties of ERQs

Because of high obscuration, it is possible that none of the observed fluxes of ERQs provides a reliable measure of the intrinsic bolometric accretion luminosity. For ease of comparison with previous work, here we use νL_ν at rest-frame $6 \mu\text{m}$ (obtained by power-law interpolating between the observed WISE fluxes) as a measure of quasar luminosity, listed in Table 1. While this is the most isotropic measure of accretion luminosity currently available to us, theoretical models of quasar obscuration predict that mid-infrared luminosity is not isotropic and can be strongly suppressed in type 2 objects (Pier & Krolik 1992). At low redshifts, type 2 quasars have noticeably redder infrared colors than type 1s (Liu et al. 2013b): at a fixed $12 \mu\text{m}$ luminosity, 5 μm luminosities are a factor of 2–3 lower in type 2s than in type 1s. In type 1 quasars, the bolometric luminosity can be estimated by multiplying $\nu L_\nu[6 \mu\text{m}]$ by a factor of ~ 7 –9 (Richards et al. 2006). In our objects, we use the same estimate, but with a caveat that it could be a lower limit on the actual bolometric luminosity if the $6 \mu\text{m}$ luminosity of ERQs is suppressed by optical depth effects.

In Figure 6, we show the infrared luminosity versus both the observed and estimated intrinsic (absorption-corrected) X-ray luminosities of ERQs in comparison with several other studies. As has been demonstrated by multiple studies, type 1 (unobscured) quasars do not follow the near-linear relationship between infrared and intrinsic X-ray luminosities characteristic of lower luminosity Seyfert galaxies ($\nu L_\nu[6 \mu\text{m}] < 10^{45} \text{ erg s}^{-1}$; Gandhi et al. 2009). Extremely luminous type 1 quasars with $\nu L_\nu[6 \mu\text{m}] = 10^{47} \text{ erg s}^{-1}$ show no signs of intrinsic

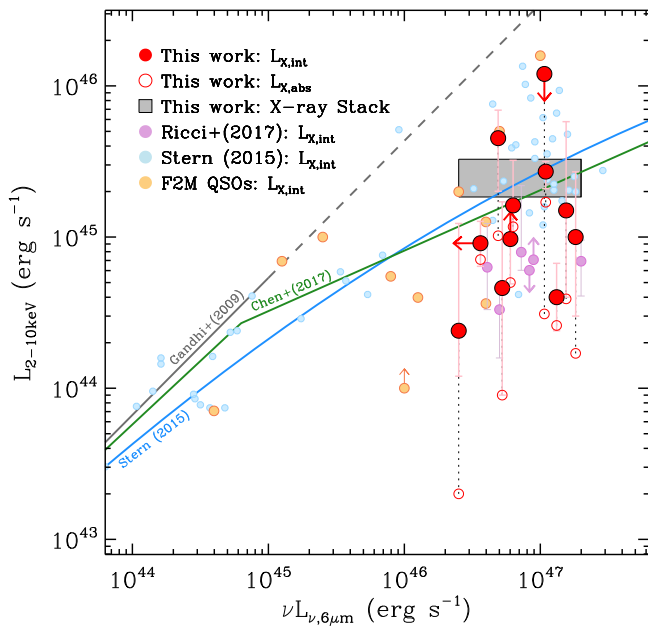


Figure 6. Relationship between infrared and intrinsic (absorption corrected) 2–10 keV X-ray luminosities for individual ERQs in this work (red circles). Absorbed (observed) X-ray luminosities for the ERQs are shown with open symbols. Additionally, we show the ERQ X-ray stack (gray shaded region) and further compare with individual type 1 quasars from Stern (2015), HotDOGs from Ricci et al. (2017), and a sample of lower redshift ($z \sim 0.1\text{--}0.7$) reddened FIRST-2MASS quasars presented in Urrutia et al. (2005), LaMassa et al. (2016), and Glikman et al. (2017), shown with blue, purple, and orange circles, respectively. The near-linear relation from Gandhi et al. (2009) was originally derived for sources with $\nu L_{\nu}[12.3 \mu\text{m}] < 10^{45} \text{ erg s}^{-1}$ and is shown here corrected to $6 \mu\text{m}$ following Chen et al. (2017) with an extrapolation to higher luminosities shown with a dotted line. Subsequent work demonstrated that the relationship is sub-linear for extremely luminous quasars (blue for Stern 2015, green for Chen et al. 2017). The intrinsic (absorption-corrected) X-ray luminosities of individual ERQs and of the stacked detection are consistent with the relationships derived for type 1 quasars.

X-ray absorption. However, they lie an order of magnitude below the extrapolation of the linear lower luminosity relationship (Stern 2015; Martocchia et al. 2017), calling for revised sub-linear infrared-to-X-ray relationships by Stern (2015) and Chen et al. (2017), producing lower L_X at fixed $L_{6 \mu\text{m}}$. Red quasars of comparable infrared luminosities (Banerji et al. 2012, 2015; Feruglio et al. 2014; LaMassa et al. 2016; Glikman et al. 2017; Martocchia et al. 2017) are consistent with the same lower X-ray-to-infrared ratios of extremely luminous type 1 quasars, once the X-ray luminosities are corrected for the effects of X-ray obscuration.

While the origin of this sub-linear relationship between X-ray and infrared luminosities is not well understood, one hypothesis is that X-ray-to-infrared ratios are anti-correlated with Eddington ratios (Leighly et al. 2007; Lusso et al. 2010; Jin et al. 2012) because a massive accretion flow disrupts the X-ray-emitting corona, suppressing the observed X-ray emission, resulting in apparently X-ray-weak quasars. At lower luminosities, Williams et al. (2004) also find that near-Eddington sources (narrow-line Seyfert 1 galaxies) have an appreciable X-ray-weak sub-population. In the recent *NuSTAR* serendipitous source survey, the source with the lowest X-ray-to-infrared ratio was a narrow-line Seyfert 1 object (Lansbury et al. 2017).

The observed (absorbed) X-ray luminosities of ERQs lie well below the type 1 relationship, with the median object

being a factor of ~ 5 underluminous in X-rays as compared to the Stern (2015) and Chen et al. (2017) relationships. However, at the high column densities found in ERQs the apparent luminosities are strongly affected by intervening absorption. As discussed above, the absorption correction at 2–10 keV energies is very sensitive to the assumed column, which is not well constrained for most of our sources. Therefore, our absorption-corrected luminosities are reported in Figure 6 with large uncertainties (which may in fact be underestimated since all of our calculations assume the same intrinsic spectral slope $\Gamma = 1.8$). Keeping in mind that our absorption-corrected X-ray luminosities should be considered estimates, we find that the median intrinsic absorption-corrected X-ray luminosity for our sample is $L_{2-10 \text{ keV, int}} = 10^{45} \text{ erg s}^{-1}$, which is on average still a factor of ~ 2 below the Stern (2015) and Chen et al. (2017) relationships, but consistent with them given the individual uncertainties.

Furthermore, in Figure 6 we show the absorption-corrected X-ray luminosity of the stacked spectrum, where we have a better-constrained spectral fit, a better N_H estimate than in individual sources, and therefore a lower uncertainty on the absorption-corrected luminosity. We find that the intrinsic X-ray luminosity of the stack is almost perfectly in line with the $L_X\text{--}L_{\text{IR}}$ relations from Chen et al. (2017) and Stern (2015). This demonstrates that the apparent luminosities of ERQs are likely suppressed by absorption, as opposed to intrinsic weakness in X-rays. We conclude that in the sample of ERQs observed so far there is no evidence for an appreciable population of X-ray-weak ERQs.

Two effects might bias our measurements. One is that, as discussed above, the $6 \mu\text{m}$ luminosity can be suppressed by obscuration. Since we do not correct for this absorption, the apparent $6 \mu\text{m}$ luminosities of our sources could be biased low, with higher intrinsic luminosities. A correction for this effect would move ERQ points and their stack in Figure 6 to the right, making them deficient in X-rays. However, because the $L_{\text{IR}}\text{--}L_X$ relationship is rather flat at these luminosities, that correction would have to be large—as much as an order of magnitude—to result in an appreciable X-ray deficit. In Section 4.4 we argue that the hot dust emission in ERQs is not as strongly affected by absorption as in type 2 quasars, so a large absorption correction in the mid-infrared is unnecessary, and even in type 2 quasars this correction is likely small (Mateos et al. 2015; Lansbury et al. 2017). Another possible bias may arise in X-ray luminosities. Although we do incorporate the off-axis detections in the serendipitous *XMM* point-source catalog and in archival *Chandra* observations, we do not consider upper limits on X-ray fluxes from off-axis *XMM* observations (Section 2). Therefore, our sample may be still marginally biased toward higher X-ray luminosities, and we could be missing a few sources with upper limits on their X-rays in Figure 6.

4.2. Comparison with HotDOGs

In terms of the SEDs, luminosities, and redshifts, ERQs are similar to HotDOGs (Eisenhardt et al. 2012; Assef et al. 2015; Tsai et al. 2015); indeed, the ERQ sample has overlap with HotDOGs targeted for follow-up X-ray observations (Section 2). The minor differences in their SEDs reflect selection effects: unlike HotDOGs which are selected based on *WISE* colors alone (see Eisenhardt et al. 2012), ERQs must

be bright enough ($r < 21.85$ mag) and must have appropriate colors for follow-up spectroscopy in BOSS, as well as be spectroscopically confirmed optical quasars. Therefore, the optical continuum of ERQs is typically stronger than that of HotDOGs, likely because ERQs have a more favorable geometry for scattering, which dominates the observed optical continuum (Alexandroff et al. 2017).

Follow-up observations have revealed physical similarities between these two populations. Near-infrared spectroscopy of ERQs reveals high equivalent widths and blueshifts of the [O III] $\lambda 5007$ line (Zakamska et al. 2016; S. Perrotta et al. 2017, in preparation), indicating a prevalence of strong outflows of ionized gas. While little equivalent data are available for HotDOGs, one published [O III] $\lambda 5007$ spectrum (Wu et al. 2017) also shows velocity dispersion above that seen in the local luminous quasar population (Zakamska & Greene 2014). With bolometric luminosities well in excess of 10^{47} erg s $^{-1}$ (Tsai et al. 2015; Zakamska et al. 2016), and estimated black hole masses of $\sim 10^9 M_{\odot}$ (Wu et al. 2017), ERQs and HotDOGs are likely near-Eddington, obscured quasars powered by some of the most massive black holes at $z \sim 3$, and are capable of launching extended ionized gas outflows.

X-ray properties of several HotDOGs were studied by Stern et al. (2014), Assef et al. (2016), and Ricci et al. (2017), and Vito et al. (2018) recently produced a systematic X-ray study and analysis of a significant sample of HotDOGs. These authors find that HotDOGs are a heavily obscured population with a deficiency in their absorption-corrected X-ray luminosities compared to type 1 quasars of similar bolometric luminosity. Another HotDOG presents a high equivalent width Fe K α line and shows other signs of Compton-thick obscuration (Piconcelli et al. 2015). The population of HotDOGs display X-ray-derived obscuration values extremely similar to those of ERQs.

X-ray luminosities of HotDOGs from Ricci et al. (2017) are shown in Figure 6 for comparison with ERQs and other populations. These luminosities have been corrected for absorption using the best available spectral information for individual sources, but despite this correction HotDOGs lie under the best-fit relationships for unobscured luminous sources, with a median absorption-corrected luminosity a factor of ~ 3 below the type 1 relationships. This finding is in contrast to what we see in the ERQ population, which (at the same $6 \mu\text{m}$ luminosity) appears to have X-ray luminosities consistent with the type 1 relationships. Ricci et al. (2017) suggest that HotDOGs (or some fraction of them) could be intrinsically X-ray weak, similarly to several other populations with known strong outflows, as discussed below. The absorption correction at column densities of $N_{\text{H}} = 10^{24}$ cm $^{-2}$ is highly uncertain, so one of the HotDOGs shown in Figure 6 has an order of magnitude luminosity uncertainty even though it has a high-quality *NuSTAR* spectrum (Ricci et al. 2017), and it thus could be appreciably more luminous than shown (with the quoted luminosity of $\log L_{2-10 \text{ keV}} [\text{erg s}^{-1}] = 44.9_{-0.14}^{+0.86}$). Future observations will allow us to determine whether the X-ray differences between ERQs and HotDOGs persist in larger samples and in better-quality data, despite the close similarity of all other observables between these two populations.

4.3. Comparison with BAL Quasars, Type 1 Quasars, and ULIRGs

While it is not yet clear exactly which selection criterion of ERQs—red color or high line equivalent width—predominantly correlates with the presence of outflows, ERQs show outflow activity on a wide range of scales. ERQs are twice as likely as type 1 quasars to show broad absorption-line (BAL) signatures and a hundred times more likely to show ≥ 2500 km s $^{-1}$ blueshifts in their C IV $\lambda 1550$ emission relative to the estimated systemic redshift (Hamann et al. 2017). Therefore, in addition to signatures of [O III] $\lambda 5007$ outflows, which are likely extended on scales of hundreds of parsecs or more, ERQs show a high prevalence of circumnuclear outflows. In this section we compare the X-ray properties of ERQs with those of other quasars with known nuclear and large-scale outflows.

Initial studies of X-ray emission from BAL quasars found that these objects were highly absorbed in X-rays ($N_{\text{H}} \geq 10^{23}$ cm $^{-2}$) and concluded that this absorption was likely the reason BAL quasars were under-represented in soft X-ray surveys (Gallagher et al. 1999, 2002, 2006; Green et al. 2001). Recent studies have confirmed that BAL quasars are faint X-ray sources, but *NuSTAR* observations have not found a significant population of Compton-thick BAL quasars, suggesting that a third of BAL quasars are intrinsically X-ray weak (Luo et al. 2013, 2014), much weaker than non-BAL type 1 quasars of the same optical luminosity. Optical emission-line diagnostics also indicate that in BALs the ionizing emission is softer than an average quasar spectrum (Richards et al. 2011; Baskin et al. 2013, 2015). These observations are consistent with the theoretical paradigm that X-rays can over-ionize circumnuclear gas, lowering its opacity and suppressing production of radiatively driven winds (Murray et al. 1995; Proga et al. 2000; Proga & Kallman 2004; Sim et al. 2010).

Type 1 quasars with outflows seen in emission lines also tend to have weak X-rays (Leighly et al. 2007). Quasars with the strongest outflow signatures in their C IV tend to have relatively weak emission lines (Richards et al. 2011). In a population of such weak-lined quasars examined by Luo et al. (2015), half of the objects had X-ray luminosities over an order of magnitude below the Chen et al. (2017) infrared versus X-ray relationship. A similar trend is seen in local ultra-luminous infrared galaxies (ULIRGs). Although not all ULIRGs are powered by a luminous active nucleus (Sanders & Mirabel 1996), those that are tend to be highly obscured in X-rays (Nardini & Risaliti 2011). Intriguingly, ULIRGs that show strong signatures of radiatively driven outflows have X-ray luminosities that are up to two orders of magnitude lower than expected from their bolometric luminosity (Teng et al. 2014, 2015).

Thus, the theoretical paradigm in which radiatively driven outflows can only arise when X-rays are relatively weak finds confirmation in known quasars with outflows, as 30%–50% of them are significantly weaker in X-rays than they should be based on relationships shown in Figure 6. It is therefore puzzling that ERQs—quasars with known powerful high-velocity ionized gas outflows—do not appear to follow the same trend. Instead, we find that when their luminosities are corrected for obscuration, they seem to be in line with these relationships.

Our X-ray observations of ERQs raise the possibility that they and other quasars (non-ERQs) with outflows are not

drawn from the same populations. Hence, ERQs may not simply be the obscured analogs of BAL quasars or weak-line type 1 quasars. An alternative explanation is that ERQ winds are not driven by the same mechanism as the line-driven winds which require low X-rays—for example, ERQ winds could instead be driven by radiation pressure on dust (Keating et al. 2012; Thompson et al. 2015), a phenomenon which Ishibashi et al. (2017) suggest may be associated specifically with red quasars. Another possibility is that the combination of high X-ray luminosities and strong outflow activity is due to orientation effects and we would see the weak X-ray population along some other directions (which are also the directions along which gas can be accelerated radiatively), but such objects would not be selected by the same color and magnitude cuts as ERQs. Finally, it is possible that the statistics of the current sample are just too limited and that X-ray-weak ERQs are yet to be discovered.

4.4. Comparison with Type 2 Quasars

The color selection that unveiled the ERQ population (Ross et al. 2015) was initially designed to identify high-redshift analogs of the low-redshift obscured quasar populations. In this section we compare X-ray properties of ERQs with those of low-redshift type 2 quasars i.e., optically selected quasars that have no broad component in their $H\alpha$ emission line (Zakamska et al. 2003; Reyes et al. 2008). The lack of broad $H\alpha$ in these systems suggests optical extinction over >10 mag, sufficient to block all direct light from the broad-line region at the $H\alpha$ wavelength. In contrast, ERQs routinely show broad $H\alpha$ components (Zakamska et al. 2016; Wu et al. 2017), limiting the amount of extinction toward the broad-line region. This difference between $H\alpha$ properties of type 2 quasars and ERQs can be due to a difference in column densities (larger in type 2s and lower in ERQs) or relative spatial scales of obscuration (larger than the broad-line region in type 2s, versus smaller than the broad-line region in ERQs). Viewing angle could also produce this difference (closer to edge-on in type 2s than in ERQs), but that should also result in the differences in the intervening X-ray column density.

Additionally, the SEDs of ERQs rise steeply at a few μm (Hamann et al. 2017) and may peak at 5–10 μm (Tsai et al. 2015). Dust which is emitting at a few μm is confined to scales of a few pc if it is in thermal equilibrium with the quasar radiation (Barvainis 1987). In contrast, type 2 quasar SEDs are well represented by smoothly rising power-laws from 3 to 20 μm , peaking at yet longer wavelengths (Mateos et al. 2013; Hickox et al. 2017), suggesting that the hot-dust-emitting region is obscured by material which is optically thick in the mid-infrared, so that this emission is reprocessed to longer wavelengths. Again, these observations can be explained by a higher level of obscuration in type 2s than in ERQs (e.g., due to a difference in viewing angles). Alternatively, the X-ray absorber could be more compact than the hot-dust-emitting region in ERQs and relatively more extended in type 2s.

X-ray observations allow us to test these hypotheses for the differences between optical and infrared properties of type 2 quasars and ERQs. In an excellent agreement with the geometric unification model (Antonucci 1993), low-redshift optically selected type 2 quasars have high levels of obscuration. Jia et al. (2013) study a large sample of type 2 quasars at $z < 1$ with *Chandra* and *XMM* and find that among the 3/4 of the population detected in the X-rays, the median

column density is $N_{\text{H}} = 10^{22.9} \text{ cm}^{-2}$. The non-detected sources are likely to be Compton-thick rather than X-ray weak, as seen in a few objects with *NuSTAR* observations (Lansbury et al. 2015), with the overall Compton-thick fraction in the type 2 population estimated between 36% and 76% (Jia et al. 2013; Lansbury et al. 2015). Thus X-ray absorption properties are indistinguishable between type 2 quasars and ERQs.

The broad-line region and the hot dust region are less obscured in ERQs than they are in type 2 quasars, yet the X-ray-absorbing column densities are similar in these two populations. This similarity rules out the dominant role of viewing angle effects: to make the broad-line region and the hot dust region more visible in ERQs than in type 2, the intervening column density detected in X-rays must be smaller in ERQs, which is not what we see in this population. We conclude that X-ray absorption is confined on scales smaller than a few pc in ERQs, whereas it must be more extended in type 2 quasars. In Section 4.5 we discuss the astrophysical implications of this observation.

Type 2 quasars at $z < 1$ studied by (Jia et al. 2013) have lower X-ray luminosities than ERQs (median $\sim 10^{43} \text{ erg s}^{-1}$ for detected sources), but they are also less bolometrically luminous than ERQs. So instead of directly comparing the X-ray luminosities of these two populations we compare each of them to the type 1 infrared-to-X-ray relationships (Stern 2015; Chen et al. 2017) at their respective luminosities. Both populations are hard to detect in the X-rays, but in both cases stacking of weak X-ray sources reveals that their X-ray faintness is due to obscuration and not to intrinsic X-ray weakness. Both $z < 1$ type 2 quasars and ERQs are consistent with the luminosity-dependent infrared-to-X-ray relationships.

4.5. Physical Properties of ERQ Winds

Spectropolarimetric observations of ERQs indicate that essentially all of the observed rest-frame ultraviolet continuum is due to scattered light, as suggested by the high continuum polarization at these wavelengths (Alexandroff et al. 2017). Furthermore, the kinematic structure of the polarization of ultraviolet emission lines is consistent with scattering produced on scales ~ 10 pc in quasi-equatorial winds with large covering factors (Veilleux et al. 2016; Alexandroff et al. 2017; N. L. Zakamska & R. Alexandroff 2017, in preparation) moving at several thousand km s^{-1} .

The SED of ERQs steeply rises at 1–5 μm (Hamann et al. 2017), indicating a dominant role of emission from hot dust ($T \simeq 500$ K, resulting in a blackbody spectrum peak at 6 μm). Taking a $10^{47} \text{ erg s}^{-1}$ ionizing luminosity, we find that these temperatures are established at $\sim 12 \times (L/10^{47} \text{ erg s}^{-1})^{1/2} (T/500 \text{ K})^{-2}$ pc from the quasar, i.e., on scales that are similar to those of the scattering wind. The difference in the shape of the infrared SED between ERQs (steeply rising to 5 μm and then flat) and type 2 quasars (continuously and slowly rising to 30 μm) suggests that the hot dust emission region cannot be strongly affected by obscuration, otherwise we would have seen the smoothing of the wavelength dependence of the SED due to radiative transfer effects. Therefore, the high X-ray column must also accumulate on similar scales.

A natural geometry suggested by these observations is that in which the observed X-ray absorption occurs in the same wind as that seen in spectropolarimetric observations. Compton-thick absorption due to a circumnuclear wind is seen for example in Mrk 231 (Braitto et al. 2004), though this interpretation is

disputed (Teng et al. 2014), and in PDS456 (Reeves et al. 2009). A wind with a column density of $N_{\text{H}} = 10^{24} \text{ cm}^{-2}$, confined to scales $r < 20 \text{ pc}$, with a covering factor of $\Omega = 0.2$ and a velocity of $v = 2000 \text{ km s}^{-1}$ (Alexandroff et al. 2017) has a mass outflow rate

$$\dot{M} = 58 M_{\odot} \text{ yr}^{-1} \times \left(\frac{r_{\text{in}}}{1 \text{ pc}} \right) \left(\frac{N_{\text{H}}}{10^{24} \text{ cm}^{-2}} \right) \left(\frac{v}{2000 \text{ km s}^{-1}} \right) \left(\frac{\Omega}{0.2} \right). \quad (3)$$

Here we have assumed that the wind is in a steady-state with a $\propto 1/r^2$ density profile, so that the column density measurement is weighted toward r_{in} , its launching distance. Similar mass outflow rates have been inferred for some X-ray absorbing winds (Chartas et al. 2009; Nardini et al. 2015), though such winds propagate with much higher velocity ($\geq 0.1c$) and are located at much smaller distances (~ 100 gravitational radii).

The kinetic power of such wind is

$$\dot{E}_{\text{kin}} = \frac{\dot{M} v^2}{2} = 7 \times 10^{43} \text{ erg s}^{-1}, \quad (4)$$

which constitutes a relatively small fraction of the bolometric output. Defining the Eddington mass outflow rate to be $\dot{M}_{\text{Edd}} = L_{\text{Edd}}/0.1c^2$, we find that such outflow is mildly super-Eddington, with

$$\frac{\dot{M}}{\dot{M}_{\text{Edd}}} = 2.6 \times \left(\frac{M_{\text{BH}}}{10^9 M_{\odot}} \right)^{-1}. \quad (5)$$

The ratio of outflowing to inflowing mass in near-Eddington accretion is poorly known, but initial simulations suggest that the inflow rates may be a factor of a few higher than the outflow rates (Volonteri et al. 2015). Thus ERQs might be accreting at ~ 10 times the Eddington limit, but with only mildly super-Eddington emerging luminosities, reflecting the low radiative efficiency of super-Eddington accretion (Jiang et al. 2014; Sądowski et al. 2014).

For the same fiducial parameters, the ratio of the momentum flux of the wind to the available momentum of the photons is

$$\frac{\dot{P}_{\text{wind}}}{\dot{P}_{\text{rad}}} = \frac{\dot{M} v c}{L} = 0.22. \quad (6)$$

Therefore, enough photons are in principle available for such wind to be radiatively driven, but the available momentum of the photons would need to be converted to the wind very efficiently. In particular, our fiducial values assume a wind covering fraction of $\Omega = 0.2$, so for an emitter which is isotropic on scales $< 1 \text{ pc}$ all photon momentum over this covering factor would need to be converted to the wind.

5. Conclusions

In this paper, we present X-ray observations of 11 ERQs at $z = 1.5\text{--}3.2$ selected from the SDSS and *WISE* databased on their high infrared-to-optical flux ratios and high equivalent width C IV emission lines (Ross et al. 2015; Hamann et al. 2017). These objects are among the most luminous quasars at the peak epoch of quasar activity, with directly measured $\nu L_{\nu}[6 \mu\text{m}]$ reaching $10^{47} \text{ erg s}^{-1}$ and with inferred bolometric luminosities close to $10^{48} \text{ erg s}^{-1}$

(Hamann et al. 2017). Because these values are in excess of the Eddington limit for a $M_{\text{BH}} = 10^9 M_{\odot}$ black hole ($1.3 \times 10^{47} \text{ erg s}^{-1}$), we hypothesize that these objects are close-to-Eddington or super-Eddington accretors and therefore are likely capable of launching powerful outflows as suggested by numerical simulations (Jiang et al. 2014; Sądowski et al. 2014).

Among the most striking properties of ERQs is the routine occurrence of the [O III] $\lambda 5007$ emission line with extreme outflow signatures—strong blueshifted asymmetries and widths reaching unprecedented FWHM $> 5000 \text{ km s}^{-1}$ (Zakamska et al. 2016; S. Perrotta et al. 2017, in preparation). This emission line can only originate in regions of relatively low density ($\ll 10^6 \text{ cm}^{-3}$) well outside of the region of gravitational influence of the black hole (likely on scales of hundreds of pc) and is therefore evidence of feedback of quasar activity onto the host galaxy. ERQs also show outflow activity on scales of a few to a few tens of pc, as evidenced by the absorption signatures in C IV $\lambda 1550$ and by the shape of the C IV emission line (Hamann et al. 2017), as well as by spectropolarimetric observations of the rest-frame ultraviolet emission lines (Alexandroff et al. 2017). While it is not yet known why the combination of the ERQ selection criteria is so successful in identifying objects with strong [O III] outflows, it may not be surprising because our photometric selection cuts yield objects with extremely high luminosities (Hamann et al. 2017) and the high C IV equivalent width requirement may preferentially select objects with outflows (Alexandroff et al. 2017).

The X-ray observations presented here enable us to further test these ideas. Because of the penetrating power of X-rays, we can use X-ray observations to estimate the amount of intervening absorption. Only two objects are detected with sufficient counts to enable spectral fitting; both are absorbed with best-fit $N_{\text{H}} \sim 6 \times 10^{22} \text{ cm}^{-2}$. We further use hardness-ratio analyses to find that ERQs are in general a strongly X-ray-absorbed population, with estimated column densities in the range $\sim (0.5\text{--}24) \times 10^{23} \text{ cm}^{-2}$ and an implied $\sim 50\%$ Compton-thick fraction, where we count as potential Compton-thick candidates the five sources with the hardest HR_2 hardness ratios and estimated column densities at or above 10^{24} cm^{-2} . The stack of *Chandra* observations reveals a spectrum with a best-fit column of $N_{\text{H}} = 8 \times 10^{23} \text{ cm}^{-2}$, which also supports high levels of average obscuration in the ERQ sample. Finding high levels of obscuration is not surprising, as signs of obscuration are present at other wavelengths and are encoded in the ERQ selection itself.







Due to the steep rise of the infrared SED of ERQs at a few μm and due to the detection of broad $\text{H}\alpha$ components in their near-infrared spectra, we postulate that the observed high X-ray absorption must accumulate on scales similar to or smaller than the warm-dust-emitting region and some of the broad-line region. Further assuming that X-ray absorption is associated with the wind seen in spectropolarimetric observations (Alexandroff et al. 2017) we arrive at estimates for the mass, energy, and momentum outflow rates of such wind presented in Section 4.5. We find that the momentum outflow rates are consistent with the availability of photons, but the efficiency of momentum transfer would need to be quite high. While the energy outflow rate is a small fraction of the Eddington limit, the mass outflow rate is marginally super-Eddington, supporting our hypothesis that ERQs are near-Eddington or super-Eddington accretors.




Another major conclusion of our analysis is that the intrinsic X-ray luminosities of ERQs are largely in agreement with those of type 1 quasars of the same infrared power (Stern 2015). Thus, ERQs do not appear to be intrinsically weak in X-rays, unlike some of the BAL quasars (Luo et al. 2014), possibly some of the HotDOGs (Ricci et al. 2017), and some other populations of quasars with known outflows. It has been hypothesized that X-rays may play a critical role in enabling or disabling powerful radiatively driven outflows (Proga et al. 2000; Sim et al. 2010; Richards et al. 2011; Luo et al. 2014): high X-ray luminosities may indicate high accretion rates which are conducive to initiating outflows, but overly high X-ray luminosities may over-ionize the surrounding gas, lower its opacity, and suppress outflows. Our observations suggest that X-ray ionization does not appreciably suppress wind activity in ERQs.

ERQs may represent a somewhat different population from BAL quasars, a third of which show weak intrinsic X-ray luminosities and which are known to be driven via opacity in bound-bound transitions of partially ionized gas. The lack of an appreciable X-ray-weak population among ERQs may be a geometric orientation effect, and perhaps the ionizing spectrum seen by the outflowing gas is different from that seen by the observer. Alternatively, it is possible that the mechanism for wind production is not the same in ERQs and in BAL quasars—for example, winds in ERQs could be launched by radiation pressure on dust, not gas (Keating et al. 2012; Thompson et al. 2015; Ishibashi et al. 2017). Observations of a larger sample of ERQs will indicate whether there exists an appreciable population of X-ray-weak ERQs.

We thank the anonymous referee for their careful reading of the manuscript, which has helped us to clarify and strengthen several aspects of the analyses. N.L.Z. is grateful to the Institute for Advanced Study for support through the Junior Visiting Professor program and hospitality during multiple subsequent visits which enabled completion of this project, and to Johns Hopkins University for support via the Catalyst Award. Support for this work was also provided in part by the National Aeronautics and Space Administration through *Chandra* Award Numbers GO5-16107X and GO6-17100X issued by the *Chandra* X-ray Observatory Center, which is operated by the Smithsonian Astrophysical Observatory for and on behalf of the National Aeronautics Space Administration under contract NAS8-03060. R.J.A. was supported by FONDECYT grant number 1151408. The work of D.S. was carried out at the Jet Propulsion Laboratory, California Institute of Technology, under a contract with NASA. The scientific results reported in this paper are based to a significant degree on observations made by the *Chandra* X-ray Observatory and data obtained from the *Chandra* Data Archive.

ORCID iDs

Nadia L. Zakamska  <https://orcid.org/0000-0001-6100-6869>
 Rachael M. Alexandroff  <https://orcid.org/0000-0003-2830-0913>
 Roberto J. Assef  <https://orcid.org/0000-0002-9508-3667>
 Dominika Wylezalek  <https://orcid.org/0000-0003-2212-6045>
 William N. Brandt  <https://orcid.org/0000-0002-0167-2453>
 George B. Lansbury  <https://orcid.org/0000-0002-5328-9827>

Gordon Richards  <https://orcid.org/0000-0002-1061-1804>
 Daniel Stern  <https://orcid.org/0000-0003-2686-9241>
 Michael A. Strauss  <https://orcid.org/0000-0002-0106-7755>

References

- Alam, S., Albareti, F. D., Allende Prieto, C., et al. 2015, *ApJS*, **219**, 12
 Alexander, D. M., Brandt, W. N., Hornschemeier, A. E., et al. 2001, *AJ*, **122**, 2156
 Alexandroff, R., Strauss, M. A., Greene, J. E., et al. 2013, *MNRAS*, **435**, 3306
 Alexandroff, R. M., Zakamska, N. L., Barth, A., & Hamann, F. 2017, *MNRAS*, submitted
 Antonucci, R. 1993, *ARA&A*, **31**, 473
 Assef, R. J., Eisenhardt, P. R. M., Stern, D., et al. 2015, *ApJ*, **804**, 27
 Assef, R. J., Walton, D. J., Brightman, M., et al. 2016, *ApJ*, **819**, 111
 Banerji, M., Alaghband-Zadeh, S., Hewett, P. C., & McMahon, R. G. 2015, *MNRAS*, **447**, 3368
 Banerji, M., McMahon, R. G., Hewett, P. C., et al. 2012, *MNRAS*, **427**, 2275
 Barvainis, R. 1987, *ApJ*, **320**, 537
 Baskin, A., Laor, A., & Hamann, F. 2013, *MNRAS*, **432**, 1525
 Baskin, A., Laor, A., & Hamann, F. 2015, *MNRAS*, **449**, 1593
 Boyle, B. J., & Terlevich, R. J. 1998, *MNRAS*, **293**, L49
 Braito, V., Della Ceca, R., Piconcelli, E., et al. 2004, *A&A*, **420**, 79
 Brusa, M., Bongiorno, A., Cresci, G., et al. 2015, *MNRAS*, **446**, 2394
 Canalizo, G., & Stockton, A. 2001, *ApJ*, **555**, 719
 Cano-Díaz, M., Maiolino, R., Marconi, A., et al. 2012, *A&A*, **537**, L8
 Cash, W. 1979, *ApJ*, **228**, 939
 Chartas, G., Saez, C., Brandt, W. N., Giustini, M., & Garmire, G. P. 2009, *ApJ*, **706**, 644
 Chen, C.-T. J., Hickox, R. C., Goulding, A. D., et al. 2017, *ApJ*, **837**, 145
 Cutri, R. M., Wright, E. L., Conrow, T., et al. 2013, *yCat*, **2328**, 1
 Dawson, K. S., Schlegel, D. J., Ahn, C. P., et al. 2013, *AJ*, **145**, 10
 Eisenhardt, P. R. M., Wu, J., Tsai, C.-W., et al. 2012, *ApJ*, **755**, 173
 Eisenstein, D. J., Weinberg, D. H., Agol, E., et al. 2011, *AJ*, **142**, 72
 Feruglio, C., Bongiorno, A., Fiore, F., et al. 2014, *A&A*, **565**, A91
 Fiore, F., Feruglio, C., Shankar, F., et al. 2017, *A&A*, **601**, A143
 Gallagher, S. C., Brandt, W. N., Chartas, G., et al. 2006, *ApJ*, **644**, 709
 Gallagher, S. C., Brandt, W. N., Chartas, G., & Garmire, G. P. 2002, *ApJ*, **567**, 37
 Gallagher, S. C., Brandt, W. N., Sambruna, R. M., Mathur, S., & Yamasaki, N. 1999, *ApJ*, **519**, 549
 Gandhi, P., Horst, H., Smette, A., et al. 2009, *A&A*, **502**, 457
 Glikman, E., LaMassa, S., Piconcelli, E., Urry, M., & Lacy, M. 2017, *ApJ*, **847**, 116
 Goulding, A. D., Forman, W. R., Hickox, R. C., et al. 2012, *ApJS*, **202**, 6
 Green, P. J., Aldcroft, T. L., Mathur, S., Wilkes, B. J., & Elvis, M. 2001, *ApJ*, **558**, 109
 Greene, J. E., Alexandroff, R., Strauss, M. A., et al. 2014, *ApJ*, **788**, 91
 Greene, J. E., Zakamska, N. L., Ho, L. C., & Barth, A. J. 2011, *ApJ*, **732**, 9
 Hamann, F., Zakamska, N. L., Ross, N., et al. 2017, *MNRAS*, **464**, 3431
 Hamann, F. W., Barlow, T. A., Chaffee, F. C., Foltz, C. B., & Weymann, R. J. 2001, *ApJ*, **550**, 142
 Harrison, C. M., Alexander, D. M., Mullaney, J. R., & Swinbank, A. M. 2014, *MNRAS*, **441**, 3306
 Hickox, R. C., Myers, A. D., Greene, J. E., et al. 2017, *ApJ*, **849**, 53
 Hopkins, P. F., Hernquist, L., Cox, T. J., et al. 2006, *ApJS*, **163**, 1
 Ishibashi, W., Banerji, M., & Fabian, A. C. 2017, *MNRAS*, **469**, 1496
 Jia, J., Ptak, A., Heckman, T., & Zakamska, N. L. 2013, *ApJ*, **777**, 27
 Jiang, Y.-F., Stone, J. M., & Davis, S. W. 2014, *ApJ*, **796**, 106
 Jin, C., Ward, M., & Done, C. 2012, *MNRAS*, **425**, 907
 Keating, S. K., Everett, J. E., Gallagher, S. C., & Deo, R. P. 2012, *ApJ*, **749**, 32
 King, A. 2003, *ApJL*, **596**, L27
 LaMassa, S. M., Ricarte, A., Glikman, E., et al. 2016, *ApJ*, **820**, 70
 Lansbury, G. B., Alexander, D. M., Del Moro, A., et al. 2014, *ApJ*, **785**, 17
 Lansbury, G. B., Gandhi, P., Alexander, D. M., et al. 2015, *ApJ*, **809**, 115
 Lansbury, G. B., Stern, D., Aird, J., et al. 2017, *ApJ*, **836**, 99
 Leighly, K. M., Halpern, J. P., Jenkins, E. B., et al. 2007, *ApJ*, **663**, 103
 Liu, G., Zakamska, N. L., Greene, J. E., Nesvadba, N. P. H., & Liu, X. 2013a, *MNRAS*, **430**, 2327
 Liu, G., Zakamska, N. L., Greene, J. E., Nesvadba, N. P. H., & Liu, X. 2013b, *MNRAS*, **436**, 2576
 Luo, B., Brandt, W. N., Alexander, D. M., et al. 2013, *ApJ*, **772**, 153
 Luo, B., Brandt, W. N., Alexander, D. M., et al. 2014, *ApJ*, **794**, 70
 Luo, B., Brandt, W. N., Hall, P. B., et al. 2015, *ApJ*, **805**, 122
 Lusso, E., Comastri, A., Vignali, C., et al. 2010, *A&A*, **512**, A34

- Martocchia, S., Piconcelli, E., Zappacosta, L., et al. 2017, arXiv:1708.00452
- Mateos, S., Alonso-Herrero, A., Carrera, F. J., et al. 2013, *MNRAS*, **434**, 941
- Mateos, S., Carrera, F. J., Alonso-Herrero, A., et al. 2015, *MNRAS*, **449**, 1422
- Moe, M., Arav, N., Bautista, M. A., & Korista, K. T. 2009, *ApJ*, **706**, 525
- Mullaney, J. R., Alexander, D. M., Fine, S., et al. 2013, *MNRAS*, **433**, 622
- Murray, N., Chiang, J., Grossman, S. A., & Voit, G. M. 1995, *ApJ*, **451**, 498
- Nandra, K., & Pounds, K. A. 1994, *MNRAS*, **268**, 405
- Nardini, E., Reeves, J. N., Gofford, J., et al. 2015, *Sci*, **347**, 860
- Nardini, E., & Risaliti, G. 2011, *MNRAS*, **415**, 619
- Page, K. L., Reeves, J. N., O'Brien, P. T., & Turner, M. J. L. 2005, *MNRAS*, **364**, 195
- Pâris, I., Petitjean, P., Aubourg, É, et al. 2014, *A&A*, **563**, A54
- Pâris, I., Petitjean, P., Ross, N. P., et al. 2017, *A&A*, **597**, A79
- Park, T., Kashyap, V. L., Siemiginowska, A., et al. 2006, *ApJ*, **652**, 610
- Piconcelli, E., Jimenez-Bailón, E., Guainazzi, M., et al. 2005, *A&A*, **432**, 15
- Piconcelli, E., Vignali, C., Bianchi, S., et al. 2015, *A&A*, **574**, L9
- Pier, E. A., & Krolik, J. H. 1992, *ApJ*, **401**, 99
- Proga, D., & Kallman, T. R. 2004, *ApJ*, **616**, 688
- Proga, D., Stone, J. M., & Kallman, T. R. 2000, *ApJ*, **543**, 686
- Reeves, J. N., O'Brien, P. T., Braitto, V., et al. 2009, *ApJ*, **701**, 493
- Reeves, J. N., & Turner, M. J. L. 2000, *MNRAS*, **316**, 234
- Reyes, R., Zakamska, N. L., Strauss, M. A., et al. 2008, *AJ*, **136**, 2373
- Ricci, C., Assef, R. J., Stern, D., et al. 2017, *ApJ*, **835**, 105
- Richards, G. T., Kruczek, N. E., Gallagher, S. C., et al. 2011, *AJ*, **141**, 167
- Richards, G. T., Lacy, M., Storrie-Lombardi, L. J., et al. 2006, *ApJS*, **166**, 470
- Rosen, S. R., Webb, N. A., Watson, M. G., et al. 2016, *A&A*, **590**, A1
- Ross, N. P., Hamann, F., Zakamska, N. L., et al. 2015, *MNRAS*, **453**, 3932
- Sądowski, A., Narayan, R., McKinney, J. C., & Tchekhovskoy, A. 2014, *MNRAS*, **439**, 503
- Sanders, D. B., & Mirabel, I. F. 1996, *ARA&A*, **34**, 749
- Silk, J., & Rees, M. J. 1998, *A&A*, **331**, L1
- Sim, S. A., Proga, D., Miller, L., Long, K. S., & Turner, T. J. 2010, *MNRAS*, **408**, 1396
- Stark, A. A., Gammie, C. F., Wilson, R. W., et al. 1992, *ApJS*, **79**, 77
- Stern, D. 2015, *ApJ*, **807**, 129
- Stern, D., Lansbury, G. B., Assef, R. J., et al. 2014, *ApJ*, **794**, 102
- Teng, S. H., Brandt, W. N., Harrison, F. A., et al. 2014, *ApJ*, **785**, 19
- Teng, S. H., Rigby, J. R., Stern, D., et al. 2015, *ApJ*, **814**, 56
- Thompson, T. A., Fabian, A. C., Quataert, E., & Murray, N. 2015, *MNRAS*, **449**, 147
- Tsai, C.-W., Eisenhardt, P. R. M., Wu, J., et al. 2015, *ApJ*, **805**, 90
- Urrutia, T., Lacy, M., Gregg, M. D., & Becker, R. H. 2005, *ApJ*, **627**, 75
- Veilleux, S., Meléndez, M., Sturm, E., et al. 2013, *ApJ*, **776**, 27
- Veilleux, S., Meléndez, M., Tripp, T. M., Hamann, F., & Rupke, D. S. N. 2016, *ApJ*, **825**, 42
- Vito, F., Brandt, W. N., Stern, D., et al. 2018, *MNRAS*, **474**, 4528
- Volonteri, M., Silk, J., & Dubus, G. 2015, *ApJ*, **804**, 148
- Wang, S., Liu, J., Qiu, Y., et al. 2016a, *ApJS*, **224**, 40
- Wang, T., Ferland, G. J., Yang, C., Wang, H., & Zhang, S. 2016b, *ApJ*, **824**, 106
- Williams, R. J., Mathur, S., & Pogge, R. W. 2004, *ApJ*, **610**, 737
- Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, *AJ*, **140**, 1868
- Wu, J., Jun, H. D., Assef, R. J., et al. 2017, arXiv:1703.06888
- Yaqoob, T. 2012, *MNRAS*, **423**, 3360
- Yuan, S., Strauss, M., & Zakamska, N. L. 2016, *MNRAS*, **462**, 1603
- Zakamska, N. L., & Greene, J. E. 2014, *MNRAS*, **442**, 784
- Zakamska, N. L., Hamann, F., Pâris, I., et al. 2016, *MNRAS*, **459**, 3144
- Zakamska, N. L., Strauss, M. A., Krolik, J. H., et al. 2003, *AJ*, **126**, 2125